Finding Cosmic Inflation

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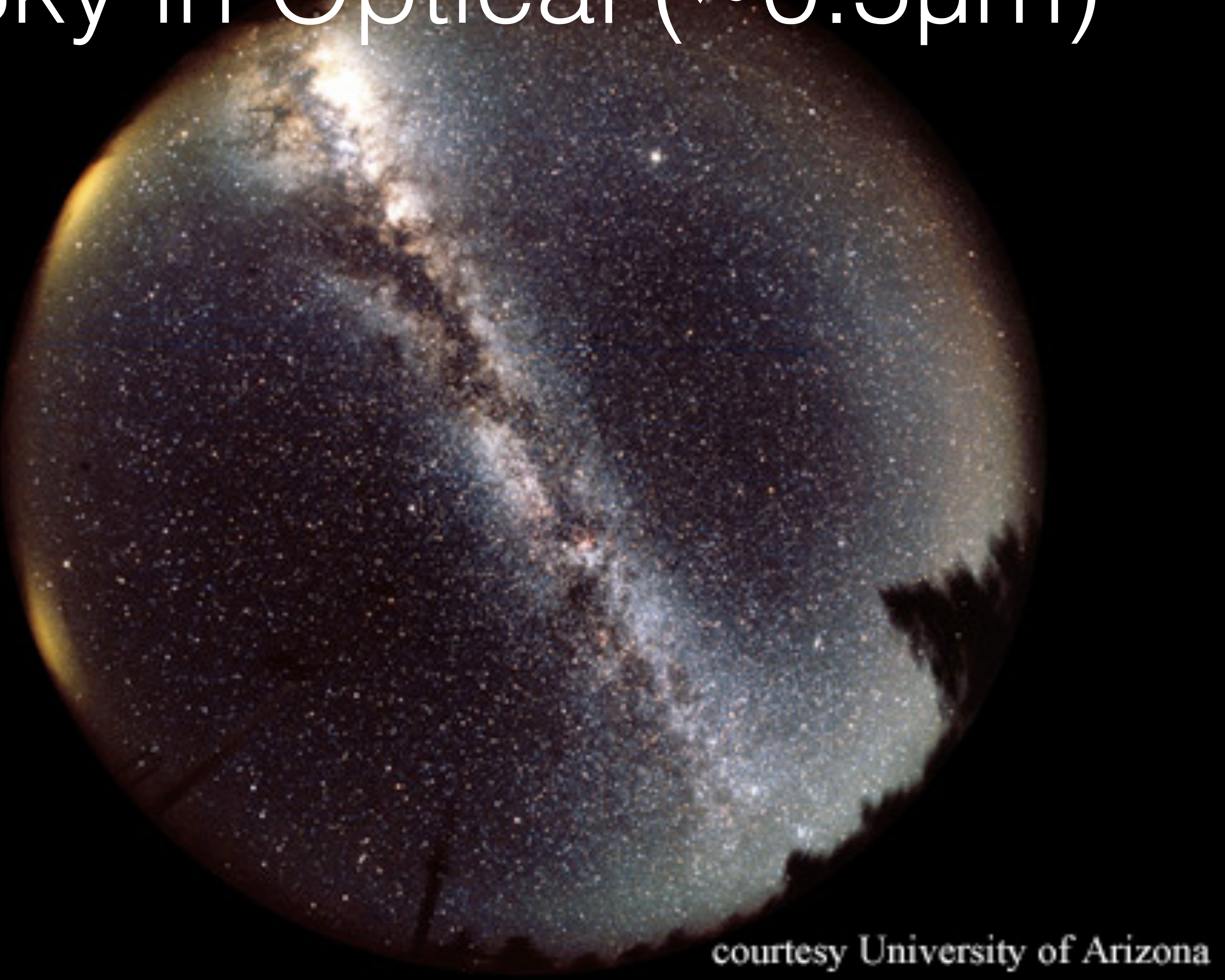
Breakthrough in Cosmological Research

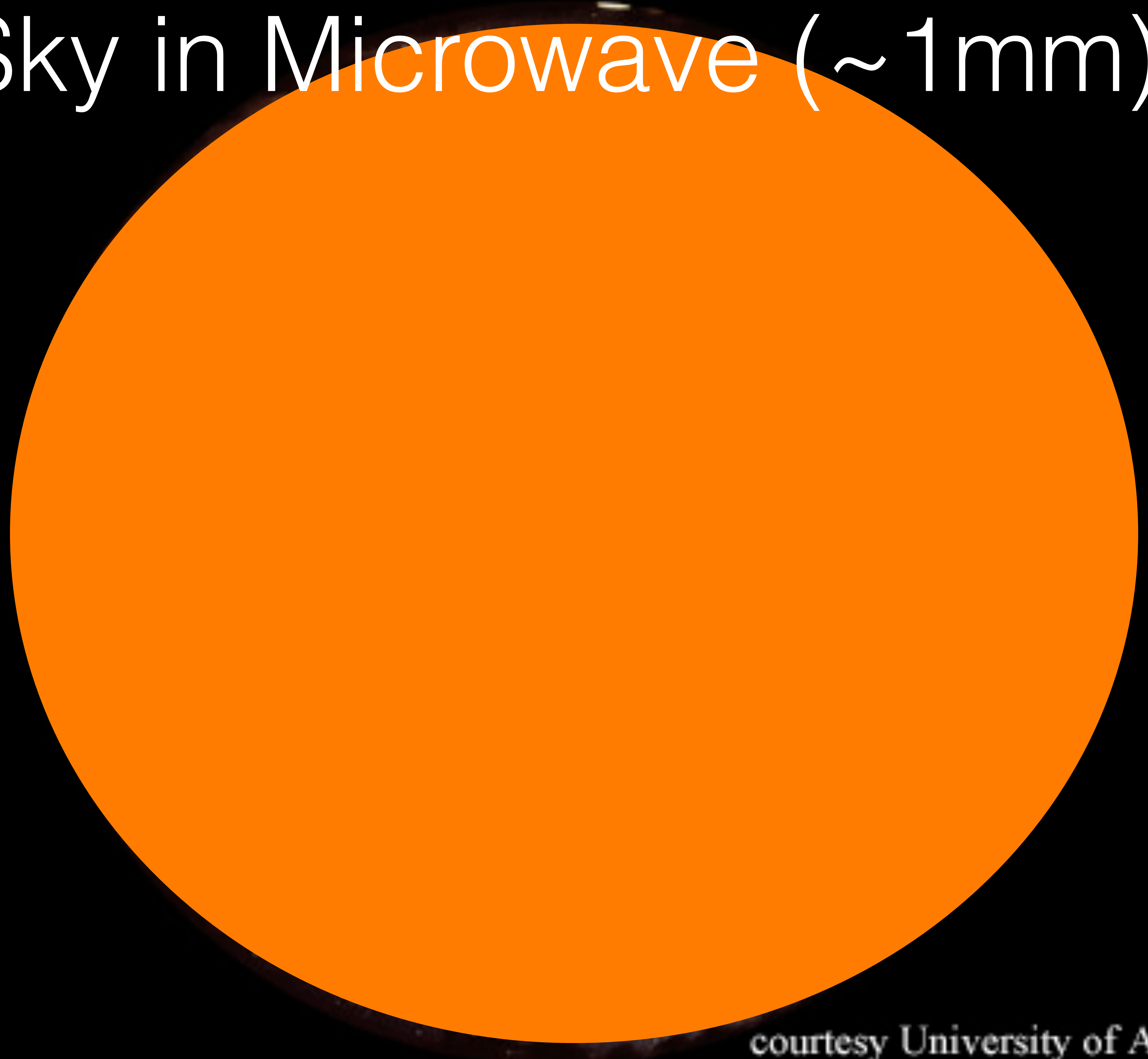
- We can actually **see** the physical condition of the universe when it was very young


From “Cosmic Voyage”

Sky in Optical (~0.5μm)

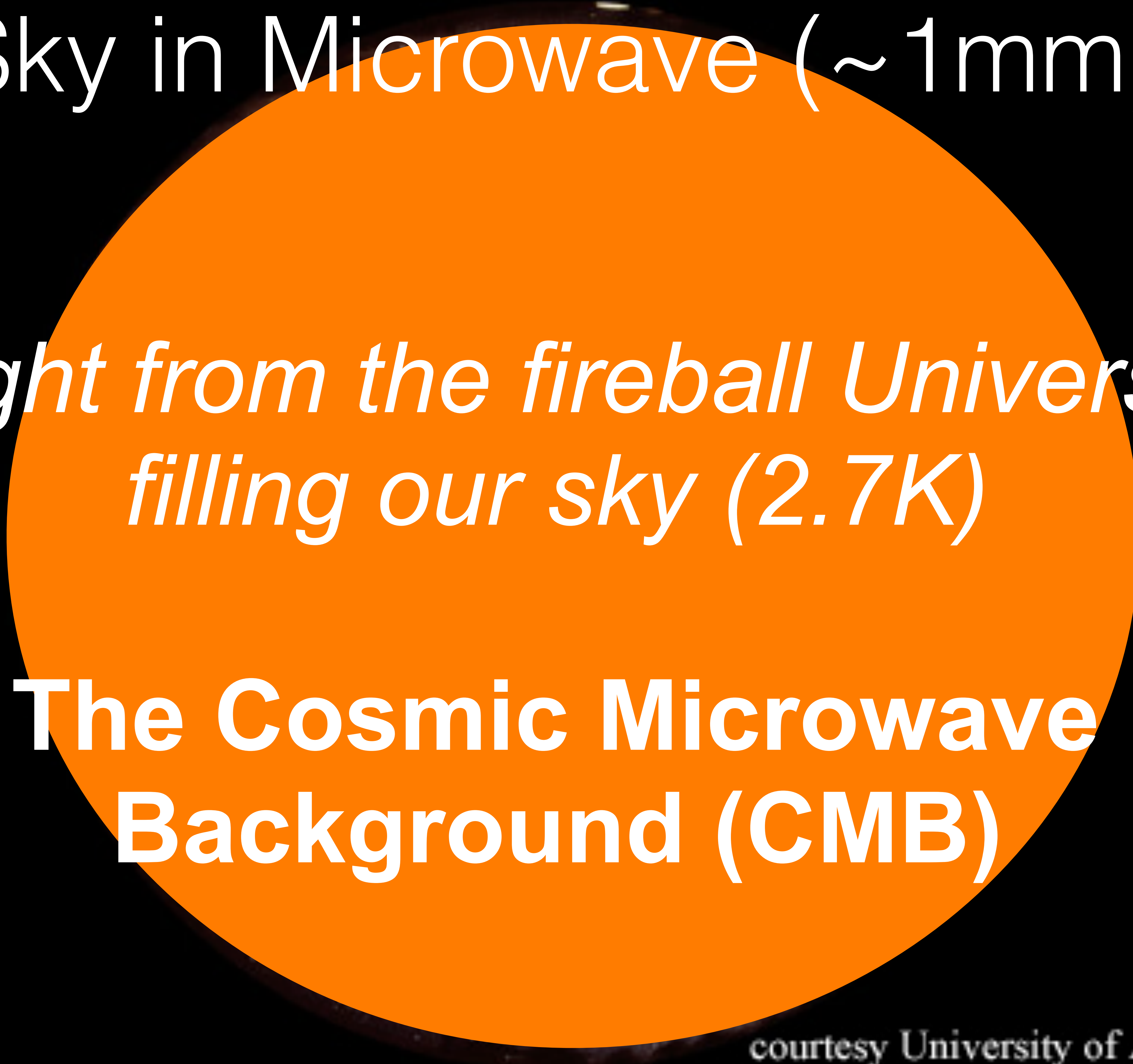


Sky in Microwave (~1mm)



courtesy University of Arizona

Sky in Microwave (~1mm)



Light from the fireball Universe filling our sky (2.7K)

The Cosmic Microwave Background (CMB)

