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## Evidence for a compact Wolf-Rayet progenitor for the Type Ic supernova PTF 10vgv

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### ABSTRACT

We present the discovery of PTF 10vgv, a Type Ic supernova detected by the Palomar Transient Factory, using the Palomar 48-inch telescope (P48). *R*-band observations of the PTF 10vgv field with P48 probe the supernova emission from its very early phases (about two weeks before *R*-band maximum), and set limits on its flux in

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the week prior to the discovery. Our sensitive upper limits and early detections constrain the post-shock-breakout luminosity of this event. Via comparison to numerical (analytical) models, we derive an upper-limit of  $R \lesssim 4.5 R_{\odot}$  ( $R \lesssim 1 R_{\odot}$ ) on the radius of the progenitor star, a direct indication in favor of a compact Wolf-Rayet star. Applying a similar analysis to the historical observations of SN 1994I, yields  $R \lesssim 1/4 R_{\odot}$  for the progenitor radius of this supernova.

*Subject headings:* supernovae: general — supernovae: individual (PTF 10vgv)

## 1. Introduction

Core-collapse supernovae (SNe) are believed to originate from evolved, massive progenitors (initial mass  $\gtrsim 8\text{--}10 M_{\odot}$ ) whose iron core undergoes gravitational collapse. Among them, Type II-Plateau (II-P) SNe show prominent hydrogen in their spectrum and a plateau in the optical light curves. Type IIb SNe have hydrogen in the spectrum initially, and a H-deficient spectrum at later times. Finally, Types Ib and Ic show no evidence for hydrogen at any time. The H-deficient/H-poor core-collapse SNe are thought to be produced by progenitors stripped of their hydrogen (SN Ib) and possibly helium (SN Ic) envelopes prior to exploding (for a review, see Filippenko 1997). Due to the stiff dependence of mass loss on luminosity/mass, a sequence of increasing main-sequence mass may be pictured going from progenitors of SNe II-P, IIb, Ib, and Ic (Heger et al. 2003; Crowther 2007; Georgy et al. 2009). Rotation, metallicity, and binarity also affect the mass loss (e.g., Meynet et al. 1994; Meynet & Maeder 2000).

The viability of the standard explosion mechanism for stars of increasing mass is challenging, given that their higher mass cores are more bound, and their SN shocks subject to a very high accretion rate (e.g., Burrows et al. 2007). Binary-star evolution has been studied as a channel to circumvent this caveat (e.g., Utrobin 1994; Woosley et al. 1994; Young et al. 1995; Blinnikov et al. 1998; Fryer et al. 2007; Yoon et al. 2010; Smith et al. 2011). The basic ingredient of this scenario is mass loss through transfer onto a companion. In this case, the mass-loss luminosity scaling does not apply, and much lower mass progenitors can explode as H-poor cores.

Recently, Dessart et al. (2011) published simulations of SN light curves resulting from explosions of SN IIb/Ib/Ic progenitors. All SNe show a  $\sim 10$ -day-long post-breakout plateau with a luminosity of  $(1 - 5) \times 10^7 L_{\odot}$ . Analytical estimates for the early-time ( $t \lesssim 1 - 2$  d since explosion) post-breakout emission have been provided by Rabinak & Waxman (2011) and Nakar & Sari (2010).

In this paper we present the discovery of type Ic SN, PTF 10vgv, detected by the Palomar

Transient Factory<sup>1</sup> (PTF; Law et al. 2009; Rau et al. 2009) (§2). We report its spectral classification (§3) and the radio follow-up observations (§4). We constrain the radius of the stellar progenitor of this SN by comparing our tight pre-discovery upper-limits with the predictions of several models (Dessart et al. 2011; Rabinak & Waxman 2011; Nakar & Sari 2010) (§5).

