

Asymmetries in the Type II_n SN 2010jl[★]

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ABSTRACT

Aims. We study possible signs of asymmetry in the luminous Type II_n SN 2010jl, to obtain independent information on the explosion geometry.
Methods. We obtained optical linear spectropolarimetry of SN 2010jl two weeks after the discovery, in the spectral range 3700–8800 Å.
Results. The object exhibits a continuum polarization at a very significant and almost constant level (1.7–2.0%). Marked line depolarization is seen at the positions of the strongest emission features, like H α and H β . This implies that the line forming region is well above the photosphere. The continuum polarization level (1.7–2.0%) indicates a substantial asphericity, of axial ratio ≤ 0.7 . The almost complete depolarization seen at H α suggests a very low level of interstellar polarization ($\leq 0.3\%$). This rules out the presence of relevant amounts of dust in the progenitor environment at the time of our observations. From a polarimetric point of view, SN 2010jl appears to be very similar to the two other well studied Type II_n SNe 1997eg and 1998S, establishing a strong link within this class of objects.

Key words. supernovae: general - supernovae: individual: SN 2010jl - ISM: dust, extinction - techniques: polarimetry

1. Introduction

Type II_n Supernovae (SNe) are characterized by intense, composite-profile emission lines, and by the absence of the broad P-Cygni absorption troughs that are the distinguishing feature of *classical* Type II SNe (Schlegel 1990). The early spectra of these objects are explained in terms of a strong interaction between the fast expanding SN ejecta and a dense, slow moving circumstellar medium, fed by a significant mass loss undergone by the progenitor star (Chugai 1997). However, many questions about these objects, the mass loss history of their progenitors, the geometry of the explosion and the circum-stellar material (CSM) remain unanswered. Even the explanation for the origin of the broad emission seen at the base of the intermediate and narrow components has been questioned (Chugai 2001). With its unique capability of retrieving geometrical information, spectropolarimetry constitutes an independent way of attacking these open issues.

The bright SN 2010jl was discovered on November 3.5, 2010 (Newton & Puckett 2010). A spectrum obtained on

November 5 revealed it was a Type II_n (Benetti et al. 2010), with the typical, multi-component emission lines of hydrogen, and weaker helium features. Based on archival images of the explosion site, Smith et al. (2010) concluded that the progenitor of this luminous was probably a massive star, with an initial mass larger than 30 M_{\odot} . Soon after the discovery, we started an optical/near-IR follow-up using a number of ground-based facilities. The results will be presented and discussed elsewhere. In this letter we focus on the spectropolarimetry of SN 2010jl obtained about two weeks after its discovery, with the aim of detecting possible signs of asymmetries in the ejecta and/or in the CSM that could provide additional constraints on the geometry of the explosion environment.

2. Observations and data reduction

We observed SN 2010jl on November 18.2 UT, i.e. 14.7 days after its discovery, using the Calar Alto Faint Object Spectrograph (CAFOS), mounted at the 2.2 m telescope in Calar Alto, Spain (Meisenheimer 1998). In this multi-mode instrument, equipped with a 2K \times 2K SITE-1d CCD (24 μ m pixels, 0.53 arcsec/pixel), polarimetry is performed by introducing into the optical path a Wollaston prism (18'' throw) and a super-achromatic half-wave plate (HWP). To reduce some

