

# A wide Chandra view of the core of the Perseus cluster

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26 May 2011

## ABSTRACT

We present new Chandra images of the X-ray emission from the core of the Perseus cluster of galaxies. The total observation time is now 1.4 Ms. New depressions in X-ray surface brightness are discovered to the north of NGC 1275, which we interpret as old rising bubbles. They imply that bubbles are long-lived and do not readily breakup when rising in the hot cluster atmosphere. The existence of a 300 kpc long NNW–SSW bubble axis means there cannot be significant transverse large scale flows exceeding  $100 \text{ km s}^{-1}$ . Interesting spatial correlations are seen along that axis in early deep radio maps. A semi-circular cold front about 100 kpc west of the nucleus is seen. It separates an inner disturbed region dominated by the activity of the active nucleus of NGC 1275 from the outer region where a subcluster merger dominates.

**Key words:** X-rays: galaxies — galaxies: clusters — intergalactic medium — galaxies: individual (NGC 1275)

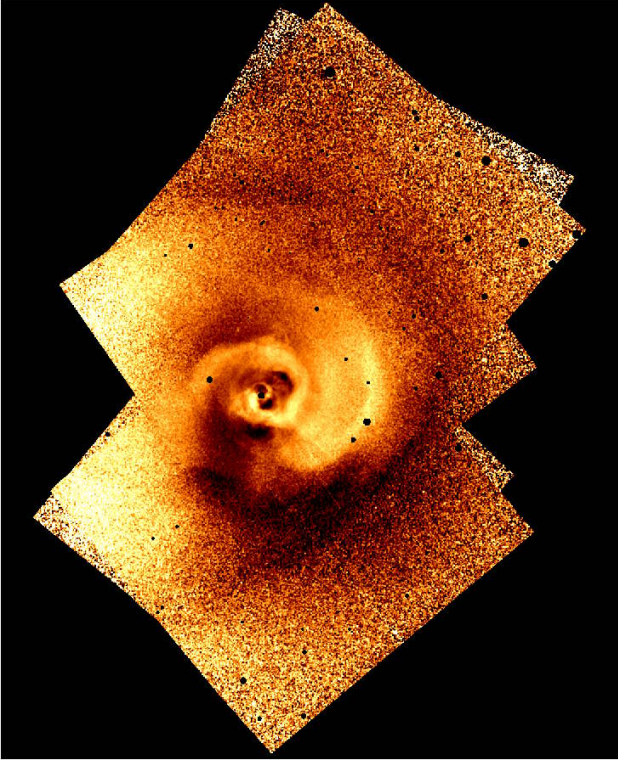
## 1 INTRODUCTION

New large-scale Chandra data on the core of the Perseus cluster, A 426, are presented here. They follow our earlier 900 ks ACIS-S image which covered a  $8 \times 8$  arcmin region centred on the nucleus of the dominant galaxy NGC 1275 (Fabian et al 2000; Fabian et al 2006; Sanders & Fabian 2007). That image, with hints from neighbouring CCD chips, indicated that bubbles, filaments and other fine-scale structure extended over a wider field. The new data extend the area imaged in detail to a region over 30 arcmin by 15 arcmin. XMM observations of the Perseus cluster, of lower spatial resolution but covering an even wider area than that discussed here, have been presented and discussed by Churazov et al (2003).

The Perseus cluster is the X-ray brightest cluster in the Sky (Forman et al 1973), having peaked emission first resolved with the Copernicus satellite (Fabian et al 1974). Einstein X-ray imag-

ing showed a dip in the X-ray emission to the NW of the nucleus of the central galaxy, NGC 1275, (Branduardi-Raymont et al 1981; Fabian et al 1981). ROSAT HRI data (Böhringer et al 1993) showed that this dip was an outer bubble in the hot X-ray emitting gas and resolved two inner bubbles coincident with the FRI radio source, 3C84 (Pedlar et al 1990). Churazov et al (2000) showed that the bubbles are inflated by the radio jets and rise buoyantly in the surrounding hot gas. The NW bubble was formed a few  $10^7$  yr ago and is now moving outward. Estimates of the likely energies and timescales indicated the power in the bubbling process to be of the order of  $10^{45} \text{ erg s}^{-1}$ . Chandra images resolve the bubbles in detail (Fabian et al 2000) and show subtle ripples centred on the inner bubbles which are interpreted as weak shocks or sound waves produced by the bubbles (Fabian et al 2003a). The shape of the NW bubble resembles the cross-section of a rising spherical cap bubble in water on Earth (Fabian et al 2003b), despite being over  $10^{22}$  times larger. Rising bubbles drag out cooler X-ray emitting gas (as in the Virgo cluster, Churazov et al 2001) from the centre and cold

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**Figure 1.** Mosaic of new images only, exposure corrected and adaptively smoothed. The light streaks to the SW direction are due to the CCD readout of the bright nucleus. The separate ACIS-I pointings are evident, each has a side of 16 arcmin.

gas (the  $H\alpha$  filaments of NGC 1275: Lynds 1970; Conselice et al 2001; Fabian et al 2003b; Fabian et al 2008).

In this paper we present evidence for further structures along the bubbling axis which may be older, outer bubbles, or merged bubbles.

The cluster is assumed to have a redshift of 0.0183, corresponding to an angular scale of  $0.37 \text{ kpc arcsec}^{-1}$  ( $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ).

## 2 THE NEW IMAGES

The new ACIS-I images have been mosaicked together into a single image 0.5–7 keV shown in Fig. 1. The average at each radius has been subtracted for display purposes. The central structure is associated with the bubbles formed by the radio jets from the active nucleus and is detailed in our earlier work. The separate ACIS-I chips are  $16 \times 16$  arcmin and the whole image is 35.5 arcmin from N to S (788 kpc).

We have added the earlier 900 ks ACIS-S images to make the final composite, Fig. 2. The nucleus of most member galaxies of the Perseus cluster are detected (see Santra et al 2007). The point source in the image to the WSW of the nucleus of NGC 1275 (about halfway to the edge) corresponds to the nucleus of the optically second-brightest cluster galaxy, NGC 1272 (see Fig. 3). It sits on a rim of X-ray bright cluster emission. This rim is seen as a step at 6–7 arcmin radius from the nucleus of NGC 1275 in Fig. 4 and 5, and is roughly circular in shape with a centre displaced 1 arcmin to the NW of that nucleus. We return to the large scale structure to the W in Section 6.

Fig. 3 right was produced with the 0.5 m telescope at the Blackbird Observatory, located in the south central Sacramento Mountains of New Mexico USA. The image represents over 31 hr of cumulative exposure time through broadband clear, red, green, blue and 6 nm narrow band filters. Each sub-exposure was reduced following standard procedures for bias correction and flat fielding. Master dark and bias frames were created by combining 10 dark sub-exposures each produced at the same exposure length and camera temperature settings used for the luminance and the filtered images. A master flat was produced by combining 10 separate sky flat exposures for each filter. The clear, red, green, blue and 6 nm narrowband filtered exposures were separately combined (using a median procedure) to produce individual master data sets. Each master data set was then projected into a gray scale image, intensity scaled, assigned to its appropriate primary color then digitally combined with the other data sets to produce a full color picture. The 6 nm narrowband data set was used to supplement the intensity of the red broadband color channel; the clear data set was used to intensify luminance information for the combined data sets. Post processing methods that resulted with the final image followed standard procedures that are also employed to create Hubble Heritage images (Rector et al 2007).