## 2. Discovery and *R*-band photometry

On 2010 September 14.1446 (UTC times are used throughout), we discovered a Type Ic SN, PTF 10vgv, via the automated Oarical software (Bloom et al. 2011). The SN was visible at a magnitude of  $R \approx 19.9$  (Table 1 and Figure 1) in an image (60 s exposure) taken with the Palomar Oschin Schmidt 48-inch telescope (P48). It was not seen in previous images of the same field taken on 2010 September 12.4830, down to a limiting magnitude of  $R > 20.2$ . The SN J2000 position is  $\alpha = 22^{\text{h}}16^{\text{m}}01.17^{\text{s}}$ ,  $\delta = +40^{\circ}52'03.3''$  (Corsi et al. 2010a), at an angular distance of  $\sim 5''$  from the galaxy SDSS<sup>2</sup> J221601.54+405206.5. P48 observations were obtained with the Mould-*R* filter (Table 1 and Figure 2). A high-quality image produced by stacking several images of the same field was used as a reference and subtracted from the individual images. Photometry was performed with an aperture of  $2''$  radius relative to the *r*-band magnitudes of ten SDSS reference stars in the field (Figure 1), applying color corrections (Corsi et al. 2011; Ofek et al. 2011). Aperture corrections were applied to account for systematic errors as well as errors introduced by the subtraction process (Corsi et al. 2011).

Table 1:: P48 observations of PTF 10vgv in *R*-band.

Start time JD-2455453.6446 (d)	Exposure (s)	Mag <sup>a</sup> [mag]
-6.776	600	< 21.2 <sup>b</sup>
-6.776	60	< 20.8 <sup>b</sup>
-6.732	60	< 21.1 <sup>b</sup>
-5.732	60	< 20.9 <sup>b</sup>
-5.688	60	< 20.8 <sup>b</sup>
-3.811	60	< 20.6 <sup>b</sup>
-3.811	360	< 20.8 <sup>b</sup>
-3.766	60	< 20.6 <sup>b</sup>
-2.814	60	< 20.8 <sup>b</sup>

<sup>1</sup><http://www.astro.caltech.edu/ptf/>

<sup>2</sup>Sloan Digital Sky Survey (York et al. 2000).

-2.768	60	< 21.4 <sup>b</sup>
-1.706	60	< 20.4 <sup>b</sup>
-1.662	60	< 20.2 <sup>b</sup>
0.000	60	19.897 ± 0.079
0.044	60	19.788 ± 0.081
0.996	60	18.489 ± 0.053
1.043	60	18.455 ± 0.037
1.994	60	17.742 ± 0.032
2.068	60	17.803 ± 0.035
3.019	60	17.283 ± 0.033
3.064	60	17.270 ± 0.034
4.075	60	16.904 ± 0.024
4.122	60	16.879 ± 0.023
5.073	60	16.676 ± 0.025
5.117	60	16.660 ± 0.034
6.065	60	16.496 ± 0.026
6.108	60	16.464 ± 0.046
7.078	60	16.360 ± 0.031
7.122	60	16.349 ± 0.025
8.058	60	16.297 ± 0.025
8.106	60	16.297 ± 0.030
9.148	60	16.224 ± 0.054
9.193	60	16.238 ± 0.049
10.149	60	16.230 ± 0.039
10.193	60	16.204 ± 0.050
11.060	60	16.222 ± 0.045
11.103	60	16.223 ± 0.042
12.126	60	16.224 ± 0.033
12.170	60	16.234 ± 0.027
13.146	60	16.267 ± 0.022
13.190	60	16.284 ± 0.020
14.068	60	16.324 ± 0.025
14.143	60	16.338 ± 0.034
15.096	60	16.391 ± 0.035
15.152	60	16.384 ± 0.11
17.134	60	16.507 ± 0.090
17.180	60	16.554 ± 0.028
18.215	60	16.707 ± 0.033