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known instrumental problems (see Patat & Romaniello 2006), we used 4 half-wave plate angles (0, 22.5, 45 and 67.5 degrees). All spectra were obtained with the low-resolution B200 grism coupled with a 1.0 arcsec slit, giving a spectral range 3300-8900 Å, a dispersion of ~ 4.7 Å px⁻¹, and a FWHM resolution of 14.0 Å (640 km s⁻¹ at H α). For each HWP angle we obtained two exposures of 20 minutes each. The resulting signal-to-noise ratio per pixel on the single beam is ~ 180 at 5000 Å. Data were bias and flat-field corrected, and wavelength calibrated using standard tasks within IRAF. Stokes parameters (Q , U), linear polarization degree (P), and position angle (θ) were computed by means of specific routines written by us. Finally, polarization bias correction and error estimates were performed following the prescriptions described by Patat & Romaniello (2006). The HWP zero-point angle chromatism was derived using the polarization standard star BD+59d389 ($P(V)=6.7\pm 0.2\%$, $\theta=98.1$ degrees; Schmidt et al. 1992), observed on November 18.8 UT with the same instrumental setup. Instrumental polarization was analyzed using the unpolarized star HD 14069 ($P(V)=0.02\pm 0.02\%$; Schmidt et al. 1992), and found to vary very slowly as a function of wavelength between 0.3% and 0.4%. Once this correction is applied to the data of BD+59d389, the resulting linear polarization degree is consistent with the published values to within 0.1%. To increase the signal-to-noise ratio, the final Stokes parameters were binned in ~ 51.2 Å wide bins (11 pixels). This results in a nominal RMS error in the polarization of 0.1% at 5000Å, where the continuum signal-to-noise ratio per resolution element reaches its maximum (580). Flux calibration was achieved by observing a spectrophotometric standard star with the polarimetric optics inserted. Wavelengths were corrected to the rest frame adopting a recession velocity of 3214 km s⁻¹ for the host galaxy (Koribalski et al. 2004).

3. Reddening and interstellar polarization

The extinction arising within the Milky Way along the line of sight to SN 2010jl is very small ($E_{B-V}=0.027$, Schlegel, Finkbeiner & Davis 1998). The inspection of a high resolution spectrum ($\lambda/\Delta\lambda \sim 46,000$) obtained with the SARG Echelle Spectrograph at the TNG on November 7, 2010 shows very weak Na I D lines at about zero velocity. The equivalent width (EW) of the D_2 component is less than 100 mÅ, implying $E_{B-V} < 0.02$ (Munari & Zwitter 1997). Therefore, the limit on interstellar polarization (ISP) within the Galaxy is $P_{MW} \leq 0.2\%$ (Serkowski, Matheson & Ford 1975; Whittet et al. 1992). The compilation of Heiles (1999) contains two stars within 3 degrees from the position of UGC 5189A. Both have $P=0.05\%$, in full agreement with the low Galactic extinction along the line of sight. Interstellar Na I D lines are detected in the high resolution spectrum at ~ 3120 km s⁻¹, possibly resulting from 2 or 3 separate components within 50 km s⁻¹. The total EW of the D_2 line is 0.20 ± 0.02 Å, implying $E_{B-V}=0.03$ for a MW dust/gas mixture (Munari & Zwitter 1997). Although there is a mounting evidence that the Serkowski law is not universal (Leonard & Filippenko 2001; Leonard et al. 2002; Maund et al. 2007; Patat et al. 2009, 2010), we reckon it is very plausible the ISP in the host galaxy is small, most likely below 0.3%.

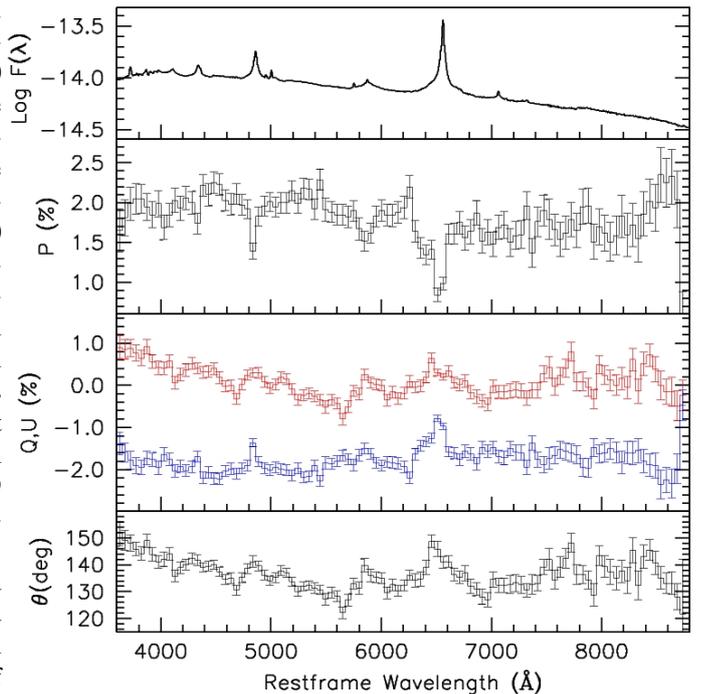


Fig. 1. Spectropolarimetric data for SN 2010jl. From top to bottom: unbinned total flux spectrum ($\log F(\lambda)$ in $\text{erg s}^{-1} \text{cm}^{-2} \text{\AA}^{-1}$), linear polarization degree, Stokes parameters Q (upper) and U (lower), and polarization position angle on the plane of the sky.