## 3 NEW STRUCTURES

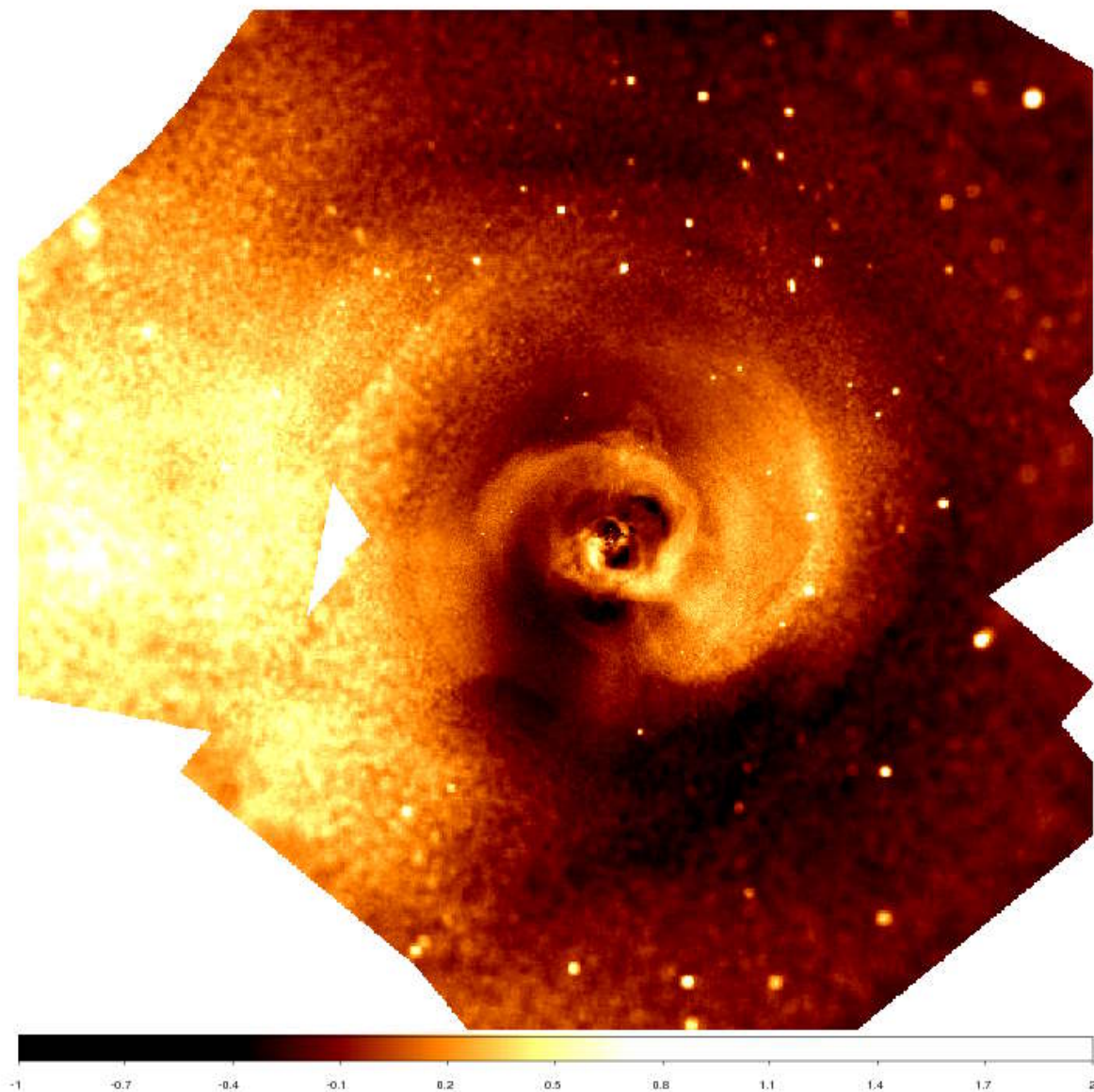
Overlays of the new composite X-ray image with the  $H\alpha$  filaments around NGC 1275 and deep radio data at 49 cm (Sijbring 1993) are shown in Figs. 7 and 8. As noted by Fabian et al (2006) there is a looplike X-ray structure at the end of the long Northern  $H\alpha$  filament. This could be fallback of the outer parts of the cooler gas dragged out to the N by earlier bubbles, when the optical filament was created.

Further structures to the N are then expected and indeed observed. The radio image shows an extension at the end of the optical filament above which there is a spur leading to a dark patch in the X-ray image, corresponding to a drop in flux, which we call the Northern trough. (Burns et al 1994 published a deep 330 MHz VLA radio image of Perseus which shows similar outer structure, lending credence to the details being real.) Just to the W of the radio spur lie two elliptical-shaped dips in the X-ray image which we identify as old outer bubbles. The outer of these was found originally on an outer chip of the deepest ACIS-S observations (Sanders & Fabian 2007). These structures are indicated in Fig. 9 and shown in profile in Fig. 10.

A bay is seen to the SSE of NGC 1275 (Fabian et al 2006), along the bubble axis (Figs. 8 and 9). The radio contours indicate a coincident minimum in radio flux.

### 3.1 The Northern Trough and Bubbles

The N trough (structure labelled A in Fig. 6) lies 9 arcmin (200 kpc) N of the nucleus. We have studied it through profiles made by spectral analysis along a sector to the N (Fig. 12). It is about 10 percent deep in surface brightness and likely lies along, or close to, an equipotential. It is also along the main jet axis to the N where within the inner 40 kpc or so there is a pronounced optical  $H\alpha$  filament (Figs. 3 and 7). If such a filament has been dragged out by



**Figure 2.** Final composite image of the 1.4 Ms exposure in the 0.5–7 keV band, adaptively smoothed with a kernel with a fixed number of counts (225 in each of 3 energy bands, which have then been summed) and with the average at each radius subtracted. The colour shows fractional variation. The image is 25.6 arcmin from N to S. The optically second brightest galaxy in the cluster, NGC 1272, lies along the bright rim 5 arcmin (111 kpc) W of the nucleus.

rising large bubbles (see Fabian et al 2003; Hatch et al 2006) then it is plausible that the trough could be the remains of those bubbles<sup>1</sup>.

We can start to explore a bubble origin by comparing the volume of the trough with that of one of the observed bubbles near the centre. A rough estimate of the volume of the trough can be made assuming that it is an ellipsoid with its major axis, of length 2.5 arcmin, i.e. 55 kpc, in the plane of the sky and radius 0.7 arcmin, i.e. 15.5 kpc. This provides a 10 per cent dip in surface brightness, as observed, if the integrated depth along the line of sight is equivalent to the central value multiplied by the radius from the nucleus.