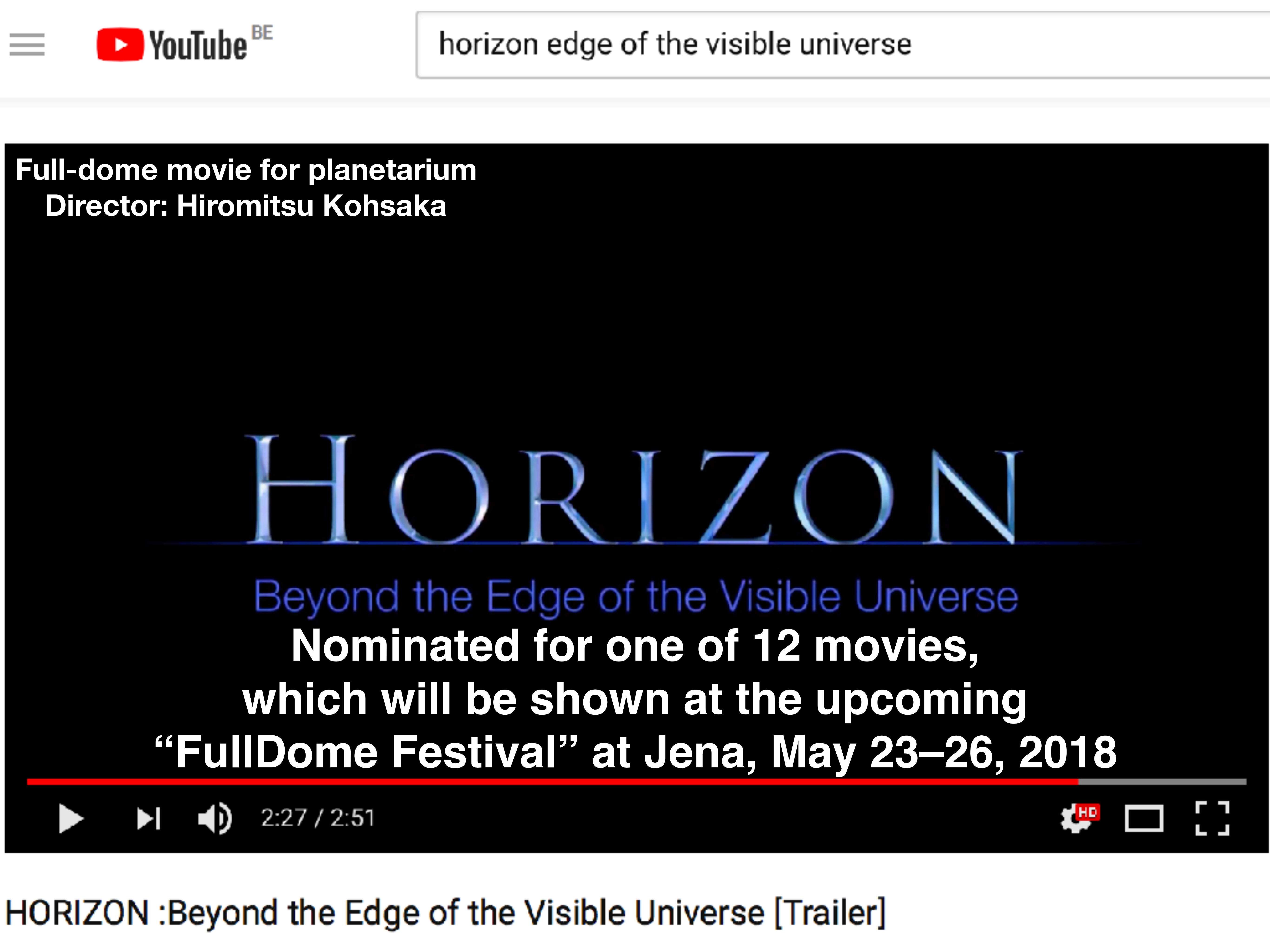
courtesy University of Arizona

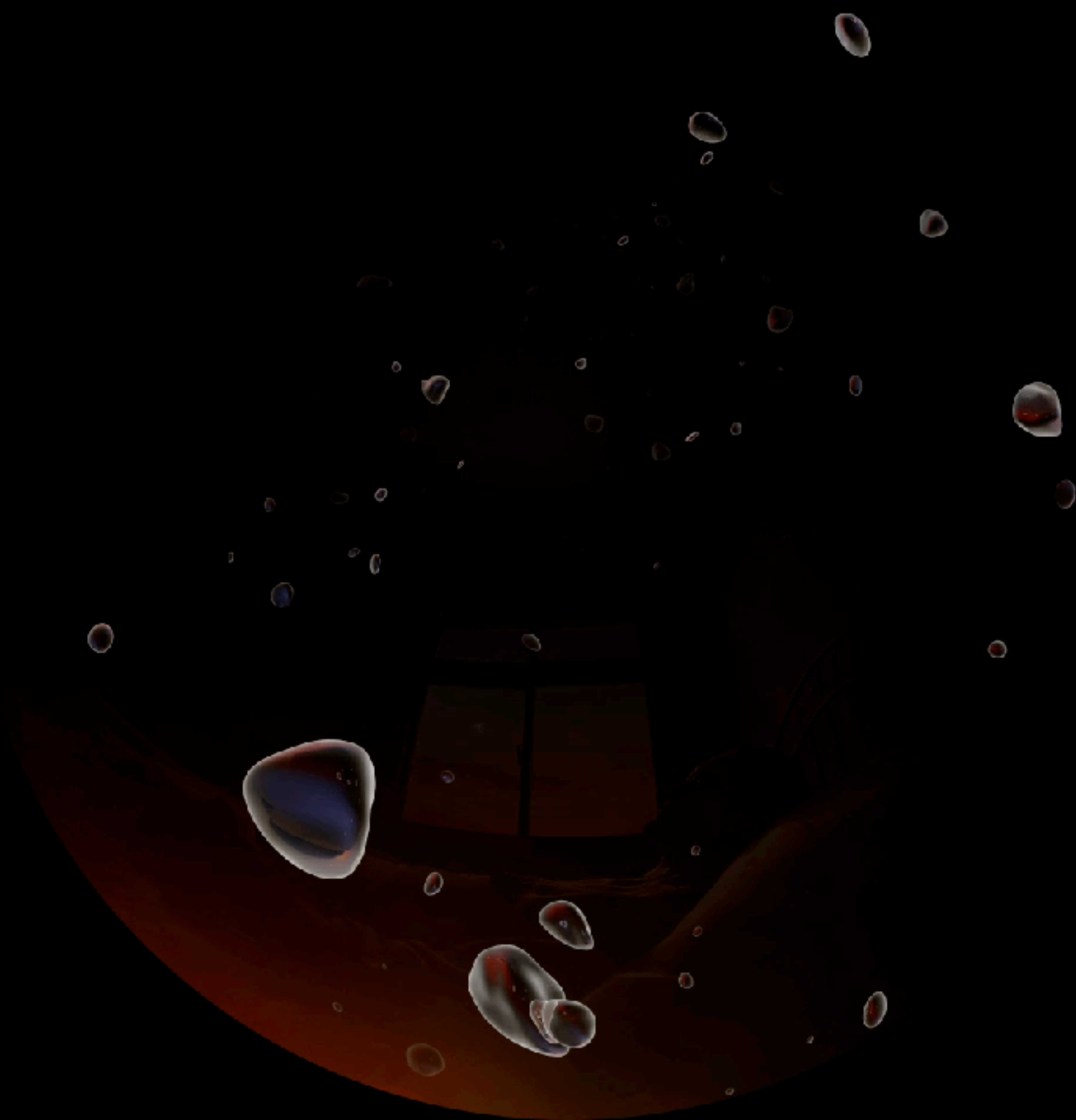
410 photons

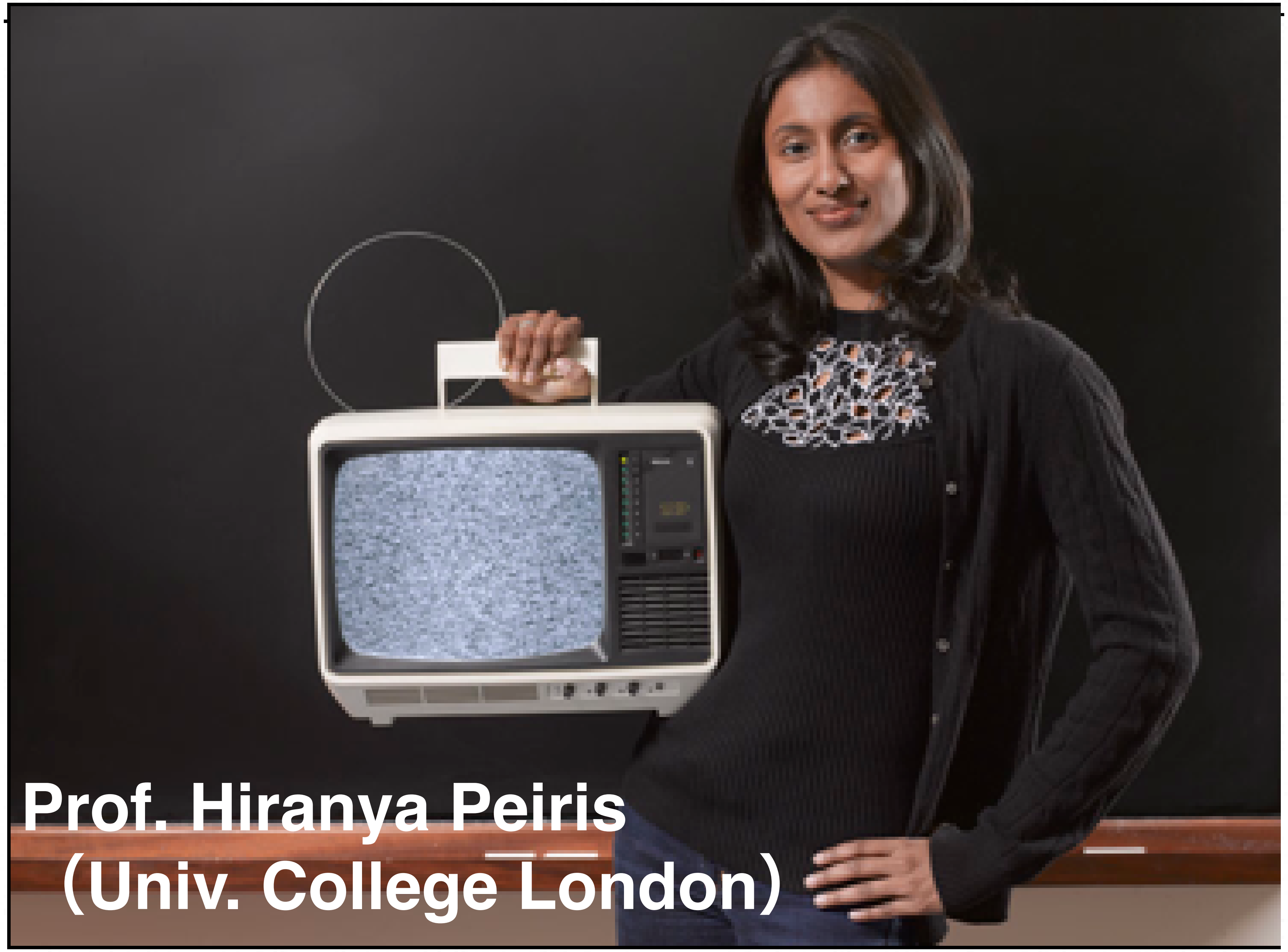
per

cubic centimeter!!

courtesy University of Arizona

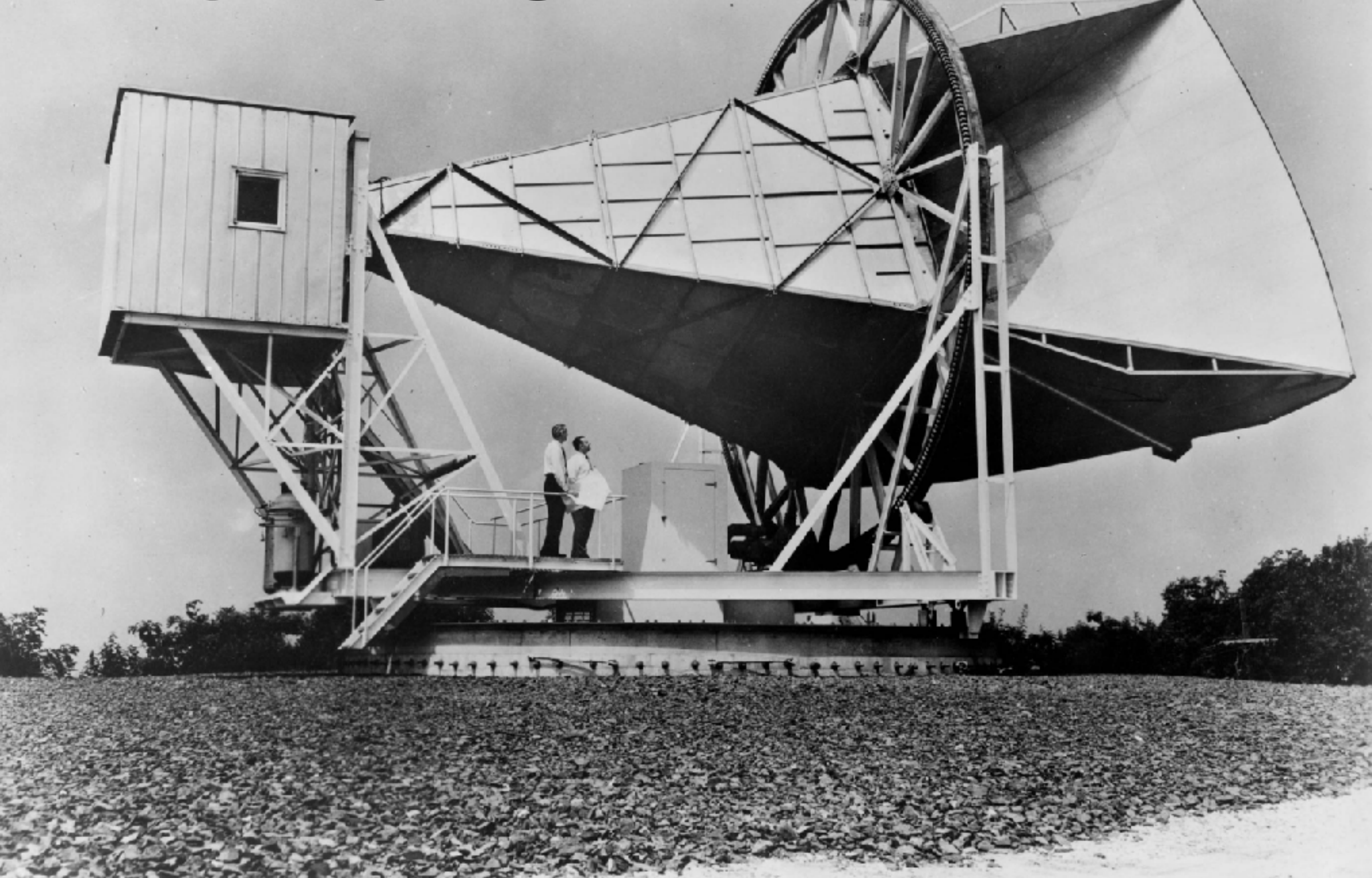

YouTube BE
horizon edge of the visible universe
Full-dome movie for planetarium
Director: Hiromitsu Kohsaka
HORIZON
Beyond the Edge of the Visible Universe
Nominated for one of 12 movies,
which will be shown at the upcoming
“FullDome Festival” at Jena, May 23–26, 2018
2:27 / 2:51
HD
HORIZON :Beyond the Edge of the Visible Universe [Trailer]



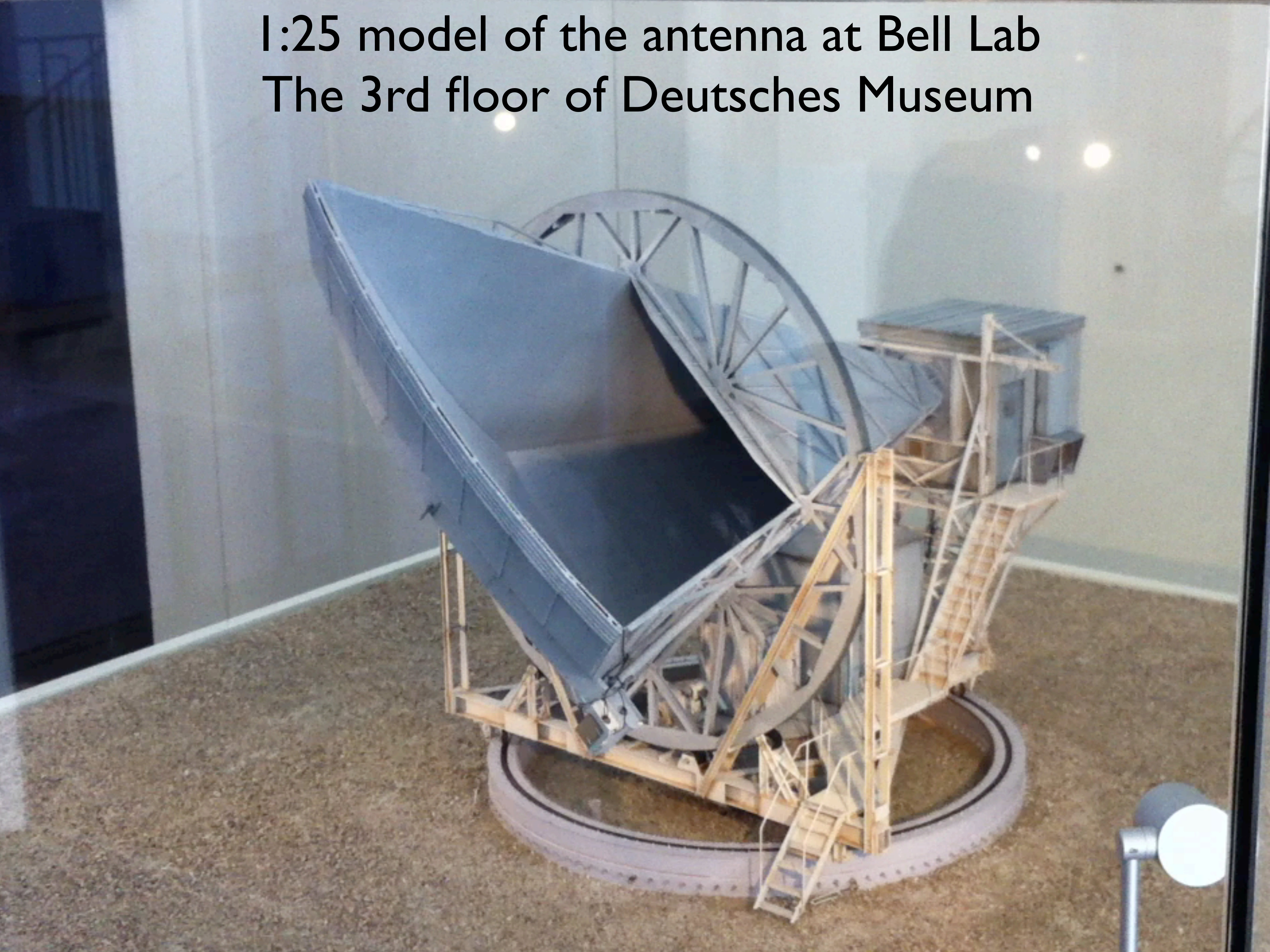


All you need to do is to detect radio waves. For example, 1% of noise on the TV is from the fireball Universe

