

ON THE PROPERTIES OF INNER COOL DISKS IN THE HARD STATE OF BLACK HOLE X-RAY TRANSIENT SYSTEMS

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ABSTRACT

The formation of a cool disk in the innermost regions of black hole X-ray transient systems in the low hard state is investigated. Taking into account the combined cooling associated with the Compton and conductive energy transport processes in a corona, the radial structure of a disk is described for a range of mass accretion rates. The mass flow in an optically thick inner region can be

maintained by the condensation of matter from a corona with the disk temperature and luminosity varying continuously as a function of the accretion rate. Although such a disk component can be present, the contribution of the optically thick disk component to the total luminosity can be small since the mass flow due to condensation in the optically thick disk underlying the corona can be significantly less than the mass flow rate in the corona. The model is applied to the observations of the low quiescent state of the black hole source GX 339-4 at luminosities of $\sim 0.01L_{Edd}$ and is able to explain the temperature of the thermal component at the observed luminosities. Since conductive cooling dominates Compton cooling at low mass accretion rates, the luminosity corresponding to the critical mass accretion rate above which a weak thermal disk component can be present in the low hard state is estimated to be as low as $0.001L_{Edd}$.

Subject headings: accretion, accretion disks — black hole physics — X-rays: stars — X-rays: binaries — stars:individual (GX 339-4)

1. Introduction

The class of objects known as transient X-ray binary systems containing black holes is one of the few examples of astrophysical laboratories available for investigating the nature of accretion flows and the structure of accretion disks surrounding compact objects. Their luminosities provide us with an opportunity to investigate a given system over a wide range of mass accretion rates onto the central black hole. Indeed, these systems have been observed in distinct states generally categorized as low/hard and high/soft states (see McClintock & Remillard 2006). In this categorization, the high/soft state is characterized by a strong thermal component corresponding to a color temperature of order 1 keV. These states are observed during the outburst phase in which the luminosity can approach the Eddington luminosity, $L_{Edd} \sim 10^{39}$ ergs s^{-1} for a $10 M_{\odot}$ black hole. On the other hand, the low/hard state is typically described by a power law with a photon power law index ~ 1.7 . This state is observed during the quiescent stages, where the luminosities can be as low as $10^{-7}L_{Edd}$ (Corbel et al. 2006). Hence, it is generally accepted that these states are distinguished by the spectrum and a range of luminosity with a transition at $\sim 1 - 4\%$ of the Eddington luminosity (Maccarone 2003) from the low/hard state to the high/soft state.

Given these observational diagnostic clues, progress in our theoretical understanding has advanced in parallel. Specifically, the spectrally soft state is believed to represent the emission from a cool optically thick and geometrically thin accretion disk extending to the innermost stable circular orbit ($3R_S$ for a non-rotating black hole, where $R_S = 2GM/c^2$ is

the Schwarzschild radius) as described in the seminal study by Shakura & Sunyaev (1973, hereafter SS). In contrast, the spectrally hard state is thought to arise from a geometrically thick, hot optically thin coronal or advectively dominated accretion flow (ADAF; Narayan & Yi 1994; 1995a, b) located in the innermost regions of the disk near the black hole. Here, the SS accretion disk is truncated at a distance of several hundred Schwarzschild radii from the black hole, leading to a transition from a disk dominated by a geometrically thin outer structure to a geometrically thick inner structure Esin et al. (1997, 1998, 2001). Although a complete description of the truncation of the optically thick disk has yet to be worked out in detail, the mechanisms suggested to facilitate this transition have been investigated in the studies by Honma (1996) and Manmoto & Kato (2000) based on a radial conductive energy transport process and by Meyer et al. (2000), Spruit & Deufel (2002), and Dullemond & Spruit (2005) based on a vertical evaporative process.

Recently, additional observational clues concerning the nature of the accretion flow and geometry in the low quiescent state have been found pointing to the possible existence of an inner cool optically thick disk structure in black hole transient X-ray binary systems. Although the observational signatures are difficult to extract because of the low X-ray count rates, the evidence for such a structure has been inferred from the source GX 339-4 (see Miller et al. 2006a, b; Tomsick et al. 2008) where a weak, soft, thermal emission component ($\lesssim 0.3$ keV) has been detected and blurred reflection and iron features are indicated in fits to the X-ray spectrum for luminosities as low as $\sim 0.01L_{\text{Edd}}$. One of the goals of current observational studies has been to extend such studies to other sources (e.g., J 1817-330, see Rykoff et al. 2007; and J1753.5-0127, see Miller et al. 2006a; Ramadevi & Seetha 2007; Soleri et al. 2008) and to determine the luminosity range for which such a thermal component exists so that a full description of the accretion flow geometry in the various states of transient X-ray binary systems can be constructed.

Given the existence of this weak thermal component, the accretion disk structure in the quiescent state would be described by three distinct regions. Specifically, a small inner cool, optically thick disk, a geometrically thick hot coronal region at intermediate radial scales, and a very extensive geometrically thin disk nearly filling the Roche lobe of the black hole would contribute to the weak thermal X-ray component, the power law X-ray component, and the infrared and optical emission respectively. Such a configuration was already contemplated in early work by Różańska & Czerny (2000) and more recently by Liu et al. (2006), Meyer et al. (2007) and Mayer & Pringle (2007). To maintain such a remnant cool disk in the innermost regions (in the presence of mass transport by turbulent viscous stresses) gas must condense from an overlying corona.

In this study we focus on the general description of the structure and properties of the

disk in its innermost regions during the quiescent state of a black hole X-ray transient binary system, taking into account the effects of both thermal bremsstrahlung and comptonization. This study differs from our recent study (Liu et al. 2007) where only one such process was considered dominant for any given structure. In §2, we describe the theoretical model and the modifications required in considering more than one cooling process in any given region of the disk. The solutions to the modified set of equations are presented for a range of mass accretion rates and viscosity parameters in §3. The numerical results are compared to the observations of GX 339-4 in §4 and their implications are discussed in the last section.

2. Theoretical Model

Early investigations (Meyer & Meyer-Hofmeister 1994, Meyer, Liu, & Meyer-Hofmeister 2000; Liu et al. 2002) have shown that a hot corona overlying a cool disk leads to the evaporation of cool gas into a corona, through which mass is continuously supplied for accretion onto a central compact object. The evaporation is driven by vertical thermal conduction, and the accretion of evaporated mass through the corona is a result of viscous transport. The properties of the evaporation fed corona are a function of the distance from the accreting compact object. In the case of a black hole transient X-ray binary system in the quiescent state, the mass flow rate in the corona reaches a maximum at a distance of $\sim 300R_S$ by efficient evaporation and decreases inward as a result of condensation of coronal gas to a cooler disk.

When the mass supply rate is sufficiently low, the optically thick outer cool disk is truncated in the region where the integral of the evaporation rate over the radial extent is greater than the mass inflow rate, leading to the formation of a gap between outer and inner disk, filled by a coronal flow/ADAF. The radial extent of the coronal gap region is dependent on the mass supply rate to the outer disk with a larger coronal gap region corresponding to a lower mass accretion rate. For black hole X-ray binary systems in the quiescent state, the accretion rate can be very low, and the outer disk is sufficiently distant from the central black hole that it can be neglected in describing the innermost disk regions. In this latter region, the dominant accretion flow lies within the corona/ADAF, accompanied by a weak mass inflow within an underlying cool disk. The condensation of coronal gas to the disk follows from energy (and pressure) balance between this inner cool disk and its overlying corona, providing the only source of mass in the inner optically thick disk. As shown by Liu et al. (2007) such a configuration can be maintained for a range of accretion rates and, hence, luminosities.

In this work, we generalize the description of the inner disk and corona configuration by

considering Compton scattering and conduction, contributing simultaneously, to the cooling of the disk. The treatment of both cooling processes leads to a modification of the temperature of the electrons in the corona, and, hence, the condensation rate. In addition, we also include contributions from both inverse Compton emission from the corona and bremsstrahlung emission from the radiating/transition layer to the total luminosities. This treatment of the cooling processes is in contrast to our previous study (Liu et al. 2007) where always only one of the dominant emission mechanism was taken into account (but similar in our neglect of non-thermal cooling processes). As a consequence of the new treatment, the physical characteristics of the inner regions (e.g., the effective temperature of the inner cool disk) vary continuously with variations in the luminosity. To facilitate comparisons with Liu et al. (2007), the effect of X-ray irradiation of the hot coronal gas on the cool disk is neglected; however, an approximate estimate of its effect is given and its implications discussed in §5.

