

Cosmological simulations of the formation of the stellar haloes around disc galaxies

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ABSTRACT

We use the Galaxies-Intergalactic Medium Interaction Calculation (GIMIC) suite of cosmological hydrodynamical simulations to study the formation of stellar spheroids of Milky Way-mass disc galaxies. The simulations contain accurate treatments of metal-dependent radiative cooling, star formation, supernova feedback, and chemodynamics, and the large volumes that have been simulated yield an unprecedentedly large sample of ≈ 400 simulated $\sim L_*$ disc galaxies. The simulated galaxies are surrounded by low-mass, low-surface brightness stellar haloes that extend out to ~ 100 kpc and beyond. The diffuse stellar distributions bear a remarkable resemblance to those observed around the Milky Way, M31 and other nearby galaxies, in terms of mass density, surface brightness, and metallicity profiles. We show that in situ star formation dominates the stellar spheroids (by mass) at radii of $r \lesssim 30$ kpc, whereas accretion of stars dominates at larger radii and this change in origin induces a change in slope of the surface brightness and metallicity profiles, which is also present in the observational data. We also show that the in situ mass fraction of the spheroid is linked to the formation epoch of the system. Dynamically older systems have, on average, larger contributions from in situ star formation, although there is significant system-to-system scatter in this relationship. Thus, in situ star formation likely represents the solution to the longstanding failure of pure accretion-based models to reproduce the observed properties of the inner spheroid.

Key words: Galaxy: evolution — Galaxy: formation — Galaxy: halo — Galaxy: kinematics and dynamics — galaxies: dwarf — galaxies: evolution — galaxies: formation — galaxies: haloes — galaxies: kinematics and dynamics — Local Group — dark matter

1 INTRODUCTION

The stellar haloes of normal disc galaxies are expected to be excellent repositories of fossil evidence from the epoch of galaxy formation, even though these components contain only a small fraction of the total light (less than 1%, for the Milky Way). This is because the stellar haloes are widely believed to have been formed through the disruption of in-

falling satellite dwarf galaxies over the lifetime of the galaxy. Recent advances in observational techniques have allowed the detection of faint stellar haloes around external galaxies, such as M31 (Ferguson et al. 2002; Guhathakurta et al. 2005; Irwin et al. 2005; Ibata et al. 2007) and other nearby disc galaxies (Mouhcine et al. 2005a,b; de Jong et al. 2007; Ibata et al. 2009). Most stellar haloes have been shown to contain evidence of merger events in the form of tidal stellar streams. The number and physical properties of these streams appear to be in general agreement with the predic-

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tions of the Λ CDM model (Bell et al. 2008; Gilbert et al. 2009a,b; McConnachie et al. 2009; Starkenburg et al. 2009; Martínez-Delgado et al. 2010).

In spite of this, questions have been raised about the compatibility with the standard paradigm of the seemingly large differences in the properties of the stellar haloes and satellite systems of M31 and the Milky Way. These two systems have similar masses but M31 appears to have experienced a much more active merger history in the recent past, as suggested by its more disturbed stellar disc, larger bulge, younger and more metal-rich stellar populations, larger and more numerous tidal streams and surviving satellite galaxies (Guhathakurta et al. 2005; Irwin et al. 2005; Brown et al. 2006a,b; Gilbert et al. 2006; Kalirai et al. 2006; McConnachie & Irwin 2006; Gilbert et al. 2007; Ibata et al. 2007; Koch et al. 2008; Kalirai et al. 2009).

On the theoretical side, the detection of extended stellar components around external galaxies has reinvigorated the efforts of modelling stellar haloes in the context of Λ CDM cosmogony. Several theoretical studies suggest that the main properties of stellar haloes can be explained entirely by the accretion and disruption of satellites (Bullock & Johnston 2005; Font et al. 2006a,b,c; De Lucia & Helmi 2008; Font et al. 2008; Johnston et al. 2008; Gilbert et al. 2009b; Cooper et al. 2010). These studies have been carried out using a combination of high-resolution dark matter-only simulations and simple semi-analytic prescriptions for assigning light (i.e., stellar mass) to infalling satellites. Recent studies using such ‘hybrid’ methods (e.g., Cooper et al. 2010) have taken advantage of the ultra-high resolution achieved by the latest generation of dark matter only Milky Way simulations (Springel et al. 2008; see also Diemand, Kuhlen & Madau 2007), which allowed them to explore the properties of even the faintest substructures (satellites and streams) surviving to the present day.

However, even though a large fraction of the stellar halo¹ of normal disc galaxies may be explained by the accretion scenario, there remain some stellar populations in the Milky Way and other nearby galaxies that have not so far been accounted for in the context of accretion-only models. This has been known for a long time (Eggen, Lynden-Bell & Sandage 1962), but the evidence has become more convincing in the past few years. For example, the metal-rich halo stars in the Solar neighbourhood have been shown to have a bi-modal $[\alpha/\text{Fe}]$ pattern (Nissen & Schuster 2010). In particular, the high $[\text{Fe}/\text{H}]$ -high $[\alpha/\text{Fe}]$ population seems to be the product of very rapid chemical evolution that cannot be sustained in the potential wells of even the most massive accreted satellites (Font et al. 2006a). Instead, it has been suggested that these stars may have formed “in situ” in the halo (e.g., Zolotov et al. 2010) or they may have been stirred up from the disc by repeated interactions with satellite galaxies (Purcell et al. 2010). Similarly, models that track only the accreted component of the stellar halo (e.g., Font et al. 2008) have difficulties matching the intermediate age (~ 8 Gyr), metal-rich populations ($[\text{Fe}/\text{H}] > -1$) detected in the halo of M31 (see Brown et al

2007). In this paper, we show that the above features are explained naturally with stellar populations formed in situ.

There is also an indication of a spatial differentiation in the stellar halo of the Milky Way: the inner region appears to be more metal-rich and have a prograde motion with respect to the disc, while the outer halo is metal-poor and appears to have a net retrograde rotation (e.g., Carollo et al. 2007, 2010; de Jong et al. 2010) or no rotation, depending on the adopted the local standard of rest (Deason et al. 2010a) and/or the adopted distance scale (Schoenrich et al. 2010). Measurements of the local halo density profile are better fit assuming two sub-components, one flattened and the other spherical (Juric et al. 2008). These results suggest that the stellar haloes of galaxies may have a dual nature, perhaps along the lines of the scenario proposed originally by Searle & Zinn (1978): while the outer halo formed primarily by accretion/mergers, some of the inner halo has formed through dissipational collapse. However, in the current paradigm, the dissipative component would not have formed through monolithic collapse (Eggen, Lynden-Bell & Sandage 1962), but rather as a result of gas-rich mergers at early times. Gas-dynamical simulations of disc galaxies formed in a hierarchical scenario produce stellar systems with a significant in situ dissipative component (Abadi et al. 2006; Brook et al. 2004; Zolotov et al. 2009, 2010; Oser et al. 2010). It remains to be determined what the relative contributions of in situ and accreted stars are as a function of radius and system mass and, importantly, what the system-to-system scatter is in these contributions. The scatter in the in situ and accreted components may be at the heart of the reported large differences in the properties of the extended stellar distributions of the Milky Way and M31.

Another important aspect that we address here, which may be related to the in situ vs. accretion question, is the origin of large-scale metallicity gradients in stellar haloes. The emergence of metallicity gradients is a well-known consequence of dissipative collapse/mergers. The evidence in favour of metallicity gradients in local disc galaxies is mounting. For example, in the Milky Way, an underlying metallicity gradient in the halo is suggested by the net difference in the mean metallicity between the inner and outer halo (Eggen, Lynden-Bell & Sandage 1962; Searle & Zinn 1978; Carollo et al. 2007, 2010; de Jong et al. 2010). The halo of M31 also shows evidence for a gradient, where the inferred metallicity may drop by as much as 1 dex from the inner regions of the galaxy out to ≈ 100 kpc (see Fig. 5 below). NGC 891 represents another local Milky Way-analog for which there is evidence of a stellar metallicity gradient (see Ibata et al. 2009; although note that metallicity measurements have only been made out to 15 kpc for this system).

Models that form haloes by non-dissipative accretion and mergers have so far not been able to produce significant metallicity gradients (Font et al. 2006a; De Lucia & Helmi 2008; Cooper et al. 2010). This suggests that either these models do not capture all the relevant physical processes (e.g., in situ star formation or ejection of disc stars by interaction with satellites) or that they do not model the entire range of possible merger histories for Milky Way-mass galaxies in the cosmological context. In this study we use high-resolution cosmological hydrodynamical simulations with a sophisticated treatment of chemodynamics to

¹ By volume, but not by mass as we argue later in this paper.

show that metallicity gradients are expected to be ubiquitous in the haloes of Milky Way-mass disc galaxies. The prominence of these gradients is shown to be tied to the fraction of the stellar mass that is formed in situ, which is itself determined by the merger histories of the galaxies.

The paper is organised as follows. In Section 2 we describe the simulations and the selection of the sample Milky Way-mass disc galaxies. In Section 3 we compute the average properties of stellar haloes of these galaxies, such as their density profiles, metallicity gradients (both $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$) and metallicity distribution functions (MDFs). In Section 4 we quantify the contribution to the stellar spheroid from in situ star formation and accretion/disruption of satellites. Finally, we summarize and discuss our main findings in Section 5.

2 SIMULATIONS

We use the Galaxies-Intergalactic Medium Interaction Calculation (GIMIC) suite of cosmological hydrodynamical simulations, which were carried out by the Virgo Consortium and are described in detail by Crain et al. (2009), hereafter C09 (see also Crain et al. 2010). We will therefore present only a brief summary of the simulations here, focusing on the details most relevant for the present study. This suite of simulations is ideal for the present study for several reasons: (i) it simulates large volumes of the universe that contain numerous Milky Way-mass disc galaxies, allowing us to robustly quantify both the mean trends and the scatter; (ii) it incorporates accurate physical prescriptions for metal-dependent radiative cooling, star formation, mass and energy feedback from Type Ia and Type II supernovae, as well as enrichment due to stellar evolution (AGB stars), which results in the formation of realistic disc galaxies; and (iii) the relatively high numerical resolution is adequate for resolving the stellar haloes of individual Milky Way-mass galaxies and their most massive satellite galaxies (i.e., the main contributors to the accreted component of the stellar halo).

The suite consists of a set of hydrodynamical re-simulations of five nearly spherical regions ($\sim 20h^{-1}$ Mpc in radius) extracted from the Millennium Simulation (Springel et al. 2005a). The regions were selected to have overdensities at $z = 1.5$ that represent $(+2, +1, 0, -1, -2)\sigma$, where σ is the root-mean-square deviation from the mean on this spatial scale. The 5 spheres are therefore environmentally diverse in terms of the large-scale structure that is present within them. For the purposes of the present study, however, we will only select systems with total ‘main halo’ (i.e., the dominant subhalo in a friends-of-friends, hereafter FoF, group) masses similar to that of the Milky Way, irrespective of which GIMIC volume it is in (i.e., irrespective of the large-scale environment). C09 found that the properties of systems of fixed main halo mass do not depend significantly on the large-scale environment (see, e.g., Fig. 8 of that study). We plan to study the properties of such galaxies as a function of their *local* environment in a future study.

The cosmological parameters adopted for GIMIC are the same as those assumed in the Millennium Simulation and correspond to a Λ CDM model with $\Omega_m = 0.25$, $\Omega_\Lambda = 0.75$, $\Omega_b = 0.045$, $\sigma_8 = 0.9$ (where σ_8 is the rms amplitude of

linearly evolved mass fluctuations on a scale of $8h^{-1}$ Mpc at $z = 0$), $H_0 = 100h$ km s $^{-1}$ Mpc $^{-1}$, $h = 0.73$, $n_s = 1$ (where n_s is the spectral index of the primordial power spectrum). The value of σ_8 is approximately 2-sigma higher than has been inferred from the most recent CMB data (Komatsu et al. 2009), which will affect the abundance of Milky Way-mass systems somewhat, but should not significantly affect their individual properties.