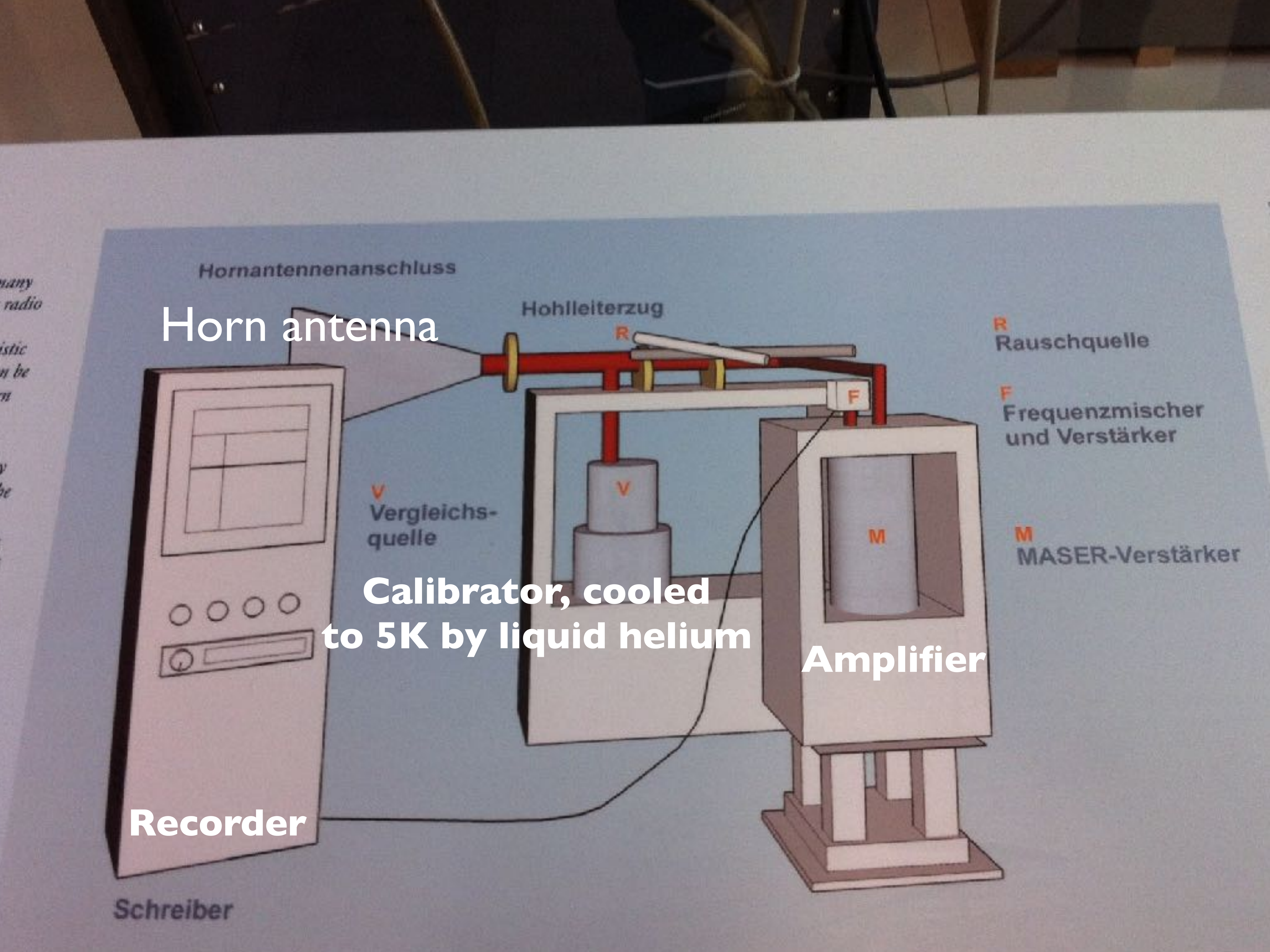
1965



1:25 model of the antenna at Bell Lab
The 3rd floor of Deutsches Museum

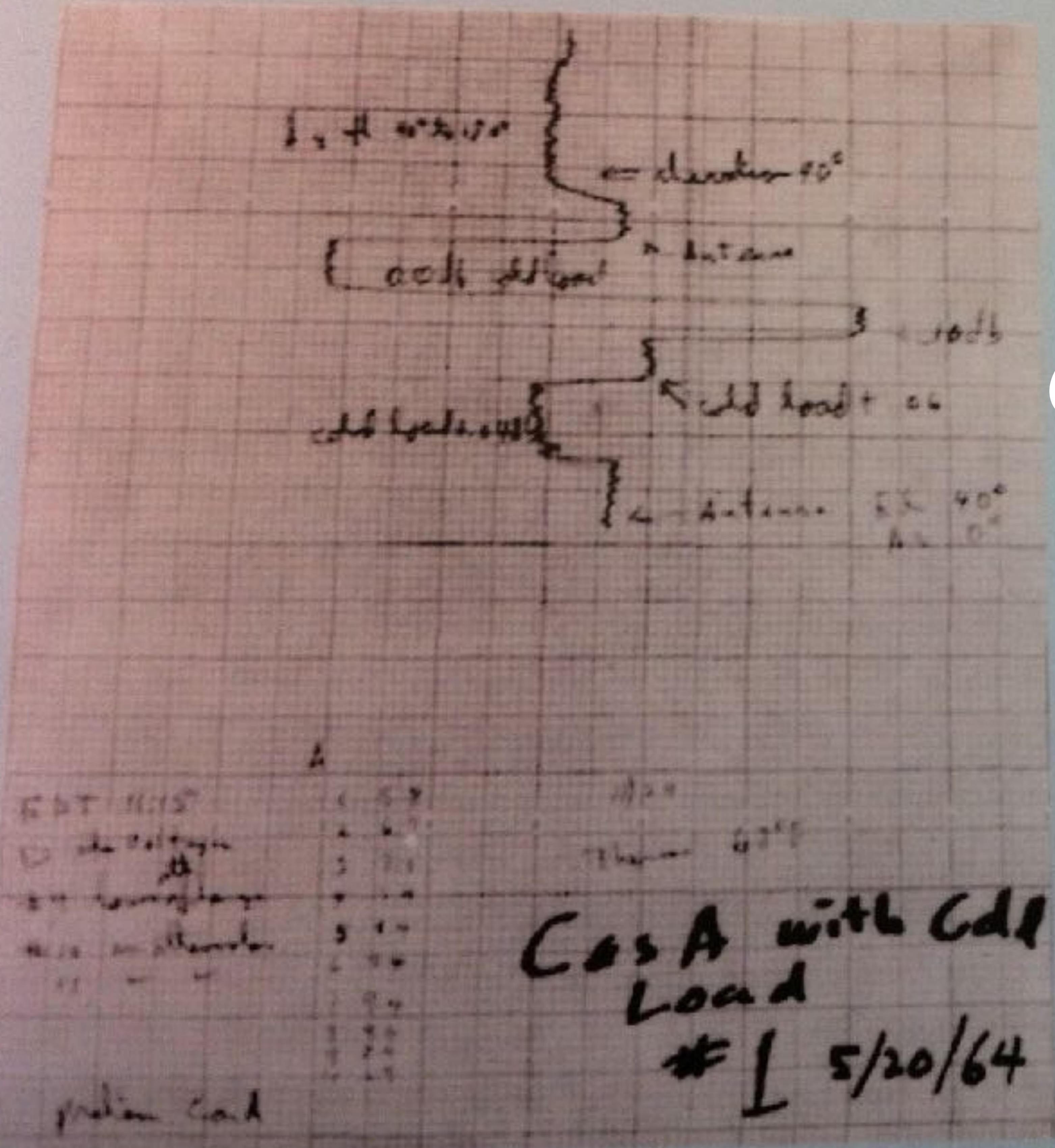


The real detector system used by Penzias & Wilson
The 3rd floor of Deutsches Museum


Hornantennenanschluss
Horn antenna
Hohlleiterzug
R
R
Rauschquelle
F
F
Frequenzmischer
und Verstärker
V
V
Vergleichs-
quelle
M
M
MASER-Verstärker
Calibrator, cooled
to 5K by liquid helium
Amplifier
Recorder
Schreiber

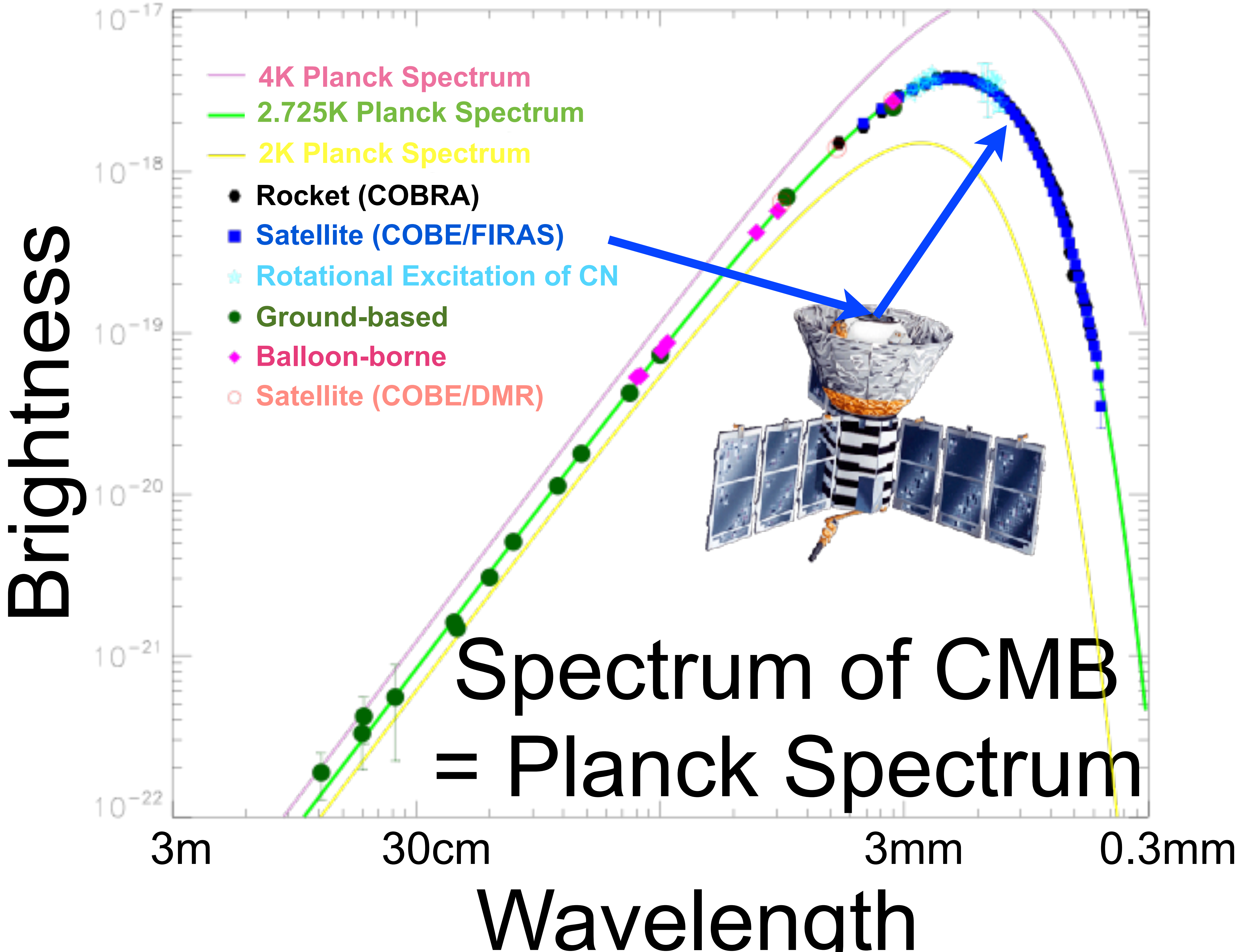
May 20, 1964 CMB Discovered

6.7–2.3–0.8–0.1 = 3.5±1.0 K

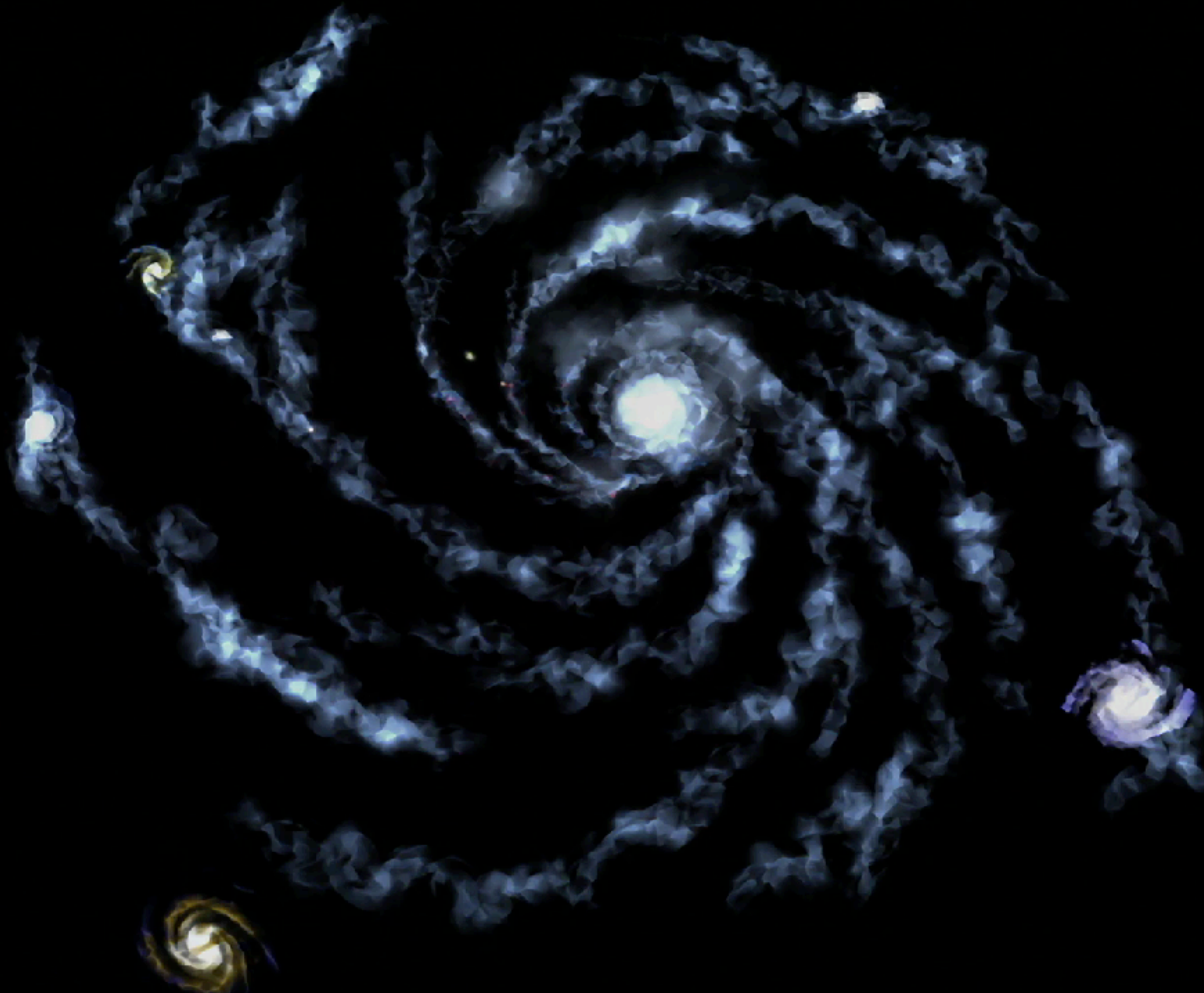


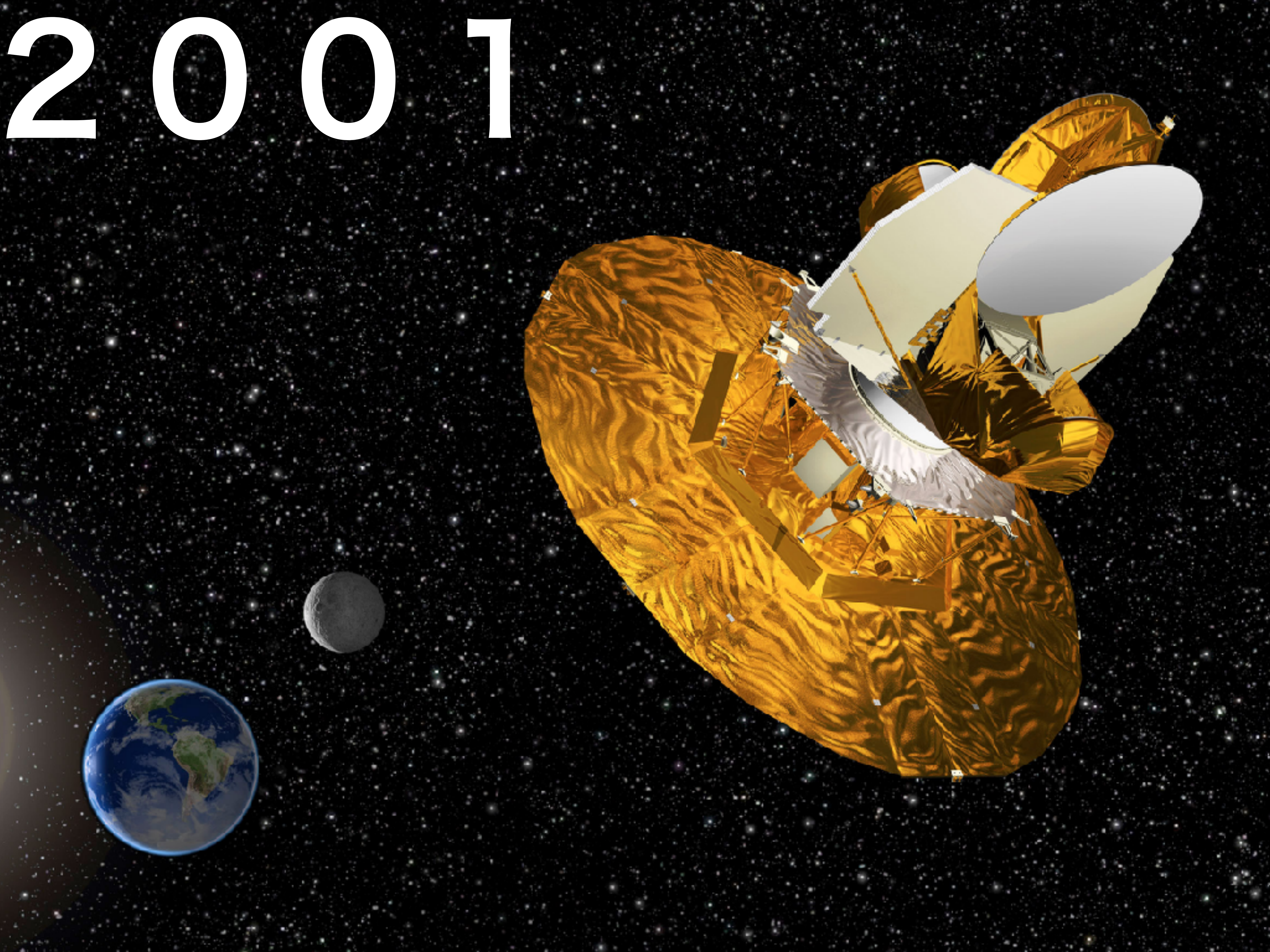
Schreiberaufzeichnung der ersten Messung des Mikrowellenhintergrundes am 20.5.1964

Recording of the first measurement of cosmic microwave background radiation taken on 5/20/1964.