18.979	60	$16.777 \pm 0.042$
19.023	60	$16.778 \pm 0.035$
19.974	60	$16.864 \pm 0.038$
20.053	60	$16.885 \pm 0.041$
24.035	60	$17.321 \pm 0.031$
24.078	60	$17.345 \pm 0.026$
25.050	60	$17.427 \pm 0.029$
25.095	60	$17.427 \pm 0.029$
26.050	60	$17.529 \pm 0.031$
26.096	60	$17.530 \pm 0.036$
27.060	60	$17.519 \pm 0.030$
27.109	60	$17.627 \pm 0.032$
28.068	60	$17.717 \pm 0.028$
28.133	60	$17.695 \pm 0.034$
29.195	60	$17.678 \pm 0.027$
29.239	60	$17.799 \pm 0.028$
31.015	60	$17.901 \pm 0.046$
31.059	60	$17.903 \pm 0.030$
32.047	60	$17.946 \pm 0.031$
32.092	60	$17.965 \pm 0.046$
33.055	60	$18.231 \pm 0.055$
33.106	60	$18.034 \pm 0.053$
43.949	60	$18.563 \pm 0.046$
43.992	60	$18.638 \pm 0.036$
44.950	60	$18.649 \pm 0.042$
44.993	60	$18.537 \pm 0.031$
45.949	60	$18.680 \pm 0.054$
45.991	60	$18.679 \pm 0.055$
48.023	60	$18.796 \pm 0.041$
48.191	60	$18.802 \pm 0.051$
53.122	60	$18.906 \pm 0.046$
53.215	60	$19.008 \pm 0.060$
60.096	60	$19.140 \pm 0.057$
60.139	60	$18.985 \pm 0.051$
61.952	60	$19.268 \pm 0.072$
61.996	60	$19.131 \pm 0.043$
62.952	60	$19.301 \pm 0.071$
62.997	60	$19.240 \pm 0.064$

63.950	60	$19.321 \pm 0.063$
63.994	60	$19.178 \pm 0.060$
64.982	60	$19.300 \pm 0.066$
65.026	60	$19.210 \pm 0.085$
65.980	60	$19.311 \pm 0.067$
66.023	60	$19.423 \pm 0.079$
73.059	60	$< 19.9^c$
77.061	60	$19.788 \pm 0.094$
77.106	60	$19.65 \pm 0.11$
78.946	60	$19.879 \pm 0.084$
78.990	60	$19.961 \pm 0.087$
80.029	60	$< 17.9^c$
83.975	60	$20.10 \pm 0.11$
84.019	60	$20.07 \pm 0.11$
86.022	60	$19.776 \pm 0.075$
90.968	60	$20.242 \pm 0.18$
91.019	60	$< 20.3^c$
91.973	60	$19.95 \pm 0.11$

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### 3. Spectral classification

After rapidly identifying PTF 10vgv, we triggered our follow-up programs (Gal-Yam et al. 2011). On 2010 September 16 and October 1, we observed PTF 10vgv with the dual-arm Kast spectrograph (Miller & Stone 1993) on the 3 m Shane telescope at Lick Observatory (Figure 3). We used a 2'' wide slit, a 600/4310 grism on the blue side, and a 300/7500 grating on the red side, yielding full width at half-maximum intensity (FWHM) resolutions of  $\sim 4 \text{ \AA}$  and  $\sim 10 \text{ \AA}$ , respectively. All observations were aligned along the parallactic angle to reduce differential light losses (Filippenko 1982). Respective exposure times and air masses were 1800 s and 1.03 for the first epoch, and 2100 s and 1.00 for the second epoch. The spectra were reduced using standard

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<sup>a</sup>Magnitudes are not corrected for Galactic extinction, and are calibrated to the SDSS  $r$  band (SDSS is estimated to be on the AB system within  $\pm 0.01$  mag in the  $r$  and  $i$  bands).

<sup>b</sup> $3\sigma$  upper limit computed by simulating stars at the position of PTF 10vgv, to account for the presence of the underlying host galaxy.

<sup>c</sup> $3\sigma$  upper limit.

techniques (e.g., Foley et al. 2003) based on IRAF and IDL routines. Using the Kast spectra we derive a redshift of  $z = 0.0142 \pm 0.0002$  (using  $H\beta$ , O III,  $H\alpha$ , N II, and S II lines) for PTF 10vgv.

On 2010 September 27, in between the two epochs of the Kast observations, we observed PTF 10vgv with the Low Resolution Spectrograph (LRS) mounted on the Hobby-Eberly Telescope (HET), using the gr300 grating and GG385 filter. We applied bias- and flat-field corrections using daytime calibration frames, and removed cosmic rays using the IRAF task “L.A. Cosmic” (van Dokkum 2001). The spectrum was extracted and wavelength-calibrated using the “apall” and “identify” IRAF tasks, respectively, and had exposure time 450 s at a mean airmass of 1.26.