The simulations were evolved to $z = 0$ using the TreePM-SPH code GADGET-3 (last described in Springel 2005). The GADGET-3 code has been substantially modified to incorporate new baryonic physics. Radiative cooling rates for the gas are computed on an element-by-element basis by interpolating within pre-computed tables generated with CLOUDY (Ferland et al. 1998) that contain cooling rates as a function of density, temperature, and redshift and that account for the presence of the cosmic microwave background and photoionisation from a Haardt & Madau (2001) ionising UV/X-Ray background (see Wiersma et al. 2009a). This background is switched on at $z = 9$ in the simulation, where it ‘reionises’ the entire simulation volume. Star formation is tracked in the simulations following the prescription of Schaye & Dalla Vecchia (2008). Gas with densities exceeding the critical density for the onset of the thermogravitational instability is expected to be multiphase and to form stars (Schaye 2004). Because the simulations lack both the physics and the resolution to model the cold interstellar gas phase, an effective equation of state (EOS) is imposed with pressure $P \propto \rho^{4/3}$ for densities $n_H > n_*$ where $n_* = 0.1$ cm $^{-3}$. As described in Schaye & Dalla Vecchia (2008), gas on the effective EOS is allowed to form stars at a pressure-dependent rate that reproduces the observed Kennicutt-Schmidt law (Kennicutt 1998) by construction. The timed release of individual elements by both massive (Type II SNe and stellar winds) and intermediate mass stars (Type Ia SNe and asymptotic giant branch stars) is included following the prescription of Wiersma et al. (2009b). A set of 11 individual elements are followed in these simulations (H, He, C, Ca, N, O, Ne, Mg, S, Si, Fe), which represent all the important species for computing radiative cooling rates.

Feedback from supernovae is incorporated using the kinetic wind model of Dalla Vecchia & Schaye (2008) with the initial wind velocity, v_w , set to 600 km/s and the mass-loading parameter (i.e., the ratio of the mass of gas given a velocity kick to that turned into newly formed star particles), η , set to 4. This corresponds to using approximately 80% of the total energy available from supernovae for a Chabrier (2003) IMF, which is assumed in the simulation. This choice of parameters results in a good match to the peak of the star formation rate history of the universe (C09; see also Schaye et al. 2010) and reproduces a number of X-ray/optical scaling relations for normal disc galaxies (Crain et al. 2010).

The GIMIC simulations have been run at three levels of numerical resolution, ‘low’, ‘intermediate’, and ‘high’, to test for numerical convergence. The low-resolution simulations have the same mass resolution as the Millennium Simulation (which, however, contains only dark matter) while the intermediate- and high-resolution simulations have 8 and 64 times better mass resolution, respectively. In the case of the high-resolution simulations, only the -2σ volume was run to $z = 0$, owing to the computational expense

of running such large volumes at this resolution. For the purposes of the present study, we therefore follow the approach of Crain et al. (2010) and adopt the intermediate-resolution simulations in the main analysis and reserve the high-resolution -2σ simulation to test the numerical convergence of our results, which is presented in the Appendix.

The intermediate-resolution runs therefore have a dark matter particle mass $M_{dm} \simeq 5.30 \times 10^7 h^{-1} M_\odot$ and an initial gas particle mass of $M_g \simeq 1.16 \times 10^7 h^{-1} M_\odot$, implying that it is possible to identify systems with masses similar to those of the classical dwarf galaxies around the Milky Way (with $M_{tot} \sim 10^9 M_\odot$). This is important for the present study, as it is currently believed that the disruption of such massive satellites is the primary contributor, by mass, to the accreted stellar halo. The Plummer equivalent force resolution (i.e., gravitational softening) of the intermediate-resolution runs is fixed in comoving space at early times, but becomes fixed at $1 h^{-1}$ kpc in physical space at $z \leq 3$. Note that for all these simulations the remainder of the $(500h^{-1} \text{ Mpc})^3$ Millennium volume is also simulated, but with dark matter only and at much lower resolution. This ensures that the influence of the surrounding large-scale structure is accurately accounted for.

2.1 Selection of Milky Way-mass systems

Below we describe the selection criteria for our sample of Milky Way-mass systems.

First we define what we mean by a ‘Milky Way-mass’ system. This is a system with a present-day total (gas+stars+dark matter) mass within r_{200} (i.e., the radius which encloses a mean density of 200 times the critical density of the universe at $z = 0$) of $7 \times 10^{11} M_\odot < M_{200} < 3 \times 10^{12} M_\odot$. For reference, this range roughly spans the mass estimates in the literature for both the Milky Way and M31 (Kochanek 1996; Evans & Wilkinson 2000; Battaglia et al. 2005; Karachentsev & Kashibadze 2006; Guo et al. 2010), although we are not attempting to reproduce the properties of either of these systems in detail. Systems and their substructures are identified in the simulations using the SUBFIND algorithm of Dolag et al. (2009), that extends the standard implementation of Springel et al. (2001) by including baryonic particles in the identification of self-bound substructures. For each FoF system that is identified, a spherical overdensity (SO) mass with $\Delta = 200$ is computed (i.e., M_{200}). The gas+stars+dark matter associated with the most massive subhalo of a FoF system are considered to belong to the ‘galaxy’, while the gas+stars+dark matter associated with the other substructures are classified as belonging to satellite galaxies. As the current study is focused on examining the extended stellar distributions of normal disc galaxies, we exclude all bound substructures (i.e., satellites) from the analyses presented in this paper, with the exception of Fig. 2 (described below).

Galaxies identified in this way are then classified morphologically as disc- or spheroid-dominated, based on their dynamics. We follow the approach of Abadi et al. (2003) in separating the disc from the spheroid. In particular, we align the total angular momentum of the stellar disc with the z-axis, calculate the angular momentum of each star particle in the x-y plane, J_z , and compare it with the angular momentum, J_{circ} , of a star particle on a co-rotating circular orbit

with the same energy. Disc particles are selected using a cut of $J_z/J_{\text{circ}} > 0.8$ as in Zolotov et al. (2009). Visually, this cut is quite successful in isolating the disc, but in a small number of cases we found that particles well above/below the disc mid-plane were also assigned to the disc. For this reason, we also apply a spatial cut such that disc particles cannot exist more than two softening lengths above or below the disc mid-plane. Note that we have also tried assigning disc particles based on the fraction of stellar kinetic energy in ordered rotation, as defined in Sales et al. (2010), but found no significant differences from the method outlined above.

We classify galaxies according to the ratio of disc stellar mass to total stellar mass (D/T) within 20 kpc. We identify and select disc galaxies using a cut $D/T > 0.3$, although very similar results are obtained for other cuts in D/T (e.g., we have also tried 0.2 and 0.4 and none of the main results of our study are changed).

Unlike Zolotov et al. (2009), we have elected not to further decompose the spheroid into bulge and halo components. While separating the disc from the spheroid is relatively straightforward to do for both the observations and the simulations, separating the bulge from the stellar halo is more challenging, as there is no sharp distinction in the dynamics or the spatial distribution of the two components. A variety of techniques have been used in theoretical and observational studies to distinguish these components, which can complicate the comparison of any single component between different studies (at least in the inner regions where the bulge and halo overlap spatially). Our approach is to avoid this complication by simply reporting on the total bulge+halo component, which can readily be deduced from observations.

Lastly, we point out that we have not imposed any explicit constraints on the merger histories of the simulated galaxies. This is in contrast with many of the previous cosmological modelling studies of Milky Way-type galaxies, which usually select their simulation candidates from systems that have not undergone any major mergers in the recent past (e.g., Diemand et al. 2005; Okamoto et al. 2005; Cooper et al. 2010). This condition is meant to ensure that these systems develop stellar discs that survive until the present day and that resemble the Milky Way’s disc (Tóth & Ostriker 1992). Note that these previous studies were based on hybrid methods applied to dark matter-only simulations and therefore did not have the benefit of being able to check which histories actually give rise to discs at the present day. Importantly, more than 70% of Milky Way-mass haloes have experienced major mergers (defined as $M_{\text{sat}}/M_{\text{disc}} \geq 1/3$, where M_{sat} is the total mass of the satellite and M_{disc} is the stellar mass of the disc) in the last 10 Gyr (Stewart et al. 2008; Boylan-Kolchin et al. 2010), implying that the majority of merger histories may be excluded from studies that select only quiescent histories. The constraints applied to previous hybrid studies appear to be overly conservative, given that the majority of the Milky Way-mass systems in GIMIC are disc-dominated systems (see Fig. 1 of Crain et al. 2010). The resiliency of galactic discs to mergers likely stems from the fact that the galaxies contain a significant dissipational gaseous component (that was not taken into account in, e.g., Tóth & Ostriker 1992), which is able to maintain (or re-establish) the disc (see Hopkins et al. 2009). In any case, for the present study, which is based on

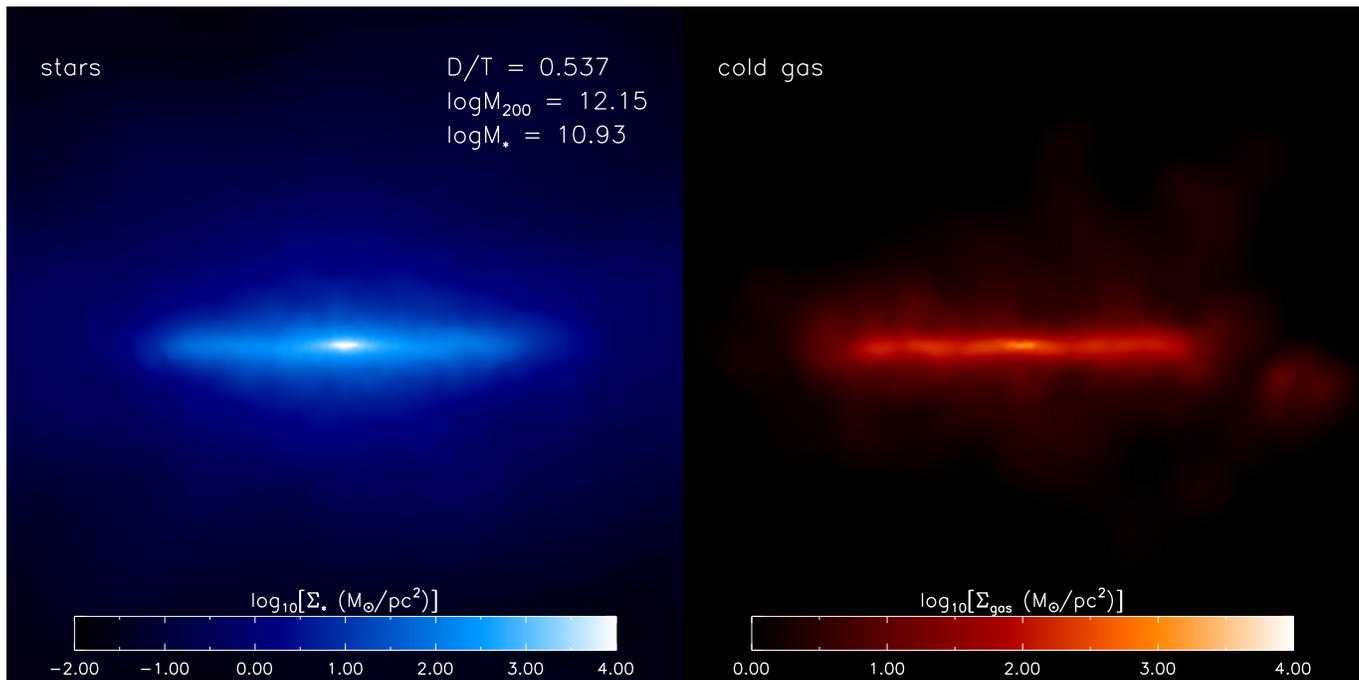


Figure 1. Surface mass density maps of all the stars (*left*) and gas (*right*) in one of the sample Milky Way-mass disc galaxies at $z = 0$ in the high-resolution simulation. The panel is 100 kpc on a side.

Table 1. Global properties of the simulated disc galaxies. Presented is the median of the property in 3 different total mass (M_{200}) bins, with the error bars representing the 5th and 95th percentiles. M_V is total V-band magnitude, $M_*(< r_{200})$ is the total stellar mass within r_{200} , $v_{\text{rot}}(R_\odot)$ is the rotation velocity of the stellar disc at the Solar radius, $[\text{Fe}/\text{H}]_{r < 30\text{kpc}}$ is the mean stellar metallicity within 30 kpc, $[\text{Fe}/\text{H}]_{r > 30\text{kpc}}$ is the mean stellar metallicity beyond 30 kpc, and $n_{\text{gal,bin}}$ is the number of simulated galaxies in the total mass bin.