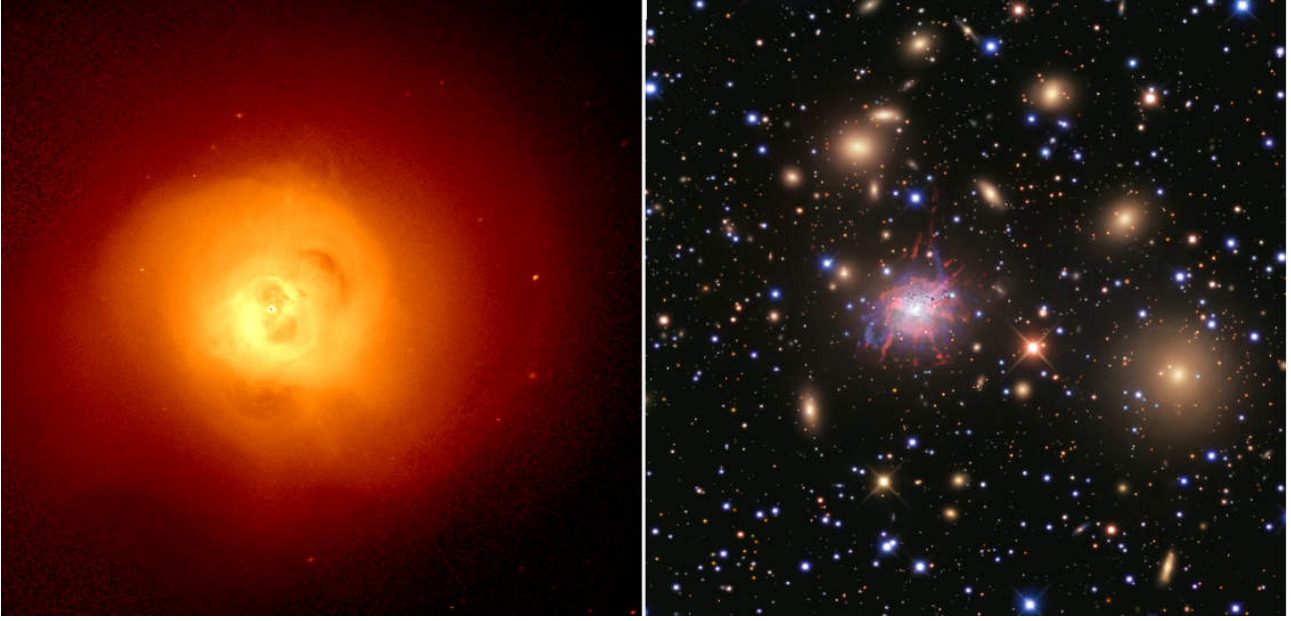
<sup>1</sup> The possibility that the jet has changed direction or precessed is discussed by Dunn, Fabian & Sanders (2006). The mean axis over the past  $\sim 5 \times 10^8$  yr remains along the NNW-SSW direction.

This volume is about  $5.6 \times 10^4 \text{ kpc}^3$ , or about 21.5 times that of the inner N bubble, which has a radius of 8.5 kpc.

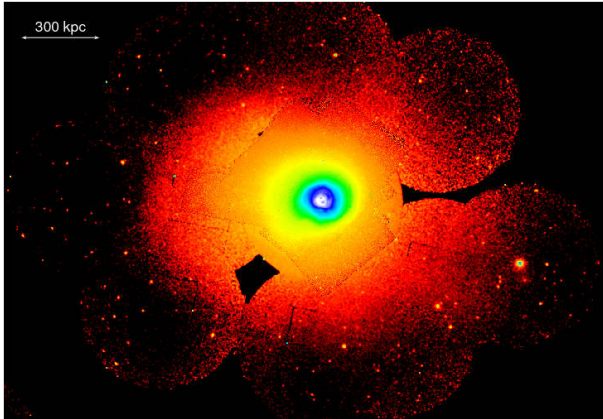
The external pressure drops by a factor of 3 between the centre and 10 arcmin radius where the trough resides (Fig. 12), which is a pressure of about  $n_e T_e \sim 7 \times 10^5 \text{ K cm}^{-3}$ . Adiabatic expansion then gives a factor of about 2 increase in volume, so we are looking at the accumulation of about 11 bubbles similar to the present inner ones (assuming the contents are preserved during the rise). The energy content of the trough,  $4PV = 1.2 \times 10^{60} \text{ erg}$ . This is one and two orders of magnitude smaller than the events that have occurred in the Hydra A (Nulsen et al 2005) and MS 0735.6+7421 (McNamara et al 2009) clusters, respectively.

The possibility arises that the trough represents the accumulation of the last 500 Myr or so of bubbles (a similar structure is discussed for A2204 by Sanders, Fabian & Taylor 2009). Bubbles rise





**Figure 3.** Matching X-ray and optical images of the core of the Perseus cluster Left: Chandra composite (from Fig. 2, but without subtraction of the mean at each radius); Right: optical from Blackbird Observatory (see text for details). The images are 11.8 arcmin from N to S. NGC 1272 is the bright elliptical galaxy 5 arcmin WSW of NGC 1275.



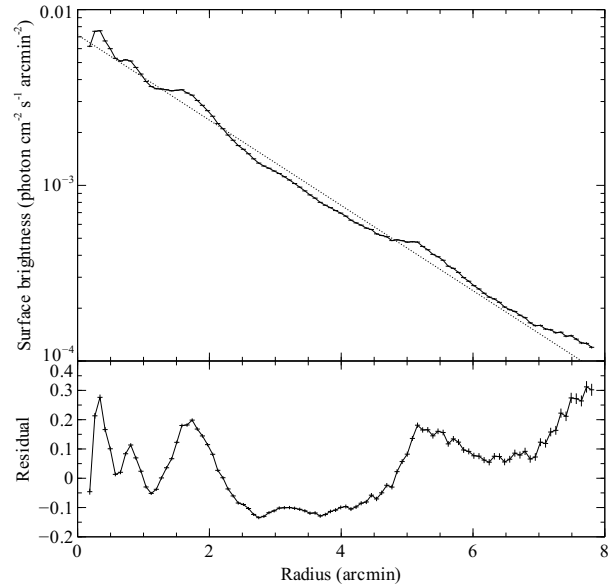
**Figure 4.** Joint Chandra and XMM image.

and are trapped at some radius, in this case at about 220 kpc. Perhaps they become neutrally buoyant there due to mixing with surrounding gas, or the magnetic structure (possibly azimuthal there; Quataert et al 2008) traps them. There also seems to be an overall structure at and just within that radius to the W, possibly due to motion of the core relative to the outer cluster gas (see e.g. Churazov et al 2003).

The two X-ray surface brightness dips to the SW of the trough (Figs 9 and 10), which we identify as rising bubbles, have volumes of approximately  $10^4 \text{ kpc}^3$  each, corresponding to about twice that of the current inner bubbles.

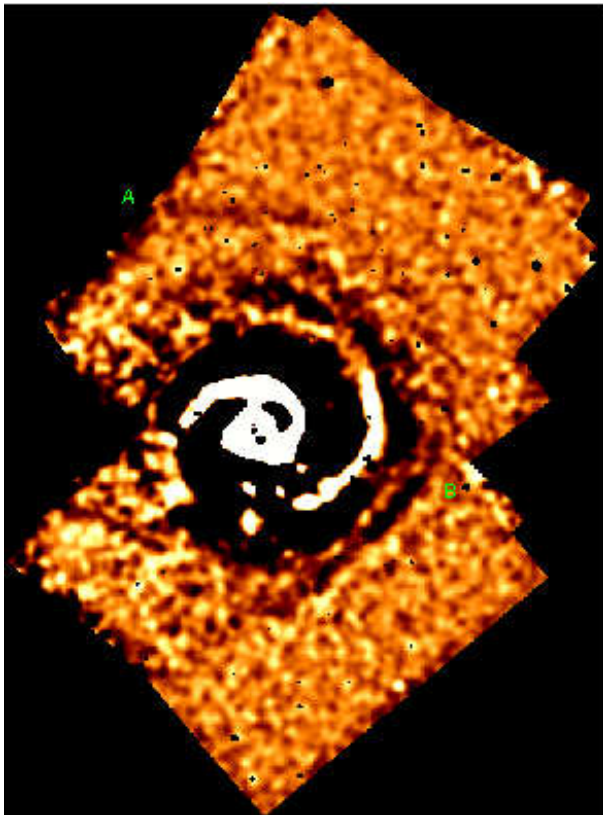
The bay to the South may result from the accumulation of Southward rising bubbles in analogy to the Northern trough. It has a sharp, curved Northern edge and the interior is hotter than the outer parts (Fig. 11). Perhaps there has been some mixing and heating taking place between the relativistic and thermal intracluster gases. It lies much closer to the nucleus of NGC 1275.