On 2010 October 30 and December 07, we observed PTF 10vgv with the Double Beam Spectrograph (DBSP; Oke & Gunn 1982) on the Palomar 200-inch telescope (P200; Figure 3). We used the 600/4000 and the 158/7500 gratings for the blue and red cameras, respectively, with a D55 dichroic, resulting in a spectral coverage of  $\sim 3500\text{--}9500 \text{ \AA}$ . The spectra were reduced using a custom pipeline combining IRAF and IDL scripts. Respective exposure times and air masses were 600 s and 1.1 for the first epoch, and 350 s and 1.04 for the second epoch.

We measured the velocity of the Si II absorption at  $6355 \text{ \AA}$ , which traces reasonably closely the position of the photosphere (e.g., Tanaka et al. 2008), using the spectra of PTF 10vgv taken on September 16, September 27, and October 1. The velocities are  $16 \times 10^3 \text{ km s}^{-1}$ ,  $9 \times 10^3 \text{ km s}^{-1}$ , and  $6 \times 10^3 \text{ km s}^{-1}$ , respectively, for the three epochs. These are comparable to those of the “normal” SN Ic 1994I at similar epochs (Sauer et al. 2006),  $\sim 0.7$  times those measured for SN 2006aj (associated with X-ray flash 060218; Mazzali et al. 2006), and smaller than those of the gamma-ray burst (GRB)-associated SN 1998bw (Iwamoto et al. 1998) and SN 2003dh (Mazzali et al. 2003, see also Figure 5 in Corsi et al. 2011). The broad-line SN Ic 2002ap also showed higher velocities ( $\geq 16 \times 10^3 \text{ km s}^{-1}$  at  $\sim 1$  week after the explosion; Gal-Yam et al. 2002; Mazzali et al. 2002).

We thus classify PTF 10vgv as a normal Type Ic SN. A cursory examination of PTF 10vgv spectra suggests that the blending of lines in this SN is stronger than in both SN 2006aj and SN 1994I, indicating that in PTF 10vgv there may be significantly more mass at  $v \approx 2 \times 10^4 \text{ km s}^{-1}$ . In Figure 3, we compare our spectra of PTF 10vgv with the one of SN 1994I around maximum light.

#### 4. Radio Follow-up Observations

Starting on 2010 October 7.16, we observed the position of PTF 10vgv (along with the necessary calibrators) with the Expanded Very Large Array (EVLA; Perley et al. 2009) in its C configuration, at 4.495 GHz and 7.915 GHz, for a total time of 30 min (Corsi et al. 2010b). We detected no radio emission from the position of PTF 10vgv, down to  $3\sigma$  limits of  $120 \mu\text{Jy}$  at 4.495 GHz and

102  $\mu\text{Jy}$  at 7.915 GHz. Based on this, we estimate the 5 GHz spectral luminosity of PTF 10vgv to be  $\lesssim 5 \times 10^{26} \text{ erg s}^{-1} \text{ Hz}^{-1}$ , or  $\sim 100$  times below the radio luminosity of the GRB-associated SN 1998bw (Kulkarni et al. 1998) on a similar timescale. This supports the idea that PTF 10vgv is a normal SN Ic, rather than an engine-drive SN (Berger et al. 2003; Soderberg et al. 2006). We reobserved PTF 10vgv with the EVLA in its BnA configuration starting on 2011 May 12.52, for a total time of 1 hr and at a central frequency of 8.46 GHz. No radio sources were detected in the error circle of PTF 10vgv down to a  $3\sigma$  limit of 30  $\mu\text{Jy}$ . EVLA data were reduced and imaged using the AIPS software package.