$\log M_{200}$ bin (M_\odot)	M_V (mag.)	$M_*(< r_{200})$ ($10^{10} M_\odot$)	$v_{\text{rot}}(R_\odot)$ (km/s)	$[\text{Fe}/\text{H}]_{r < 30\text{kpc}}$	$[\text{Fe}/\text{H}]_{r > 30\text{kpc}}$	$n_{\text{gal,bin}}$
11.85 – 12.05	$-21.37_{+1.46}^{-0.98}$	$3.46_{-2.15}^{+6.32}$	184_{-52}^{+76}	$-0.49_{-0.28}^{+0.27}$	$-1.13_{-0.19}^{+0.27}$	127
12.05 – 12.25	$-22.17_{+1.53}^{-0.47}$	$8.18_{-6.32}^{+5.66}$	243_{-93}^{+58}	$-0.35_{-0.37}^{+0.18}$	$-1.12_{-0.17}^{+0.16}$	154
12.25 – 12.50	$-22.45_{+0.69}^{-0.50}$	$12.82_{-8.85}^{+8.83}$	280_{-89}^{+85}	$-0.30_{-0.25}^{+0.13}$	$-1.15_{-0.18}^{+0.23}$	128

hydrodynamic simulations of large volumes, there is no need to place explicit constraints on the merger histories since we know which galaxies have prominent discs at $z = 0$.

Applying the system mass and morphology criteria described above yields a total sample of 412 systems in the 5 intermediate-resolution GIMIC volumes. This greatly exceeds the sample size of all previous studies of this type (i.e., those based on hydrodynamic re-simulations or hybrid methods), albeit at lower resolution than some previous studies (such as Zolotov et al. 2009). The large number of systems in our sample allows us to robustly quantify the system-to-system scatter present in the properties of stellar haloes. Our work is therefore complementary to studies such as Abadi et al. (2006) and Zolotov et al. (2009), which are based on high-resolution simulations of a small number of systems. (But note we demonstrate in the Appendix that the properties of our simulated stellar haloes are robust to increased numerical resolution.)

As an example, Fig. 1 shows the stellar and cold gas mass distributions of a typical simulated disc galaxy. An

extended, flattened stellar distribution is clearly visible out to large radii.

3 PROPERTIES OF STELLAR HALOES

3.1 Global properties

Before analysing the detailed properties of the stellar haloes of the simulated disc galaxies, we briefly discuss the overall global properties of the sample galaxies. The motivation for this is simple: if the simulated galaxies were to look wholly unlike observed normal disc galaxies, then that would limit what can be concluded about the formation of their stellar haloes.

In Table 1 we present the typical properties of the simulated disc galaxies, split into several halo mass (M_{200}) bins. In particular, we list total V-band magnitude (M_V ,

in AB mags.), total stellar mass² within r_{200} [$M_*(< r_{200})$], disc rotation speed at the solar radius [$v_{rot}(R_\odot)$, where R_\odot is assumed to be 8.5 kpc] and mean stellar metallicity within and beyond a galacto-centric distance of 30 kpc ($[Fe/H]_{r < 30 \text{ kpc}}$). (The choice of 30 kpc is motivated below, where we show how the physical nature of the bulge+halo component changes with radius.) The values listed in Table 1 are medians for galaxies in the mass bin, while the error bars represent the 5th and 95th percentiles. We also list the number of galaxies in each mass bin ($n_{gal,bin}$).

The V-band luminosity of star particles in the simulations is calculated as follows. Each star particle is treated as a simple stellar population (SSP). The ages and metallicities of star particles are used to compute the spectral energy distribution (SED) by interpolating over the GALAXEV models of Bruzual & Charlot (2003). The V-band luminosity is obtained by integrating the product of the SED with the V-band filter transmission function. We ignore the effects of dust attenuation in this calculation.

As already noted, the GIMIC simulations successfully reproduce the peak of the observed star formation rate density history of the Universe, as well as the X-ray-optical scaling properties of normal disc galaxies (Crain et al. 2010). Another notable success of the simulations is that they produce large numbers of disc galaxies (see Fig. 1 of Crain et al. 2010). Realistic spiral discs have been notoriously difficult to achieve in the past with cosmological hydrodynamic simulations. We will present a detailed analysis of the galaxy discs in a future study (McCarthy et al., in prep.).

The GIMIC simulations, however, do not reproduce the bright end of the global galaxy luminosity function, $M_V < -22$ (see C09). This is perhaps not surprising, as the simulations lack feedback from supermassive black holes, which is thought to be crucial for regulating star formation in the most massive systems (e.g., Springel et al. 2005b; Bower et al. 2006; Croton et al. 2006; Booth & Schaye 2009). The match to the observed abundance of $\sim L_*$ galaxies is, however, reasonable. One should also recall that the luminosity function depends not only on the luminosity as a function of halo mass (which we quote below), but also on the number density of dark matter haloes, which depends on the adopted cosmological parameters. This implies that one can obtain correct properties of individual galaxies but still not reproduce the observed luminosity function, if the wrong cosmological parameters are adopted. Therefore, a better approach to assessing whether the simulated galaxies are realistic, is to compare the mass-to-light ratio (or total luminosity) as a function of halo mass with the observations, which we do immediately below.

For comparison, M31 has an absolute visual magnitude of $M_V \approx -21.0$, $M_* \approx 5.6 \times 10^{10} M_\odot$, and $v_{rot} \approx 250$ km/s. The absolute magnitude was calculated using the *GALEX* apparent V magnitude of Gil de Paz et al. (2007) and assuming a distance to M31 of 785 kpc. The total stellar mass was calculated using B-V and V of Gil de Paz et al. (2007)

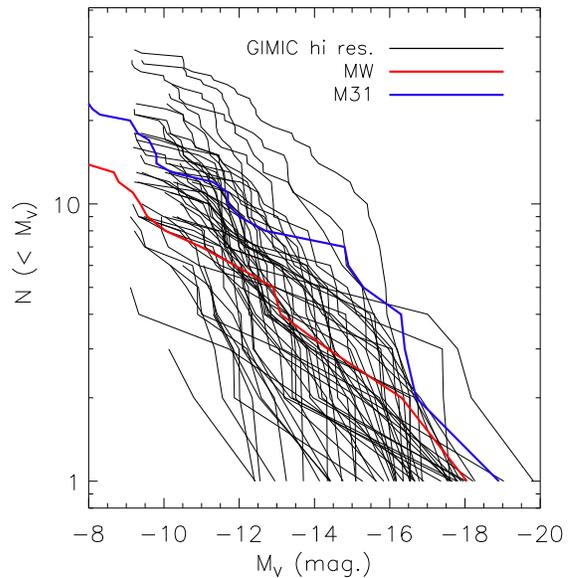


Figure 2. The satellite luminosity functions of 50 simulated galaxies in the high-resolution GIMIC simulation. The solid black curves show the luminosity functions of individual simulated galaxies. The solid, red and blue histograms show the satellite luminosity functions for the Milky Way and M31, respectively. The Milky Way data is taken from Mateo (1998) and the M31 data is taken from McConnachie et al. (2009). Overall, the simulations produce luminosity functions that are similar to those observed for the Milky Way and M31, suggesting that the accreted component of the halo in the simulations should have a realistic mass fraction.

to infer the stellar M/L using the relation of Bell & de Jong (2001) and scaling the result by 0.7 to convert from the Salpeter IMF assumed by Bell & de Jong (2001) to a Chabrier IMF, which is adopted in the GIMIC simulations. The rotation velocity, which has been corrected for inclination, was obtained from the HyperLeda data base³. The Milky Way has a total stellar mass of $M_* \approx 5.5 \times 10^{10} M_\odot$ (Flynn et al. 2006)⁴ and $v_{rot} = 250$ km/s (Reid et al. 2009). The properties of M31 and the Milky Way are quite typical of the simulated galaxies, particularly those with halo masses of $M_{200} \lesssim 2 \times 10^{12} M_\odot$. We will comment in detail on the metallicity of the stellar halo of both observed and simulated systems below.

The above comparison demonstrates that the simulated disc galaxies have global observable properties similar to those of the Milky Way and M31 which, in and of itself, is an important achievement that will be described in more detail in a future study (McCarthy et al. in prep).

Since the present work concerns itself with the properties of the extended stellar distributions of disc galaxies (i.e., the spheroids), it is also important to check whether

² When referring to stellar masses throughout the paper we are always referring to the *current* stellar mass, which is less than the initial stellar mass due to stellar mass loss. Typically, about 45% of the initial mass is lost due to stellar evolution, which is dictated by the Chabrier IMF adopted in the simulations and the generally old stellar populations we are studying here.

³ <http://leda.univ-lyon1.fr>.

⁴ Flynn et al. (2006) directly measure the stellar M/L locally and apply it to the Galaxy to estimate the total stellar mass. The measured value of M/L is consistent with the predictions of a Chabrier IMF.

the simulations yield reasonable satellite luminosity functions and radial distributions, since a significant fraction of the spheroid (by volume at least) is thought to have been assembled through the tidal disruption of massive satellites. In Fig. 2 we compare the satellite luminosity functions (LFs) of simulated galaxies at $z = 0$ in the high-resolution GIMIC simulation⁵. The LF is calculated for all satellite galaxies within r_{200} . Encouragingly, the satellite systems of the Milky Way and M31 fall within the system-to-system scatter of the simulated satellite systems. Recently, Deason et al. (2010b) examined the cumulative radial distribution of the 10 brightest satellites around Milky Way-mass galaxies in the intermediate-resolution GIMIC simulations. They find that the Milky Way and M31 distributions fall within the system-to-system scatter of the simulated galaxies. Taken together, these results imply that the simulated spheroids should have realistic contributions from the accretion and disruption of satellites.

We now proceed to our analysis of the stellar spheroids of simulated disc galaxies.

3.2 Structural properties

Below we present the radial distributions of stellar mass density, ρ_* , and surface brightness, Σ_V , as well as the radial profiles of the metallicity $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$ of our sample of simulated galaxies.

3.2.1 Stellar mass density and surface brightness profiles

Fig. 3 shows the median spherically-averaged stellar mass density (left panel) and azimuthally-averaged V-band surface brightness (right panel) radial profiles of the bulge+halo component (solid black curves), as well as the 1- and 2-sigma scatter about the median (shaded regions). The solid red curves show the median profiles for the *total* stellar (disc+bulge+halo) components. The dashed green curve represents a power-law in ρ_* with index of -3.5 in the left panel, which corresponds to a power-law in surface brightness (in linear units) with index of -2.5 , assuming the stellar mass-to-light ratio is independent of radius. Converting to mags. per arcsec², implies a power-law with index 5.25, which is represented by the dashed green curve in the right panel.

First, extended (out to $r \sim 100$ kpc and beyond) stellar distributions are a ubiquitous feature of the simulated galaxy populations. The radial distributions of mass density and surface brightness have relatively simple forms, but as a comparison of the solid black and dashed green curves demonstrates, no single power-law distribution can match the profiles over all radii. A power-law with index of -3.5 (in ρ_*) provides an excellent fit to the outer regions, $r \gtrsim 30$ kpc, of the median bulge+halo profile, but the same power-law under-predicts the mass density/surface brightness at

smaller radii. As we will discuss in Section 4, the change in the nature of the radial distributions at radii of $r \approx 30$ kpc is driven by a change in the relative contributions of stars deposited by tidally disrupted satellites and stars that formed in situ.

We can quantify the system-to-system scatter in the slope of the profile at large radii by fitting power-laws to the mass density distributions of all ≈ 400 disc galaxies in our sample. We find a median slope of -3.51 with -3.05 , -3.20 , -3.89 , and -4.16 as the 5th, 14th, 86th, and 95th percentiles of the distribution of slopes, respectively. Although it usually does not fit the profiles as well as a power-law, we have also fit a de Vaucouleurs profile to the outer regions (as done, e.g., by Ibata et al. 2007 for the outer regions of M31). In this case we find a median effective radius, r_e , of 17.80 kpc with 6.71, 10.05, 34.74, and 54.98 kpc as the 5th, 14th, 86th, and 95th percentiles, respectively.