2.1. Condensation caused by either conductive cooling or Compton cooling

The energy balance in the transition layer between the disk and corona depends on the ADAF lying above the disk (Narayan & Yi 1995b). As derived in our earlier work (Meyer et al. 2007; Liu et al. 2007) the condensation rate (mass flow rate in vertical direction) per unit area is determined as

$$\dot{m}_z = \frac{\gamma - 1}{\gamma} \beta \frac{-F_c^{\text{ADAF}}}{\frac{\mathfrak{R}T_i}{\mu_i}} \left(\sqrt{C} - 1 \right), \quad (1)$$

where

$$C \equiv \kappa_0 b \left(\frac{0.25 \beta^2 p_0^2}{k^2} \right) \left(\frac{T_{\text{cpl}}}{F_c^{\text{ADAF}}} \right)^2$$

with F_c^{ADAF} the conductive flux from the upper ADAF arriving at the interface of the transition layer, T_{cpl} the coupling temperature and $p_0 = (2/\sqrt{\pi})p$ the pressure in the transition layer,

$$\begin{aligned} F_c^{\text{ADAF}} &= -(\kappa_0 K n_i n_e T_i)^{1/2} T_{\text{em}}, \\ T_{\text{cpl}} &= 1.98 \times 10^9 \alpha^{-4/3} \dot{m}^{2/3} \text{K}. \end{aligned} \quad (2)$$

Pressure, ion temperature and electron number density, p , T_i and n_e are quantities of the ADAF,

$$T_i + 1.077 T_e = 1.98 \times 10^{12} r^{-1} \text{K}, \quad (3)$$

$$p = 1.87 \times 10^{16} \alpha^{-1} m^{-1} \dot{m} r^{-5/2} \text{ gcm}^{-1} \text{ s}^{-2},$$

$$n_e = 5.91 \times 10^{19} \alpha^{-1} m^{-1} \dot{m} r^{-3/2} \text{ cm}^{-3}.$$

The meaning of the further quantities in the equations is: β is the ratio of gas pressure to total pressure, γ the ratio of specific heats, $\gamma = (8 - 3\beta)/(6 - 3\beta)$, μ_i the molecular weight of ions (1.23 for the assumed standard chemical abundance) and k is the Boltzmann constant. $\kappa_0 = 10^{-6} \text{ erg s}^{-1} \text{ cm}^{-1} \text{ K}^{-7/2}$ is the thermal conductivity coefficient in the relation for the conductive flux, $F_c = -\kappa_0 T_e^{5/2} dT_e/dz$ (Spitzer 1962). The coefficient $b = 10^{-26.56} \text{ g cm}^5 \text{ s}^{-3} \text{ K}^{-1/2}$ is used for the description of the free-free radiation (Sutherland & Dopita 1993). The energy transfer rate from ions to electrons is taken as $q_{ie} = K n_i n_e T_i T_e^{-3/2}$ with $K = 1.64 \times 10^{-17} \text{ g cm}^5 \text{ s}^{-3} \text{ deg}^{1/2}$ (Stepney 1983) and $n_i = n_e/1.077$ is the ion number density. Following convention, we take α as the viscosity coefficient. The quantities m , \dot{m} , and r are expressed in units of a solar mass, the Eddington rate, $\dot{M}_{\text{Edd}} = L_{\text{Edd}}/0.1c^2$, and the Schwarzschild radius, R_S , respectively. T_{em} is the maximal electron temperature in the corona, to be determined for either conduction or Compton cooling, as derived in the earlier work (Liu et al. 2007).

Once T_{em} is known the condensation rate can be determined by Eqs.(1)–(3) at any radius for a given black hole mass, mass accretion rate and viscosity.

In the case of inefficient Compton cooling in the corona, vertical conduction is the dominant cooling mechanism. Here, by collisional heat input from the ions ($dF_c/dz = q_{ie}$), a vertical temperature gradient forms which results in a maximal temperature at the upper boundary of the ADAF and a conductive flux at the lower boundary,

$$T_{\text{em}} = 2.01 \times 10^{10} \alpha^{-2/5} \dot{m}^{2/5} r^{-2/5} \text{ K}, \quad (4)$$

$$F_c^{\text{ADAF}} = -6.52 \times 10^{24} \alpha^{-7/5} m^{-1} \dot{m}^{7/5} r^{-12/5}.$$

On the other hand, for a sufficiently high radiation flux from the underlying disk, the disk radiation propagating through the corona can lead to efficient Compton cooling of the corona. The electrons are cooled by Compton scattering to a temperature (determined by $q_{ie} = q_{\text{cmp}}$) lower than the value given by conductive cooling (Eq.(4)). As a result, the heat flux at the transition layer decreases and the condensation rate increases, as determined by the energy balance in the transition layer between cooling by bremsstrahlung emission and heating by the conductive flux from the upper ADAF and by the enthalpy flux of the condensing gas. The maximal temperature in the corona and the corresponding conductive flux to the transition/radiating layer are expressed as

$$T_{\text{em}} = 3.025 \times 10^9 \alpha^{-2/5} m^{-2/5} \dot{m}^{2/5} r^{1/5} \left[1 - \left(\frac{3}{r} \right)^{1/2} \right]^{-2/5} \left(\frac{T_{\text{eff,max}}}{0.3 \text{ keV}} \right)^{-8/5} \text{ K}, \quad (5)$$

$$F_c^{\text{ADAF}} = -9.816 \times 10^{23} \alpha^{-7/5} m^{-7/5} \dot{m}^{7/5} r^{-9/5} \left[1 - \left(\frac{3}{r} \right)^{1/2} \right]^{-2/5} \left(\frac{T_{\text{eff,max}}}{0.3 \text{ keV}} \right)^{-8/5} \text{ erg s}^{-1} \text{ cm}^{-2}.$$

The effective temperature of the disk, T_{eff} , corresponds to the mass flow in the disk accumulated by the condensation of matter from the corona. T_{eff} reaches its maximum value $T_{\text{eff,max}} = 1.335 \times 10^7 m^{-1/4} \dot{m}_{\text{cnd}}^{1/4}$ at a distance $r_{\text{tmax}} = 3 \frac{49}{36} R_s$ (for details see Liu et al. 2007).

2.2. Condensation caused by the two processes: conductive cooling and Compton cooling operating simultaneously

The above describes the formulation of our model as outlined in a previous study (Liu et al. 2007) when one of the cooling processes (either thermal conduction or Compton scattering) is dominant. However, in general, both processes can contribute to the cooling simultaneously. In particular, for a certain range of mass accretion rates, the rate of cooling associated with the conduction and Compton processes are comparable. In this case, neither the Compton dominated model nor the conduction dominated model provides an adequate description for the disk/corona configuration. In this case, both mechanisms must be treated together. However, construction of the detailed vertical structure for the condensation region cannot be expressed in analytical form. In the spirit of the analysis of Liu et al. (2007), we modify the formulae describing the conduction dominant and Compton dominant cases in such a way that we obtain a smooth transition when both mechanisms are of comparable importance. Our approach is described below.

For given values of α , m , and \dot{m} at a given distance, r , the importance of Compton or conduction cooling is determined by comparing the electron temperature given by considering the operation of Compton cooling alone (i.e. $q_{\text{ie}} = q_{\text{cmp}}$) and conductive cooling alone (i.e. $q_{\text{ie}} = dF_c/dz$). The process that results in a lower temperature is the more efficient cooling mechanism and, thus, is taken as the dominant one. To determine the strength of the neglected cooling process compared to that of the dominant process, we define the ratio, λ , between the Compton cooling rate, q_{cmp} , and the conduction cooling rate, dF_c/dz ,

$$\lambda \equiv \frac{q_{\text{cmp}}}{dF_c/dz}. \quad (6)$$

The Compton cooling rate is given by

$$q_{\text{cmp}} = \frac{4kT_e}{m_e c^2} n_e \sigma_T c \frac{a T_{\text{eff}}^4(r)}{2} = f_1(\alpha, m, \dot{m}) \dot{m}_{\text{cnd}} T_e r^{-9/2} \left[1 - \left(\frac{3}{r} \right)^{1/2} \right], \quad (7)$$

where a is radiation density constant, $\sigma = \frac{ac}{4}$, σ_T Thomson cross section, $\sigma T_{\text{eff}}^4(r) = (3GM\dot{M}_{\text{cnd}}/8\pi R^3) [1 - (3R_S/R)^{1/2}]$, and f_1 is some function only depending on the ADAF parameters.

dF_c/dz can be approximated by taking the value of the flux at the upper boundary ($F_c = 0$) and at the lower boundary, $F_c^{\text{ADAF}} = -(k_0 K n_i n_e T_i)^{1/2} T_{\text{em}}$, to give

$$dF_c/dz \propto -F_c^{\text{ADAF}}/H \propto -F_c^{\text{ADAF}} r^{-1}.$$

Replacing F_c^{ADAF} by Eqs.(2) and (3) we have

$$dF_c/dz = f_2(\alpha, m, \dot{m}) r^{-3} T_{\text{em}}. \quad (8)$$

Thus, from Eqs.(6)–(8) and $T_e = T_{\text{em}}$, λ can be approximately expressed as

$$\lambda = f(\alpha, m, \dot{m}) \dot{m}_{\text{cnd}} r^{-3/2} \left[1 - \left(\frac{3}{r} \right)^{1/2} \right]. \quad (9)$$

The function f again only depends on the ADAF parameters. The value of λ is very small when conduction is the overwhelmingly dominant cooling process, reaching unity when the Compton process is equally important, and is much larger than unity for cooling dominated by Compton scattering. In the critical case when the cooling process switches from conduction dominant to Compton dominant, $\lambda_c = 1$, the temperature inferred from the Compton cooling process and the conduction cooling process is the same, which gives (by Eq.(4)=Eq.(5)),

$$\dot{m}_{\text{cnd}} r_c^{-3/2} \left[1 - \left(\frac{3}{r_c} \right)^{1/2} \right] = \frac{1}{2.464 \times 10^4}. \quad (10)$$