The evolution of rising bubbles in cluster gas has been stud-

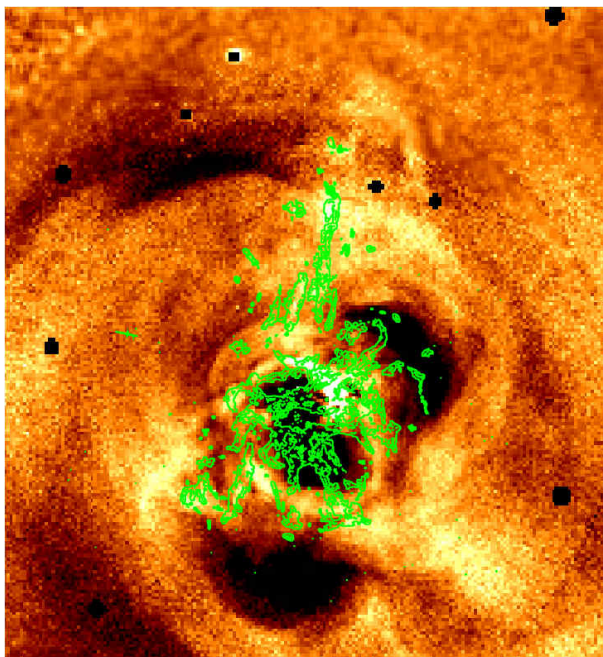


**Figure 5.** Surface brightness profile in the 0.5–7 keV band to the West of the nucleus of NGC 1275.

ied and simulated by many authors (e.g. Diehl et al 2008; Liu et al 2008). Bubbles blown by a jet are not Rayleigh-Taylor unstable because the upper surface of the bubble is not at rest relative to the hot gas above them. The expansion of the bubbles means that the hot gas continuously flows around them. The growth time of the Kelvin-Helmholtz instability is comparable to the flow time. Whether they break up or not depends on the amplitude and scale of velocity perturbations in the hot gas. The stability of a large gas bubble rising through liquid has been studied by Batchelor (1987). Rising air bubbles in water can be surprisingly large. The scale size of disruptive perturbations depends on surface tension (which in



**Figure 6.** Unsharp masked image using larger regions than for Fig. 1.



**Figure 7.** Inner part of the X-ray image with H $\alpha$  contours (from Conselice et al 2001) overlaid. The image is 6.8 arcmin (150 kpc) from N to S.

the case of cluster bubbles means magnetic field), viscosity and gravitational acceleration. Lyutikov (2006) has proposed that magnetic draping can provide a surface tension. Reynolds et al (2005) have discussed and simulated the role of viscosity in rising bubbles. Higher viscosity leads to stable, long-lived, bubbles.

Diehl et al (2008) show that the bubbles seen at larger radii in several clusters tend to be much larger than those at smaller radii. Larger even than expected by adiabatic expansion. In Hydra A, where 3 pairs of expanding bubbles are identified (Wise et al 2007) there is a progressive steep increase of bubbles radius with distance from the nucleus (see also Randall et al 2011 for a sequence of bubbles in an elliptical galaxy).

Diehl et al (2008) discuss bubbles which are continuously fed energy from the nucleus, which seems unlikely for Perseus. We note that if the rising speed of a bubble depends on its size and if a merger always occurs if two bubbles meet, then a sequence of bubbles of slightly varying size will at larger radii become a set of fewer larger bubbles. The radius at which this happens will depend on the steepness of the bubble size-speed relation and the initial size variation amplitude. It will happen both if larger ones are faster or slower. This could well explain our observations of the outer Perseus bubbles.

It is not clear what is happening to the S of the nucleus. The early observations showed an inner and an outer bubble there. At larger radii there is a sharp edge to part of an elliptical-shaped structure which we term the Bay. This could represent the end point for bubbles rising in that direction. The temperature there (Fig. 11) is hotter. This could either be because there is hotter gas at that radius, or it could be because relativistic plasma has displaced gas and we are only seeing the projected outer hotter intracluster gas. The pseudo-pressure maps (Fig. 13 and 14) do however peak in this region so the first option is preferred. The energy in this cavity, assuming it contains only thermal gas and a prolate-ellipsoidal shape, is  $1.7 \times 10^{60}$  erg, slightly more than for the Northern trough.

#### 4 TEMPERATURES, PRESSURES AND METALS

The new data have been contour-binned (Sanders 2007) into regions each with 22,500 counts ( $S/N = 150$ ). The spectra within each bin have been fitted separately resulting in the temperature and metallicity maps shown in Fig. 13. The higher metallicity gas does stretch to the N in the direction of the trough, but no more detailed correspondence is evident. We also show a pseudo-pressure map (obtained from the square root of the mean surface brightness in each region multiplied by the temperature).

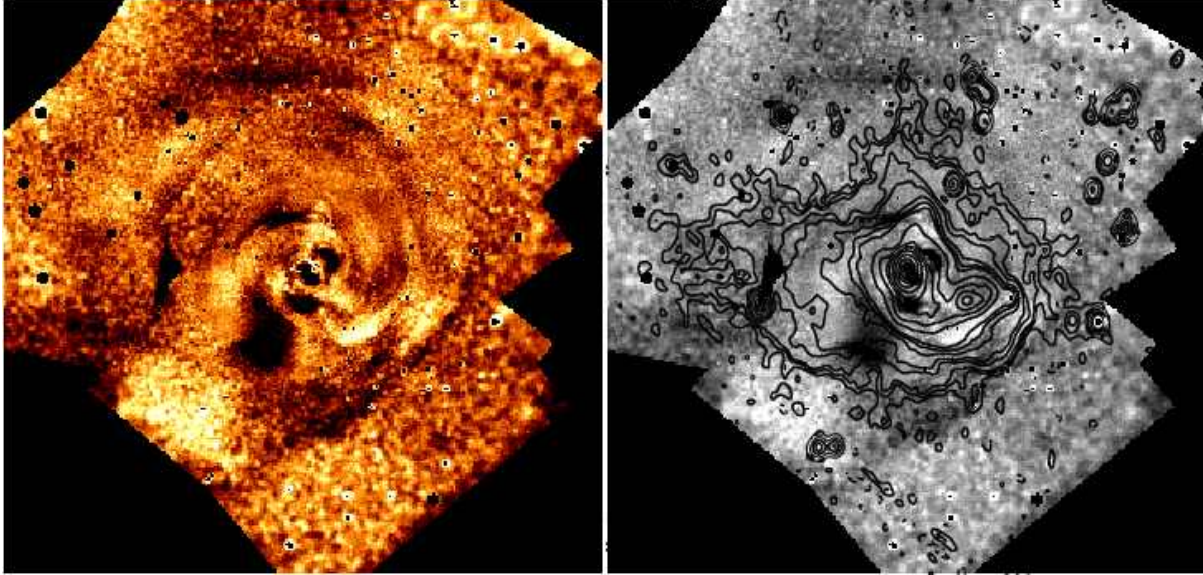
A semicircular cold front is seen to the W. Across this front the temperature jumps from 5.5-6 keV to 7 to 8 keV, and the metallicity drops from 0.55-0.6 to 0.4-0.45. The pressure declines relatively smoothly across the region (see also Figs. 16 and 17 in the Appendix). The smooth pressure change while the temperature jumps is characteristic of a cold front (see Section 6 for further discussion).