In the inner regions, $r \lesssim 30$ kpc, we find that a de Vaucouleurs profile usually fits the mass density distribution very well. Here we find a median r_e of 1.83 kpc with 0.73, 0.95, 4.90, and 7.54 kpc as the 5th, 14th, 86th, and 95th percentiles, respectively. A power-law fit to $r \lesssim 30$ kpc yields a median slope of -2.98 and -2.47 , -2.58 , -3.30 , and -3.36 as the 5th, 14th, 86th, and 95th percentiles, respectively. This is shallower than in the outer regions, but note that the mass density profile has a higher normalisation at small radii, so that a power-law with slope -3.5 extrapolated to the inner regions underpredicts the mass density there.

In general, our results confirm those of several previous theoretical studies on the shape of the mass density profile at larger radii. For example, Bullock & Johnston (2005), who used a hybrid semi-analytic plus idealised N-body approach to study the formation of the stellar halo, find that a power-law with index of -3.5 describes the shape of their haloes very well at radii of $\approx 50 - 100$ kpc. More recently, Cooper et al. (2010) applied a similar approach to the very high-resolution Aquarius cosmological N-body simulations. These authors did not report on a power-law fit to their simulated profiles at large radii, but an examination of their Fig. 4 suggests a slightly steeper logarithmic slope of ≈ -4 . The shape of the bulge+halo component at large radii is also similar to that found by Abadi et al. (2006), who performed cosmological SPH re-simulations of a small number of Milky Way-mass systems. In particular, these authors found logarithmic slopes of -2.9 at 100 kpc and -3.5 at the virial radius of their simulated galaxies.

While there is generally good agreement at large radii between our results and those of previous theoretical studies, some differences become apparent at small radii. In particular, Bullock & Johnston (2005) find that their profiles become increasingly flat at small radii, so that the logarithmic slope approaches -1 for $r \sim 10$ kpc, whereas our spheroid component shows no evidence of flattening at small radii - on the contrary, the profile actually steepens somewhat at radii of $r \approx 30$ kpc before leveling off to a slope of ≈ -3 at $r \lesssim 20$ kpc. This dramatic flattening in the Bullock & Johnston (2005) model is almost certainly in part due to the neglect of in situ star formation in their hybrid N-body model, but another potentially important factor is the treatment of the galaxy potential well, which does not feel the back-reaction due to the orbiting satellites in their idealised N-body simulations. For example, Cooper et al.

⁵ We use the high-resolution simulation here as it extends down to lower luminosities due to the smaller particle mass (although the accreted spheroid is assembled primarily from the most massive satellites). We have checked that there is good convergence in the LFs between the intermediate- and high-resolution runs for the most massive satellites ($M_V < -12$; see also the Appendix).

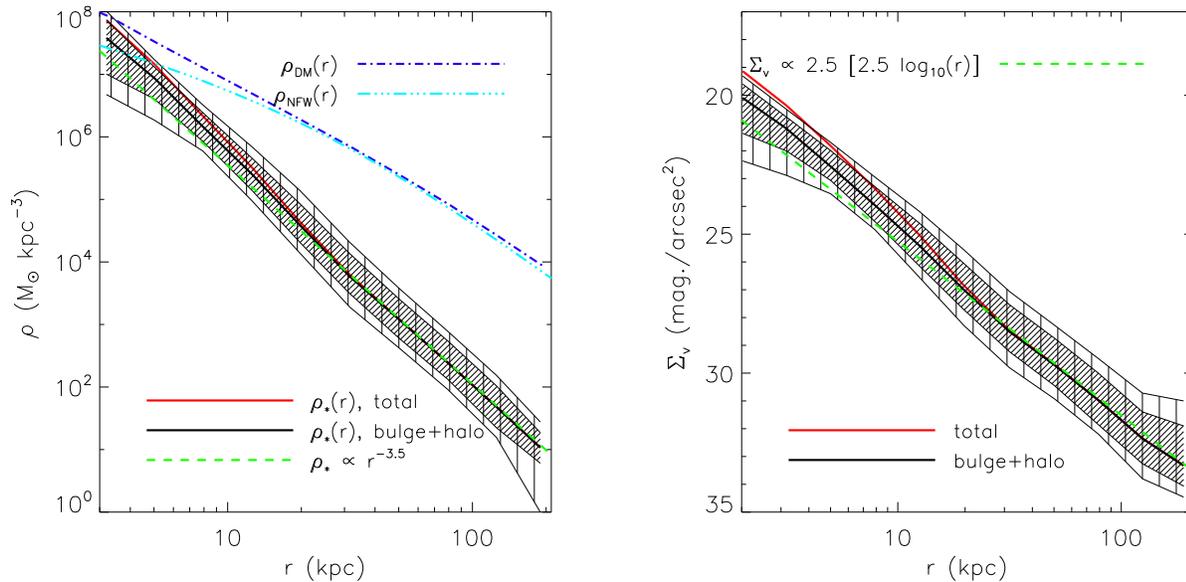


Figure 3. *Left:* Median spherically-averaged mass density profiles. The solid red curve is the median stellar mass density profile for all stars, whereas the solid black curve is the median stellar mass density profile for just the spheroidal (bulge+halo) component. The finely hatched region encloses the 14th and 86th percentiles of the stellar bulge+halo mass density profile, while the sparsely hatched region encloses the 5th and 95th percentiles. The dashed green curve represents a power-law profile with $\rho_* \propto r^{-3.5}$ with a normalisation chosen to match the median stellar halo mass density profile at large radii. The dot-dashed blue curve represents the median dark matter mass density profile, while the dot-dashed cyan curve represents a Navarro, Frenk & White (1996) mass density profile with $M_{200} = 1.3 \times 10^{12} M_\odot$, which corresponds to the median halo mass of our sample of simulated galaxies. *Right:* Median azimuthally-averaged surface brightness profile in the V-band for the randomly-oriented simulated galaxies. The solid red curve is the median surface brightness profile for all stars, whereas the solid black curve is the median profile for the bulge+halo. The finely (respectively sparsely) hatched region encloses the 14th and 86th (respectively 5th and 95th) percentiles of the bulge+halo surface brightness profile. The dashed green curve represents a power-law profile (functional form given in the legend) with a normalisation chosen to match the median surface brightness profile at large radii. No single power-law can match the simulated bulge+halo profiles over all radii. At radii of $r \lesssim 30$ kpc, a ‘bump’ is present, signaling the transition from accretion-dominated to in situ-dominated stars.

(2010) (see also Diemand et al. 2005), who applied a similar hybrid method to cosmological N-body simulations with ‘live’ galaxy potential wells, find a relatively steep distribution at small radii, with a logarithmic slope of ≈ -3 . This suggests that the treatment of the potential well is indeed relevant. Consistent with this hypothesis is that when we fit a power-law to the density profile *due to accreted stars only* at radii of $r \leq 30$ kpc, we still find a steep logarithmic slope of ≈ -2.9 (i.e., similar to that found by Cooper et al. 2010). The fact that our *total* (accreted+in situ) profiles show evidence for an increase in normalisation at small radii (at $r \approx 30$ kpc), whereas Cooper et al. (2010) find, if anything, a slight flattening, suggests that in situ star formation also plays an important role at small radii. We demonstrate that this is indeed the case in Section 4.

It is interesting to compare the mass density/surface brightness profiles of the simulated galaxies with those inferred for local galaxies. We first consider M31, which may represent the best test of the models, as its relatively close location and our external view point allow observations of low surface brightness levels down to ~ 32 mag arcsec⁻² and hence enable the tracing of the radial profiles out to large projected radii.

In Fig. 4 we compare the V-band surface brightness profiles of the simulated galaxies to that derived for M31.

To date the surface brightness profile of M31 has mainly been derived along (or close to) the minor axis. For the purposes of comparison we therefore also measure the minor axis profile of our simulated galaxies. This is done by orienting the simulated galaxies to an edge-on configuration (with $x-y$ plane being parallel to the plane of the disc) and measuring the surface brightness as a function of the distance z from the disc midplane within a cylindrical distance ($R = \sqrt{x^2 + y^2}$) of 5 kpc from the system center. Unfortunately, the mass resolution of our simulations prevents us from measuring the minor axis profile on a system-by-system basis⁶, so we have stacked all 412 edge-on galaxies to derive a mean minor-axis surface brightness profile. This is represented by the solid black curve in Fig. 4. For comparison, we show the minor axis surface brightness measurements of Pritchet & van den Bergh (1994), Ibata et al. (2007), and Gilbert et al. (2009a).

Quite remarkably, the shape of the stacked simulated profile matches that of the observed surface brightness profile of M31 over almost two decades in radius (from $r \approx 2$

⁶ This was possible for the azimuthally-averaged profiles in Fig. 3 because of the much larger surface area used for measuring the profiles.

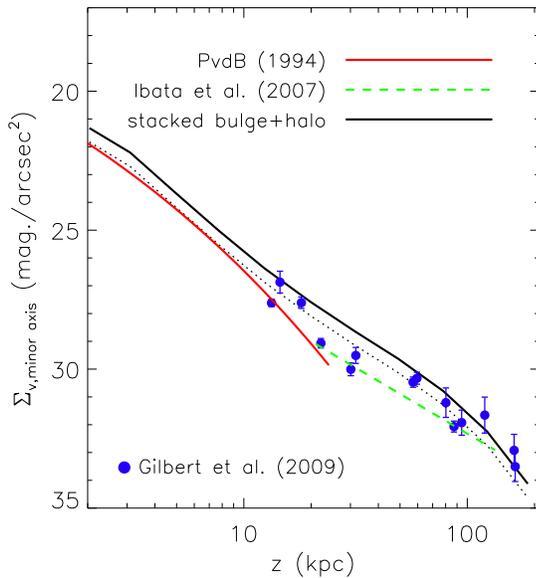


Figure 4. Comparison of the minor axis V-band surface brightness profile for the simulated galaxies with that of M31. The solid black curve represents the stacked profile derived from all 412 simulated disc galaxies. The dotted black curve is the stacked simulated profile shifted down by $0.5 \text{ mags arcsec}^{-2}$. The solid red curve represents the de Vaucouleurs profile fit to the minor axis profile of the M31 “bulge” derived by (Pritchert & van den Bergh 1994, PvdB). The dashed green line represents the power-law fit to the outer halo of M31 by Ibata et al. (2007). The solid blue circles represent the measurements of Gilbert et al. (2009a). The shape of the stacked profile is in very good agreement with that derived for M31 over a wide range of radii.

kpc to beyond 100 kpc), including the change in slope that occurs at radii of 20 – 30 kpc.

Note that it is the *shape* of the profile, rather than its normalisation, that is the more critical test of the models, since the normalisation of the observed profile requires making an uncertain conversion between the tracer population (Red Giant star counts in this case) to the total surface brightness. Typically this is done by assuming that the spheroid has an identical stellar luminosity function to that of local dwarfs or globular clusters and therefore one can use the relationship between integrated light and the same tracers in these local systems to re-normalise the halo. Of course this procedure will only be accurate if the stellar halo indeed has a similar stellar luminosity function to the local dwarfs or globular clusters. For the data plotted in Fig. 4, Pritchert & van den Bergh (1994) used the observations of the Milky Way globular cluster 47 Tuc for the conversion, while Ibata et al. (2007) and Gilbert et al. (2009a) shifted their data vertically to map onto the profile Pritchert & van den Bergh (1994) at radii of $\approx 10 - 30$ kpc. In spite of the potential for important systematics in normalisation of the observed profile, the stacked profile agrees with the data to an accuracy of $\approx 0.5 \text{ mags arcsec}^{-2}$. This is actually better than one might have naively expected, especially given that the system-to-system scatter in the simulated profiles is $\approx 3 \text{ mags arcsec}^{-2}$ (see Fig. 3).