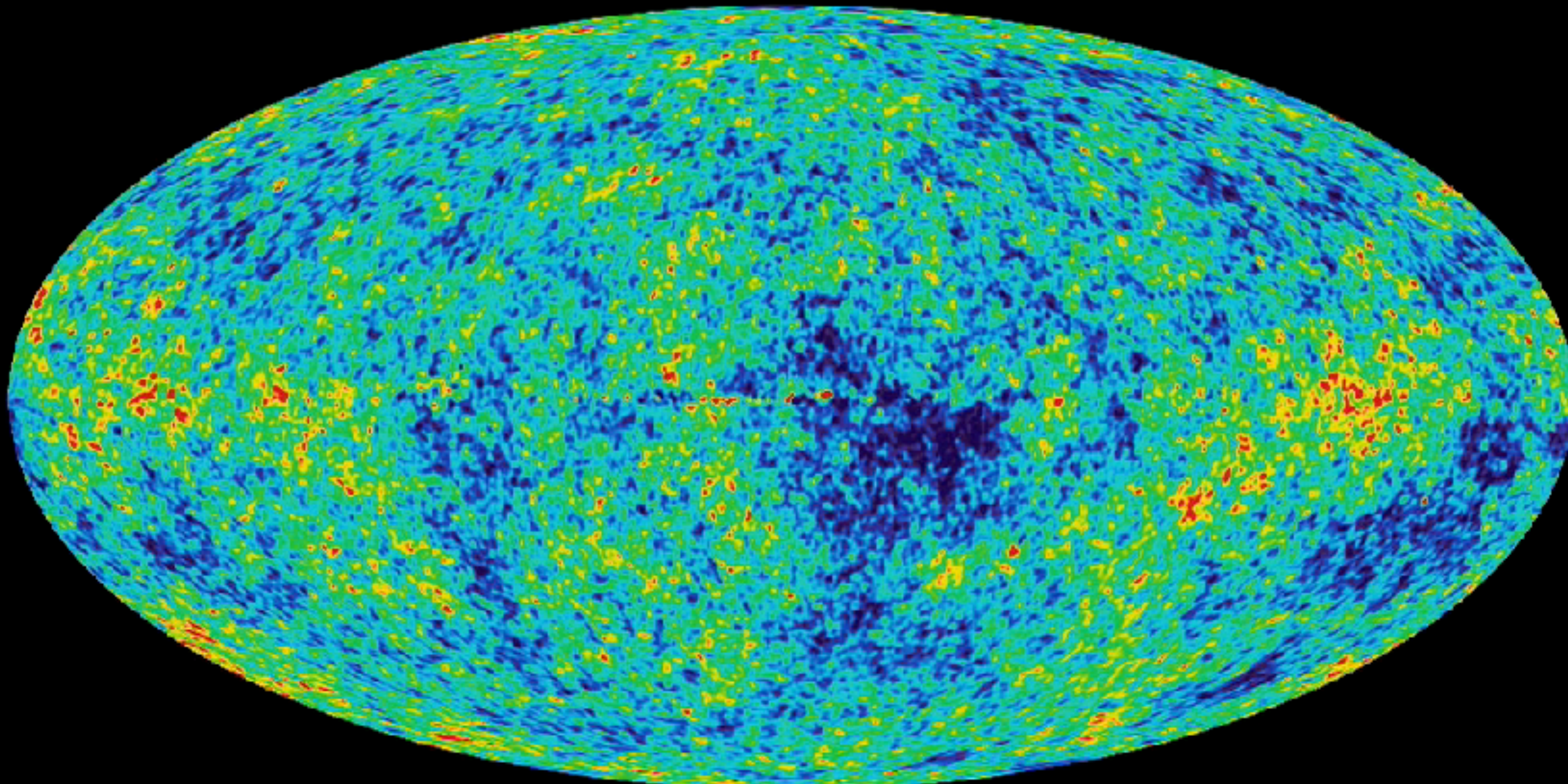
Brightness
4K Planck Spectrum
2.725K Planck Spectrum
2K Planck Spectrum
Rocket (COBRA)
Satellite (COBE/FIRAS)
Rotational Excitation of CN
Ground-based
Balloon-borne
Satellite (COBE/DMR)
10^-17
10^-18
10^-19
10^-20
10^-21
10^-22
Spectrum of CMB
= Planck Spectrum
3m
30cm
3mm
0.3mm
Wavelength




2001

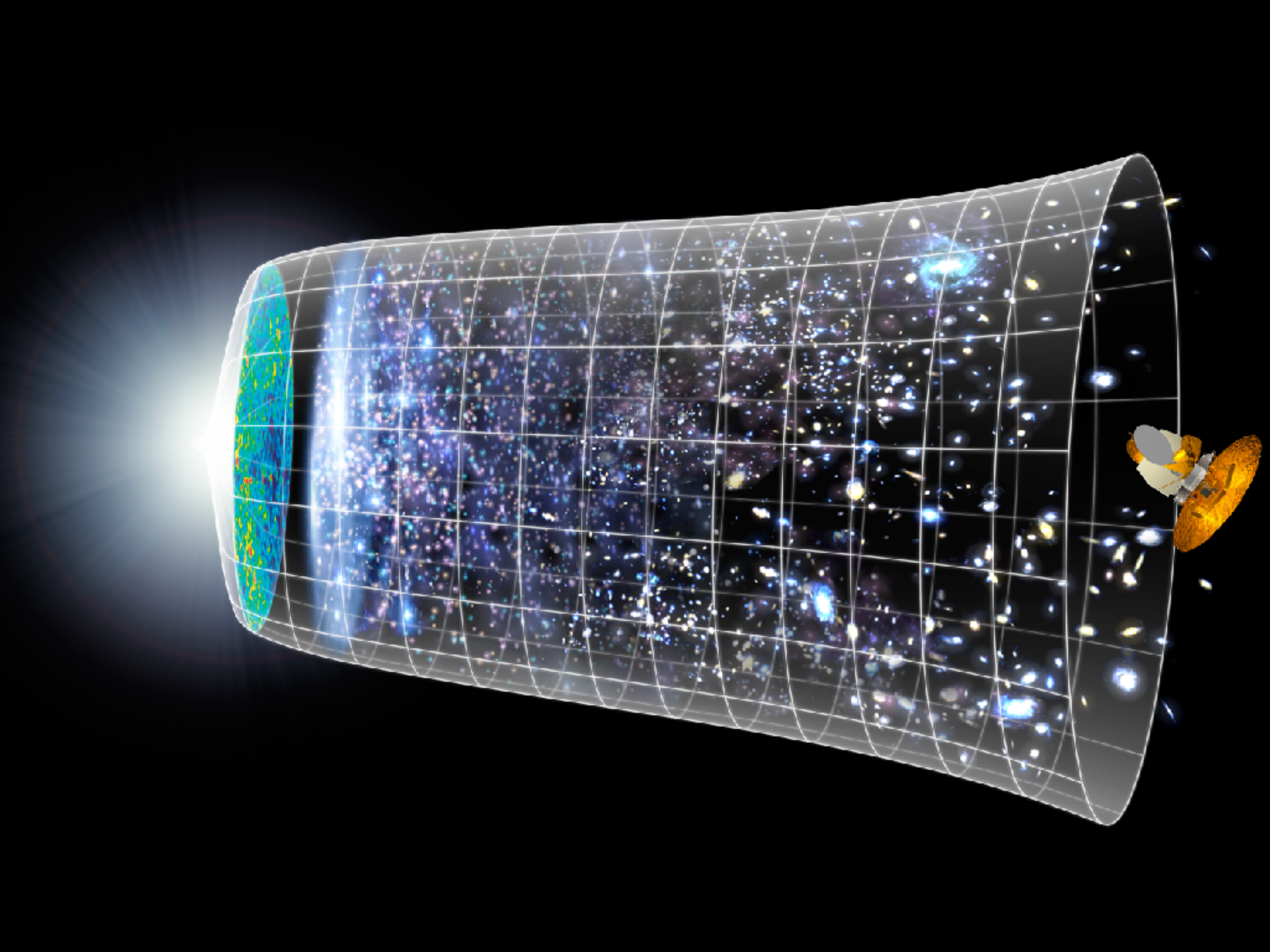
WMAP Science Team
July 19, 2002

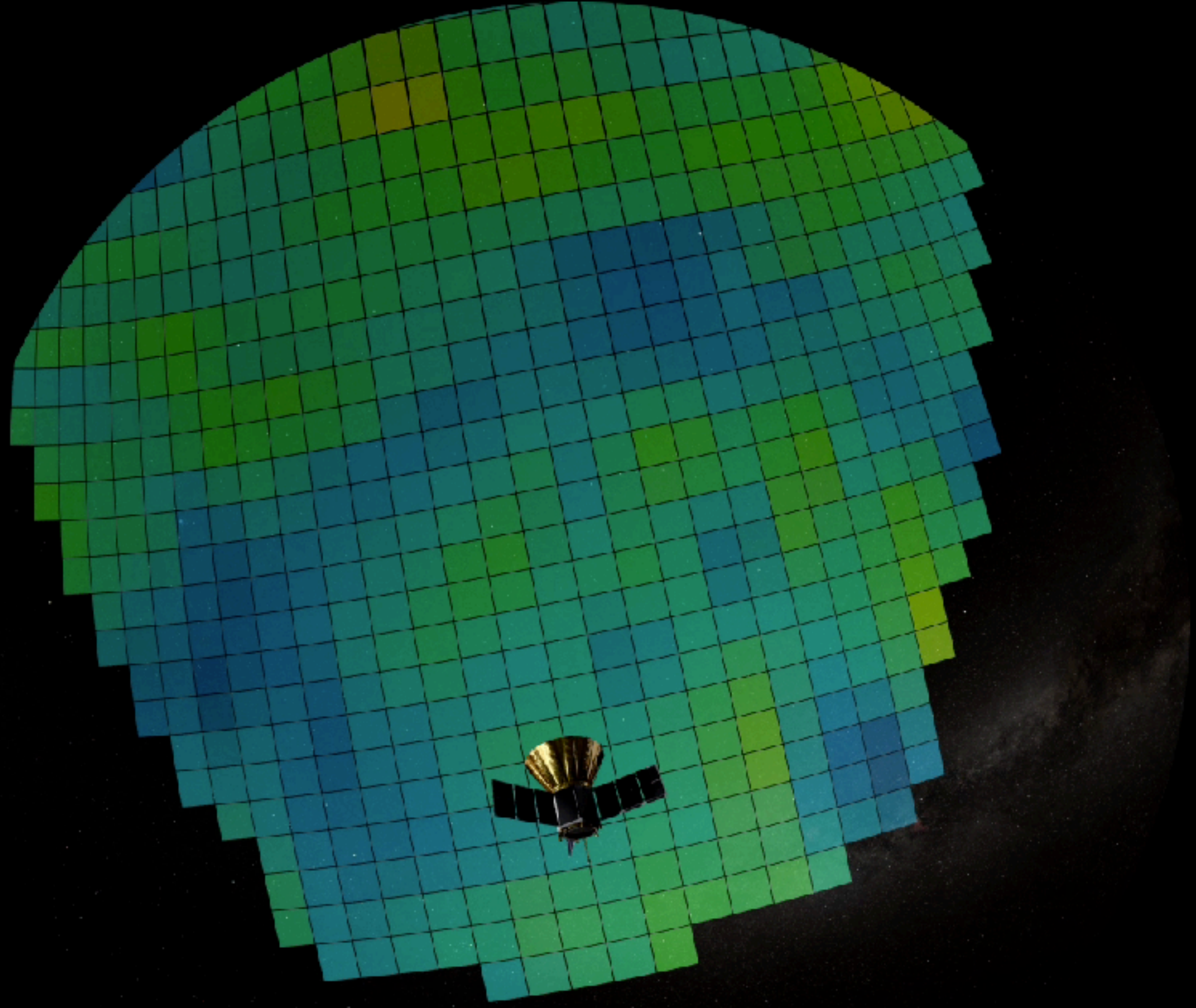
- WMAP was launched on June 30, 2001
- The WMAP mission ended after 9 years of operation



A Remarkable Story

- Observations of the cosmic microwave background and their interpretation taught us that **galaxies, stars, planets, and ourselves originated from tiny fluctuations in the early Universe**
- *But,* ***what generated the initial fluctuations?***





Mukhanov & Chibisov (1981); Hawking (1982); Starobinsky (1982); Guth & Pi (1982); Bardeen, Turner & Steinhardt (1983)

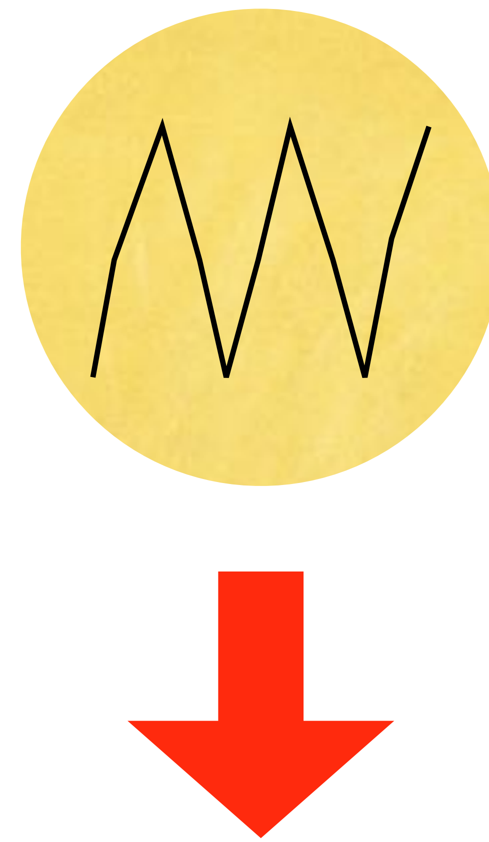
Leading Idea

- Quantum mechanics at work in the early Universe
  - "*We all came from quantum fluctuations*"
- But, how did quantum fluctuations on the *microscopic* scales become *macroscopic* fluctuations over large distances?
  - What is the **missing link** between small and large scales?

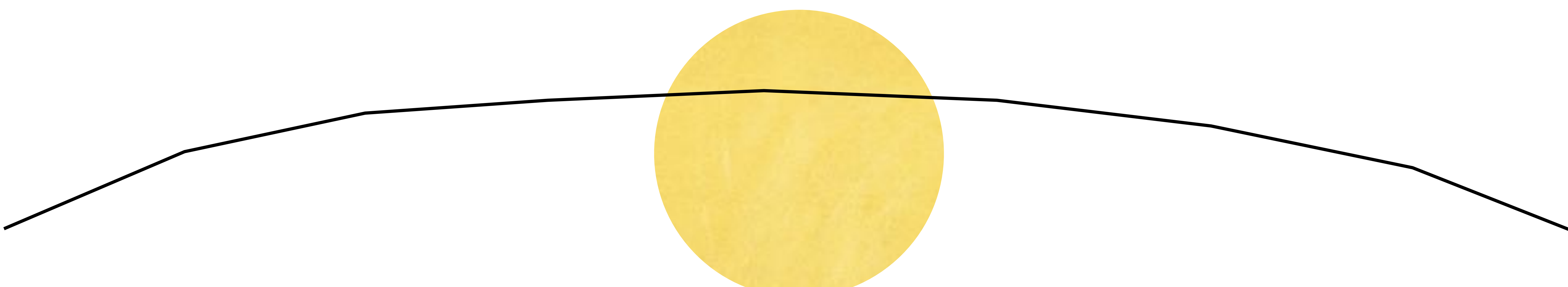
Starobinsky (1980); Sato (1981); Guth (1981); Linde (1982); Albrecht & Steinhardt (1982)

Cosmic Inflation

Quantum fluctuations on microscopic scales



Inflation!