## 5. Discussion

The measured peak magnitude of PTF 10vgv (see Table 1) corrected for Galactic extinction ( $A_R \approx 0.49$  mag; Schlegel et al. 1998) gives  $M_R = -18.20 \pm 0.05$  mag. The peak absolute magnitude of SN 1994I was  $M_R = -17.99 \pm 0.48$  (Richmond et al. 1996), while SN 2006aj had  $M_R = -18.81 \pm 0.06$  mag (Mazzali et al. 2006; Valenti et al. 2008). Since PTF 10vgv is intermediate, in terms of  $R$ -band peak luminosity, between SN 1994I and SN 2006aj, we estimate its nickel mass  $M_{56\text{Ni},10\text{vgv}}$  by interpolating between these two SNe (Sauer et al. 2006; Mazzali et al. 2006), using the scaling  $L_{\text{peak}} \propto M_{\text{Ni}} \tau_c^{-1}$  for the peak luminosity (where  $\tau_c$  is the light-curve peak width; Arnett 1982), and considering that the PTF 10vgv light curve is a factor of  $\sim 1.25$  broader than that of SN 1994I (while we take the same  $\tau_c$  for PTF 10vgv and SN 2006aj). This yields  $M_{56\text{Ni},10\text{vgv}} \approx 0.12 M_{\odot}$ .

The mass and kinetic energy of the SN ejecta scale as (Arnett 1982)  $M_{\text{ej}} \propto \tau_c^2 v_{\text{ph}}$  and  $E_K \propto \tau_c^2 v_{\text{ph}}^3$ , where  $v_{\text{ph}}$  is the photospheric velocity. Using these scalings, and considering that the photospheric velocities of PTF 10vgv are comparable to those of SN 1994I and  $\sim 0.7$  times those of SN 2006aj (§3), we estimate the ejecta mass and kinetic energy of PTF 10vgv interpolating between SN 2006aj (Mazzali et al. 2006) and SN 1994I (Sauer et al. 2006). We get  $M_{\text{ej},10\text{vgv}} = (1.5 \pm 0.3) M_{\odot}$  and  $E_{K,10\text{vgv}} = (0.9 \pm 0.3) \times 10^{51}$  erg. This estimate may be refined through spectral modeling. Our spectral analysis suggests that a different mass-velocity distribution may be realized in PTF 10vgv (§3), which may lead to a larger  $E_K$  than estimated on the basis of the light-curve properties, since the latter are mostly determined by the opacity in the inner ejecta.

Our pre-discovery upper limits can be used to constrain the radius of the stellar progenitor of PTF 10vgv via comparison with model predictions (Dessart et al. 2011; Rabinak & Waxman 2011; Nakar & Sari 2010). We apply a bolometric correction to our  $R$ -band data, computed assuming that



the SN emits as a black body at temperature  $T_{\text{phot}}$  (and neglecting redshift corrections):

$$M_{\text{bol}} - M_R = -2.5 \log_{10} \left( \frac{4\pi (10 \text{ pc})^2 F_0 \int_{\nu_1}^{\nu_2} S(\nu) d\nu}{L_{\odot}} \right) + M_{\text{bol},\odot} + 2.5 \log_{10} \left( \frac{\int_{\nu_1/kT}^{\nu_2/kT} S(x) x^3 (e^x - 1)^{-1} dx}{\pi^4/15} \right), \quad (1)$$

where  $M_{\text{bol},\odot} = 4.72$ ;  $S(\nu)$  is the P48 Mould- $R$  filter transmission;  $\nu_1 - \nu_2 = (4.1 - 5.3) \times 10^{14}$  Hz; and  $F_0$  is the photometric zero-point flux ( $F_0 = 3.631 \times 10^{-20}$  erg cm $^{-2}$ s $^{-1}$ Hz $^{-1}$  for AB magnitudes). We conservatively maximize the bolometric correction setting  $T_{\text{phot}} \approx 10^4$  K, the largest early-time ( $t \lesssim 10$  d since breakout) temperature predicted by the  $^{56}\text{Ni}$ -rich models of Dessart et al. (2011; see their Figure 2, bottom-left panel). In this way we get  $M_{\text{bol}} - M_R = -0.496$  mag.

The optical luminosity of core-collapse SNe after breakout depends on the ejecta composition (via the opacity parameter), the stellar radius, and the  $E_K/M_{\text{ej}}$  ratio. A larger  $E_K/M_{\text{ej}}$  ratio and a lower He fraction both increase the predicted luminosity, for a given stellar radius (Rabinak & Waxman 2011, Equations (25) and (29)).