We turn now to the Milky Way. Bearing in mind

that observations of the extended stellar distribution of the Milky Way are typically limited to $r \lesssim 40$ kpc, a wide variety of studies have found logarithmic slopes for the density profile of the stellar halo ranging from -2.5 to -3.75 , depending on the data and methodology that were used (Harris 1976; Zinn 1985; Saha 1985; Preston, Shtetman Beers 1991; Wetterer & McGraw 1996; Ivezić et al. 2000; Morrison et al. 2000; Yanny et al. 2000; Chiba & Beers 2000; Chen et al. 2001; Siegel et al. 2002; Vivas & Zinn 2006; Bell et al. 2008; Juric et al. 2008; Smith et al. 2009). This is quite compatible with results of the cosmological hydrodynamical simulations. Interestingly, some of the latest observations of the Milky Way halo, which are based on large samples of stars of the Sloan Digital Sky Survey (SDSS), find evidence for a change in the slope density profile at radii of around ≈ 30 kpc (Carollo et al. 2007; Juric et al. 2008), which is remarkably consistent with our simulations. Juric et al. (2008) suggest there also is a change in the shape of the halo at this radius, with the inner most regions being flattened and the outermost regions being more spherical in shape. We present a study of the shape and kinematics of the simulated galaxy haloes in McCarthy et al. (in prep.). However, a visual inspection of the simulated azimuthally-averaged and minor-axis profiles in Fig. 3 and 4 already reveals that our spheroids are indeed flattened within 30 kpc or so.

Observations of stellar haloes of other (more distant) external galaxies are more difficult. Currently, only a few stellar haloes of individual galaxies have been detected, and only down to levels of $\sim 28 - 29 \text{ mag arcsec}^{-2}$ (de Jong et al. 2007), corresponding to the inner regions. One can go deeper by stacking distant galaxies. For example, Zibetti et al. (2004) stacked more than 1000 stellar haloes of edge-on spiral galaxies observed in the SDSS and found that, out to about 25 kpc ($\Sigma_V \sim 31 \text{ mags arcsec}^{-2}$), the average density profile can be fitted relatively well with a power-law $\rho_* \sim r^{-3}$, very similar to what the simulations predict.

With our large sample of simulated galaxies (≈ 400), we are in an excellent position to characterise the system-to-system *scatter in the normalisation* of the ρ_* and Σ_V profiles, which is another potential test of the simulations. The solid black curves represent the median profiles of the bulge+halo components only. The finely-shaded regions enclose the 14th and 86th percentiles (i.e., 1σ) and the sparsely-shaded regions enclose the 5th and 95th percentiles (2σ) of the stellar bulge+halo profiles. The 2σ scatter in ρ_* is approximately an order of magnitude over the range of 10 – 100kpc, corresponding to a scatter in Σ_V of $\sim 3 \text{ mags arcsec}^{-2}$ over the same range. [In the Appendix, we test the numerical convergence of the stellar mass density profiles. We demonstrate there that our results (both median and scatter) are quite robust to a factor of eight increase in the mass resolution.]

It is also worth commenting briefly on the structure of the dark matter halo. In addition to the lines corresponding to the stellar components described above, the left panel of Fig. 3 also shows the median dark matter density distribution (as the blue curve) and an analytic NFW (Navarro, Frenk & White 1996) profile (the cyan curve). The analytic halo has a mass $M_{200} = 1.4 \times 10^{12} M_\odot$ (which corresponds to the median halo mass of our sample of simulated galaxies) and a concentration parameter of $c_{200} = 9.4$,

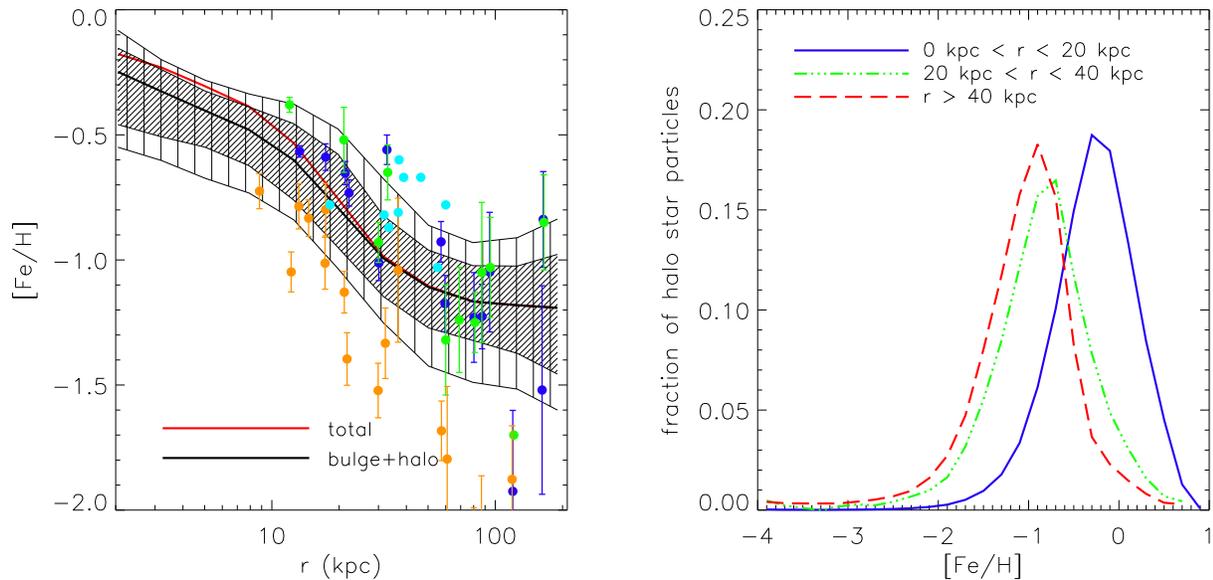


Figure 5. *Left:* Median spherically-averaged stellar metallicity profiles. The solid red curve is the median stellar metallicity profile for all stars, whereas the solid black curve is the median stellar metallicity profile for just the bulge+halo component. The finely (sparsely) hatched region encloses the 14th and 86th (5th and 95th) percentiles of the stellar halo metallicity profile. The solid blue, green, cyan, and orange circles represent M31 metallicity measurements of Gilbert et al. (2009a), Kalirai et al. (2006), Richardson et al. (2009), and Koch et al. (2008), respectively. The observational data have been scaled to the same set of solar abundances. *Right:* Median metallicity distribution function (MDF) in three radial bins. Significant negative metallicity gradients are a ubiquitous feature of the simulated galaxy populations.

which places it on the $M_{200} - c_{200}$ relation measured by Gao et al. (2008) for the Millennium Simulation. The dot-dashed cyan curve *is not a fit* to the dot-dashed blue curve.

A comparison of the dot-dashed blue curve with the solid red curve demonstrates that dark matter dominates over the stellar contribution down to very small radii (the two become comparable at a few kpc). Comparison of the dot-dashed blue and cyan curves shows that the NFW analytic form provides an excellent description of the dark matter density profile beyond about $r \approx 10$ kpc (i.e., $\approx 0.05r_{200}$). Within this radius, the dark matter profile steepens, presumably as a result of contraction of the dark matter halo due to the condensation of baryons (e.g., Blumenthal et al. 1986; Gnedin et al. 2004; Duffy et al. 2010).

Given that the NFW profile drops off as $\sim r^{-2}$ at intermediate radii and as $\sim r^{-3}$ at large radii, this implies that the stellar mass density profile drops off significantly faster than that of the dark matter, as is clearly visible in the left panel of Fig. 3. This is also consistent with the findings of previous theoretical studies of the stellar halo (e.g. Bullock & Johnston 2005; Abadi et al. 2006). The difference in the stellar and dark matter distributions likely stems from the fact that the stars in infalling satellites are more tightly bound, and thus less susceptible to tidal disruption, than their dark matter haloes. In addition, the most massive satellites (which also have the highest stellar mass fractions) will spiral more quickly into the center due to dynamical friction, which scales as mass squared. Finally, in situ star formation will occur preferentially in the central regions (due to the higher gas densities there). These

effects would produce a more centrally-concentrated stellar distribution, as seen in the simulations as well as in the observations.

3.2.2 $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$ radial profiles

The radial distribution of stellar metallicity is a particularly interesting test of the simulations, since previous simulations/models in the Λ CDM context have not been able to reproduce the significant gradients seen in M31 and the Milky Way. It is not presently clear why this is the case. It may signal a fundamental problem with galaxy formation in the Λ CDM context or that previous models of stellar haloes neglected important physics, such as in situ star formation or a sufficiently realistic implementation of chemodynamics.

In the left panel of Fig. 5 we show the median spherically-averaged metallicity profile and system-to-system scatter about the median. The metallicity, $[\text{Fe}/\text{H}]$, of each star particle⁷ is computed by taking the logarithm of the ratio of the iron-to-hydrogen mass fractions of the star particle and subtracting the logarithm of the iron-to-hydrogen mass fraction ratio of the Sun (which we assume

⁷ For consistency with the way the chemodynamics is done in the simulations, we used so-called ‘smoothed metallicities’ rather than ‘particle metallicities’. See Wiersma et al. (2009b) for discussion. Switching to ‘particle metallicities’ has the effect of shifting the profiles in the left panel of Fig. 5 down by ≈ 0.2 dex, which, as discussed later in the paper, is smaller than the uncertainties in the adopted empirical nucleosynthesis yields and Type Ia SNe rates.

is 0.00156, consistent with the recent measurements and 3D modelling of Asplund et al. 2005). A spherically-averaged metallicity profile is computed for each galaxy by computing the mean of metallicity (i.e., $\langle[\text{Fe}/\text{H}]\rangle$) of star particles in radial bins. The median profiles plotted in Fig. 5 are then computed by taking the median of all of individual metallicity profiles in radial bins.

Interestingly, the stellar metallicity distribution shows a prominent negative gradient over all radii, with the steepest decline exhibited over the range $20 \text{ kpc} \lesssim r \lesssim 40 \text{ kpc}$. This is independent of whether the disc component is included or not. The solid black curve and shaded regions indicate that *negative metallicity gradients are a ubiquitous feature of the simulated galaxy populations*. (This statement is further confirmed in the middle panel Fig. 9 in Section 4, which shows system-to-system scatter in the difference in the mean metallicity within and beyond 30 kpc.)

The solid blue, green, cyan, and orange circles in Fig. 5 represent metallicity measurements of M31 from Gilbert et al. (2009a), Kalirai et al. (2006), Richardson et al. (2009), and Koch et al. (2008), respectively⁸. The simulated profiles have a remarkably similar shape to that derived by Gilbert et al. (2009a) and Kalirai et al. (2006), which shows a drop of $\approx 0.7 - 0.8$ dex over the wide radial range of $\approx 10 - 150 \text{ kpc}$. The simulated profiles are also consistent with the measurements of Richardson et al. (2009), which probe a much smaller dynamic range. The profile derived by Koch et al. (2008), however, is somewhat steeper than the predicted profiles, particularly at very large radii. It is worth noting that Gilbert et al. (2009a), Kalirai et al. (2006), and Richardson et al. (2009) all employed similar methods in deriving the metallicity (in particular, they used spectroscopically-calibrated photometric methods), whereas Koch et al. (2008) used a new spectroscopic method that derives abundances from the calcium triplet (CaT) part of the spectrum. It is not presently clear what the origin is of the difference between the different data sets.

It is also encouraging that the normalisation of the simulation metallicity profiles agrees as well as it does (to within 0.2-0.3 dex accuracy) with the observations of M31. The simulations adopt empirical nucleosynthesis yields and Type Ia SNe rates, both of which are uncertain by a factor of a few (see Wiersma et al. 2009b), implying that there is freedom to vertically shift the simulated profiles up or down by a few tenths of a dex. Another potentially relevant factor is the selection of the observational fields. Often they are selected to have a relatively high surface brightness. Font et al. (2008) and Gilbert et al. (2009a) have shown that surface brightness tends to be positively correlated with metallicity, so it may well be the case that the current measurements are biased somewhat high. Large panoramic observations of M31 (such as PAndAS) will soon shed light on whether this bias is important.

⁸ Where appropriate we have adjusted the observational metallicity measurements (by applying a constant offset) to account for differences in the assumed solar abundances in the studies of Gilbert et al. (2009a), Kalirai et al. (2006), Richardson et al. (2009), and Koch et al. (2008) from the solar abundances of Asplund et al. (2005), which we use to normalise the simulated profiles.

For the Milky Way, metallicity measurements out to large galactocentric radii are much more challenging. However, recent results based on SDSS observations suggest a similar drop of $\approx 0.7 - 0.8$ dex from the Solar neighbourhood out to $\approx 30 - 40 \text{ kpc}$ (Carollo et al. 2007; Ivezić et al. 2008; Carollo et al. 2010; de Jong et al. 2010).