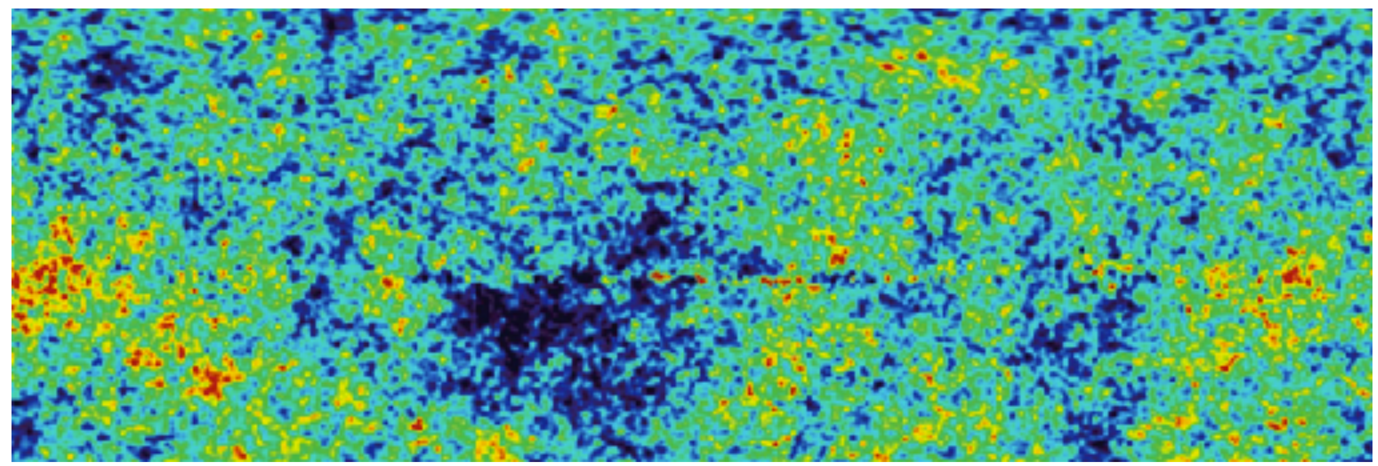
- Exponential expansion (inflation) stretches the wavelength of quantum fluctuations to cosmological scales

Key Predictions

ζ

scalar mode

- Fluctuations we observe today in CMB and the matter distribution originate from quantum fluctuations during inflation



h_{ij}

tensor mode

- There should also be *ultra long-wavelength* gravitational waves generated during inflation

Starobinsky (1979)

We measure distortions in space

- A distance between two points in space

$$d\ell^2 = a^2(t)[1+2\zeta(\mathbf{x},t)][\delta_{ij}+h_{ij}(\mathbf{x},t)]dx^i dx^j$$

- **ζ** : "curvature perturbation" (scalar mode)
  - Perturbation to the determinant of the spatial metric
- **h_{ij}** : "gravitational waves" (tensor mode)
  - Perturbation that does not alter the determinant

$$\sum_i h_{ii} = 0$$

We measure distortions in space

- A distance between two points in space

$$d\ell^2 = a^2(t)[1 + 2\zeta(\mathbf{x}, t)][\delta_{ij} + h_{ij}(\mathbf{x}, t)]dx^i dx^j$$

scale factor

- **ζ** : “curvature perturbation” (scalar mode)
  - Perturbation to the determinant of the spatial metric
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  - Perturbation that does not alter the determinant

$$\sum_i h_{ii} = 0$$

Finding Inflation

- Inflation is the accelerated, quasi-exponential expansion. Defining the Hubble expansion rate as **H(t)=dln(a)/dt**, we must find

$$\frac{\ddot{a}}{a} = \dot{H} + H^2 > 0 \quad \longrightarrow \quad \epsilon \equiv -\frac{\dot{H}}{H^2} < 1$$

- For inflation to explain flatness of spatial geometry of our observable Universe, we need to have a **sustained** period of inflation. This implies $\varepsilon = O(N^{-1})$ or smaller, where N is the number of e-folds of expansion counted from the end of inflation:

$$N \equiv \ln \frac{a_{\rm end}}{a} = \int_t^{t_{\rm end}} dt' \; H(t') \approx 50$$

Have we found inflation?

$$\epsilon \equiv -\frac{\dot{H}}{H^2}$$

- *Have we found ε << 1?*

- To achieve this, we need to map out **H(t)**, and show that it does not change very much with time

Fluctuations are proportional to H

- **Both scalar (ζ) and tensor (h_{ij}) perturbations are proportional to H**
  - Consequence of the uncertainty principle
    - [energy you can borrow] ~ [time you borrow]$^{-1}$ ~ H
  - **THE KEY**: The earlier the fluctuations are generated, the more its wavelength is stretched, and thus the bigger the angles they subtend in the sky. **We can map H(t) by measuring CMB fluctuations over a wide range of angles**

Fluctuations are proportional to H

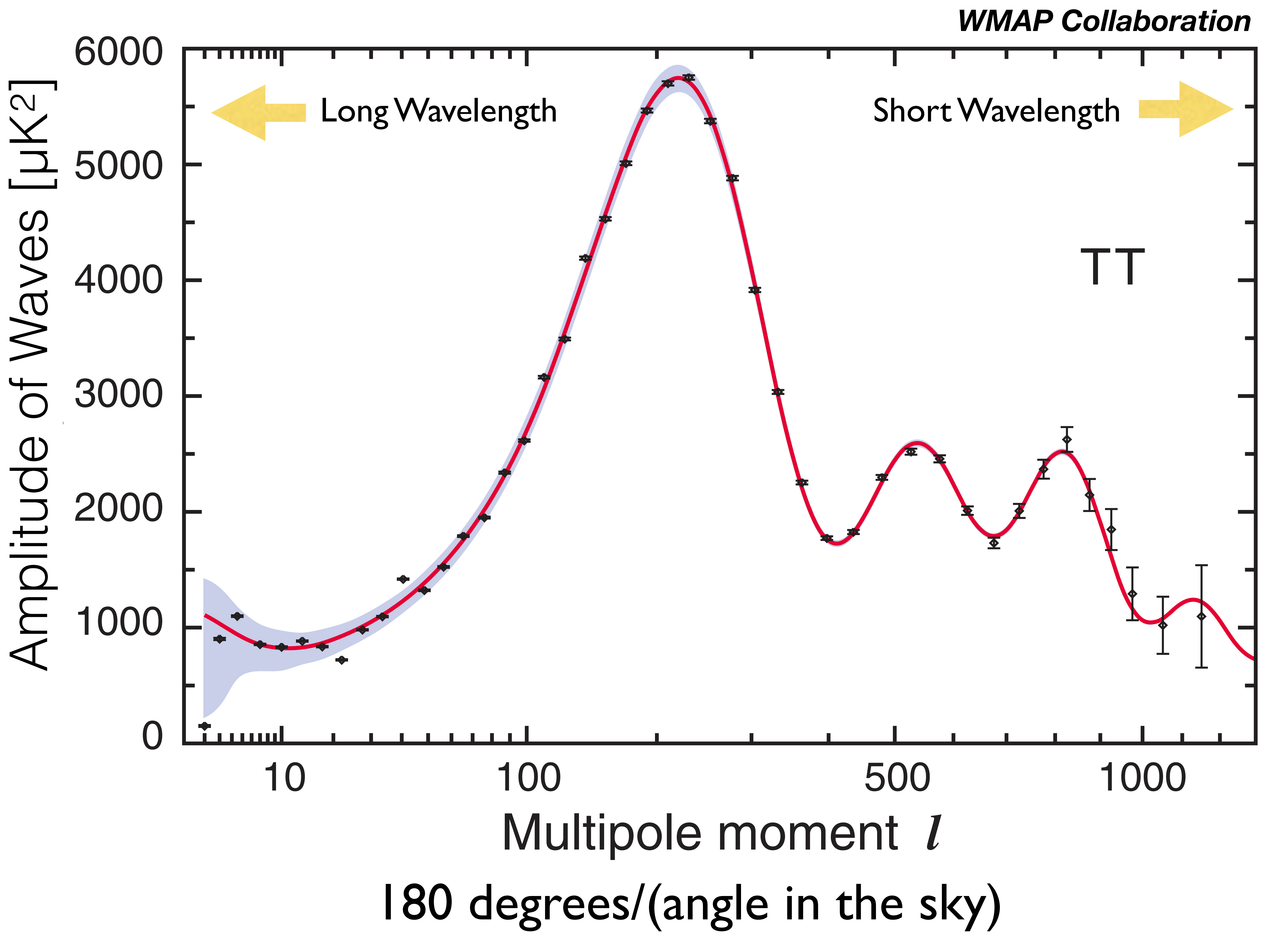
- **We can map H(t) by measuring CMB fluctuations over a wide range of angles**

1. We want to show that the amplitude of CMB fluctuations does not depend very much on angles

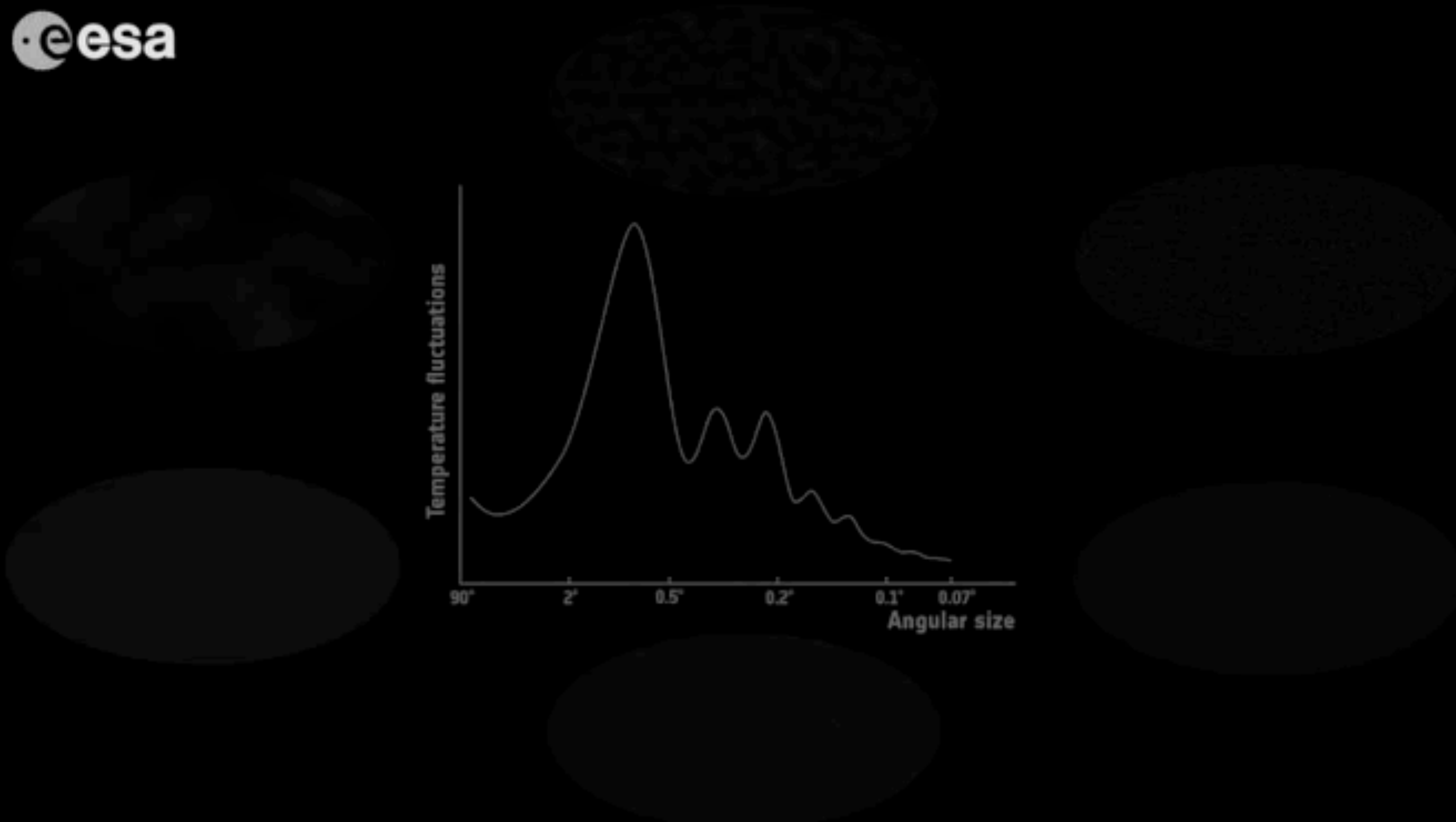
2. Moreover, since inflation must end, H would be a decreasing function of time. It would be fantastic to show that the amplitude of CMB fluctuations actually DOES depend on angles such that the small scale has ***slightly*** smaller power

Data Analysis

- Decompose temperature fluctuations in the sky into a set of waves with various wavelengths
- Make a diagram showing the strength of each wavelength

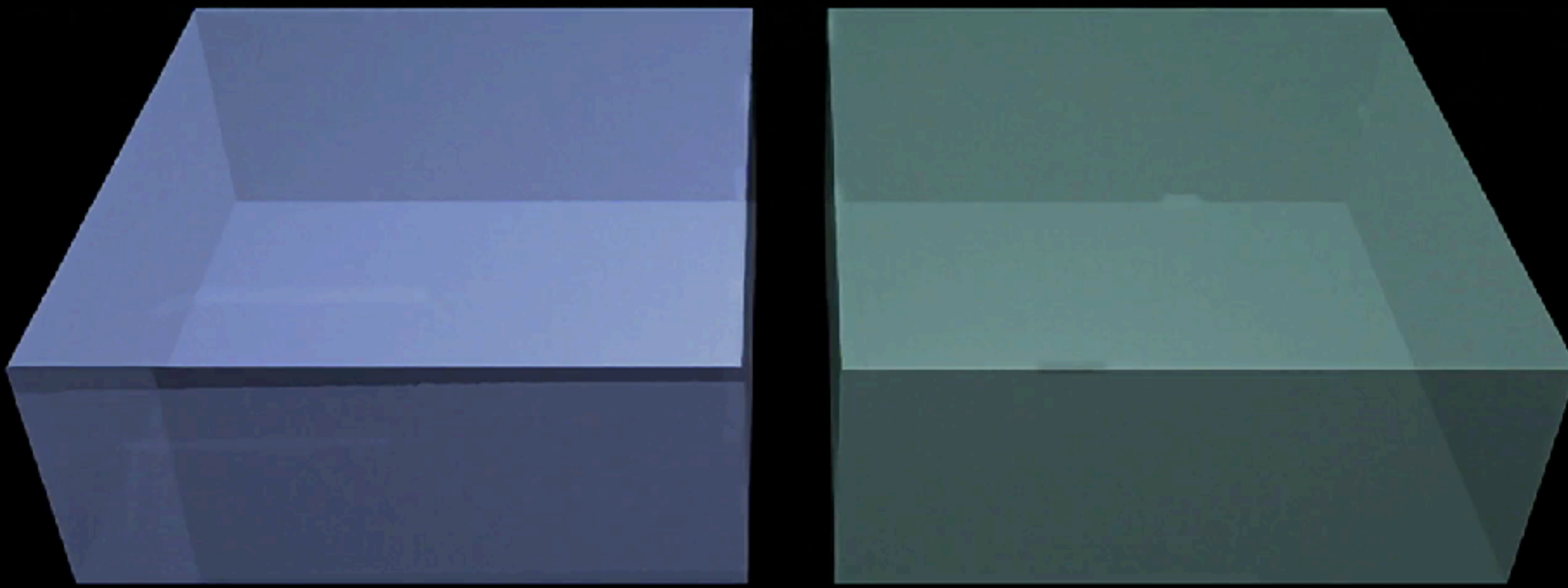

WMAP Collaboration
Long Wavelength
Short Wavelength
TT
Amplitude of Waves [μK2]
6000
5000
4000
3000
2000
1000
0
10
100
500
1000
Multipole moment l
180 degrees/(angle in the sky)