This determines $f(\alpha, m, \dot{m}) = 2.464 \times 10^4$ by the conditions defining the critical case from Eq.(9). Thus, λ can be written, in general, as

$$\lambda = 2.464 \times 10^4 \dot{m}_{\text{cnd}} r^{-3/2} \left[1 - \left(\frac{3}{r} \right)^{1/2} \right] \quad (11)$$

and used in calculating the condensation rate. To obtain a solution that includes both effects, thermal conduction and Compton cooling, we modify the energy equation, leading to a change in the electron temperature and, hence, a change in the condensation rate and luminosities from both the corona and the disk. In the case of conduction dominant cooling, the energy equation is modified as

$$q_{\text{ie}} = (1 + \lambda) dF_c/dz. \quad (12)$$

Since $q_{\text{ie}} \propto T_e^{-3/2}$ and $dF_c/dz \approx -F_c^{\text{ADAF}} r^{-1} \propto T_{\text{em}}$, the energy equation leads to a modified temperature

$$T_{\text{em}} = 2.01 \times 10^{10} \alpha^{-2/5} \dot{m}^{2/5} r^{-2/5} (1 + \lambda)^{-2/5} \text{K}. \quad (13)$$

Similarly, in the case of Compton dominant cooling, the energy equation should be modified as

$$q_{ie} = \left(1 + \frac{1}{\lambda}\right) q_{\text{cmp}}. \quad (14)$$

Note that $q_{\text{cmp}} \propto T_e$, and the modified energy equation results in a modification factor $\left(1 + \frac{1}{\lambda}\right)^{-2/5}$ to the temperature,

$$T_{\text{em}} = 3.025 \times 10^9 \alpha^{-2/5} m^{-2/5} \dot{m}^{2/5} r^{1/5} \left[1 - \left(\frac{3}{r}\right)^{1/2}\right]^{-2/5} \left(1 + \frac{1}{\lambda}\right)^{-2/5} \left(\frac{T_{\text{eff,max}}}{0.3 \text{keV}}\right)^{-8/5} \text{K}. \quad (15)$$

With the modified temperature, the heat flux and the resulting integrated condensation rate in a region between R_i and R_o can be calculated for given α , m and \dot{m} ,

$$\dot{m}_{\text{cnd}} = \int_{R_i}^{R_o} \frac{4\pi R}{\dot{M}_{\text{Edd}}} \dot{m}_z dR. \quad (16)$$

The condensation rate for Compton dominant cooling is

$$\dot{m}_{\text{cnd}} = A \left\{ 2B \left[\left(\frac{r_o}{r_i}\right)^{1/2} - 1 \right] - \int_{r_i/3}^{r_o/3} \left(1 + \frac{1}{\lambda}\right)^{-2/5} x^{1/5} (1 - x^{-1/2})^{-2/5} dx \right\}, \quad (17)$$

and for conduction dominant cooling

$$\dot{m}_{\text{cnd}} = 3.23 \times 10^{-3} \alpha^{-7} \dot{m}^3 f(x) \quad (18)$$

with

$$f(x) = \left(\frac{r_o}{r_1}\right)^{3/5} \left[6 \left(\frac{r_1}{r_o}\right)^{1/10} - 6 \left(\frac{r_1}{r_o}\right)^{1/10} x^{1/2} - 3 \int_{r_i/r_o}^1 (1 + \lambda)^{-2/5} x^{-2/5} dx \right], \quad (19)$$

where $r_1 = 0.815 \alpha^{-28/3} \dot{m}^{8/3}$ is the condensation radius determined by $C = 1$ in the case of conduction cooling before modification. The new factor $\frac{r_o}{r_1}$ is a consequence of generalizing the condensation formula to reflect the condensation rate at any condensation region. When $\lambda = 0$ and $r_o = r_1$, Eq.(19) reverts to our old form. When $\lambda \rightarrow \infty$, Eq.(17) also reverts to the previous form for the case of Compton cooling.

Note that since λ depends on the condensation rate, iteration is required until a self-consistent condensation rate is obtained. In practice, the modification of the condensation rate is very small at low accretion rate since the conduction process dominates. Similarly, the modification is also small if Compton cooling is very strong compared to conduction. We note, however, that these conditions are not necessarily satisfied for black hole X-ray binary systems in their low hard or intermediate states. In this regime, the modification factors are important since conduction and Compton scattering processes can be of comparable strength in cooling the corona.

2.3. The radial extent of Compton and conduction dominated regions

In describing the thermal state of the corona above an inner disk, (i) cooling in the corona could be conduction dominant throughout the corona, (ii) Compton dominant throughout, or (iii) conduction dominant in the outer regions and Compton dominant in the inner region, depending on the binary system parameters. Therefore, for given values of α , m and \dot{m} , an estimate of the radial distance from the black hole for the Compton dominant region is obtained by assuming that the Compton process cools the electrons to the same temperature as the conduction process,

$$r_{\text{cmp}} \left[1 - \left(\frac{3}{r_{\text{cmp}}} \right)^{1/2} \right]^{-2/3} = 846.71 \dot{m}_{\text{cnd}}^{2/3}, \quad (20)$$

The above equation either admits two solutions or no solution for r_{cmp} . In case no solution exists, Compton cooling never dominates in the corona and the modified conduction model is an adequate description for the inner disk and corona. The existence of the two solutions r_{cmp1} and r_{cmp2} determines the Compton dominant region $r_{\text{cmp1}} < r < r_{\text{cmp2}}$. If r_{cmp2} is larger than the condensation radius calculated by the conduction model, which is modified as $r_d = 0.815\alpha^{-28/3}\dot{m}^{8/3} [1 + \lambda(r_d)]^4$, the cooling is dominated by the Compton process at all radii, and we use the modified Compton model for the numerical calculations. On the other hand, if r_{cmp2} is smaller than r_d , the coronal region is divided into a conduction and Compton dominant region. In Figure 1 the radial extent of Compton and conduction dominant regions is presented. The left panel delineates the regions dominated by Compton or conduction cooling as a function of the total condensation rate. At the lowest condensation rates, conduction is dominant throughout the disk. On the other hand, at higher condensation rates, the Compton dominant region always lies interior to the conduction dominant region. The right panel reveals the relative strength of the Compton cooling rate as compared to conduction as a function of distance at given condensation rates. The modified condensation rate is calculated employing the relevant models and values of λ in each region. Their contributions are summed to determine the modified total condensation rate of the system.

The luminosity from bremsstrahlung radiation in the radiating/transition layer does not vary with the electron temperature in the corona since the density and coupling temperature are determined by α , m , and \dot{m} , independent of the coronal temperature. Thus, it has the same form for both the Compton and the conduction dominant regions and is unaffected by the modification when both cooling mechanisms are included. We integrate the bremsstrahlung emission throughout the disk-corona region yielding,

$$\frac{L_{\text{Brem}}}{L_{\text{Edd}}} = 0.0642\alpha^{-7/3}\dot{m}^{5/3} \left[1 - \left(\frac{3}{r_d} \right)^{1/2} \right]. \quad (21)$$

Since the Compton luminosity from the corona depends on the electron temperature, its determination must include the contributions from both cooling processes. In the Compton cooling dominant region, a correction factor $(1 + \frac{1}{\lambda})^{-2/5}$ is added to the integrand,

$$\frac{L_{\text{cmp}}}{L_{\text{Edd}}} = 0.392\alpha^{-7/5}m^{3/5}\dot{m}^{7/5} \left(\frac{T_{\text{eff,max}}}{0.3\text{keV}}\right)^{12/5} \int_{r_i/3}^{r_o/3} \left(1 + \frac{1}{\lambda}\right)^{-2/5} x^{-23/10} (1 - x^{-1/2})^{3/5} dx. \quad (22)$$

In the conduction dominant region, the luminosity from Compton scattering was neglected in our previous work. To take this contribution into account, the electron temperature is replaced by Eq.(4) in the Compton cooling rate, leading to a Compton luminosity from the conduction dominant region from r_i to r_o ,

$$\frac{L_{\text{cmp}}}{L_{\text{Edd}}} = 1.349\alpha^{-7/5}m\dot{m}^{7/5} \left(\frac{T_{\text{eff,max}}}{0.3\text{keV}}\right)^4 \int_{r_i/3}^{r_o/3} (1 + \lambda)^{-2/5} x^{-29/10} (1 - x^{-1/2}) dx. \quad (23)$$

The total corona luminosity is, thus, the sum of bremsstrahlung radiation (Eq.(21)) in the radiation/transition layers and Compton scattering in the ADAF above (Eqs. (22) and (23)) when both regions exist.

The disk luminosity is determined under the assumption of emission by a local blackbody and calculated as an integral over the disk as,

$$\frac{L_d}{L_{\text{Edd}}} = \int_{3R_s}^{R_d} \sigma T_{\text{eff}}^4(R) \frac{4\pi R}{L_{\text{Edd}}} dR = \int_{3R_s}^{R_d} \frac{3GM\dot{M} [1 - (\frac{3R_s}{R})^{1/2}]}{8\pi R^3} \frac{4\pi R}{L_{\text{Edd}}} dR. \quad (24)$$

Note that the mass flowing in the inner disk through any radius R originates from the condensed gas integrated from the critical condensation radius (R_d) to R . Hence, the mass accretion rate in the disk is a function of radius, which increases with decreasing radii, representing the cumulative effect of the condensation rate of gas onto the disk. The disk luminosity is then calculated from all radii with the corresponding accretion rate profile.