Finally, Fig. 14 shows the temperature and pseudo-pressure maps obtained when the mean value at that radius (from the nucleus of NGC 1275) has been subtracted, to highlight small differences.

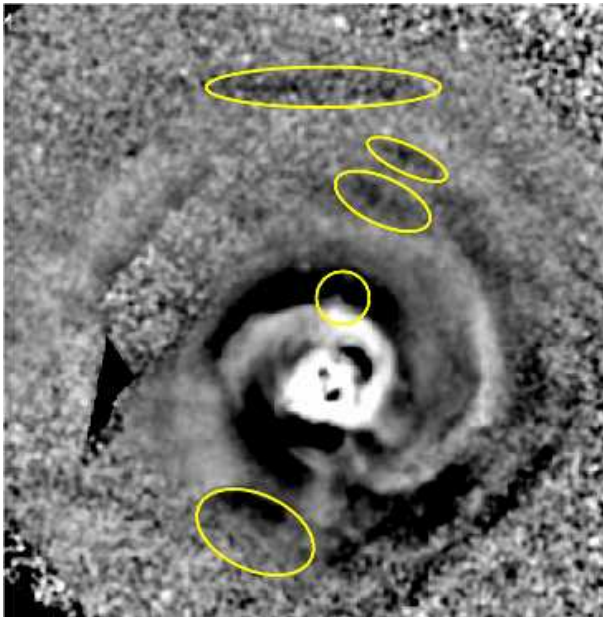
#### 5 RADIO CORRELATIONS

Minihaloes such as 3C84 are rare (Ferrari et al 2008) and may be due to turbulence (Gitti et al 2004). There is a possible association





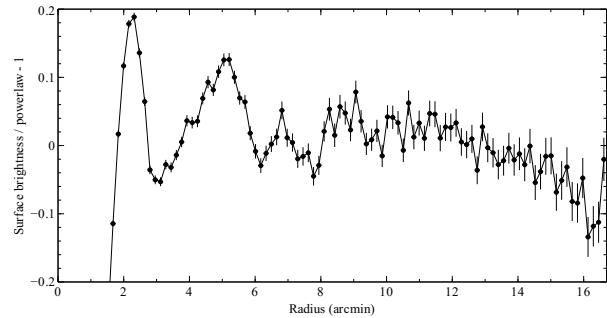
**Figure 8.** Left Chandra X-ray image residuals obtained after subtracting a smooth elliptical model. Ellipses were fitted to surface brightness contours, and an elliptical model created by interpolating between these ellipses. Right: Radio contours (from Sijbring 1993) overlaid on the X-ray image. The radio data of 3C84 were taken at 49 cm (608.5 MHz). Similar structures are seen in the 92 cm map. The radio contours stop abruptly at the edge of the outer ring of X-ray emission to the SW. A spur of radio emission to the North points towards the trough.



**Figure 9.** Structures identified in the composite image: from top to bottom –N trough, two older rising bubbles, top of N H $\alpha$  filament and S “bay” (old bubble?).

between minihaloes and gas sloshing (ZuHone & Markevitch 2011) in cluster cores (Mazzotta & Giacintucci 2008). The swirl seen in the temperature maps is indicative of gas sloshing, perhaps induced by an off-axis merger or at least the close passage of a stripped group core (Churazov et al 2003; Ascasibar & Markevitch 2006; Roediger et al 2011).

We have noted interesting correlations between structures in the radio and X-ray maps (Fig. 8). While there might be small scale turbulence in the hot gas, we doubt that there is considerable hydrodynamic turbulence due to the straightness and ordered structures



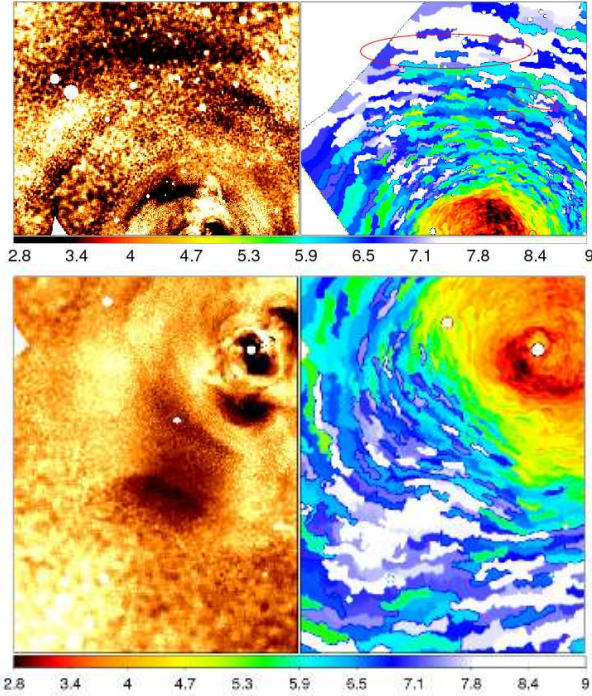
**Figure 10.** Surface brightness profile to the NNW. The dips at  $\sim 6$  and  $7.7$  arcmin are the two Northernmost bubbles.

of the H $\alpha$  filaments. Indeed the 40–50 kpc long Northern filament shows little evidence for turbulent flow.

The large scale linearity of the bubble structures shown in Fig. 8, apart from the Northern trough, is evidence against significant large scale flows or rotation of the core gas. If we conservatively take a minimum transverse motion of 50 kpc (the semi-major axis of the trough) over the time to make the structures, say  $5 \times 10^8$  yr, we obtain a maximum transverse velocity of  $100 \text{ km s}^{-1}$ .

## 6 LARGE-SCALE STRUCTURES AND COLD FRONTS

The W edge, rim and semi-circular structure seen in Figs 1, 2, 5, and 6 appears to separate the inner AGN-dominated region from the outer cluster. As outlined in Section 4, the temperature rises sharply to the W of the structure, whereas the pseudo-pressure shows no abrupt change (Figs 13 – 16). It therefore appears to be a cold-front (for a review see Markevitch & Vikhlinin 2007), and is probably related to a past subcluster merger (Churazov et al 2003 and Section 5).