Power spectrum, explained

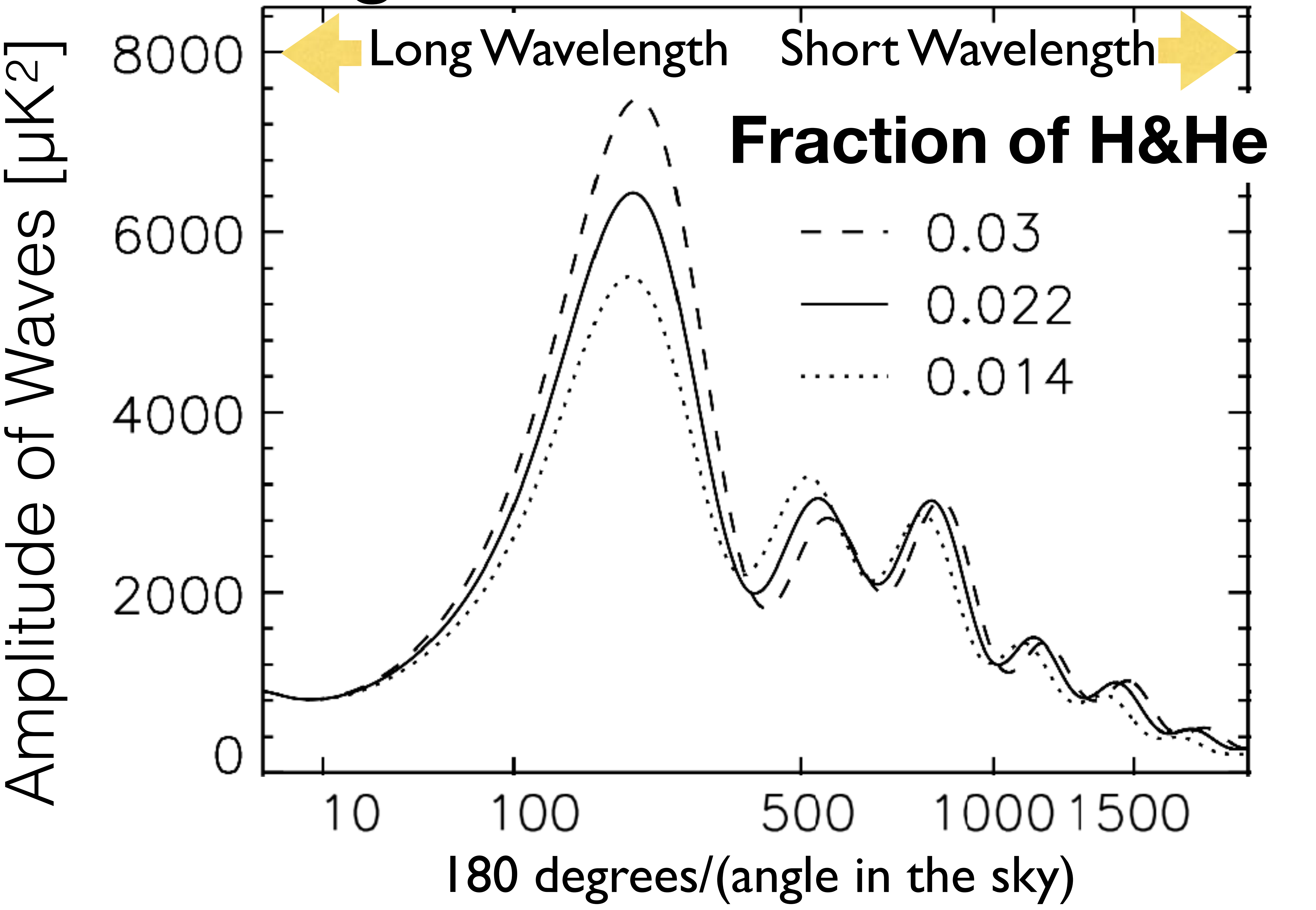


La Soupe Miso Cosmique

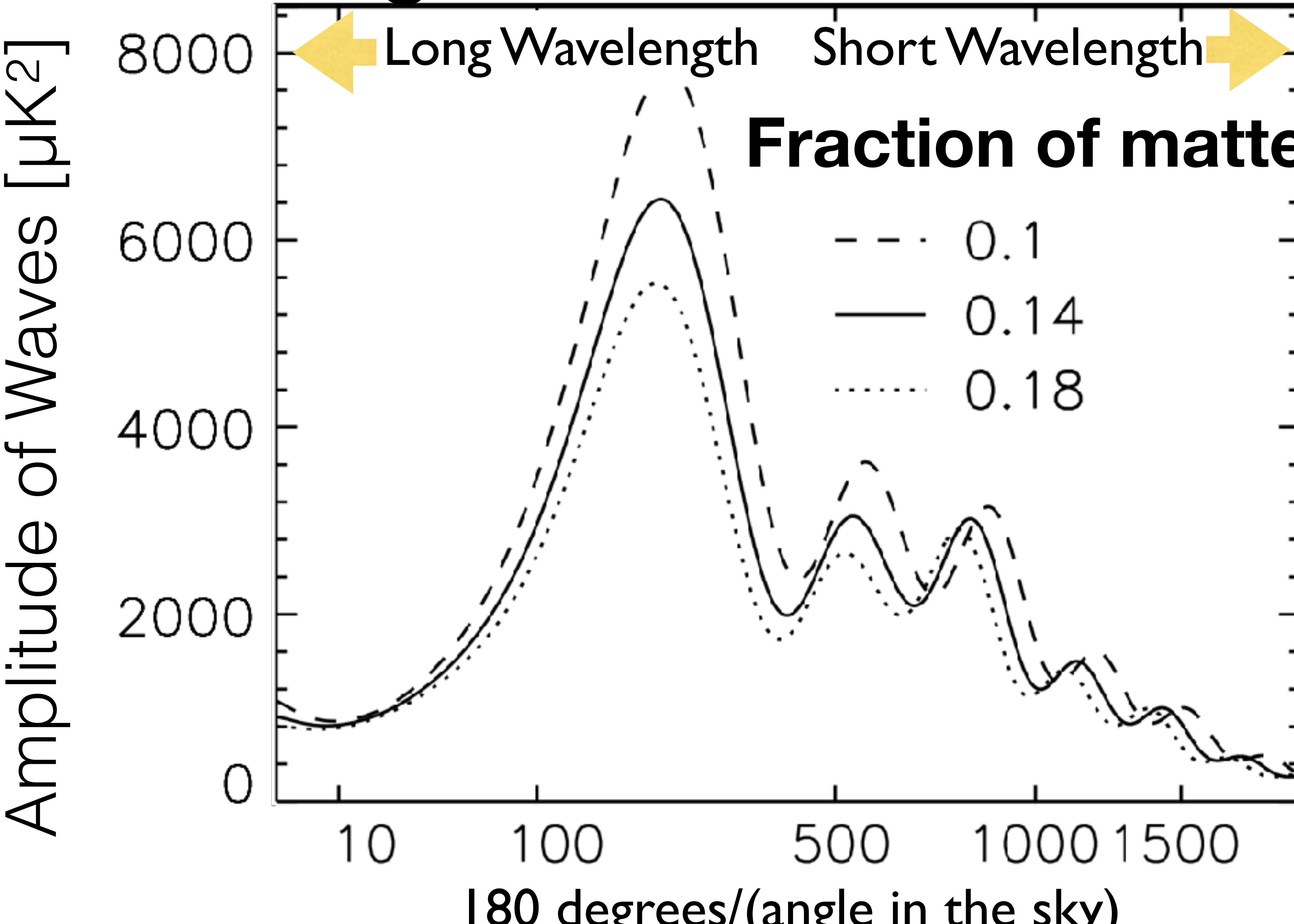
- When matter and radiation were hotter than 3000 K, matter was completely ionised. The Universe was filled with plasma, which behaves just like a soup
- Think about a Miso soup (if you know what it is). Imagine throwing Tofus into a Miso soup, while changing the density of Miso
- And imagine watching how ripples are created and propagate throughout the soup



Measuring Abundance of H&He

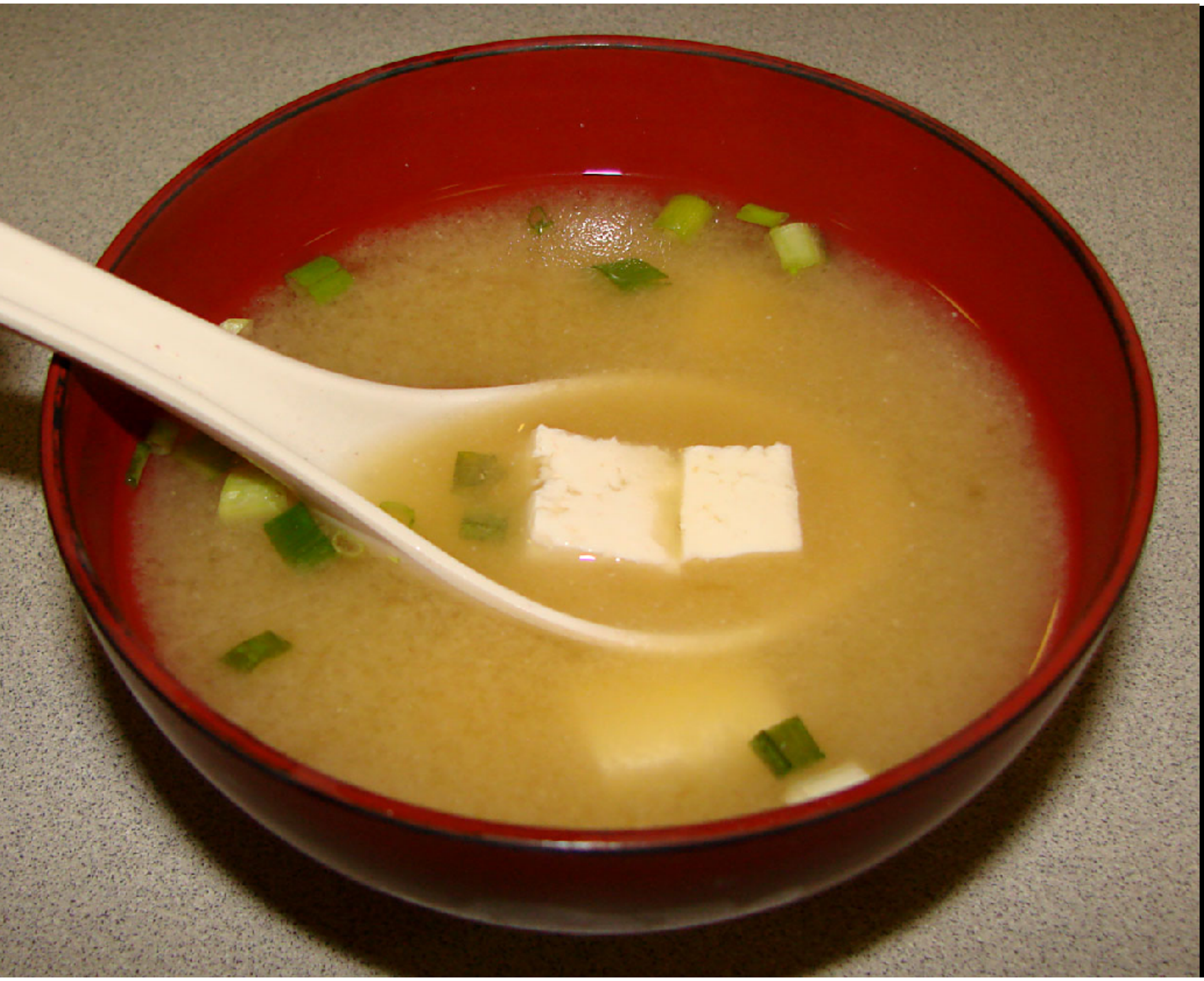


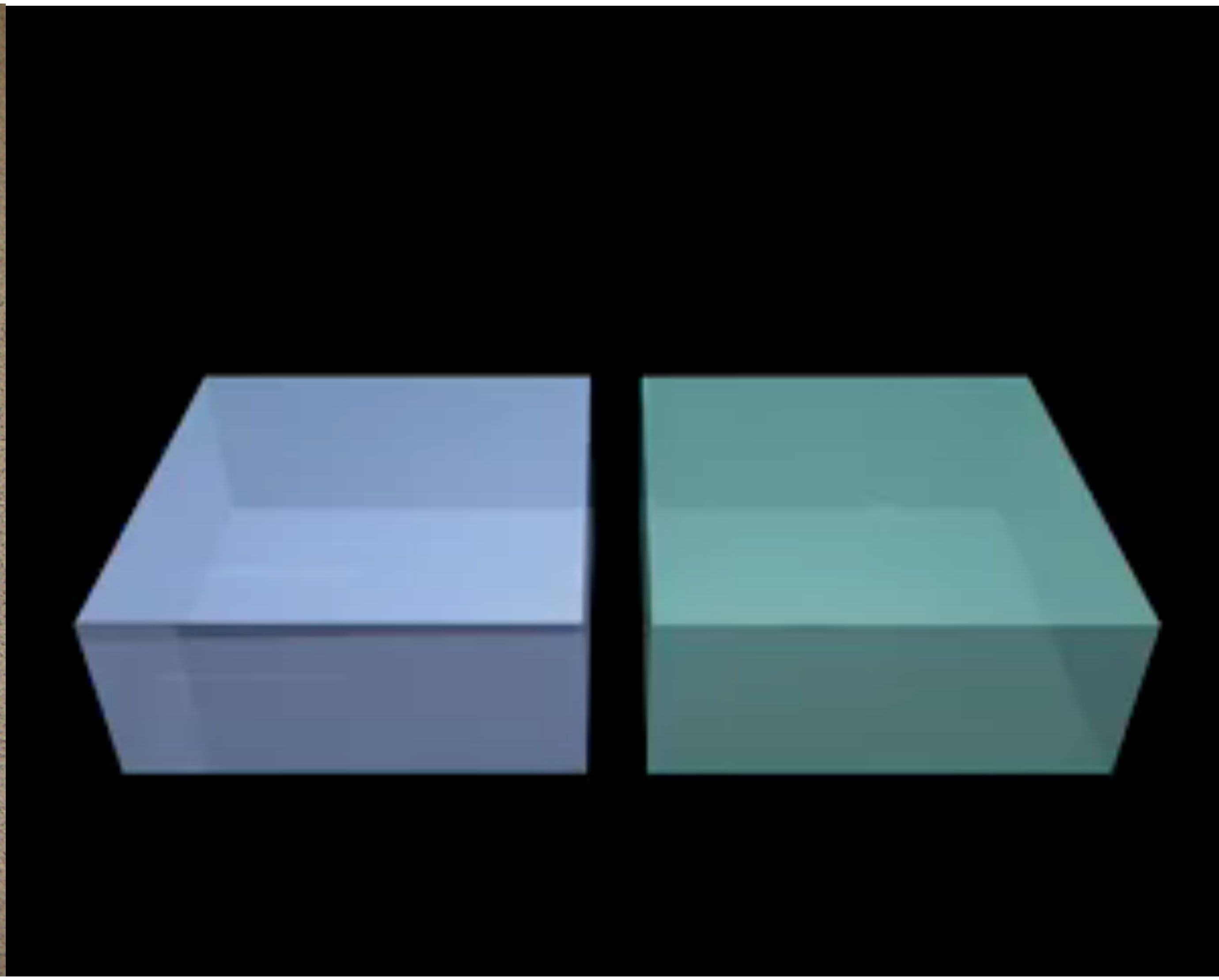
Measuring ***Total*** Matter Density

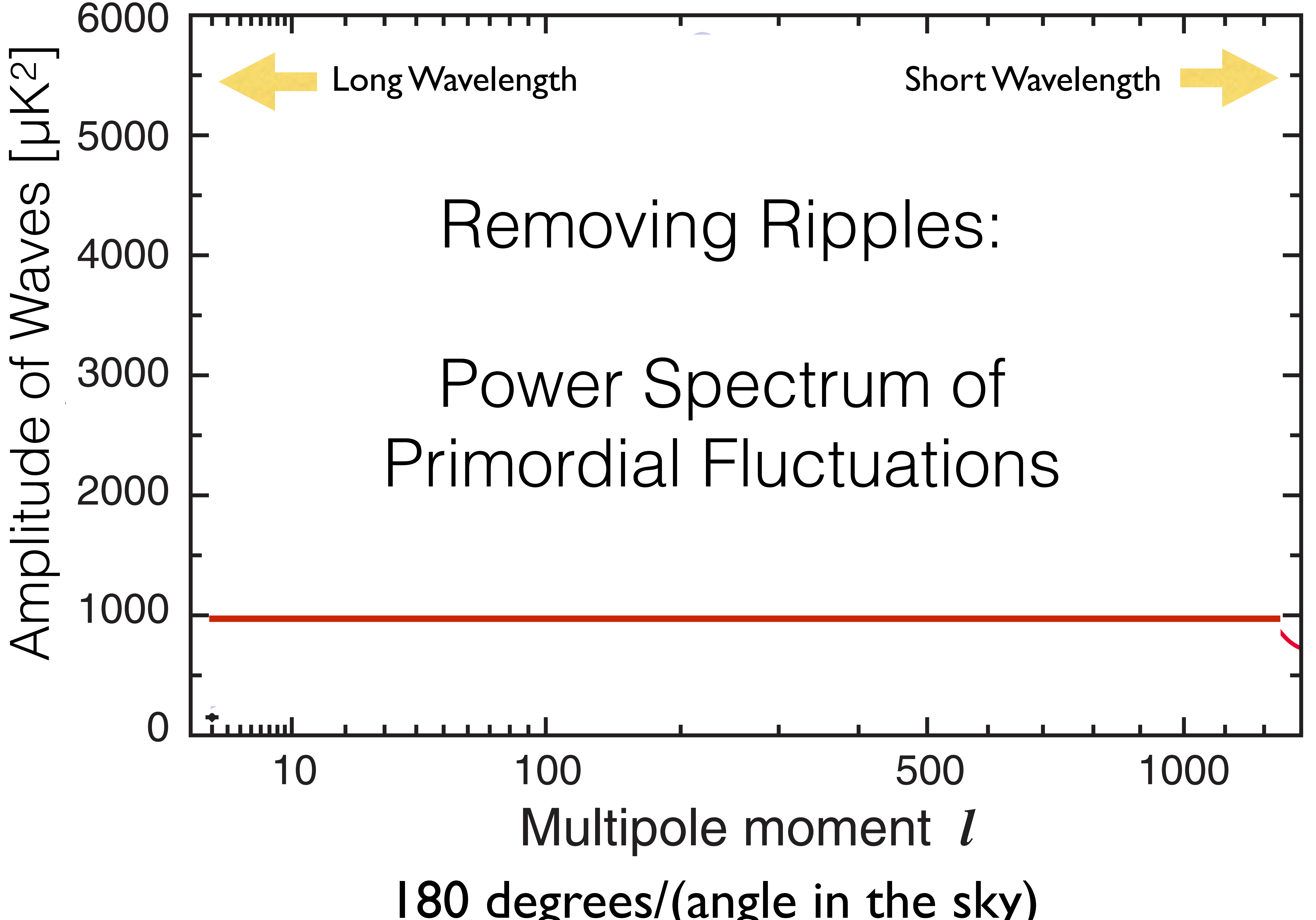


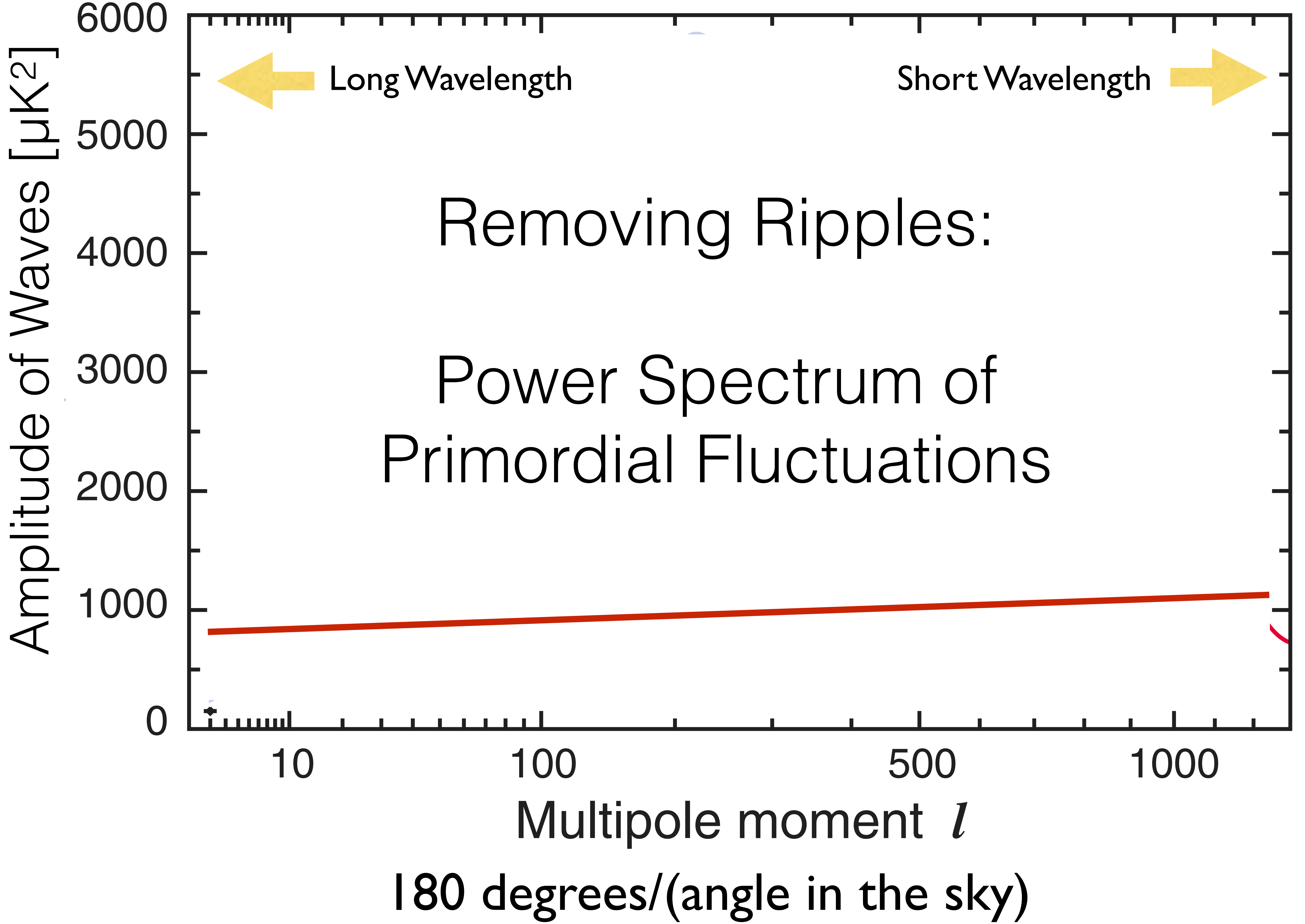
Origin of Fluctuations

- Who dropped those Tofus into the cosmic Miso soup?