In our procedure, the condensation rate is first calculated assuming conduction as the only cooling mechanism. If this condensation rate heats the disk to a temperature sufficiently high that Compton cooling is more efficient than conduction in some regions, we adopt the Compton cooling model in the Compton dominant region but retain the conduction model in the conduction region. The condensation rate is recalculated from these different regions as iterative calculations are required for the Compton dominant region. These results are modified taking into account the neglected cooling by calculating the value of λ from the inferred condensation rate. As the modified condensation rate results in a new value of λ , the process is repeated until the presumed condensation rate approaches the derived condensation rate. Finally, the contributions to the total luminosity from Compton scattering and

bremsstrahlung throughout the coronal region and multi-colour blackbody radiation from the inner disk are calculated with the derived condensation rate and correction factor λ by equations (21), (22), (23) and (24).

3. Numerical Results

To investigate the properties of the inner optically thick disk in X-ray transient binary systems, we adopt a black hole mass of $10M_{\odot}$ and the low mass accretion rate range corresponding to the transition to an ADAF. In the following, we describe the numerical results for the inner disk temperature, size, and luminosity from the inner disk as determined from the condensation rates from the corona as a function of the total mass accretion rate. To determine their sensitivity to the viscosity in the disk, we also examine their dependence on the viscosity parameter.

The luminosity from the disk and corona is illustrated as a function of accretion rate in Figure 2. It can be seen that the Eddington scaled luminosity does not scale linearly with the mass accretion rate, indicating that the conversion of gravitational energy to radiative energy by an ADAF+disk in low states is not constant nor as efficient ($\eta = 0.10$) as in a SS accretion disk. The deviation of the $L/L_{\text{Edd}} - \dot{M}/\dot{M}_{\text{Edd}}$ curve from a linear relation indicates that the energy conversion efficiency decreases with lower mass accretion rates, qualitatively consistent with an ADAF model.

The size of the inner disk is taken to be limited by the critical radius where neither evaporation nor condensation occurs. Externally, gas from the disk evaporates into the corona, while interior to this radius the coronal gas condenses onto the disk as a consequence of the disk corona interaction. For greater mass flow rates in the corona, more gas condenses onto the disk and the critical radius of the inner disk increases. To illustrate this relationship, the size is calculated for a range of mass accretion rates. The radial extent of the inner disk is modified by the inclusion of the two cooling processes since the additional cooling results in a lower heat flux down to the transition layer. However, the over-cooling by bremsstrahlung radiation is partially compensated by the heating associated with the enthalpy flux from the condensation. As a consequence, the size of the inner disk is larger than the case for which one cooling process was considered. In the left panel of Figure 3, the inner disk size is shown as a function of the luminosity. It can be seen that the inner disk extends to greater distances with increasing luminosities or mass accretion rates. In the right panel of Figure 3, the critical radii of the optically thick disk components are illustrated as a function of the mass accretion rate. For a given accretion rate, the smaller radius corresponds to the outer edge of the inner disk and the larger radius corresponds to the inner edge of the outer

disk, with their difference representing the spatial extent of the coronal gap region. Hence, the radial extent of the coronal gap region between the inner and outer optically thick disk shrinks as the luminosity increases, in agreement with the presence of a radially continuous cold disk for sufficiently high luminosities. In this case, condensation is no longer the only source for disk accretion, and the contribution of the disk to the total luminosity is much greater than for a disk fed by condensation alone.

The size of the inner disk is a function of the viscosity parameter as also seen in Figure 3 where, at a given luminosity, the disk is smaller for larger viscosity parameters. This trend results from the (vertical) pressure and energy balance between the disk and overlying corona. Specifically, for a given accretion rate, the surface density in the corona decreases as α is increased. With a lower density the heat flux to the transition layer decreases, but the bremsstrahlung radiation rate is reduced even more. The energy balance between cooling and heating results in a reduced heating rate associated with a lower enthalpy flux. The net effect results in a decreased condensation rate and, hence, a smaller disk size. We note that larger viscosity parameters, yielding smaller disk radii, would be preferable to be consistent with our neglect of X-ray irradiation effects in our present level of approximation.

The total condensation rate, which is an integral of the rate of locally condensing mass from the critical condensation radius to $r_{\text{tmax}} = 3\frac{49}{36}$, relative to the mass accretion rate is displayed as a function of the mass accretion rate in Figure 4. It is evident that a large fraction (up to 20%) of the coronal flow condenses onto an inner disk at high accretion rates. As a consequence, the rate of mass flow in the corona decreases with decreasing distance. On the other hand, all the gas accretes through the corona without condensation if the mass flow rate in the corona is very low. The relative mass flow rate in the inner disk depends sensitively on the viscosity value.

The maximal disk temperature as calculated from the rate of condensation integrated over the entire inner disk is illustrated as a function of luminosity in Figure 5. It is evident that the temperature increases as the total luminosity increases. Moreover, the luminosity from the disk component increases at higher accretion rates and luminosities as shown in Figure 6 and Figure 7. This feature, together with the increasing disk size, implies that the corona/ADAF dominant flow in the low state naturally evolves to a disk dominant flow at high states consistent with “standard” picture developed for the different spectral states of X-ray binaries, although the exact transition is likely to be affected by the inclusion of X-ray irradiation of the hot coronal gas on the cool disk.

The increasing condensation rate with accretion rate or luminosity is a direct result of the efficient cooling at high accretion rates. Specifically for higher accretion rates, the density correspondingly increases ($n \propto \dot{m}$), which greatly enhances the bremsstrahlung cooling

rate ($\propto n^2$). In the transition layer, the conductive heating rate approximately increases with the viscous heating rate ($\propto n$). To maintain energy balance, additional heating, i.e., the condensation enthalpy rate, must increase with density more rapidly than $\propto n^2$. As the condensation rate affects the luminosity of the inner disk, while the total luminosity consists of contributions from bremsstrahlung, Compton and disk radiation, the disk luminosity increases more rapidly than the total luminosity. Thus, their ratio increases with the luminosity.

The numerical results reveal that condensation occurs only for accretion rates exceeding a critical value, which is found to depend on the viscosity parameter. In Figures 3-7, the effect of condensation was plotted from the lowest accretion rates found in the numerical calculations. In fact, the corresponding luminosity at the critical accretion rate is from a pure ADAF rather than from the small disk region as shown in the figures. Therefore, the lower limit to the luminosity for which an inner disk could exist may well be as low as $0.001 L_{\text{Edd}}$, depending on the radiative efficiency of the ADAF.

4. Comparison to Observations

The geometry of the accretion flow in the low hard states of X-ray transient systems has yet to be clarified since observational indications exist that cool material can be present in the inner regions where the reflection of X-rays and the formation of broad iron lines can contribute to the spectrum. For the bright portion of hard-state systems near the transition from a spectrally hard to soft state, continuous emission and reflection components from a geometrically thin, optically thick inner disk are theoretically expected. Although the observational signatures for such a disk component are difficult to detect, evidence for such a feature has been inferred from the source GX 339-4 during its 2004 outburst (Miller et al. 2006a, b) by the presence of a weak, soft, thermal component and by the suggestion of blurred reflection and iron lines. SWIFT J1753.5-0127 also reveals such a weak disk component (Miller et al. 2006a; Ramadevi & Seetha 2007; Soleri et al. 2008), however iron features were not detected (Miller et al. 2006a). For systems at very low luminosities, the general theoretical consensus favors a disk described by a pure ADAF solution in the inner region. In support of this picture, high quality spectra for several systems in quiescence at luminosities $L/L_{\text{Edd}} \lesssim 10^{-6}$ have not revealed any evidence for a thermal component or iron features (e.g. Bradley et al. 2007; Corbel et al. 2006).

Recently, Tomsick et al. (2008) have detected a weak thermal component from the source GX 339-4 at low luminosities ($L \lesssim 0.03L_{\text{Edd}}$) based on observations they had obtained using SWIFT and RXTE. In addition, broad features due to iron K transitions that are

likely produced by reflection of hard X-rays from the optically thick material near the ISCO have also been inferred from the spectrum. Here, we compare the results from our model predictions with these observational data. Of particular relevance for comparison with our theoretical analysis are the two observations of GX 339-4 after the source underwent a transition to the hard state carried out by Tomsick et al. (2008).

Specifically, Tomsick et al. (2008) carried out two observations of GX 339-4 after the source had undergone a transition to the hard state. One spectrum was obtained a few days after the transition (called spectrum #1) and another spectrum was obtained 2-3 weeks later (called spectrum #2). The inferred luminosity of spectrum #1 in the energy range from 1 to 100 keV was 0.023 times the Eddington value based on simultaneous observations obtained with the XRT, PCA and HEXTE instruments. Spectral fitting indicated the presence of a cool, optically thick disk located near the ISCO with a temperature of 0.201 keV. A similar analysis of spectrum #2 yielded a luminosity of 0.008 L_{Edd} and a temperature of the inner disk of 0.165 keV. Here, for convenient comparison we adopt a source distance of 8 kpc and a black hole mass of 5.8 M_{\odot} , which were used by Tomsick et al. (2008) and Miller et al. (2006b).