In recent numerical simulations of core-collapse explosions of single and binary progenitors of SNe IIb/Ib/Ic, Dessart et al. (2011) predicted the existence a  $\sim 10$ -day-long ( $\sim 10$  times shorter than in SNe II-P) post-breakout<sup>3</sup> plateau, with a luminosity of  $(1 - 5) \times 10^7 L_{\odot}$  ( $\sim 10$  times smaller than in SNe II-P). This plateau has the same origin as that observed in SNe II-P<sup>4</sup>, but in the case of SNe IIb/Ib/Ic it is predicted to have a smaller duration and luminosity because of a more compact progenitor.

For PTF 10vgv we can exclude the presence of a post-breakout plateau with luminosity greater than the one of the compact progenitor model Bmi25mf5p09z1 (Figure 2, lower panel). We thus derive  $R \lesssim 4.4 R_{\odot}$  for the radius of PTF 10vgv progenitor. However, the stellar models analyzed by Dessart et al. (2011) have  $E_K/M_{\text{ej}}$  lower than we derive here, and a high surface He fraction. So the bound on the progenitor radius derived from the comparison with these models is likely over-estimated.

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<sup>3</sup>The breakout of a shock through the stellar surface is predicted to be the first electromagnetic signal marking the birth of a SN (e.g., Klein & Chevalier 1978; Chevalier 1992; Waxman et al. 2007; Nakar & Sari 2010; Rabinak & Waxman 2011).

<sup>4</sup>The plateau is associated with a cooling and recombination wave (CRW) propagating downward through the SN envelope, separating almost recombined outer layers from strongly ionized inner ones (e.g., Nadyozhin 2003). During the plateau phase, the photosphere sits on the upper edge of the CRW front, whose downward speed is approximately equal to the outward expansion velocity, thus  $R_{\text{phot}} \approx \text{const}$ . Since also  $T_{\text{phot}} \approx \text{const} \approx T_{\text{recomb}}$  (where  $T_{\text{recomb}}$  is the recombination temperature), a plateau in the luminosity is expected.

Similar limits ( $R \lesssim 5 R_{\odot}$ ) can be derived using the predictions by Nakar & Sari (2010; black line in their Figure 3). But this model is accurate only up to  $\lesssim 11$  hr after the explosion, since recombination is not treated.

The tightest constraint,  $R \lesssim 1 R_{\odot}$ , is derived from the C/O model of Rabinak & Waxman (2011), using  $M_{\text{ej},10\text{vgv}} = (1.5 \pm 0.3) M_{\odot}$  and  $E_{\text{K},10\text{vgv}} = (0.9 \pm 0.3) \times 10^{51}$  erg, as derived above.

Applying this same analytical model to the first clear detection of SN 1994I (Sauer et al. 2006, Figure 8; Richmond et al. 1996, Figure 7), we get  $R \lesssim 1/4 R_{\odot}$ , considering that  $M_{\text{ej},1994\text{I}} \approx M_{\text{ej},10\text{vgv}}$ ,  $E_{\text{K},1994\text{I}} \approx E_{\text{K},10\text{vgv}}$ , and that the luminosity of SN 1994I at the time of detection was  $\approx 3$  times smaller than the one of PTF 10vgv.

Our limits for PTF 10vgv,  $R \lesssim (1 - 5) R_{\odot}$ , are consistent with a small Wolf-Rayet star (e.g., Crowther 2007), as expected for a highly stripped SN Ic. PTF 10vgv thus provides the first constraint on the progenitor radius of a SN ever obtained from optical pre-explosion limits extending up to a week before discovery. Optical surveys with rapid cadence and relatively deep exposures (like PTF) should allow us to study many more objects in this manner.

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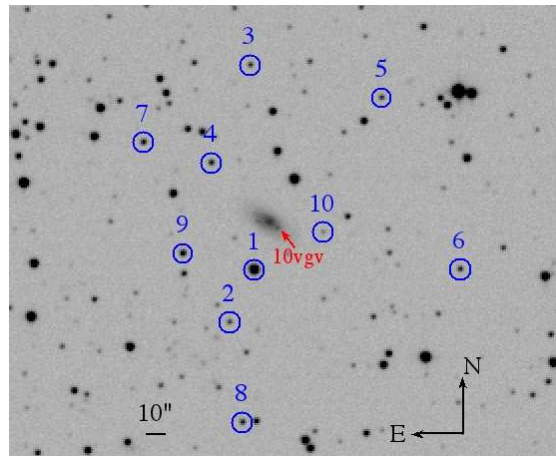


Fig. 1.— Discovery image of PTF 10vgv (marked with a red arrow) in the  $R$  band; the host galaxy is also visible. Circles of  $5''$  radius mark the positions of the ten reference stars used for calibration of the P48 photometry (see text).