Removing Ripples:
Power Spectrum of Primordial Fluctuations
Long Wavelength
Short Wavelength
Amplitude of Waves [μK²]
6000
5000
4000
3000
2000
1000
0
10
100
500
1000
Multipole moment l
180 degrees/(angle in the sky)


Long Wavelength
Short Wavelength
Removing Ripples:
Power Spectrum of
Primordial Fluctuations
Amplitude of Waves [μK²]
6000
5000
4000
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2000
1000
0
10
100
500
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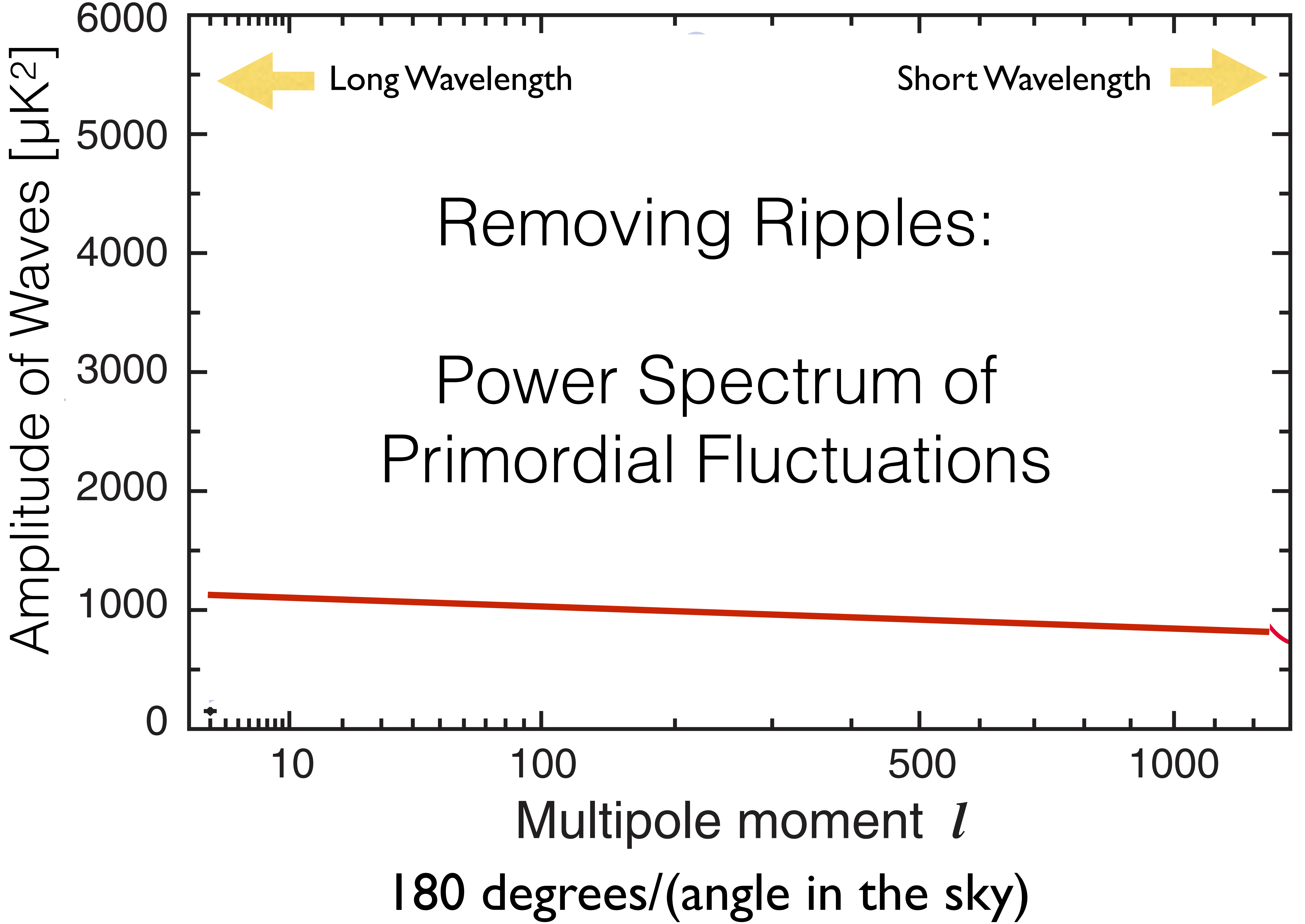
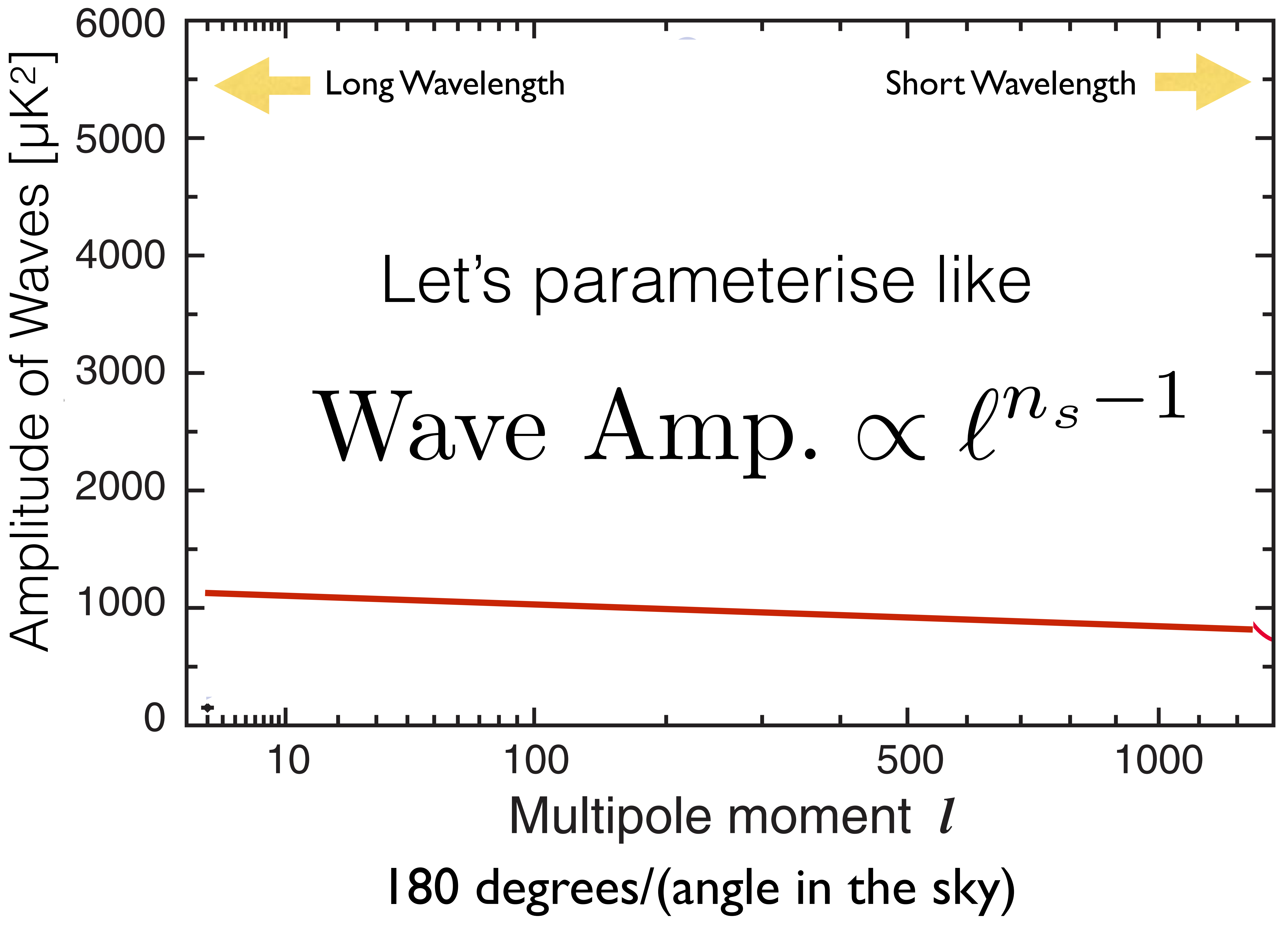
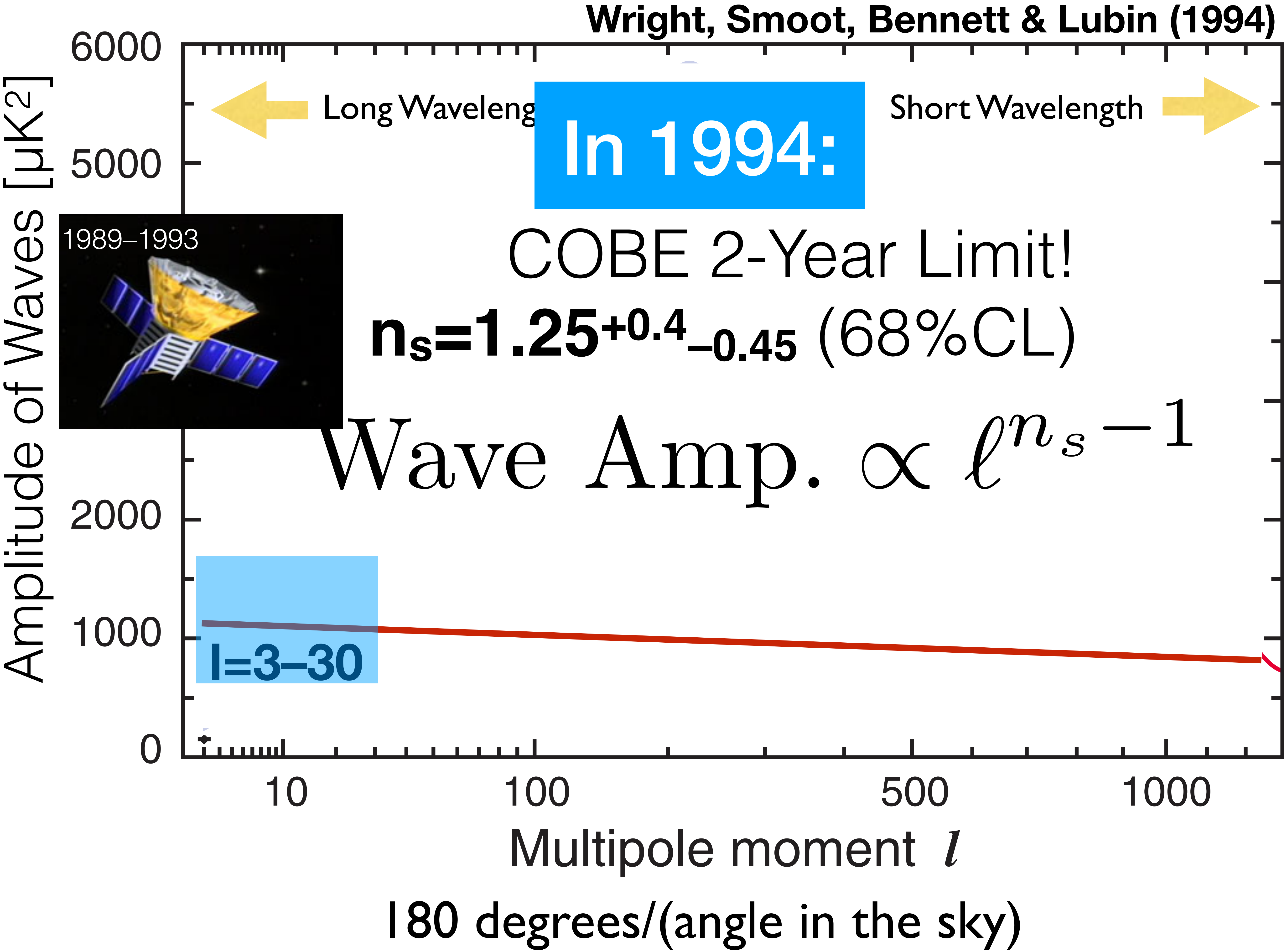
Multipole moment l
180 degrees/(angle in the sky)



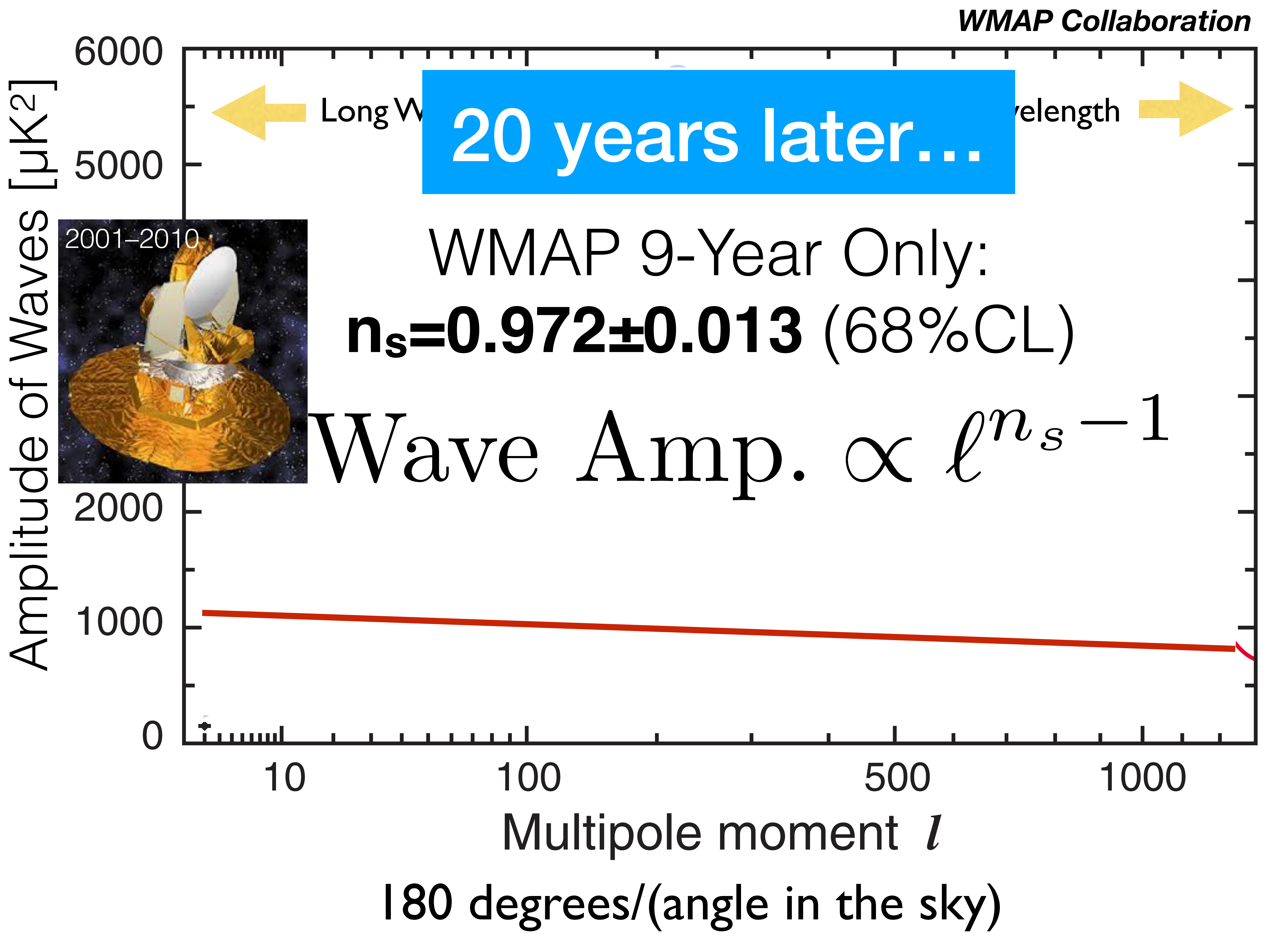
Long Wavelength
Short Wavelength
Removing Ripples:
Power Spectrum of
Primordial Fluctuations
Amplitude of Waves [μK²]
6000
5000
4000
3000
2000
1000
0
10
100
500
1000
Multipole moment l
180 degrees/(angle in the sky)