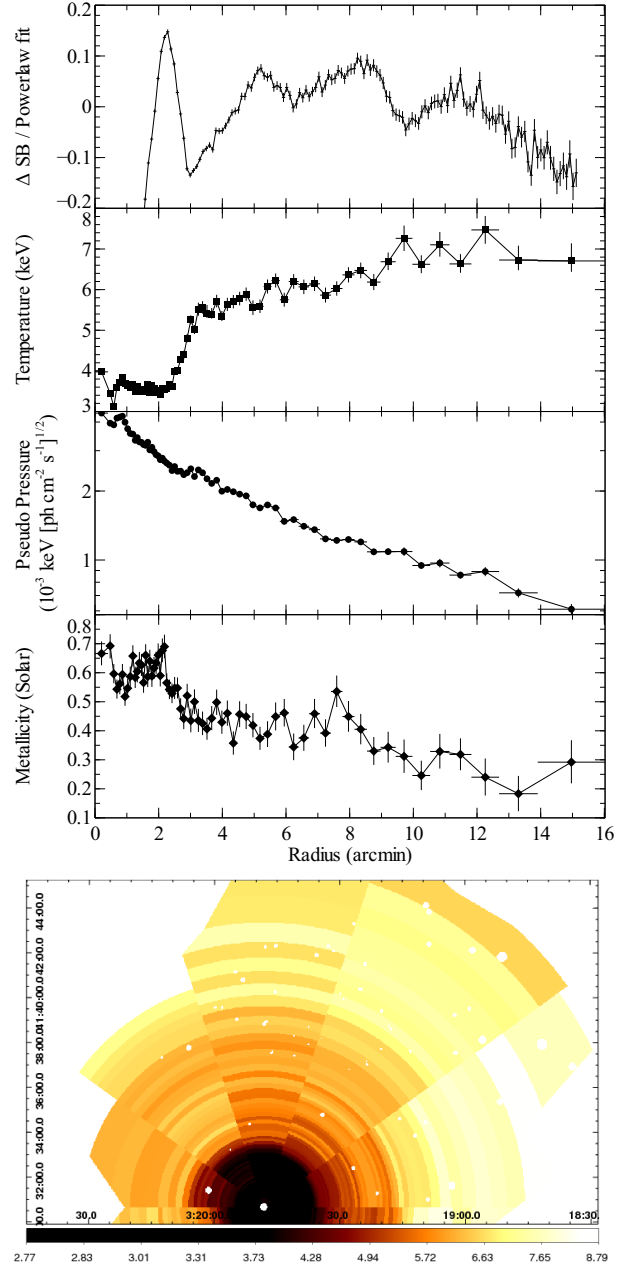
**Figure 11.** Top panel: Surface brightness and temperature maps to the N. Lower temperature gas is seen stretching S to the centre of the N trough. The two possible old bubbles coincide with higher temperature regions. The significance of this is not high since the temperatures have uncertainties of  $\sim 0.5$  keV (???). Lower panel: X-ray surface brightness with average at that radius subtracted (left) and temperature (right) to the SE, showing the “bay” which contains a “tongue” of hotter material. A sharp cold front occurs around the N side of the bay, but, being concave, is not a conventional cold front (Markevitch & Vikhlinin 2008)

Mergers producing slopping/sloshing/bathtub/seich-modes in the core of the Perseus cluster have been discussed for decades (see e.g. Allen et al 1992), since Einstein Observatory imaging showed the overall emission to be lop-sided (Branduardi-Raymont et al 1982). Such modes can persist for a considerable time (Gyrs). Cold fronts can form in the process (e.g. Birnboim, Keshet & Herquist 2010). Tangential flows below the front and magnetic fields amplification are expected (Keshet et al 2010). Once again we note the lack of evidence for significant transverse flows in the relatively straight  $H\alpha$  filament system or indeed in the NNW–SSW bubble axis, so any cold front flows do not penetrate far. Radial velocity measurements of the  $H\alpha$  filaments show smooth laminar flow (Hatch et al 2006; Salomé et al 2011). The radio mini-halo does end abruptly at the front. The metallicity also undergoes a sharp change (Fig. 13).

The Perseus cluster will be an excellent target for high resolution X-ray spectroscopy, such as anticipated with ASTRO-H (Takahashi et al 2010). This will resolve the velocities flows along the line of sight. Both the central region where the bubbles form (Heinz et al 2010) and the W cold front will of great interest.

### 6.1 The substructure profile

In order to assess the energy associated with large scale structures in the Perseus cluster, we have characterized the energy/power needed to produce observed substructure (non-radial part) in the ICM (e.g. Fig. 6).

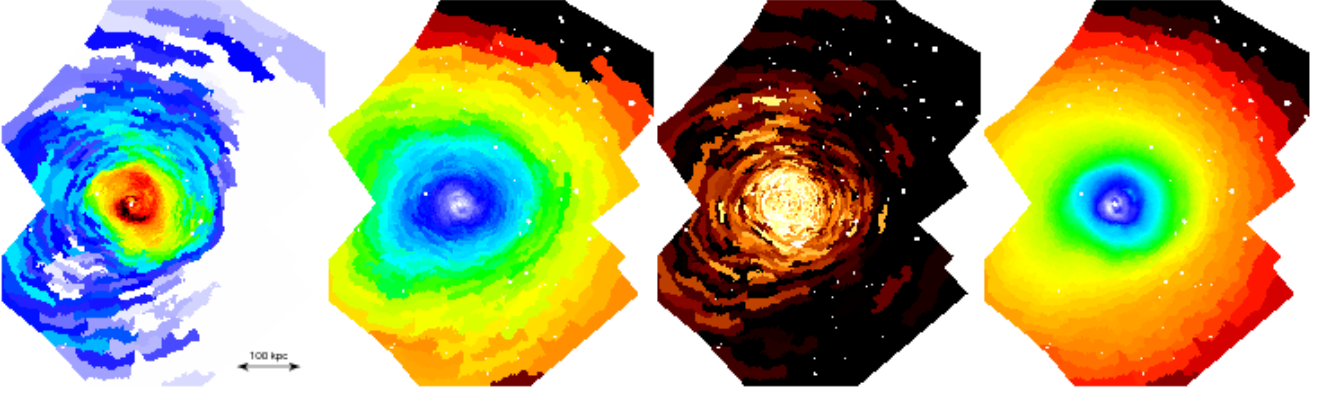


**Figure 12.** Profiles of differential surface brightness, temperature, pseudo-pressure and metallicity in the sector directly to the N (shown below).

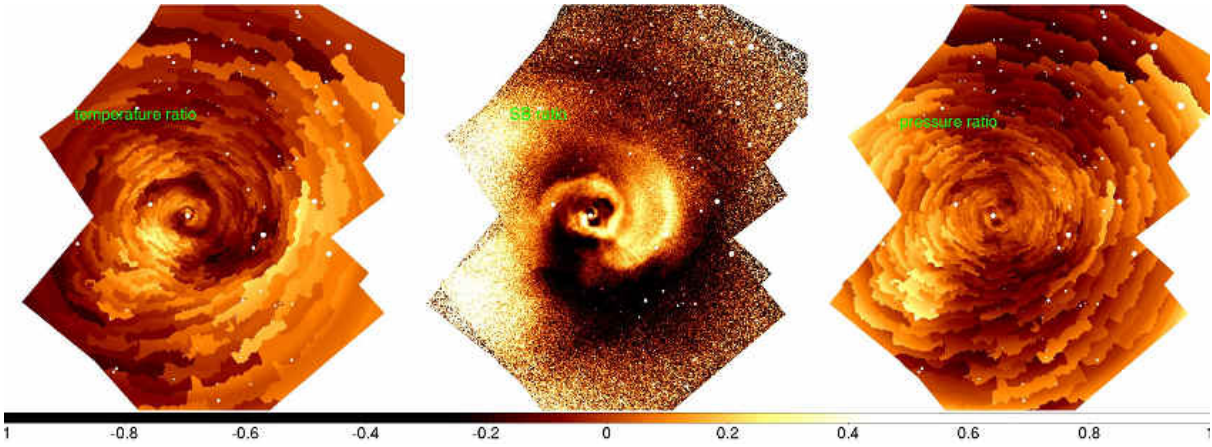
In order to cover a large radial range we use XMM-Newton data. While not providing a high resolution view on the Perseus core, they extend to larger radii than the existing Chandra data<sup>2</sup>. We first made a deprojection analysis of the data in four wedges (90 deg each) and for each wedge we calculated pressure in several radial shells spanning the range from few to 700 kpc. The thermal pressure  $nkT$  was evaluated in each radial bin and root-mean-square variation of pressure between wedges was calculated as  $\delta P \langle (P - \langle P(r) \rangle)^2 \rangle^{1/2}$ . The ratio of this quantity to the mean

<sup>2</sup> Analysis of the Chandra data gives similar results over the 10 – 300 kpc range, but is noisier due to real structure at smaller radii and to photon noise at larger radii





**Figure 13.** Distribution of (left to right) temperature, pseudo-pressure, metallicity and surface brightness. Each bin has about 22,500 counts.