In order to fit the observed luminosity and disk temperature, the accretion rate \dot{m} and viscosity parameter α were varied until agreement between the theoretical modelling and the data deduced from observations was achieved. In general, the disk can be described in one of four different configurations depending upon the chosen values of the observables. In particular, for systems characterized by very low luminosities and disk temperatures (e.g. with coronal luminosities, $L_c/L_{\text{Edd}} = 0.001$ and $T_{\text{eff}} = 0.04$ keV) in the quiescent state, condensation is unimportant and the disk is described by an ADAF without the presence of an optically thick inner disk. At the other extreme, for systems at very high luminosities, say, $L_c/L_{\text{Edd}} \sim 0.1$, the mass accretion rate exceeds the critical rate above which an ADAF does not exist. The third configuration corresponds to the case where the spatial extent of the inner optically thick disk connects to the outer optically thick disk indicating that the disk is geometrically thin throughout and is no longer condensation-fed (a limiting case of configuration 4, see below). This occurs as a consequence of a relatively high mass accretion rate and low viscosity for systems with intermediate luminosities and relatively high temperatures. For instance, taking $L_c/L_{\text{Edd}} = 0.03$ and $T_{\text{eff}} = 0.26$ keV, it is found that the accretion rate and viscous parameter are 0.035 and 0.2 respectively. However, the derived size of the disk is $417R_{\text{S}}$, implying that the coronal gap region is absent. Only for values of \dot{m} and α in the appropriate range where the derived size of the inner disk is not very large (the fourth configuration) is the fit regarded as consistent with observations.

Our fitting procedure leads to the resulting parameters for spectrum #1 and #2 as

listed in Table 1. Given $\alpha = 0.31$ and $\dot{m} = 0.0687$, we derive a coronal luminosity of $0.023 L_{\text{Edd}}$, a disk temperature of 0.203 keV, and an inner disk size of $71 R_S$. Similarly, given $\alpha = 0.23$ and $\dot{m} = 0.0299$, we derive a coronal luminosity of $0.008 L_{\text{Edd}}$, a disk temperature of 0.164 keV, and an inner disk size of $82 R_S$. These results are in good agreement with the inferred disk parameters based on the two observations of Tomsick et al. (2008). With these two fits the relative strength of the cool disk to the hot corona can be estimated. As listed in Table 1, the ratio of the luminosities associated with the thermal component, L_d , to the total luminosity, L , is $\sim 10\%$, indicating a weak disk in the two observations. In addition, the radii delineating the regions where Compton cooling or conduction is important are listed in Table 1. For the quantitative comparison of cooling via Compton scattering and thermal conduction the ratio λ is shown in Figure 8 as a function of radius for these two cases. As the radial distribution of λ only depends on the condensation rate (see Eq.11), i.e., the disk temperature listed in Table 1, the values of λ in the two fits differ by a factor of ~ 2.3 at any given distance. Since it seems plausible to have the same value of α for both cases we have also explored solutions for the same value of α (see Table 1). Here, the observed luminosities are fit, but we permit small differences between the predicted and observed temperatures. It can be seen that the fit can still be taken as consistent with observations. The distinction between the fits for $\alpha = 0.25$ in the two spectra can be seen from the extent of the Compton dominated region as shown in Table 1. Specifically, the cooling in the corona for the fit to spectrum #2 is dominated by conduction throughout the inner disk, whereas the fit to spectrum #1 is dominated by Compton scattering in the inner region. This is a consequence of the differing disk temperatures as listed in Table 1. This also leads to the greater sensitivity of the luminosities for a small change in accretion rates in the two fits. Since the disk temperature determines the energy density of the soft photon field for Compton scattering ($u \propto T_{\text{eff}}^4$), the Compton luminosity in the fit to spectrum #1 is much larger than in spectrum #2, hence, becoming the dominant contribution to the coronal luminosity. Therefore, the radiation losses are more efficient when Compton cooling becomes important.

The fits to the two observations reveal that the temperature of the inner disk is positively correlated with the luminosity. We note the inferred reflection covering factors of Tomsick et al. (2008) as well as Miller et al. (2006b) are comparable (i.e., the covering factor is inferred as 0.22 from spectrum #1 and 0.24 from spectrum #2 and 0.22 from the 2004 observations (Miller et al. 2006b)), which may suggest a comparable spatial extent of the inner disk. Such an insensitivity to the luminosity could be understood if the viscosity parameter varied (see Table 1).

5. Discussion

The structure of the inner disk region of accretion disks have been investigated for black hole X-ray transient systems in their low luminosity quiescent state within the context of an α viscosity model. Particular attention has been focused on the conditions under which a weak, soft thermal component would be present in this state. By generalizing the study of Liu et al. (2007) to include the cooling associated with both conduction and Compton processes together, we have examined the properties of an optically thick inner disk fed by the condensation of gas from a corona for a range of mass accretion rates. It has been shown that such a disk component can be present with its spatial extent, temperature, and luminosity varying continuously and decreasing with a decreasing mass accretion rate. Despite its existence and the higher conversion efficiency of gravitational energy into radiation in the cool disk, the contribution of the disk component to the total luminosity can be very low due to the low condensation rate. Throughout the quiescent state, the mass flow rate to the central black hole from the corona always exceeds the mass flow rate in the inner cool disk. Hence, the disk structure remains geometrically thick even in the innermost regions although a soft thermal component is present in the spectrum.

In our analysis, we have neglected the effect of X-ray irradiation to facilitate comparisons with our earlier work. Although the qualitative features of our model will not change, the parameters describing the quantitative comparisons will be modified. In particular, the inclusion of X-ray irradiation can affect the fitting parameters α and \dot{m} since the observed disk temperature will not solely be determined by the mass flow in the inner disk. For example, for ADAF emission from the innermost region at height $H = R_{\text{ISCO}} = 3R_s$, the flux irradiating the disk is given by

$$F_{\text{ir}}(r) = \frac{\frac{1}{2}L_c}{4\pi R^2} \frac{H}{R} = 1.716 \times 10^{26} \text{erg s}^{-1} \text{cm}^{-2} m^{-1} r^{-3} \frac{L_c}{L_{\text{Edd}}}. \quad (25)$$

Adopting a mean photon energy of 50keV, about one tenth of the irradiation flux (the ratio of the mean photon energy to electron rest mass energy) is absorbed by the disk through Compton scattering collisions with electrons in the disk. Although the K-shell absorption of the metals is negligible at 50 keV, it can become significant at lower energies since the absorption cross section is proportional to $(m_e c^2/h\nu)^{7/2}$. For a photon distribution of $n_\nu \propto \nu^{-1.6}$, the flux absorbed by K-shell absorption is found to be $\sim 30\%$ of the total irradiation flux. Thus, about 40% of the total irradiation flux is absorbed by the K-shell absorption and Compton scattering. The re-radiation of the absorbed flux at the inner disk surface yields an effective temperature of

$$T_{\text{ir}}(r) = \left(\frac{0.4F_{\text{ir}}}{\sigma} \right)^{1/4} = 2.86 \text{keV} m^{-1/4} r^{-3/4} \left(\frac{L_c}{L_{\text{Edd}}} \right)^{1/4}. \quad (26)$$

At the distance of maximal temperature, $r_{tmax} = 3\frac{49}{36}$, the effective temperature caused by irradiation alone for spectrum#1 is 0.25keV and for Spectrum#2 is 0.19keV (see §4). These temperatures are slightly higher than the temperatures generated by disk accretion and could contribute to the observationally inferred temperature. We note that this theoretical estimate is very approximate since an accurate calculation depends on the region of origin of the ADAF flux and the size of the coronal gap region. From Eqs.22 and 23 it can be seen that the corona/ADAF radiation is dominant from a region outside the ISCO. In this case, the irradiation flux incident on the innermost disk decreases with distance of the source, $F_{ir} \propto r^{-2}$. For instance, assuming the ADAF radiation is dominant from $5R_s$ rather than $3R_s$, the irradiation flux absorbed by the disk around the ISCO is 7% of the total flux. On the other hand, many coronal photons escape without interception by the disk for a large coronal gap, thereby, reducing the thermal effects associated with the irradiation. In the case for Spectrum #1 and #2 of GX339-4, the reflection covering factor inferred from observations is $\sim 20\%$ of the half sky, indicating that, indeed, a very large fraction of the coronal photons escape, in contrast to the case of Haardt & Maraschi (1991). Hence, these estimates indicate that our results are indicative and that additional freedom exists in determining \dot{m} and α when irradiation is taken into account.

For sufficiently low mass accretion rates, the cool inner disk vanishes. An estimate of the mass accretion rate for which this occurs depends on whether condensation occurs in the vicinity of the ISCO. If condensation does not occur, accretion and/or evaporation can deplete the inner disk within a viscous time. In the case of cooling dominated by conduction, the critical condensation radius is given by $r_d = 0.815\alpha^{-28/3}\dot{m}^{8/3} \geq 3$. Thus, the lower limit to the accretion rate is

$$\dot{m} = 0.006 \left(\frac{\alpha}{0.2} \right)^{7/2}. \quad (27)$$

For an accretion rate just above this lower limit, both the size of the inner disk and the condensation rate are very small, leading to a very low disk temperature. In this case, the Compton cooling is negligible, and the lower limit predicted by the conduction model is self-consistent.

For disk temperatures exceeding 0.14 keV, on the other hand, Compton cooling dominates in a corona around a $10M_\odot$ black hole. If this temperature is produced by the accretion of condensed gas, the condensation rate would be greater than $0.002 \dot{M}_{Edd}$. However, this occurs at high accretion rates and is greater than the lower limit given by conduction model for $0.1 \leq \alpha \leq 0.4$ (King et al. 2007). Thus, the lower limit to the mass accretion rate for maintenance of an inner disk is set by the conduction model, i.e. equation (27). Notwithstanding the possible importance of irradiation of the inner disk by the corona in providing some latitude in disk parameters, α and \dot{m} (see above), when fitting the observational data,

a determination of the mass flow rate below which the cool disk component vanishes could provide a constraint, within the framework of the condensation model, on the α viscosity parameter during the quiescent state.