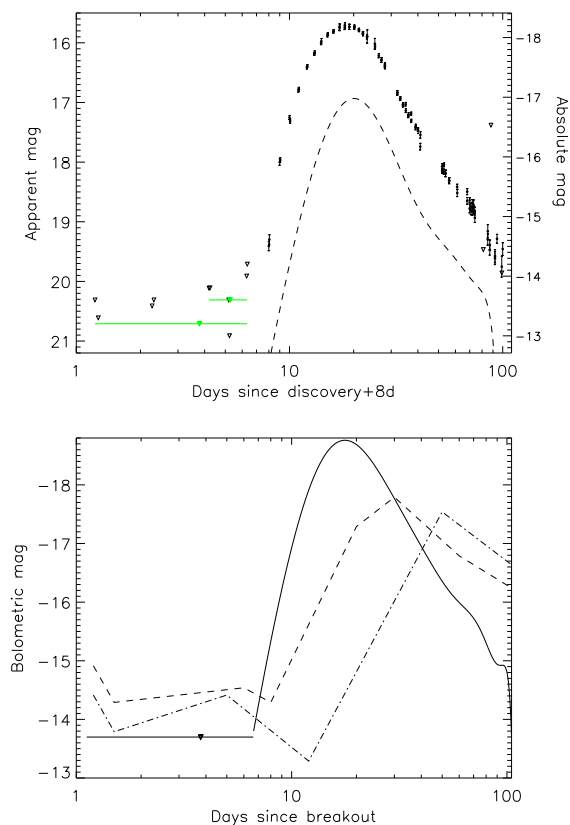


Fig. 2.— *Top*: P48  $R$ -band light curve of PTF 10vgv (black dots) corrected for Galactic extinction. P48 pre-discovery upper limits derived using 60 s exposure images are plotted as black triangles. Deeper upper limits obtained by coadding the pre-explosion images are plotted as green triangles, with the green horizontal lines indicating the time range spanned by the coadded images. For comparison, we also plot the light curve of SN 1994I (dashed line), rescaled to the redshift of PTF 10vgv. *Bottom*: Schematic representations of the bolometric light curves of models Bmi18mf3p79z=1 (dashed line) and Bmi25mf5p09z1 (dash-dotted line) of Dessart et al. (2011) are compared with the PTF 10vgv bolometric light curve (solid line). The black triangle and solid horizontal line indicate our deepest pre-explosion coadded upper limit (see upper panel), rescaled to account for the bolometric correction (and for Galactic extinction). See §5 for discussion.

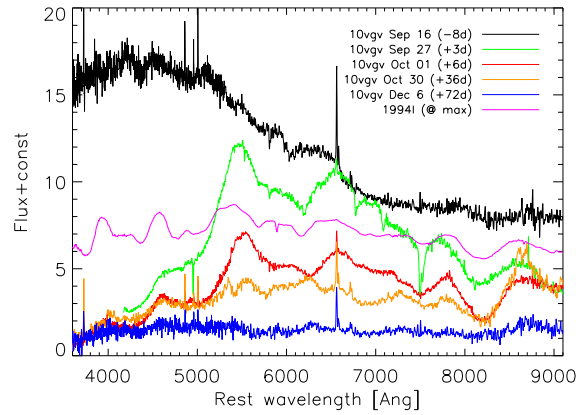


Fig. 3.— Spectra of PTF 10gvv from Lick/Kast (black and red lines), HET/LRS (green line; telluric absorption lines not removed), and P200/DBSP (orange and blue lines). The approximate epoch since the *R*-band peak is also indicated for each spectrum. For comparison, the spectrum of SN 1994I around maximum light is also shown (magenta). All data are available in digital form from the Weizmann Institute of Science Experimental Astrophysics Spectroscopy System at <http://www.weizmann.ac.il/astrophysics/wiseass/>.