As alluded to above, most previous theoretical studies, particularly those based on accretion-only hybrid models (e.g., Font et al. 2006a; De Lucia & Helmi 2008; Cooper et al. 2010), did not find significant gradients. To our knowledge, previous studies based on hydrodynamic re-simulations of disc galaxies (e.g., Abadi et al. 2006; Zolotov et al. 2010) did not examine the radial distribution of the stellar metallicity⁹. As we demonstrate in Section 4, the change in slope at $r \sim 30 \text{ kpc}$ marks the transition from accretion ($r \gtrsim 30 \text{ kpc}$) to in situ ($r \lesssim 30 \text{ kpc}$) stellar halo formation.

While the spherically-averaged profiles of the simulated galaxies display prominent gradients, at any given radius there is a wide range of stellar metallicities. To illustrate the spread in metallicities as a function of radius, we show the median stellar metallicity distribution function (MDF) for three radial bins in the right panel of Fig. 5. For all three bins the MDF is similar in shape (in particular, fitting a Gaussian distribution yields similar widths of $\sigma \approx 0.5$ dex), but the peak shifts to lower metallicities with increasing radius, which is what gives rise (mathematically at least) to the metallicity gradients in the left panel of Fig. 5. A Gaussian provides a reasonable match to the MDFs near the peak and for higher metallicities, but it systematically under-predicts the fraction of halo stars with low metallicities. The MDFs in the right panel of Fig. 5 are qualitatively similar to that found for the M31 stellar halo (Gilbert et al. in prep) and other nearby galaxies (Mouhcine et al. 2005b).

Since the simulations include enrichment from both Type II and Type Ia supernovae, as well as stellar mass loss (i.e., AGB stars), we can also make predictions for the $[\alpha/\text{Fe}]$ radial profiles, which should be measurable in galaxies like M31 in the near future. The so-called α elements (such as oxygen, magnesium, and silicon) provide an additional constraint, as they are produced primarily in massive stars. In the left panel of Fig. 6 we show the median spherically-averaged $[\alpha/\text{Fe}]$ profile and scatter, where we have used both oxygen and magnesium to represent α . We have normalised the profiles using the oxygen-to-iron and magnesium-to-iron mass fractions of the Sun (5.2887 and 0.5355, respectively, Asplund et al. 2005). Interestingly, while the metallicity (i.e. $[\text{Fe}/\text{H}]$) profile shows a strong radial dependence (Fig. 5), the $[\alpha/\text{Fe}]$ distribution does not: $[\alpha/\text{Fe}]$ rises by only ≈ 0.1 dex from the center to the outskirts. This implies that the physical mechanism responsible for producing the iron does not change appreciably with galactocentric distance.

It is interesting to elucidate the lack of a strong radial dependence in the $[\alpha/\text{Fe}]$ profiles presented in the left panel of Fig. 6. The GIMIC simulations track not only the total iron mass fraction, but also the iron mass fraction pro-

⁹ But note that hydrodynamic re-simulations have been used to study the radial distribution of metallicity in *early-type* galaxies (e.g., Kobayashi 2004, 2005; Gibson et al. 2007) and significant gradients were shown to be common in these systems as well.

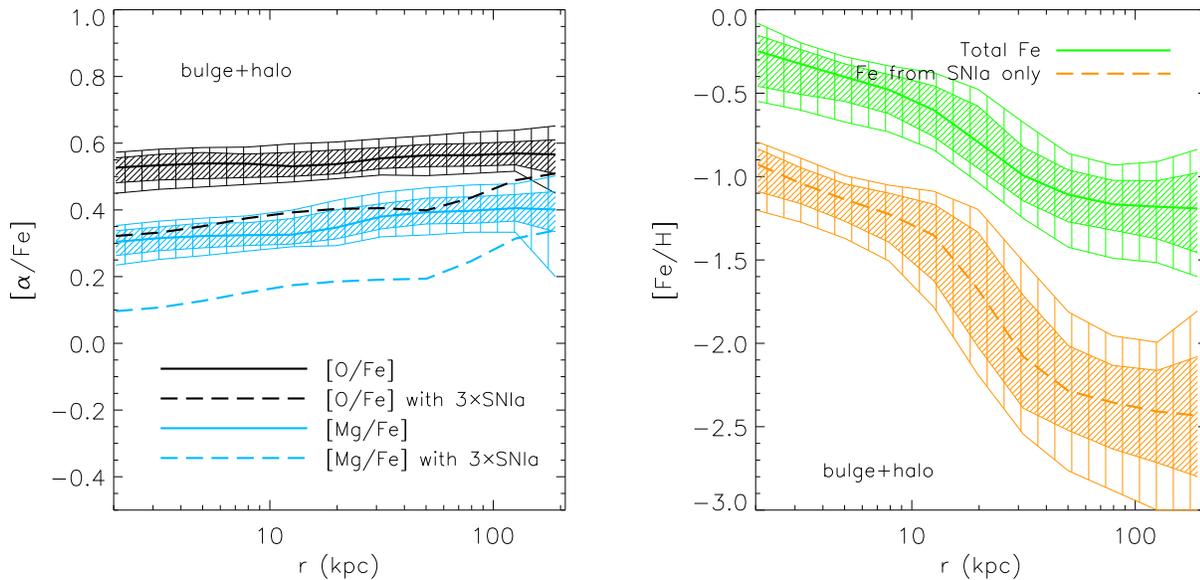


Figure 6. *Left:* Median spherically-averaged stellar $[\alpha/\text{Fe}]$ profiles. The solid black (light blue) curve is the median $[\text{O}/\text{Fe}]$ ($[\text{Mg}/\text{Fe}]$) profile for the bulge+halo component. The finely (sparsely) hatched region encloses the 14th and 86th (5th and 95th) percentiles. The dashed black (light blue) curve shows the effect on the median $[\text{O}/\text{Fe}]$ ($[\text{Mg}/\text{Fe}]$) profiles if the rate of Type Ia SNe were to be increased by a factor of 3 in the simulations. *Right:* Median spherically-averaged stellar $[\text{Fe}/\text{H}]$ profiles for the spheroidal component using the total mass of iron (green) and just the iron produced by Type Ia supernovae (orange). The finely (sparsely) hatched regions enclose the 14th and 86th (5th and 95th) percentiles. There is no significant radial variation in the simulated $[\alpha/\text{Fe}]$ profiles, owing to the dominance of metal production by Type II SNe at all radii.

duced by Type Ia supernovae only. This therefore allows us to determine, for any given star particle, the individual contribution of Type Ia to the iron mass fraction. In the right panel of Fig. 6 we show the $[\text{Fe}/\text{H}]$ profiles for the total iron mass fraction and for the iron produced solely by Type Ia supernovae. As it can be clearly seen, Type Ia supernovae make only a minor contribution to the metallicity of the stellar halo. Type II supernovae are therefore the dominant producers of metals at all radii. The very mild gradient in $[\alpha/\text{Fe}]$ is due to the relatively larger importance (though still small in an absolute sense) of Type Ia for stars located in the inner regions.

It is important to acknowledge potential caveats in the above prediction. Besides the nucleosynthesis yields, the rate of Type Ia supernovae is very uncertain. Rather than try to calculate the rate of Type Ia supernovae using a physical model, the simulations adopt empirical Type Ia rates, which are still uncertain by a factor of a few (see Fig. A6 of Wiersma et al. 2009b). If the Type Ia rate is actually a factor of a few higher than current estimates imply, this would have the effect of producing a steeper positive $[\alpha/\text{Fe}]$ gradient (see the dashed curves in the left panel of Fig. 6, which show the effect of boosting the number of Type Ia SNe by a factor of 3), since the stars in the central regions were formed more recently (see Section 4) and therefore the gas from which they formed had more time to be polluted by Type Ia supernovae.

4 PHYSICAL ORIGIN OF THE RADIAL DENSITY AND METALlicity PROFILES OF THE STELLAR HALOES

In the previous section we showed that the GIMIC simulations produce disc galaxies with stellar haloes that bear a strong resemblance to those observed around nearby galaxies, including M31 and our own Milky Way. In the present section we use the simulations to trace back the physical origin of the stellar haloes. We first quantify the relative importance of accretion vs. in situ star formation for the formation of the stellar halo. We show that the break in the mass density and surface brightness profiles near $r \sim 30$ kpc and the sharp rise in the stellar halo metallicity profile at similar radii are all caused by the transition from accretion-dominated to in situ-dominated star formation (Section 4.2). We also explore the time of star formation and assembly of the stellar halo (Section 4.3) and the degree of radial migration of the in situ component (Section 4.4).

4.1 Distinguishing accretion and in situ formation

To assess the relative importance of accretion vs. in situ star formation we must first define what we mean by these terms. We define in situ star formation as star formation that takes place in the most massive subhalo of the most massive progenitor (MMP) FoF group of the present-day system. Stars are said to be ‘accreted’ if they formed either in another FoF group or in a satellite galaxy (i.e., not the dominant subhalo) of the MMP. In practice, only a small proportion (2%) of the stellar halo is from stars that formed

in satellites of the MMP and were then stripped; most of the accreted stars were formed in other FoF groups (i.e., other haloes) before they became satellites of the MMP.

The procedure by which we identify the MMP and determine whether a given star particle was accreted or formed in situ is as follows. For a given Milky Way-mass system at $z = 0$ we select all of the dark matter particles within r_{200} belonging to the dominant subhalo. We use the unique IDs assigned to these particles to identify the FoF group in previous snapshots that contains the largest number of these particles. This FoF group is said to be the MMP and the dominant subhalo of that FoF group is assumed to be the most massive progenitor of the dominant subhalo of the system at $z = 0$. Each of the star particles that comprise the $z = 0$ stellar halo is tracked back in time to the FoF group and subhalo it belonged to, if any, when it formed from a gas particle. If, at the time of star formation, the star particle is a member of the dominant subhalo of the MMP, the star particle is said to have formed in situ. If it formed in another FoF group or in a non-dominant subhalo of the MMP, then it is said to have been accreted¹⁰.

The number of snapshots output for each of the GIMIC volumes is approximately 60, and they span the redshift range $z = 20$ to $z = 0$. The time of star formation of each star particle will, in general, not correspond to the redshift at which the snapshot data was written. We therefore must use the snapshot that is closest in time to identify where the star particle was when it formed. This could, in certain circumstances, lead to a misidentification of the location of star formation (i.e., to which subhalo it belonged when it formed). However, we have explicitly checked that this is not an important issue by re-running our tracing code using only every second snapshot. We find excellent quantitative agreement with our results from using all of the available snapshots. We therefore conclude that the redshift sampling of our snapshot data is sufficient to deduce whether a given particle formed in situ or was accreted.

We trace each of the $z = 0$ stellar halo particles back in time as far as $z = 10$ to determine whether they formed in situ or were accreted. We ignore the very small fraction (0.01%) of halo stars that formed before $z = 10$.

4.2 Stellar mass density and metallicity profiles by formation mode

In Fig. 7 we show the median spherically-averaged mass density profile of systems split according to formation mode (i.e., in situ vs. accreted). A comparison of the dot-dashed red and dashed blue curves demonstrates that in situ star formation contributes significantly to (dominates) the stellar halo within a radius of 30 (20) kpc. Beyond 30 kpc the majority of the stars were formed in satellites prior to accretion. This transition accounts for most of the deviation of the total stellar halo mass density profile from a pure power-law distribution of $r^{-3.5}$ at small radii (compare the solid black and dashed green curves in the left panel of Fig. 3).

¹⁰ We have also allowed for the possibility that the star particle may have formed outside of bound substructures but find that this mode of formation is rare, accounting for 2% of the stellar halo by mass.