In the future, global calculations of the full radial and vertical disk structure should be considered taking into account the effect of X-ray irradiation of the hot coronal gas on the inner disk and of the cooling associated with both the Compton and conductive processes in order to quantitatively confirm the picture presented in this paper. With such models in hand, a detailed comparison of the spectra from such disks with observations can be undertaken.

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Table 1: Fitting results for GX 339-4

α	\dot{m}	L_c/L_{Edd}	L/L_{Edd}	$L_d/L(\%)$	$T_{\text{eff}}(\text{keV})$	r_d	Compton region
Variable viscous parameter							
0.31	0.0687	0.023	0.0253	9.3	0.203	71	$3.3 \lesssim r \lesssim 19$
0.23	0.0299	0.008	0.009	11.2	0.164	82	$4.0 \lesssim r \lesssim 8$
Fixed viscous parameter							
0.25	0.0471	0.023	0.0269	14.4	0.221	142	$3.2 \lesssim r \lesssim 25$
0.25	0.0350	0.008	0.0088	9.0	0.156	61	no Compton region

Note. — Here L_c , L , L_d , T_{eff} , and r_d are the model predictions for the corona luminosity, total luminosity, disk luminosity, disk temperature and disk size for given viscous parameter α and accretion rate \dot{m} .

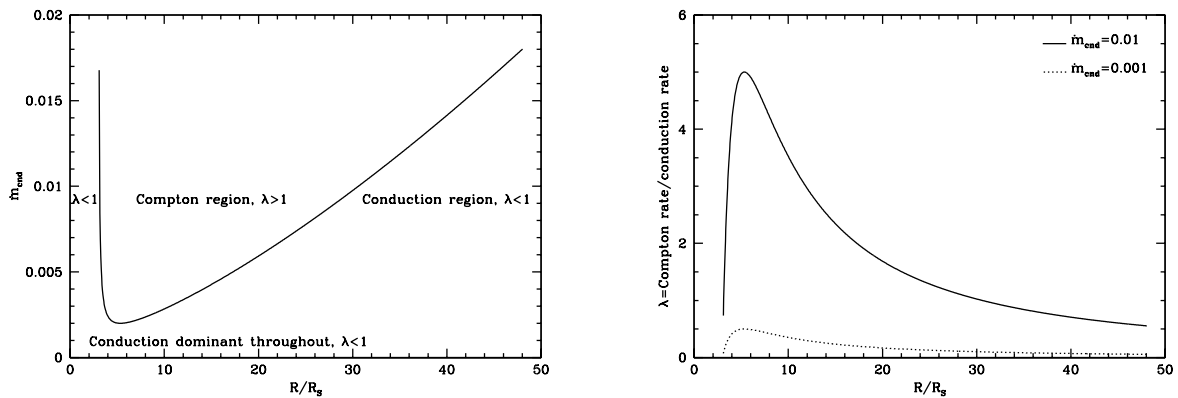


Fig. 1.— The left panel illustrates regions dominated by the Compton or conduction cooling process as a function of the total condensation rate. The Compton scattering is dominant at higher condensation rates (or higher accretion rates). The right panel shows the regions where Compton cooling is important and its relative strength, λ , as a function of distance for condensation rates of $\dot{m}_{\text{cnd}} = 0.001$ (dotted curve) and $\dot{m}_{\text{cnd}} = 0.01$ (solid curve). Between these two condensation rates, inclusion of both cooling mechanisms is necessary, especially at distances where $\lambda \sim 1$.

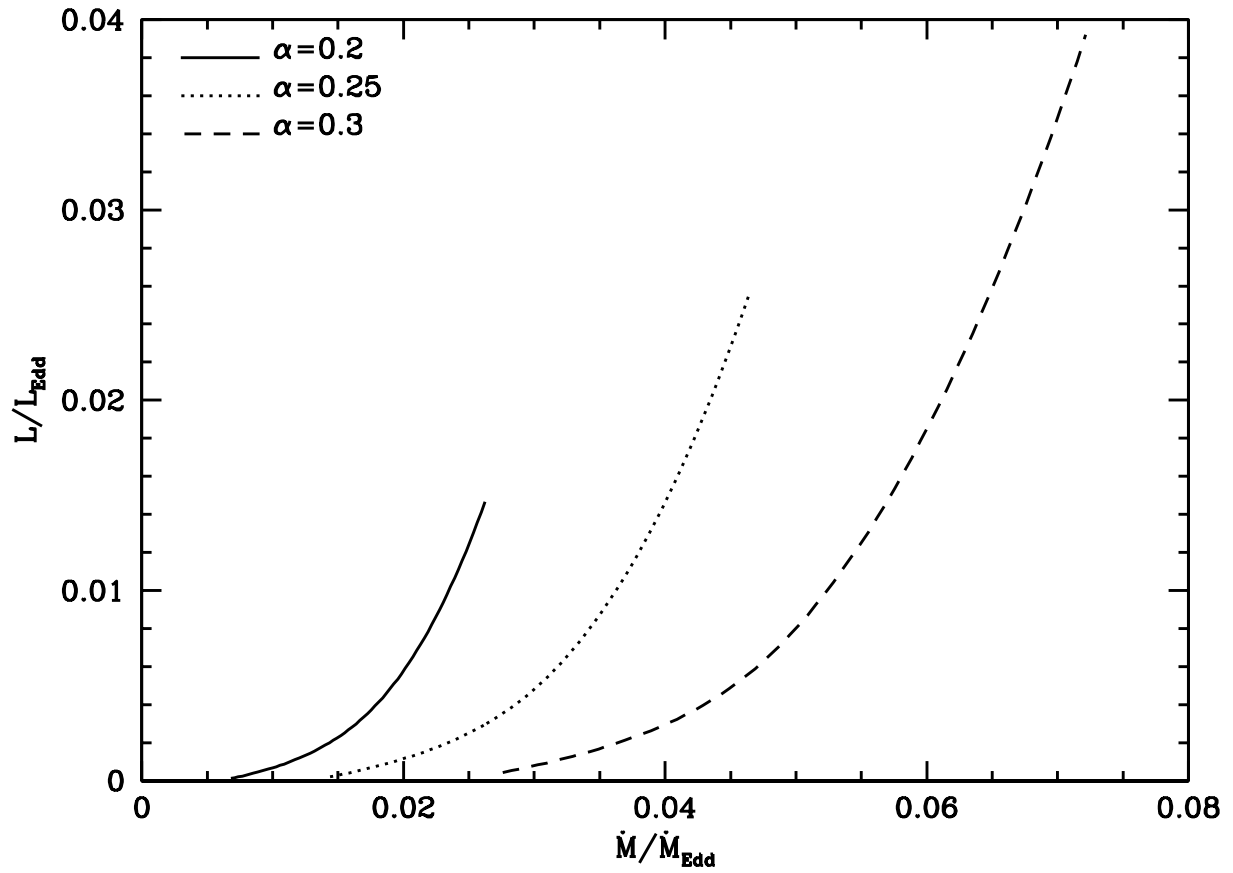


Fig. 2.— The luminosity as a function of the mass accretion rate. The total luminosity increases with accretion rate and the energy conversion efficiency in the accretion flow, as measured by $\frac{\eta}{0.1} = \frac{L}{L_{\text{Edd}}} / \frac{\dot{M}}{\dot{M}_{\text{Edd}}}$, is greater at higher accretion rates and lower values of the viscosity parameter α .

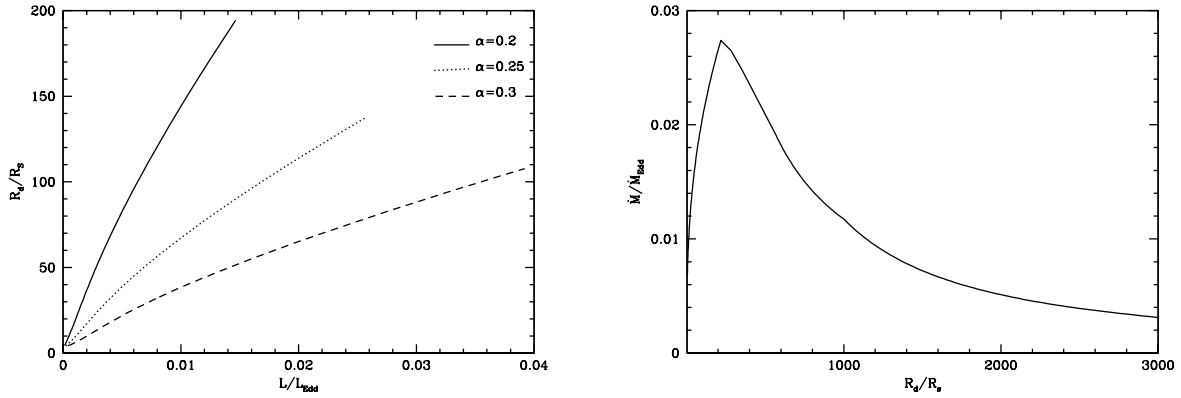


Fig. 3.— Left panel: The inner disk size, which is modified by inclusion of two cooling processes is shown as a function of the total luminosity. The size of the inner disk is large at high luminosity. At a given luminosity, the inner disk size increases with decreasing values of the viscosity parameter α . Right panel: The outer extent of the inner cool disk and the inner extent of the outer cool disk are illustrated as a function of the mass accretion rate for a given viscosity parameter, $\alpha = 0.2$. For a given accretion rate, the difference between these two radii yields the spatial extent of the coronal gap region. The curve combines the condensation model described in §2 for the inner disk and the disk evaporation model for the inward extent of the outer disk based on Liu et al. (2002).