For this reason, in the following we consider the ISP negligible. This assumption is independently bolstered by the lack of reddening evidence in the flux spectrum (see Sect. 4), and the almost complete depolarization observed at the peak of H α (see Sect. 5.2).

4. Flux spectrum

The total flux spectrum of SN 2010jl shown in Fig. 1 (top panel) is characterised by emission lines of H, He and CNO elements superimposed on a blue continuum. This is nicely matched by a blackbody of ~ 7000 K after correcting for Galactic reddening. With this correction there is no additional need for host-galaxy reddening or multiple-temperature components (see the discussion in Smith et al. 2010). This is fully in line with the low reddening implied by the measured Na I D EWs (Sect. 3). As often found in SNe II_n, a de-blending of the H α emission reveals a broad component of $\sim 10,500$ km s⁻¹ FWHM, with a blue-shift of ~ 580 km s⁻¹, and an integrated flux of 6.8×10^{-13} erg cm⁻² s⁻¹. An intermediate (~ 2400 km s⁻¹) and an unresolved component (~ 640 km s⁻¹) are also present. We note that, based on higher resolution spectra, Smith et al. (2010) resolved the narrow component at 120 km s⁻¹ FWHM. A similar width is measured in our TNG-SARG spectrum. Apart from the Balmer series of hydrogen, narrow unresolved emission lines of [O II] $\lambda 3727$, [O III] $\lambda \lambda 4959, 5007$, [N II] $\lambda 5755$, He I $\lambda 5876$ and He I $\lambda 7065$ are also detected. For the He I lines, underlying broader components can be dis-

cerned. There is also weak evidence of He II $\lambda 4686$, He I $\lambda 7281$ and [O II] $\lambda\lambda 7320, 7330$. In general, SN 2010jl is spectroscopically very similar to other well studied Type II-n events, like SN 2007rt (Trundle et al. 2009 and references therein).

5. Spectropolarimetry

5.1. Continuum polarization

The continuum appears to be polarized at a very significant level across the whole spectral range (Fig. 1). With the remarkable exceptions of regions dominated by the main emission lines (see next section), the polarization degree is roughly constant. The mean level estimated on different portions of the continuum is: $2.00 \pm 0.04\%$ (3800–4600 Å), $2.02 \pm 0.05\%$ (5000–5600 Å), and $1.67 \pm 0.03\%$ (6900–8400 Å). On average, the polarization is slightly higher in the blue than in the red. The analysis of the $Q - U$ plane reveals that the data strongly cluster around one single point, and there is no statistically significant correlation between Q and U . The averages within the spectral range 4000–8000 Å are $0.02 \pm 0.03\%$ and $-1.77 \pm 0.03\%$ for Q and U respectively (the RMS deviation from the mean value is 0.3% for both parameters). For this reason, the position angle is almost constant as a function of wavelength (see also Fig. 1, bottom panel). The average values measured on the continuum in the same ranges as above are 140.9 ± 0.9 , 132.1 ± 1.0 , and 136.0 ± 0.9 degrees respectively (with RMS deviations of 3.5, 3.3, and 5.2 degrees). The polarized signal is almost completely carried by the U component.