**Figure 14.** Temperature, surface brightness and pressure distributions. The means at each radius has been subtracted and the intensity is proportional to the fractional difference.

pressure at the same radius  $\frac{\delta P}{P} = \frac{\langle (P - \langle P(r) \rangle)^2 \rangle^{1/2}}{P}$  can be used as a crude characteristic of the magnitude of the substructure (non-radial part) in the cluster. We then calculated the total thermal energy  $E_{th}(< r)$  of the ICM within given radius

$$E_{th}(< r) = \int_0^r \frac{3}{2} nkT 4\pi r^2 dr, \quad (1)$$

where  $n$  and  $T$  are the gas density and temperature at a given radius and estimated the energy associated with the substructure  $E_d(< r)$  as

$$E_d(< r) = \int_0^r \frac{3}{2} nkT \frac{\delta P}{P} 4\pi r^2 dr. \quad (2)$$

The values of  $E_{th}(< r)$  and  $E_d(< r)$  and their ratio are plotted in the two bottom panels in Fig. 15. In the top panel the energy associated with substructure  $E_d(< r)$  is divided by the sound crossing time of the region  $t_s = r/c_s$ , where  $c_s = \sqrt{\gamma \frac{kT}{\mu m_p}}$ . The quantity

$$L_d = \frac{E_d(< r)}{t_s} \quad (3)$$

can be regarded as an estimate of the power needed to maintain the observed non-radial substructure.

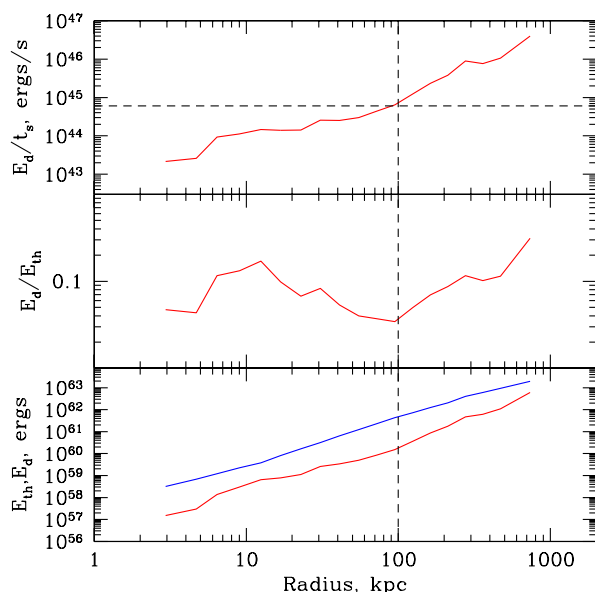
Clearly these are order of magnitude estimates. By construction they are sensitive only to non-radial part of the substructure and

only to the low angular modes. Even more uncertain is the estimate of power (top panel in Fig. 15) which uses the radial sound crossing time as an estimate of the time needed to dissolve the substructure. Nevertheless, taken at face value this figure suggests that the cluster can be broadly divided into two zones at about 100 kpc. (The bright rim to the W of the nucleus is at a radius of  $\sim 110$  kpc, see Figs. 2 and 6). The estimates of the mechanical power of the central AGN in NGC1275 gives values of order  $10^{45}$  erg s $^{-1}$  over the period of  $10^8$  yr (Fabian et al 2006). If the same estimate is applicable to the mean power during the last Gyr, then the central AGN is not capable of producing the asymmetry outside the central  $\sim 100$  kpc region. Instead the most promising explanation for the substructure beyond 100 kpc is the merger of the Perseus cluster with smaller galaxies/group along the chain of bright galaxies to the West of the core. At yet larger radii beyond 1 Mpc and out to the virial radius, the intracluster gas appears to become intrinsically clumpy (Simionescu et al 2011).

## 7 SUMMARY

Our wider image has revealed several structures consisting of dips in surface brightness along the NNW–SSE axis which are plausibly outer rising bubbles. The two new ones are not much larger in pressure-corrected volume than the inner bubbles whereas the large trough to the North is an order of magnitude larger. Bubbles must





**Figure 15.** Energy and power associated with non-radial substructure. Bottom panel shows the total ICM thermal energy within given radius  $E_{th}(< r)$  (see eqn. 1) and total energy associated with the observed substructure  $E_d(< r)$  (eqn. 2). The ratio of these quantities is plotted in the middle panel. The top panel shows an estimate of power (eqn. 3) needed to support the observed non-radial substructure. The horizontal dashed line shows an estimate of the central AGN mechanical power, while the vertical line schematically divides the cluster into “AGN-dominated” and “merger-dominated” parts.

be long-lived and we suggest that they may grow in size by merger with other, slower, rising bubbles.

A roughly semi-circular cold front is seen to the W of the nucleus of NGC 1275. The metallicity of the gas drops abruptly across the front.

We suspect that the trough and a bay to the South are where rising bubbles have accumulated. The S bay is distinctly hotter than its surroundings so the relativistic plasma may have mixed with the intracluster gas there, unlike the situation of the N or the inner bubbles. Both the N trough and the S bay lie along a continuation of the W cold front. The other structures seen at radii of 100 - 150 kpc are likely associated with a subcluster merger (Churazov et al 2003). The merger also accounts for the E–W asymmetry in surface brightness across the image (Figs. 2, 4, 14).

The energy in substructure peaks at small radii, where it is dominated by activity induced by the AGN, and at larger radii, where it is dominated by the merger.

## ACKNOWLEDGEMENTS

We acknowledge financial support from the Royal Society (ACF). JHK thanks the Cambridge Trusts and the Natural Sciences and Engineering Research Council of Canada (NSERC). GBT acknowledges support for this provided by the National Aeronautics and Space Administration through Chandra Award Numbers GO0-11139X and GO0-11138B issued by the Chandra X-ray Observatory Center, which is operated by the Smithsonian Astrophysical

Observatory for and on behalf of the National Aeronautics Space Administration under contract NAS8-03060. CSR acknowledges support from the Chandra Guest Observer Program under grant GO011138A. SWA was supported in part by the U.S. Department of Energy under contract number DE-AC02-76SF00515.