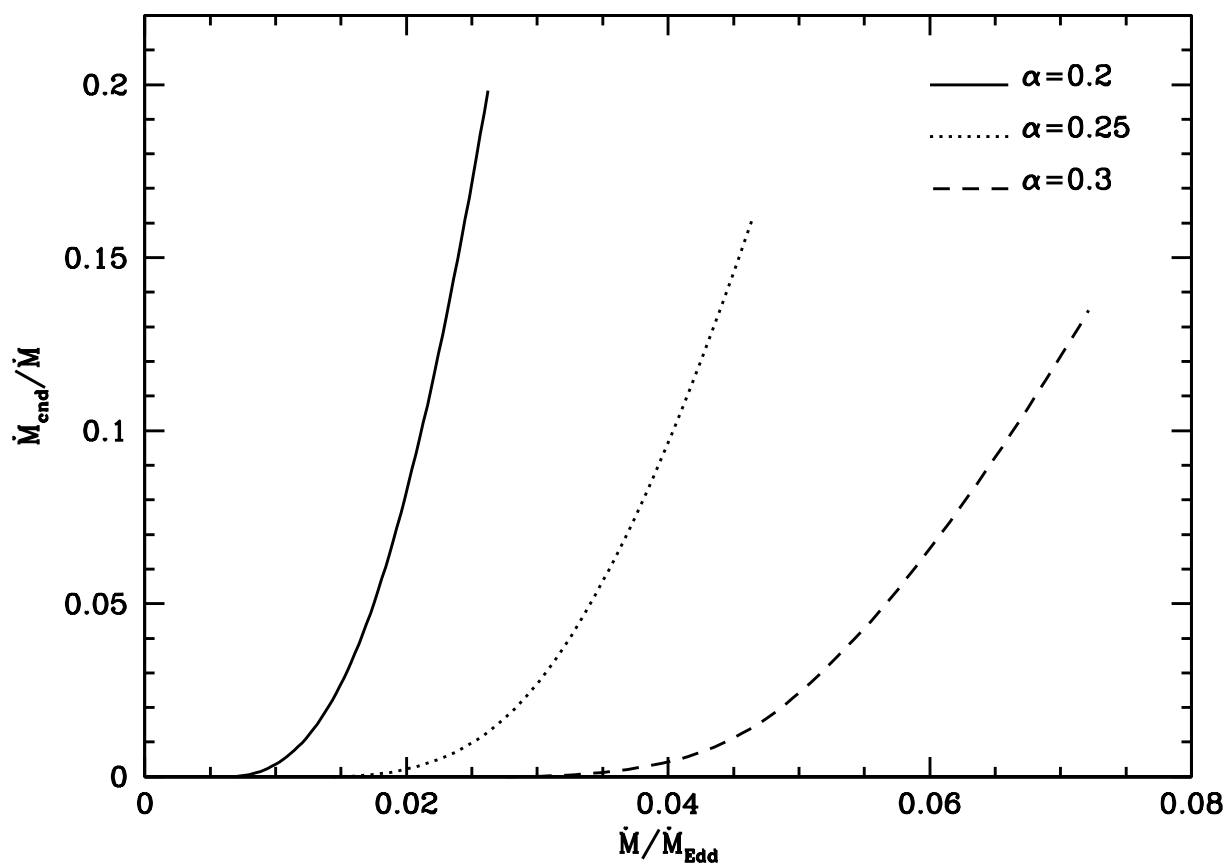


Fig. 4.— The fraction of integrated condensation flow as a function of mass accretion rate for three values of the viscosity parameter α . A greater fraction of the coronal flow condenses onto an inner disk at higher accretion rates, whereas no condensation occurs at very low accretion rates. This indicates a strong coronal flow can become weaker toward the central black hole with gas continuously condensing and accreting through a cool disk.

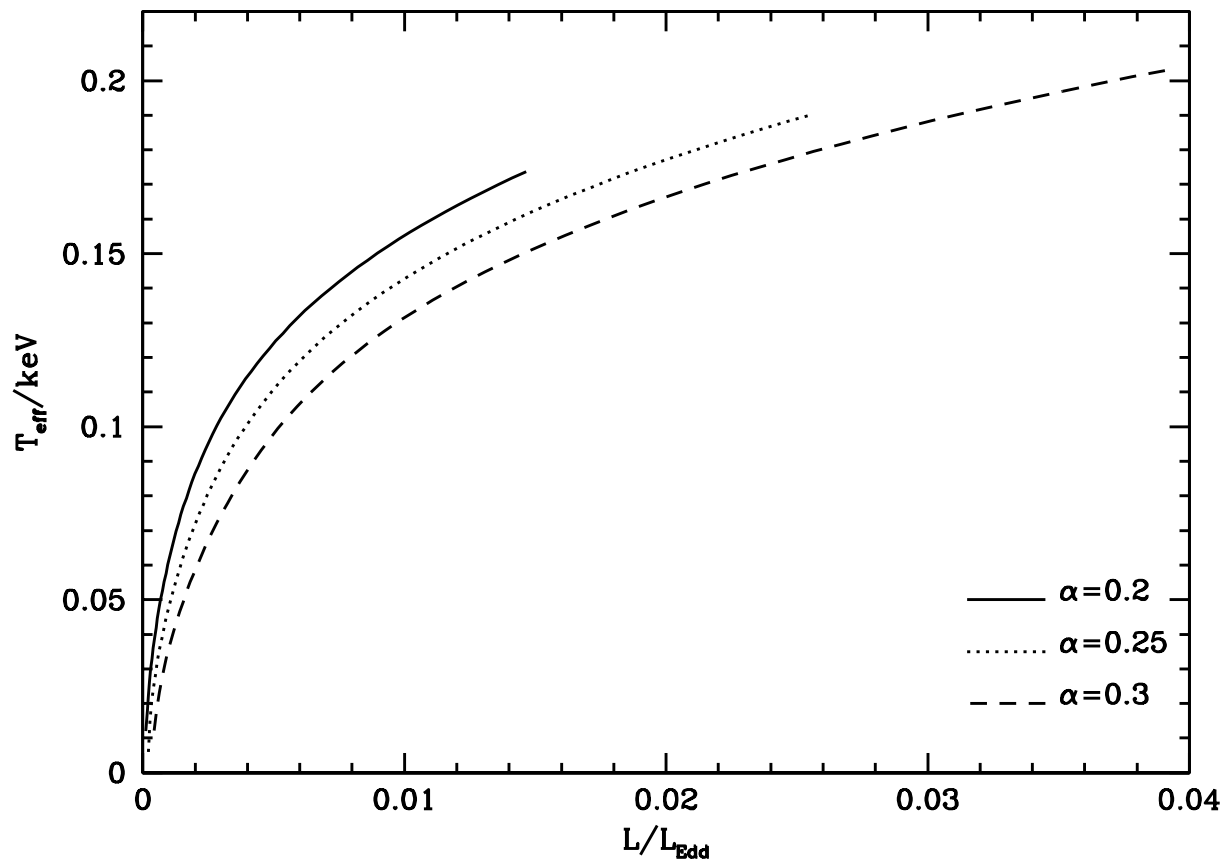


Fig. 5.— The maximal disk effective temperature as a function of luminosity for three values of the viscosity parameter α . The disk is heated by the accretion of the condensed gas to a high temperature at a high luminosity. At a given luminosity, the temperatures are insensitive to the viscosity parameters with only a slight increase for lower viscosity parameters α .