5.2. Line polarization

Significant depolarizations are observed at the positions of the main emission features H α , H β , H γ , and He I 5876 Å (Fig. 1). The level of depolarization appears to be related to the line intensity, therefore strongly favoring continuum dilution as the responsible mechanism. Given the marked enhancement in the signal-to-noise across the H α emission, this line can be studied at full spectral resolution (640 km s $^{-1}$). The polarization profile is roughly symmetric, and P departs from the continuum level ($\sim 1.7\%$) at about ± 3000 km s $^{-1}$ (i.e. where the intermediate component starts to dominate over the broad one), to reach a minimum ($0.12 \pm 0.20\%$) around -250 km s $^{-1}$, which is slightly blue-shifted with respect to the total flux profile (Fig. 2). Under the hypothesis that the emission line is unpolarized, the polarization across the profile can be computed using the semi-empirical relation proposed by McLean & Clarke (1979). The result is plotted in Fig. 2 (mid panel, thin solid curve). Although the polarization level at the center of the line is well reproduced, the wings of the line display a residual polarization. A much better match is obtained considering only the intermediate and narrow components (mid panel, dashed curve). This suggests that the broad component is polarized to some extent (see also Sect. 6). All the variation takes place along the U component (Fig. 2, bottom panel). As Q is close to 0, the points in the $Q - U$ plane move along the U axis, so that the position angle remains constant across the line profile. This is what is

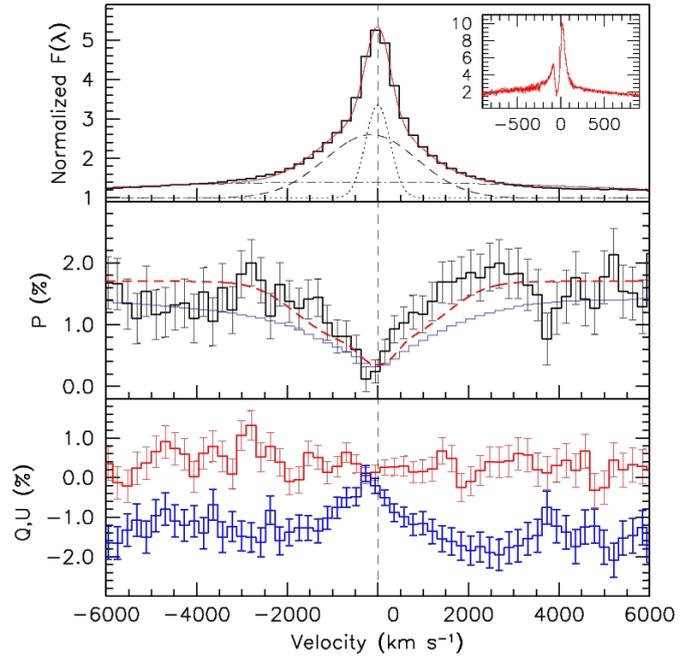


Fig. 2. Spectropolarimetry of the H α line (unbinned data). **Top panel:** continuum normalized flux profile. The thin curves trace the narrow (dotted), intermediate (dashed), and broad (dotted-dashed) components. The solid curve is the sum of the three. The inset shows the high-resolution SARG spectrum (see text). **Mid panel:** polarization degree. The thin solid line traces the depolarization expected for dilution by the full line profile, while the dashed line is the same for the intermediate and narrow components only. **Bottom panel:** Stokes parameters Q (upper) and U (lower).

expected for simple continuum depolarization dilution (see for instance Hoffman et al. 2008).

The presence of dust above the line forming region (e.g. in the circum- or inter-stellar environment) would polarize the light irrespective of its origin. i.e. affecting both the adjacent continuum and the line. Admittedly, depending on its position angle, the ISP could also act as a depolarizer. However, in order to have a null net polarization, the component of the IS Stokes vector perpendicular to the intrinsic polarization needs to have a fine tuned value, which makes this possibility rather unlikely. We note that if this were the case, then the intrinsic polarization would be higher than what is observed. The almost complete depolarization seen at the center of H α is strongly suggestive of a negligible ISP, thus confirming the expectations based on reddening considerations (Sect. 3).