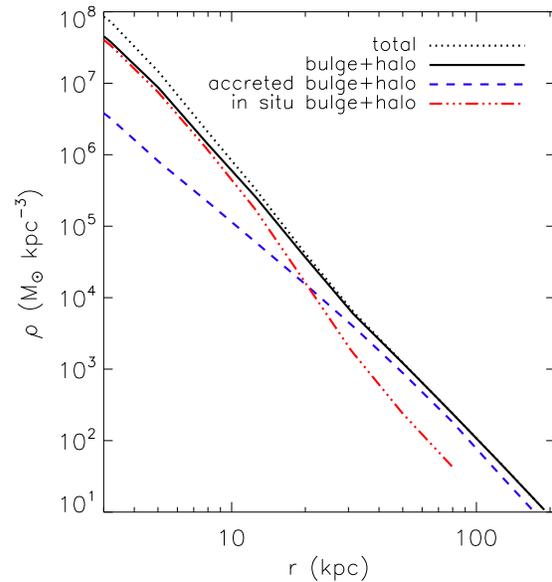


Figure 7. Median spherically-averaged stellar mass density profiles. The dotted and solid black curves represent all stars and stars in the spheroid only, respectively. The dotted black curve represents all stars (disc+spheroid). The dashed blue curve represents accreted halo stars (i.e., stars that formed outside the most massive progenitor). The dot-dashed red curve represents spheroid stars that formed in situ. In situ star formation contributes significantly to (dominates) the stellar halo within a radius of 30 (20) kpc.

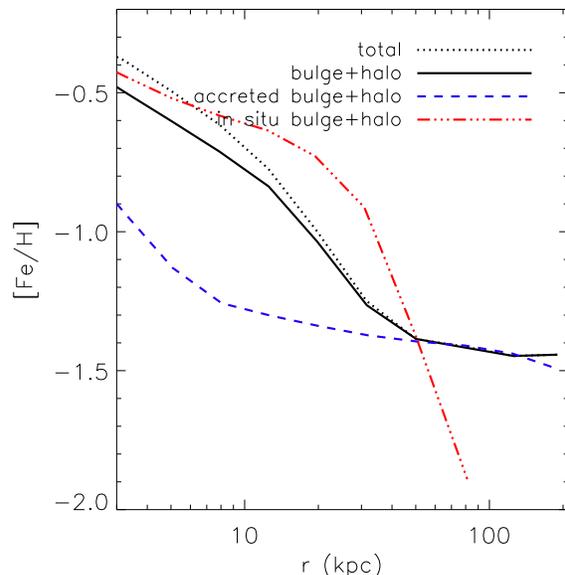


Figure 8. Median spherically-averaged metallicity profiles. The curves have the same meaning as in Fig. 7. For $r < 50$ kpc, stars formed in situ have higher metallicity and their dominance of the stellar halo by mass at small radii is what produces the strong radial variation in $[\text{Fe}/\text{H}]$.

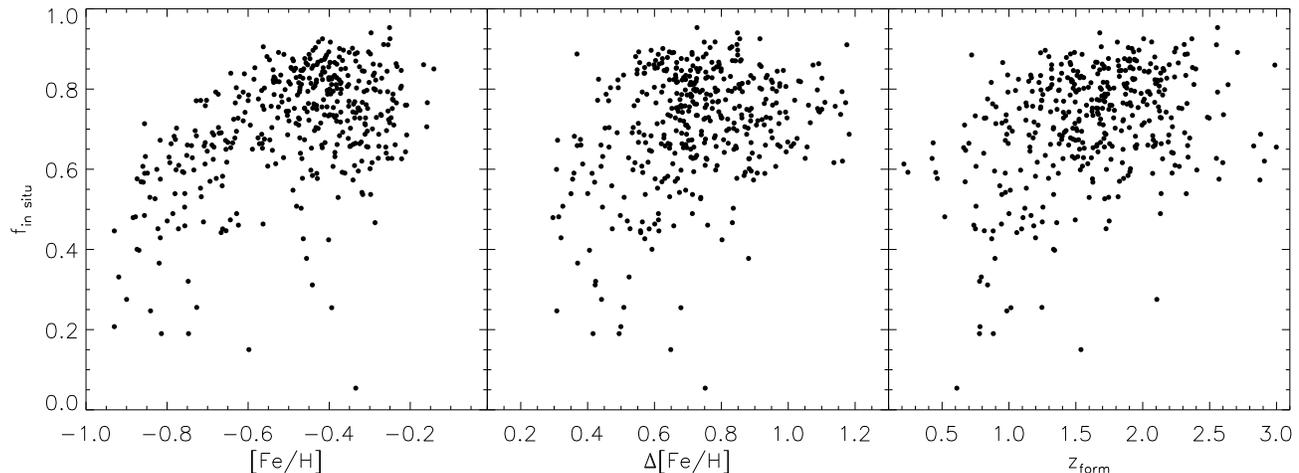


Figure 9. The current stellar mass fraction of the bulge+halo component due to in situ star formation as a function of (*Left*) mean metallicity, $[\text{Fe}/\text{H}]$, (*Middle*) the metallicity difference between the inner and outer bulge+halo, $\Delta[\text{Fe}/\text{H}]$ (calculated within and beyond 30 kpc), and (*Right*) system formation epoch, z_{form} (i.e., the redshift when half the present-day *total* mass (gas+stars+DM) is in place). Each data point represents a galaxy. Galaxies with higher in situ mass fractions generally have higher mean metallicities, more prominent metallicity gradients, and are dynamically older, although the system-to-system scatter in these trends is large.

Averaged over all of the simulated galaxies, in situ stars account for approximately 68% of the current mass of the stellar spheroids (with considerable system-to-system scatter, as shown in Fig. 9), while accreted stars constitute approximately 32%, with 28% of the mass formed in satellites prior to accretion, 2% formed in satellites after accretion, and 2% accreted smoothly. Our inferred average in situ fraction is therefore larger than that found recently by Zolotov et al. (2009), who found that 3 out of their 4 re-simulated galaxies had in situ mass fractions of $\lesssim 20\%$. However, part of this difference is likely due to the fact that we have made no attempt to distinguish between bulge and halo components (see Section 2.1 for a discussion of this point). Consistent with this is that both Abadi et al. (2006) and Oser et al. (2010) find in situ mass fractions of $\sim 40 - 50\%$ and, like us, these authors have not decomposed the bulge from the halo¹¹. In terms of our simulations, if we limit our comparison to $r > 20$ kpc (well beyond any likely bulge component), the in situ fraction drops to 20% while the accreted fraction rises to 80% (with 68% formed in satellites prior to accretion, 3% formed in satellites after accretion, and 9% accreted smoothly).

In Fig. 8 we show the median spherically-averaged stellar metallicity profile split by formation mode. Here we see that the gradient in the total stellar metallicity profile (solid black curve) is driven by the generally high metallicity of stars formed in situ combined with the fact that the in situ stars dominate the stellar halo by mass in the inner regions. The metallicity of accreted stars shows no strong variation with radius (except possibly in the central few kpc). The fact that the overall metallicity gradient is established as the result of a transition from in situ-

dominated to accretion-dominated stars, appears to resolve the longstanding problem of the inability of purely accretion-based models (e.g., Font et al. 2006a; De Lucia & Helmi 2008; Cooper et al. 2010) to reproduce the strong variation in the mean stellar metallicity with radius observed in M31 (e.g., Kalirai et al. 2006; Koch et al. 2008) and the Milky Way (e.g., Carollo et al. 2007; de Jong et al. 2010).

In Fig. 9 we show scatter plots of the in situ mass fraction as a function of the mean metallicity of the bulge+halo (left panel) and the metallicity difference between the mean metallicity within and beyond 30 kpc (middle panel). Each data point represents a simulated galaxy. Although there is considerable system-to-system scatter, there are identifiable correlations between the in situ mass fraction and the mean metallicity and metallicity difference, in the sense that galaxies with higher in situ mass fractions have higher mean metallicities and more prominent metallicity gradients. We discuss the right panel of Fig. 9 below.

4.3 The time of star formation and assembly of the stellar halo

We now examine the age of the stars and the assembly time of the spheroidal stellar component. In Fig. 10 we show profiles of the median spherically-averaged formation look-back time (i.e., the age) of stars, split by formation mode. Accreted stars (dashed blue curve) are typically quite old ($\approx 11 - 12$ Gyr), whereas most of the stars that formed in situ and which are located within 30 kpc or so, have a median age of $\approx 7 - 8$ Gyr.

The finding that stars that were formed in situ are, on average, younger than accreted stars, is consistent with our expectations based on the hierarchical growth of structure. In particular, stars that formed in situ are, by definition, those that formed in the MMP, whereas accreted stars formed in less massive and typically dynamically older satellites. For a $1 - 4 \times 10^{12} M_{\odot}$ (DM) halo the epoch of forma-

¹¹ One possible reason why Abadi et al. (2006) and Oser et al. (2010) find slightly lower in situ mass fractions than us, is that their simulations ignore metal-line cooling, which is expected to be important.

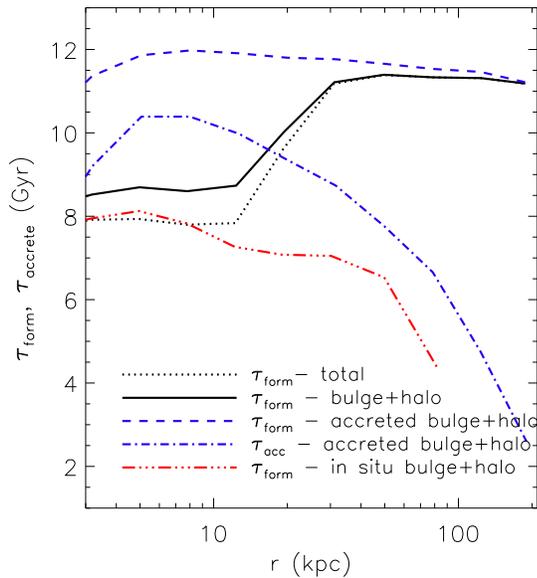


Figure 10. Median spherically-averaged stellar age (τ_{form}) profiles. The solid black curve represents all stars in the spheroidal component. The dotted black curve represents all stars (disc+spheroid). The dashed blue curve represents accreted stars in the spheroid. The triple dot-dashed red curve represents spheroid stars that formed in situ. The dot-dashed blue curve represents the time of accretion of accreted stars (i.e., the time when the system in which the stars formed became a satellite of the MMP). The accreted component is old ($\approx 11 - 12$ Gyr), but the lookback time at which it was assembled varies strongly with radius, with the outer most regions having been accreted quite recently. The in situ component is younger and was typically formed at $z \approx 1 - 1.5$.

tion, defined as the redshift where half the final dark matter mass is in place (Lacey & Cole 1993), is $z \approx 1 - 1.5$ (e.g., Wechsler et al. 2002), which corresponds roughly to the average star formation time of the stars formed in situ. Of course, baryons feel not only the effects of gravity, but also those of non-gravitational processes such as cooling and heating due to feedback (e.g., from supernovae). Feedback in particular can result in the delay (or prevention) of significant star formation. C09 have examined the star formation histories of galaxies in the GIMIC simulations and found that the peak occurs at $z \sim 1 - 3$ for normal galaxies, but at $z \sim 3 - 6$ for dwarf galaxies (which are the systems in which the vast majority of the accreted stars formed).

Returning to Fig. 9, we plot in the right hand panel the in situ mass fraction of the galaxy against the formation epoch of the system [i.e., the redshift for which the MMP had a total mass due to gas+stars+DM that is half of the final ($z = 0$) total mass]. There exists a correlation (again with considerable scatter) between in situ mass fraction and formation epoch, in that older systems have higher in situ mass fractions. Physically this makes sense, as older systems have been able to form stars quiescently for a longer period of time, while dynamically younger systems accreted relatively larger amounts of mass recently. The fact that dynamically older systems have higher in situ mass fractions is consistent with the findings of Zolotov et al. (2009), although our

improved statistics shows that the scatter in this correlation is large.

It is interesting that, even though the accreted stars are quite old (as indicated by the dashed blue curve in Fig. 10), many of them were accreted recently, as indicated by the dot-dashed blue curve in Fig. 10. In agreement with hybrid simulations of pure accretion-based halo formation (e.g., Bullock & Johnston 2005; Font et al. 2006a), we find that the inner accreted halo was formed a long time ago ($\tau_{\text{acc}} \approx 9 - 10$ Gyr), but that the outer accreted halo was assembled as recently as 5 Gyr ago. We therefore expect to see increased spatial lumpiness and more coherent velocity structure at large radii, since there has been less time since accretion for phase mixing and violent relaxation to act. We plan to examine the degree of substructure in the spheroidal component in a future paper.

4.4 Radial migration of in situ stars

We have shown that it is in situ star formation that is responsible for the change in the nature of the stellar halo at $r \lesssim 30$ kpc. Did these stars form at small radii and where then dispersed, as suggested by Zolotov et al. (2009) (see also Purcell et al. 2010)?