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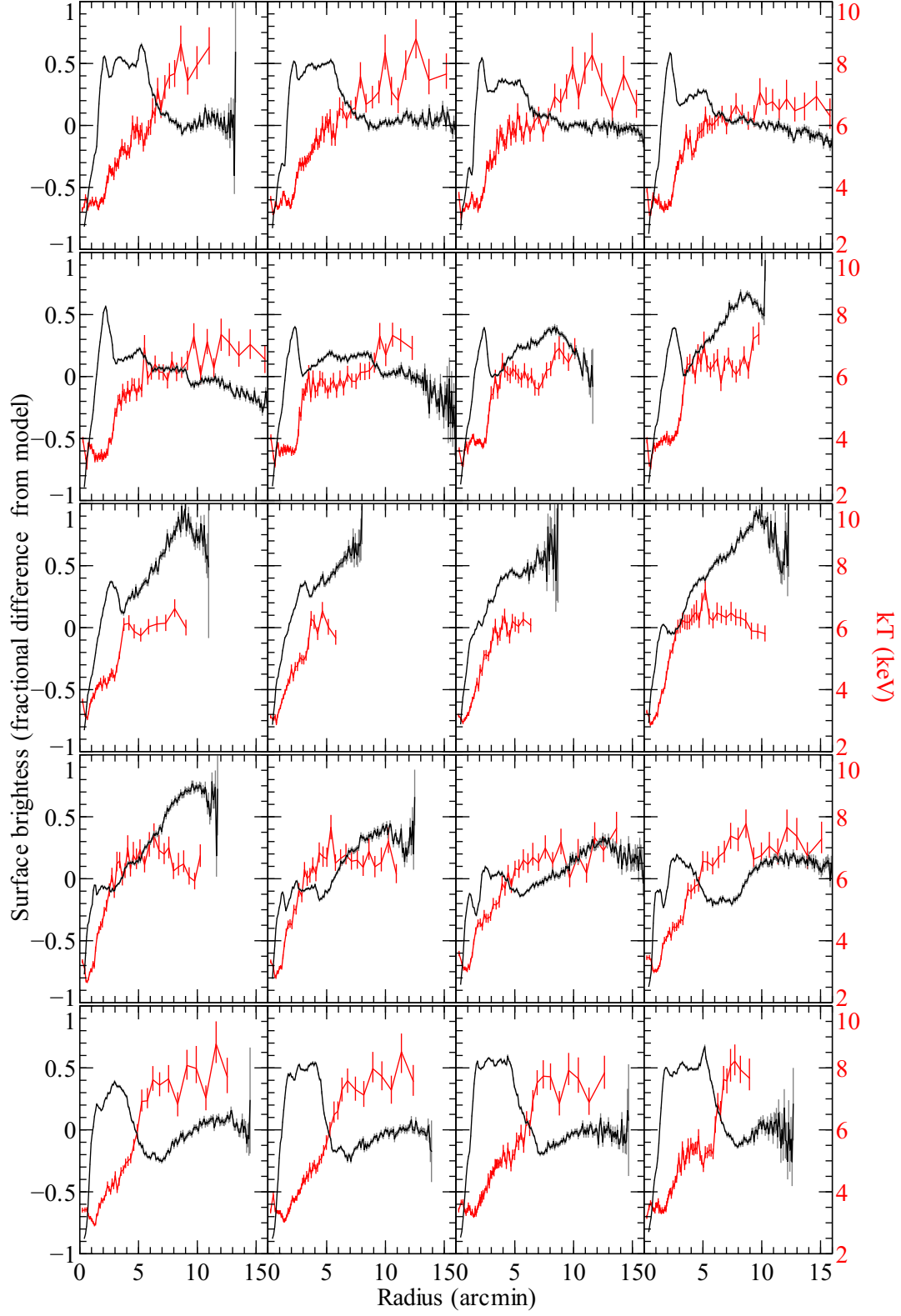
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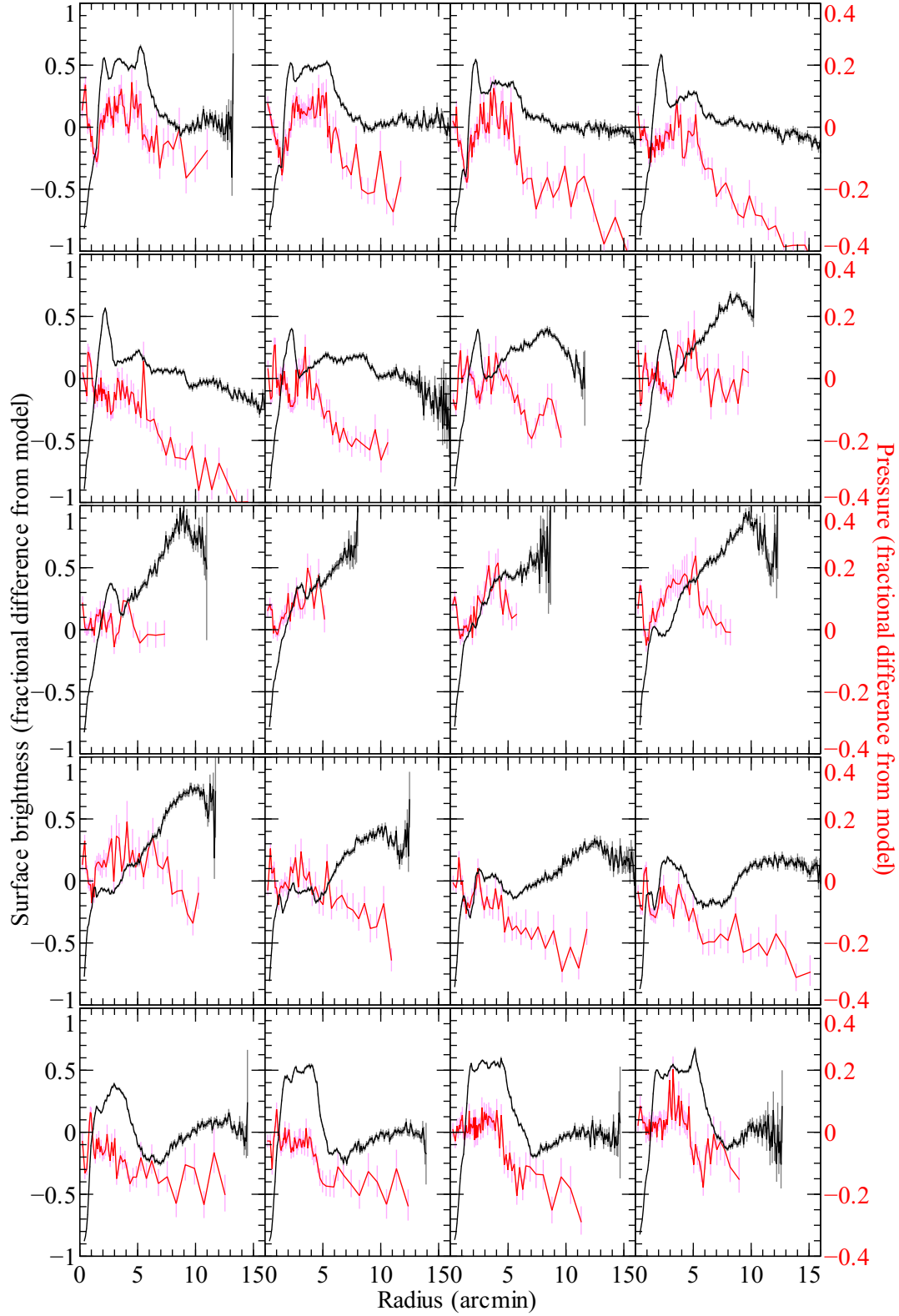
## **8 APPENDIX**

Temperature and pressure profiles about the nucleus in 18 deg steps  
anticlockwise from W.





**Figure 16.** Temperature profiles (in red) with differential surface brightness (black; subtracted from a smooth model) in 20x18 deg sectors arranged from W through N.



**Figure 17.** Similar to Fig. 4, but for Pressure (the model is a  $\beta$ -profile with a core, which explains why these profiles are somewhat different near the centre to the other profiles e.g. Fig 5, which use a power-law model at all radii).