6. Discussion

As Type II-n events are only 6-9% of all core-collapse SNe (Li et al. 2010), and the technique requires very high signal-to-noise data, spectropolarimetry of this class of objects is rare (Wang & Wheeler 2008). The only two published cases are those of SNe 1997eg (Hoffman et al. 2008: 16, 44 and 93 days after discovery), and 1998S (Leonard et al. 2000: 5

days after discovery; Wang et al. 2001: 28 and 59 days after discovery), which reached a comparable absolute magnitude ($M_V \sim -18.6$), which is significantly fainter than the peak magnitude reached by SN 2010jl ($M_V \sim -20.6$, Smith et al. 2010). Although the analysis was made difficult by the uncertainty in the ISP correction, in the earliest epochs the two objects displayed a significant and almost constant continuum polarization (2.0-2.6%), characterized by a constant position angle. Additionally, marked line depolarizations were observed at the positions of the most prominent emission lines (Leonard et al. 2000; Wang et al. 2001; Hoffman et al. 2008). Our data show that SN 2010jl conforms strikingly well to this behavior, hinting to a common geometry of the explosion or, more in general, of the environment in which the explosion takes place.

The basic physical facts that emerge from the early time data are: *i*) the wavelength-independent continuum polarization; *ii*) the constant position angle, and *iii*) the strong line depolarization. These place SNe IIn in a very special niche within the SN polarimetric zoo (see Wang & Wheeler 2008 for a general review). The most natural source of wavelength-independent linear polarization in a stellar envelope is Thomson scattering by free electrons. In this context, a non-null net polarization is indicative of an asymmetry in the continuum-emitting region. In the case of an axis-symmetric geometry, this would produce a wavelength-independent polarization angle (perpendicular to the major axis). Under the assumption of a spheroidal electron-scattering atmosphere, the level of polarization measured in SN 2010jl ($\sim 2\%$) indicates a substantial asphericity, of axial ratio ≤ 0.7 (Höflich 1991). This value is very similar to those derived for the other two known cases (Leonard et al. 2000; Hoffman et al. 2008). As first proposed by Leonard et al. (2000), the simplest explanation for the line depolarization in Type IIn SNe is the dilution of polarized continuum light by unpolarized line recombination emission, which has to take place well above the photosphere, in a region where electron scattering becomes negligible. This picture was made more complicated by the detection of a different behavior in the narrow component, suggesting a different geometry for the regions where these lines form (Leonard et al. 2000; Hoffman et al. 2008). As the narrow component is not resolved in our spectra, we cannot address this issue. However, we note that while the core of the $H\alpha$ profile (corresponding to the narrow and intermediate components) is depolarized, this is not the case for the wings at velocities larger than 3000 km s^{-1} (see Sect. 5.2).

This might be understood in the scenario proposed by Chugai (2001), where the narrow component is essentially unpolarized, while the broad wings are the result of electron scattering. In this picture, the broad component does not arise in the reverse-shocked ejecta (as commonly assumed), but rather in a dense, fully ionized CSM shell dominated by electron scattering, where the photosphere would also reside. A closer inspection of the binned $H\alpha$ profile indeed shows a polarization bump (mostly in the Q parameter), spanning the full breadth of the broad component (Fig. 1), and peaking at $\sim 0.4 \pm 0.1\%$ above the adjacent continuum polarization level. The change of θ seen across the bump suggests a different geometry for the photosphere and the broad line forming region. Also, there is

an excess of polarization on the blue wing, similar to what was seen in SN 2007eg. This possibly indicates that the receding side of the scattering region is obscured either by its approaching side or by the ejecta (Hoffman et al. 2008).

In a scenario where continuum polarization is acquired within the CSM, spectropolarimetry can hardly provide information on the asymmetry of the explosion itself, but only on the geometry of the circum-stellar environment. As a consequence, the high level of continuum polarization observed in Type IIn SNe would be interpreted as a strong deviation from sphericity in the CSM and, in turn, in the outflow of material during the mass loss episodes preceding the explosion. This conundrum can be resolved only via multi-epoch, higher spectral resolution spectropolarimetry, which with 8m-class telescopes should be possible, at least for the most luminous events. This will allow one to distinguish the different behavior of the various components, and to study their evolution with time, as the ejecta proceed into the CSM.

One important fact implied by our observations is that the low polarization level measured at the peak of $H\alpha$ ($\leq 0.2\%$) is inconsistent with the presence of substantial amounts of dust along the line of sight, including the immediate surroundings of the explosion site. Therefore, although a significant quantity of CSM must be present at the epoch of our observations (as revealed by the strong signs of interaction), this material has to have a very low dust content. Whether this is intrinsic or caused by dust evaporation by the UV radiation cannot be concluded based on the available data.

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