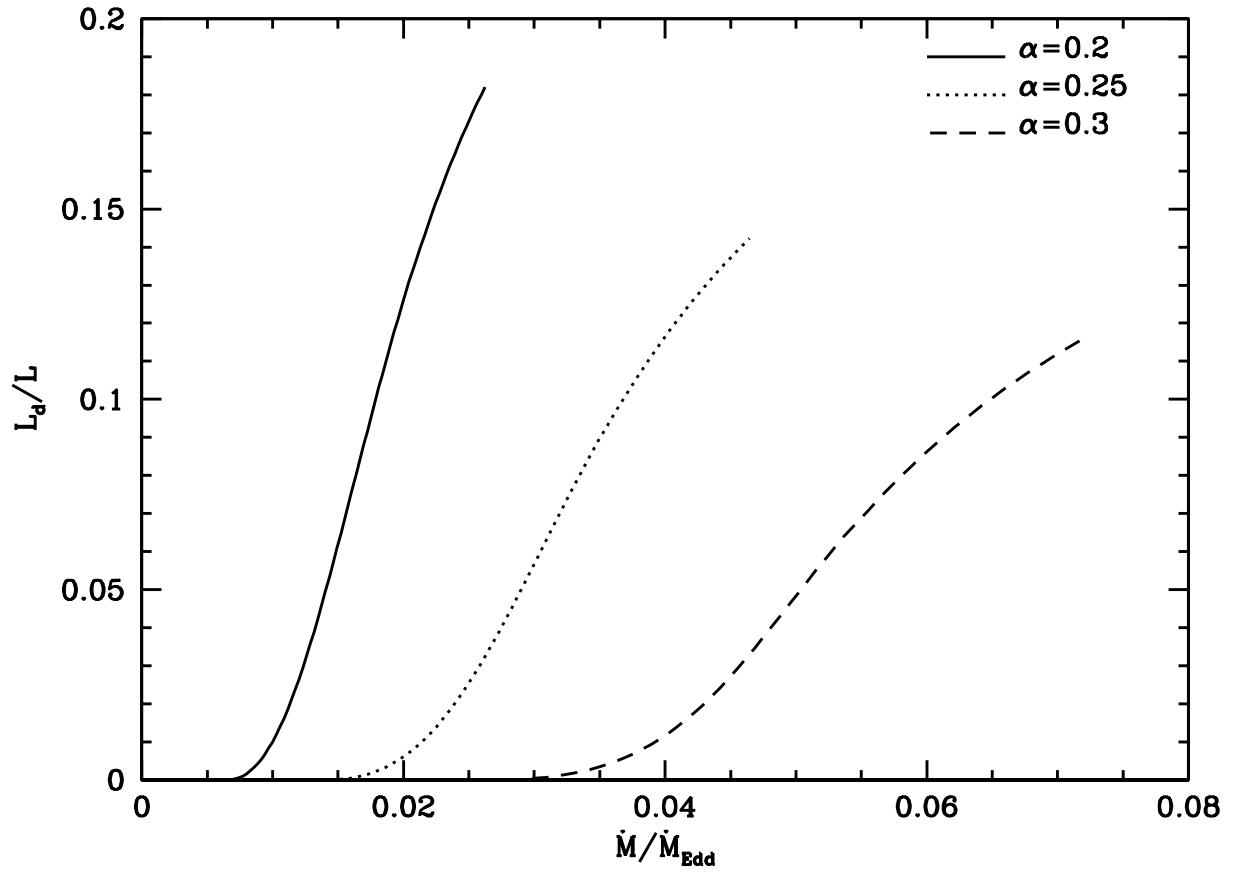


Fig. 6.— The fraction of luminosity produced in the inner disk to the total luminosity as a function of the mass accretion rate for three values of the viscosity parameter α . The contribution of the disk strengthens at higher accretion rates.

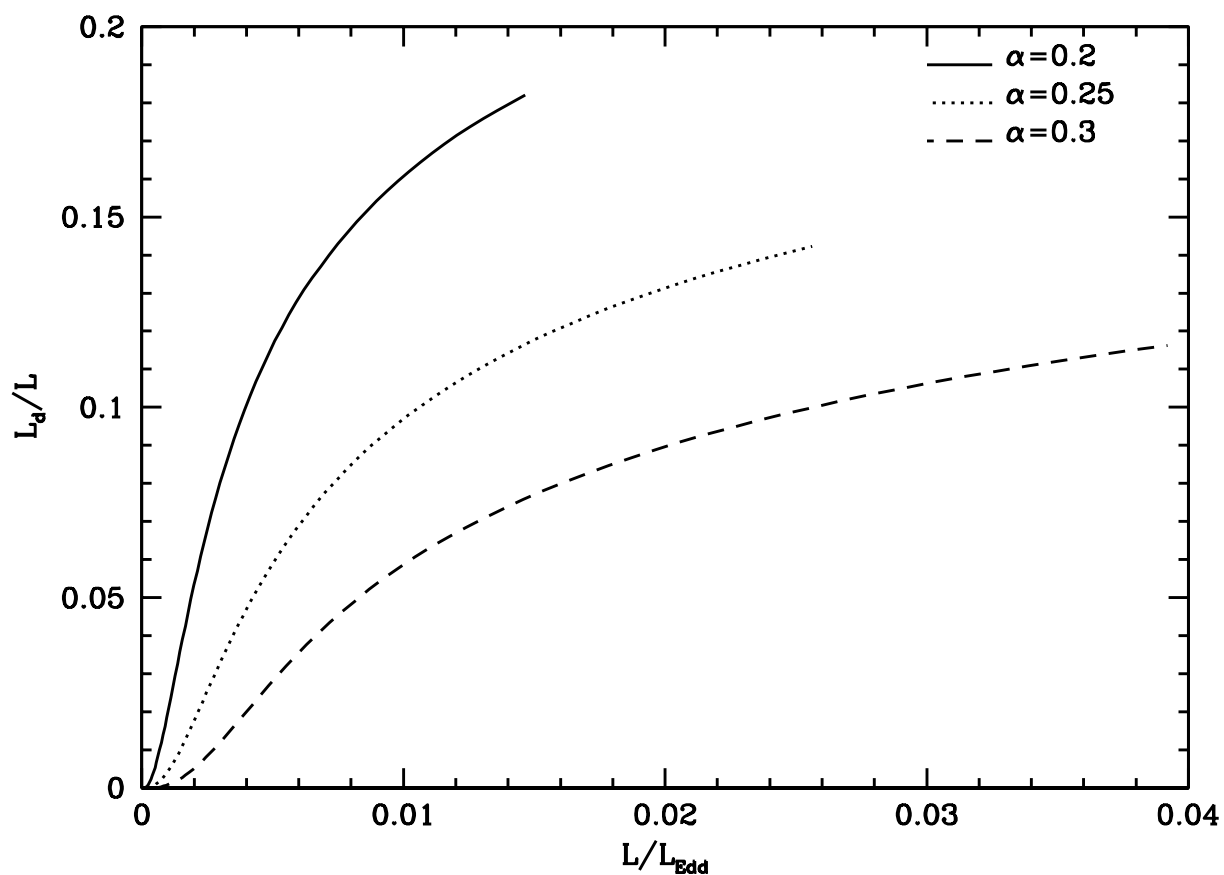


Fig. 7.— The fraction of disk luminosity to the total luminosity as a function of the total luminosity for three values of the viscosity parameter α . At low luminosities ($L/L_{\text{Edd}} \lesssim 0.01$), the relative strength of disk component increases rapidly with the total luminosity. The increase is less rapid and the ratio of luminosities tends to nearly a constant value at higher luminosities.

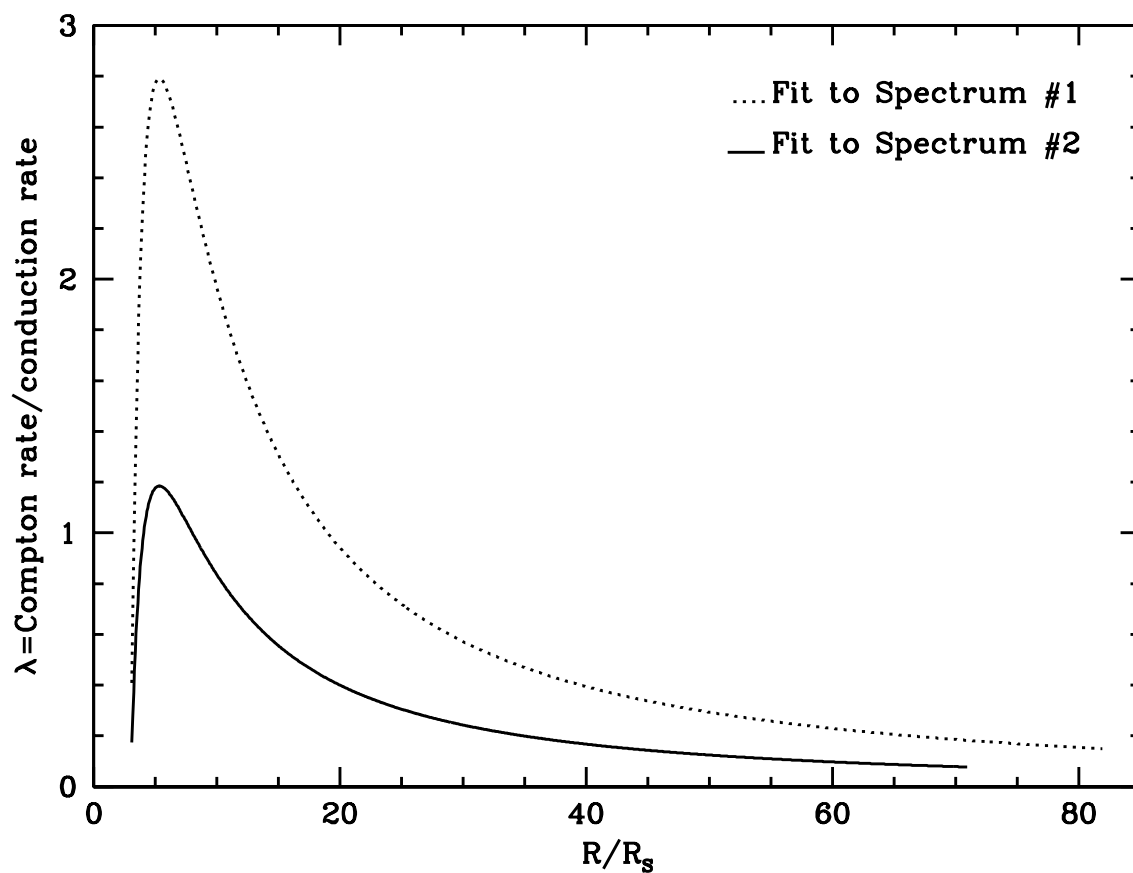


Fig. 8.— The cooling rate of the Compton process relative to that of the conduction process as a function of distance in the fits to spectrum #1 and #2. It can be seen that Compton cooling is stronger than conduction cooling in the model for spectrum #1 where a disk thermal component of 0.2 keV is fit. In the model for spectrum #2 the Compton cooling is comparable with conduction only in a small region and is much lower in the outer region.