Zolotov et al. (2009) found that the in situ stars in their re-simulations were formed from gas that was brought into the main system in smooth ‘cold flows’. In their simulations the stars formed at very small galactocentric distances of ≈ 1 kpc. These stars were then dispersed out to larger radii (as far as 20 kpc or so) by $z \approx 2$, with major mergers acting as the primary dispersal mechanism.

In Fig. 11 we show the distribution of initial (i.e., at the time of star formation, z_{form}) and final ($z = 0$) radii of stars that formed in situ in the GIMIC simulations. This distribution is computed by first creating a regular (50×50) 2D grid in initial and final radii (in log space, from $\log_{10}(r) = 0.0 - 2.5$ in steps of 0.05) and then counting the number of in situ star particles in each pixel. The map is normalised by the total number of in situ star particles. The red, yellow, green, and blue iso-density contours enclose 50%, 68%, 90%, and 95% of the total number of in situ star particles, respectively. Fig. 11 represents a stacked distribution of all the in situ stars of all simulated galaxies.

Fig. 11 shows that the vast majority of in situ stars are presently located at, and were also formed at, relatively small radii ($r \lesssim 30$ kpc). At these radial distances there is a relatively large scatter about the $r[z = 0] = r[z_{\text{form}}]$ relation, indicating that in situ stars have migrated in both directions (i.e., toward and away from the present galaxy center). A very small population of in situ stars are presently located at relatively large radii ($r \gtrsim 50$ kpc). According to Fig. 11 these stars were formed at slightly larger radii and have since migrated inward somewhat. It is conceivable that these stars formed in unresolved substructures.

Our results therefore differ from those of Zolotov et al. (2009) in terms of the radius of in situ star formation, with the in situ stars in our simulations forming at radii comparable to where they are presently located. Our in situ stellar population is also younger than that of Zolotov et al. (2009) (who found $z_{\text{form}} \approx 3$), having formed primarily at $z \sim 1 - 1.5$, which is similar to that also reported recently by Oser et al. (2010) for galaxies with masses similar to

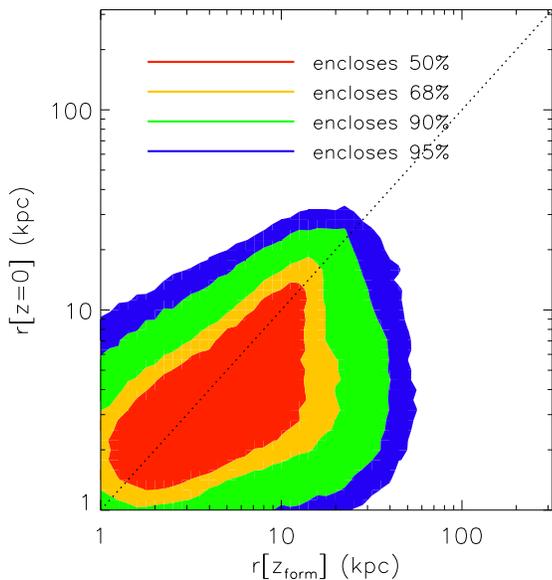


Figure 11. Distribution of initial (i.e., at redshift of star formation, z_{form}) and final ($z = 0$) radii of stars formed in situ. The red, yellow, green, and blue filled iso-density contours enclose 50%, 68%, 90%, and 95% of the total number of in situ star particles, respectively. The dotted curve represents the $r[z = 0] = r[z_{\text{form}}]$ relation. On average, in situ stars formed at radii comparable to where they are presently located, although migration in both directions (towards and away from the center) is significant.

those in our sample. The differences in the formation radii and ages of the in situ stars in our simulations and those of Zolotov et al. (2009) would be expected if the feedback in their simulations is relatively less efficient in low-mass haloes. This indeed appears to be the case, as their simulations significantly overproduce the number of bright satellite galaxies orbiting Milky Way-mass galaxies today (see Fig. 4 of Governato et al. 2007).

More importantly, the above differences imply a different formation scenario for in situ stars in our simulations compared with those of Zolotov et al. (2009). In McCarthy et al. (in prep.), we will show that the in situ stars form primarily from the cooling of *hot* gas (rather than cold flows, as advocated by Zolotov et al. 2009) during the initial assembly of the disc. This has profound implications for a wide range of properties of these stars today, including their spatial distribution, kinematics and metallicities.

5 SUMMARY AND DISCUSSION

We set out to address the question of the possible dual-nature of the extended stellar distributions around the Milky Way, M31, and other $\sim L_*$ nearby galaxies. In particular, what are the relative contributions (by mass) to the stellar spheroid from tidal stripping/disruption of satellites and ‘in situ’ star formation? And what is the origin of the metallicity gradients observed in the Milky Way and M31?

We have used the GIMIC suite of cosmological hydrodynamical simulations to address these important questions.

GIMIC is ideal for this purpose as it contains large numbers (≈ 400) of Milky Way-mass galaxies simulated at high-resolution ($M_g \simeq 1.16 \times 10^7 h^{-1} M_\odot$ and a gravitational softening of $1 h^{-1}$ kpc) and with sophisticated prescriptions for radiative cooling, chemodynamics, and supernova feedback.

Our main conclusions are as follows:

- The simulated stellar mass density of the spheroid (bulge+halo) falls off as $r^{-3.5}$ at large radii ($r \gtrsim 30$ kpc), in agreement with the findings of previous theoretical studies based on hybrid analytic+N-body methods and cosmological hydrodynamical re-simulations, as well as deep observations of M31. Fitting all ≈ 400 simulated galaxies separately, we find the 1-sigma (2-sigma) scatter in the logarithmic slope is ≈ 0.4 (≈ 0.6).

- In the inner region, the simulated mass density profile can be fit very well with a de Vaucouleurs profile with a median $r_{\text{eff}} \approx 1.8$ kpc (with considerable system-to-system scatter), and relatively well with a power-law with logarithmic slope $\approx -3.0 \pm -0.4$. This is shallower than in the outer regions, but note that the mass density profile has a higher normalisation at small radii, so that a power-law with slope -3.5 extrapolated to the inner regions underpredicts the mass density there. The increased normalisation and still fairly steep profile at small radii differs from that found previously with hybrid analytic+N-body methods, which is most likely due to their neglect of in situ star formation. The shape of our simulated mass density profile matches that inferred for M31 (see Fig. 4) and other nearby galaxies remarkably well.

- *Significant negative metallicity gradients are a ubiquitous feature of the simulated galaxy populations.* Typically $[\text{Fe}/\text{H}]$ drops by 0.6 to 0.9 dex from the inner ($r < 30$ kpc) to the outer ($r \approx 30 - 100$ kpc) regions (see middle panel of Fig. 9). This drop is compatible with that inferred for the MW (e.g., de Jong et al. 2010) and M31 (see Fig. 5).

- On average, 2/3 of the mass of the stellar spheroid (bulge+halo) is composed of stars that formed in situ, but with considerable system-to-system scatter. Limiting the comparison to $r > 20$ kpc, this fraction drops to 1/5 on average.

- The large contribution from in situ star formation in the inner regions drives the change in slope of the mass density profile at $r \simeq 30$ kpc and also gives rise to large-scale metallicity gradients.

- On average, the in situ stars were formed at radii comparable to their present distance from the Galactic center, with considerable migration in both directions (towards and away from the center; see Fig. 11). Typically, the stars were formed at $z \approx 1 - 1.5$, in a proto-disc. See McCarthy et al. (in prep.) for further discussion.

In the introduction (Section 1) we noted that the large differences in the properties of the stellar bulge+halo of the Milky Way and M31 (particularly their metallicities) has raised questions about the compatibility with the Λ CDM scenario. As we have shown, however, there is considerable system-to-system scatter in the mean metallicities (as well as in the prominence of metallicity gradients) of the simulated Milky Way-mass systems (see Fig. 9). The scatter is roughly compatible with the observed differences between the Milky Way and M31.

M31 is typically considered to be more dynamically ac-

tive than the Milky Way (Hammer et al. 2007, 2010). The correlation between the epoch of formation and the in situ mass fraction found in the simulations would therefore suggest that the Milky Way has a higher in situ mass fraction than M31 (as also argued by Zolotov et al. 2009). However, this is difficult to reconcile with the fact that we find that the in situ mass fraction is also correlated with system metallicity, in the sense that more metal rich systems have higher in situ fractions (and that M31 has a more metal rich spheroidal component than the Milky Way). However, as we have already noted, the simulations predict a large degree of scatter in these correlations, which may reconcile this issue.

It would clearly be very interesting to *measure* the in situ mass fraction of the spheroidal components of the Milky Way and Andromeda and compare it directly to the predictions of the simulations (e.g., in Fig. 9). Given their different formation histories, we expect there to be differences in the detailed spatial distributions and kinematics of the in situ and accreted stellar components, and this may hold the key to unambiguously separating these components from one another. We examine the kinematic properties, as well as the 3D spatial structure, of the in situ and accreted components in McCarthy et al. (in prep.).

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APPENDIX: NUMERICAL CONVERGENCE

Here we investigate the sensitivity of our results to numerical resolution. As noted in Section 2, the -2σ sphere has been simulated at eight times better mass resolution than the intermediate-resolution runs analysed in the present study. In Fig. 11 we compare the mass density and metallicity profiles (median + scatter) of 50 Milky Way-mass disc galaxies in the high-resolution simulation with those of the ≈ 400 galaxies in the intermediate-resolution runs. Overall, the agreement between the high-resolution and intermediate-resolution GIMIC simulations is good for both the mass density and metallicity profiles.

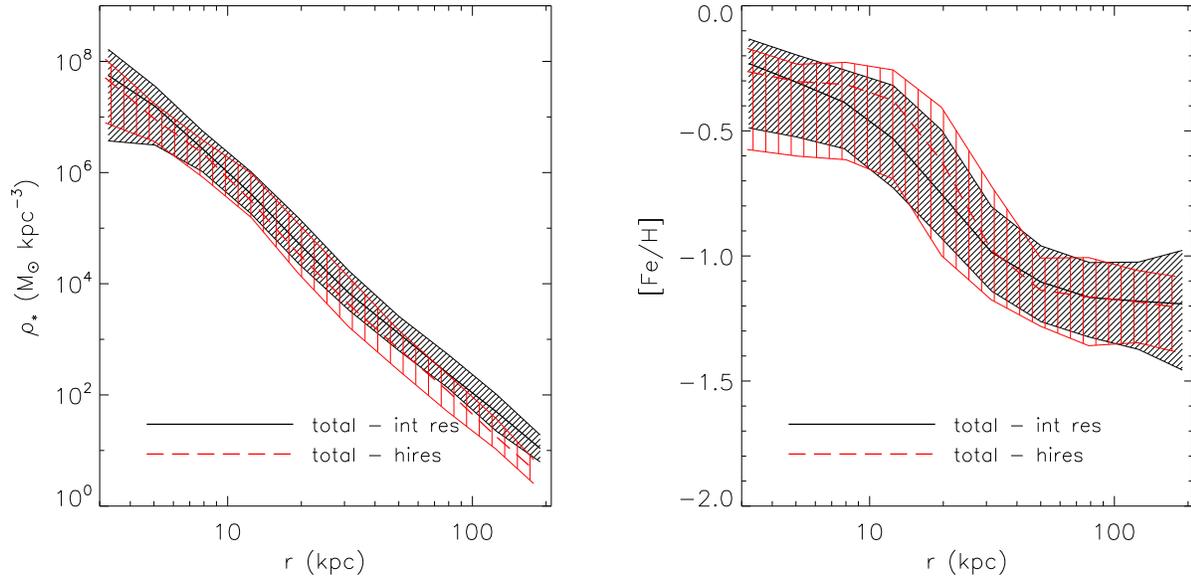


Figure 12. *Left:* Median spherically-averaged stellar mass density profiles for all Milky Way-mass systems in the high- and intermediate-resolution simulations. The scatter around the mean is indicated in each case by the different hashed regions (thinly spaced oblique lines for the intermediate-resolution run and widely spaced vertical lines for the high-resolution run). *Right:* Median spherically-averaged stellar metallicity profiles for the two runs. Overall, the agreement between the high-resolution and intermediate-resolution GIMIC simulations is good for both the mass density and metallicity profiles.