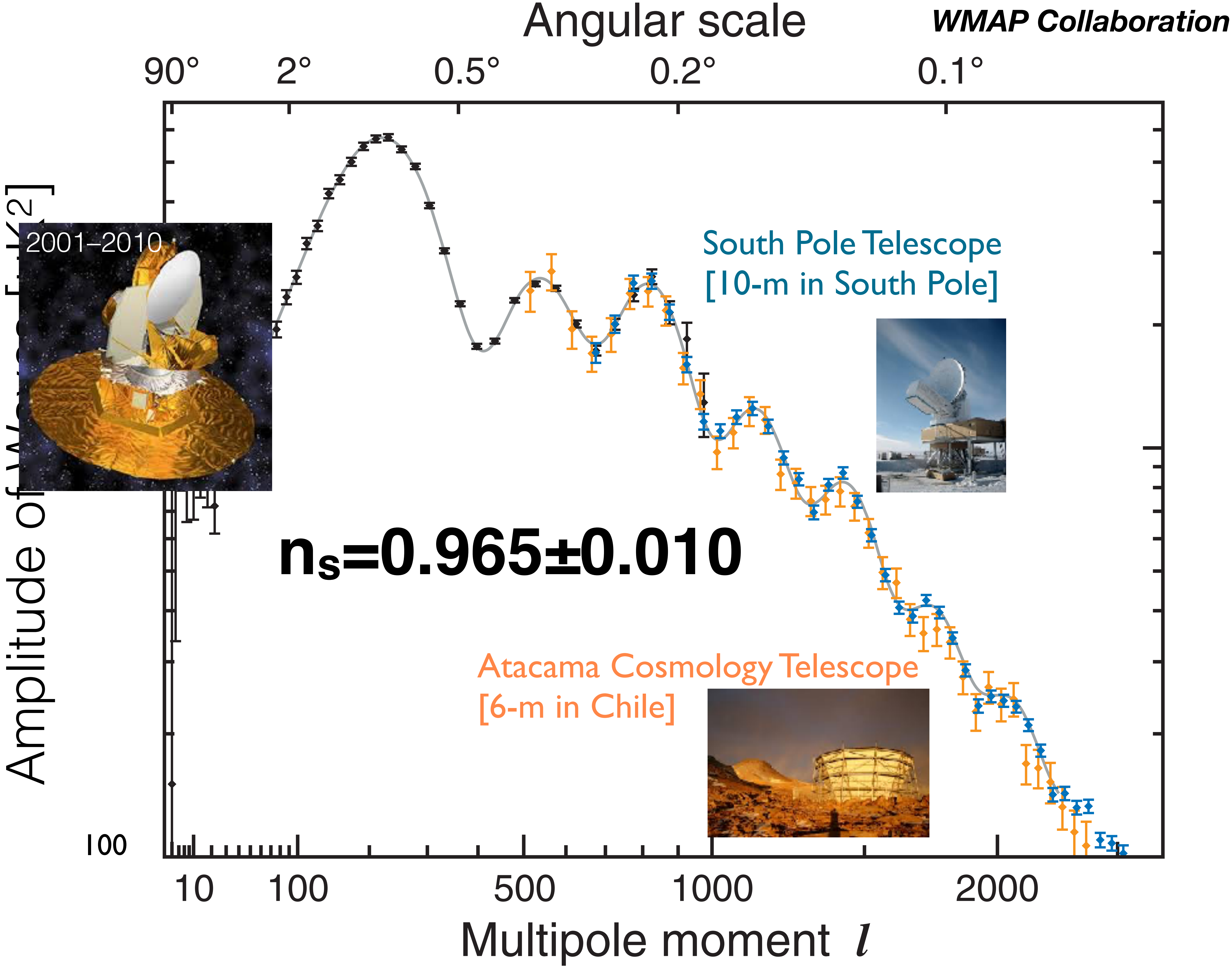


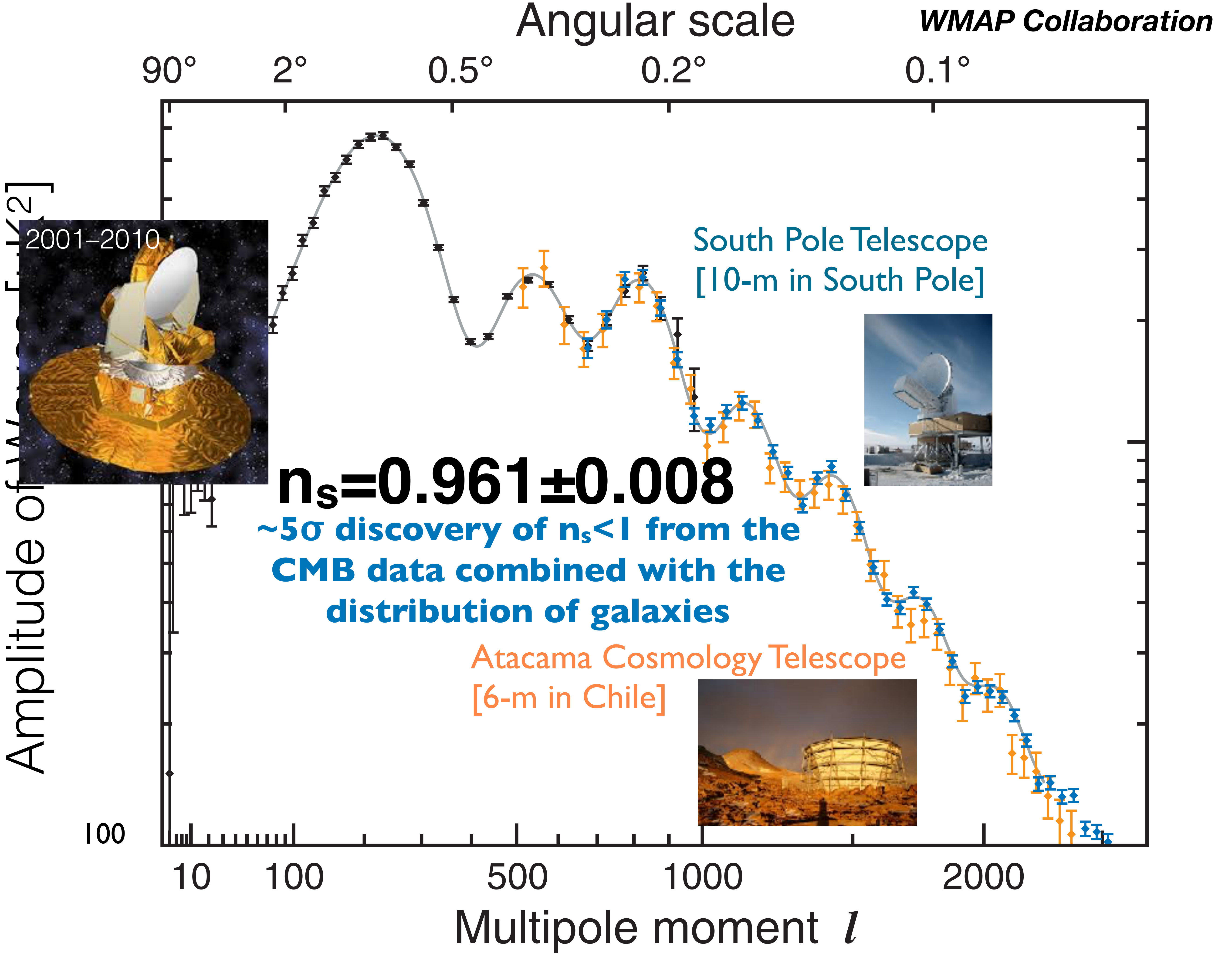
6000
5000
4000
3000
2000
1000
0
Amplitude of Waves [μK²]
Long Wavelength
Short Wavelength
Let's parameterise like
Wave Amp. ∝ ℓ^(n_s−1)
10
100
500
1000
Multipole moment l
180 degrees/(angle in the sky)



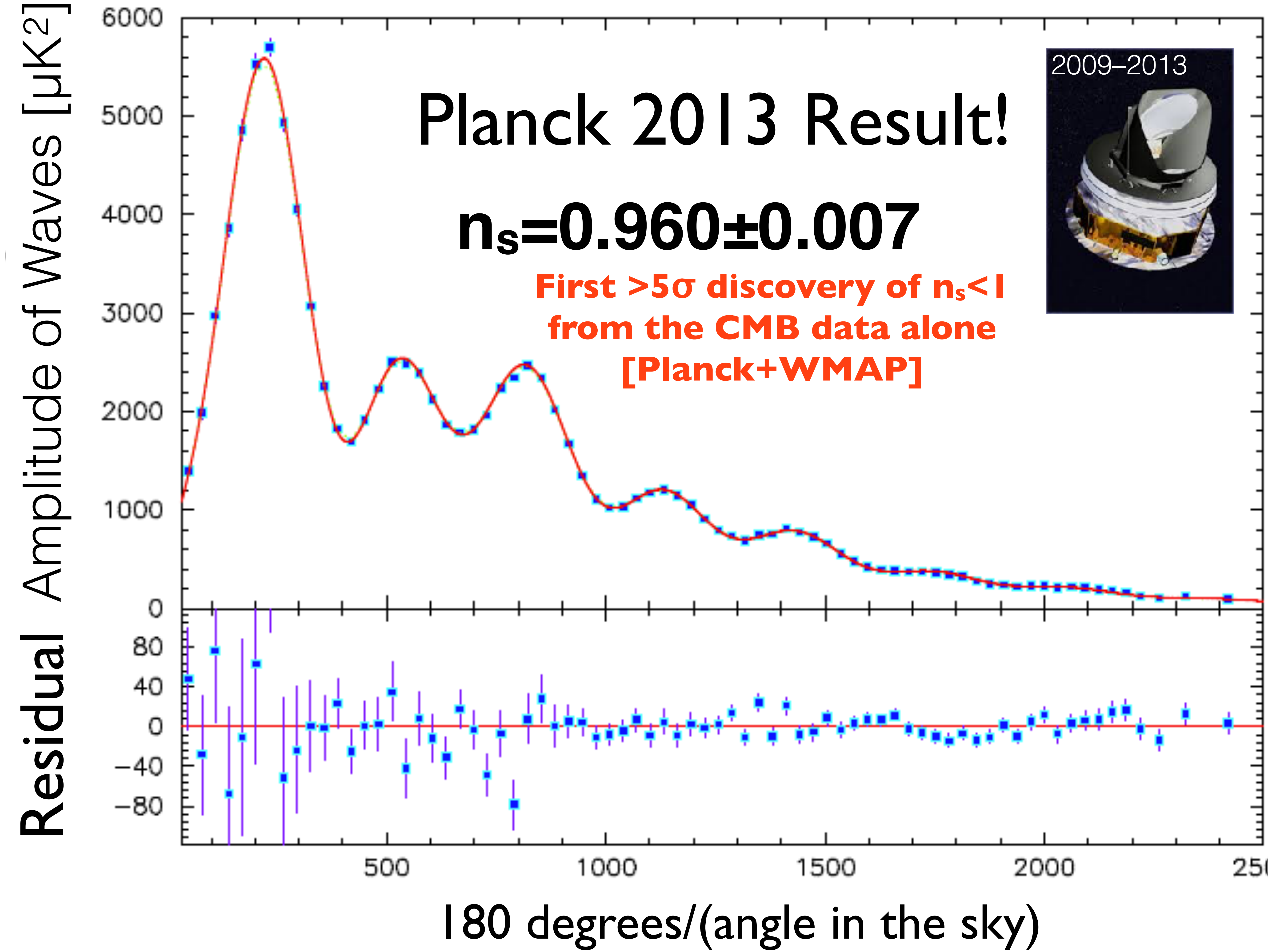
Wright, Smoot, Bennett & Lubin (1994)
In 1994:
COBE 2-Year Limit!
$n_s=1.25^{+0.4}_{-0.45}$ (68%CL)
$\text{Wave Amp.} \propto \ell^{n_s-1}$
1989–1993
Long Waveleng
Short Wavelength
l=3–30
Amplitude of Waves [μK²]
6000
5000
2000
1000
0
10
100
500
1000
Multipole moment l
180 degrees/(angle in the sky)


WMAP Collaboration
20 years later…
WMAP 9-Year Only:
n_s=0.972±0.013 (68%CL)
$\text{Wave Amp.} \propto \ell^{n_s - 1}$
2001–2010
Amplitude of Waves [μK²]
6000
5000
2000
1000
0
10
100
500
1000
Multipole moment l
180 degrees/(angle in the sky)

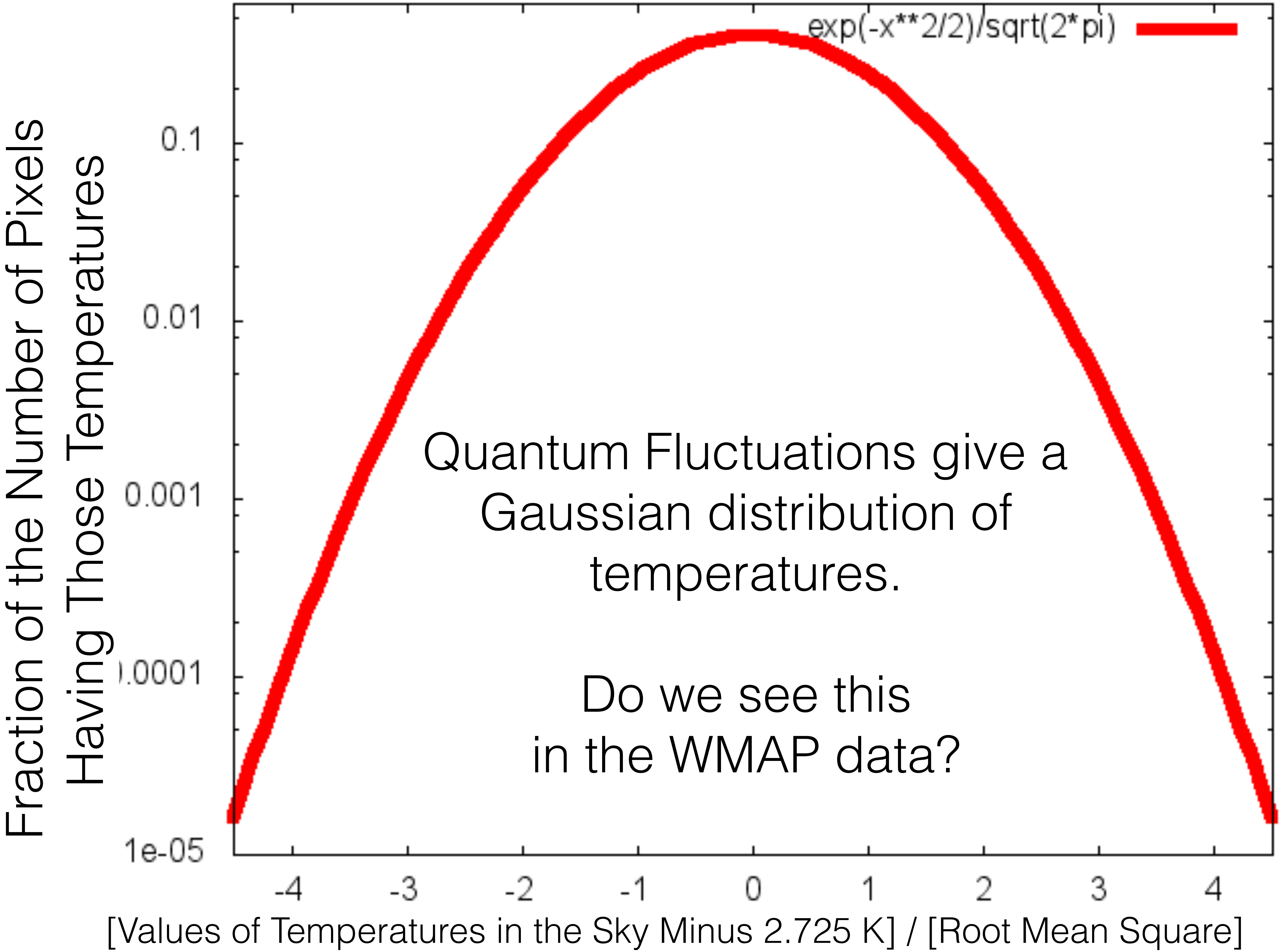

Angular scale
WMAP Collaboration
90°
2°
0.5°
0.2°
0.1°
2001–2010
South Pole Telescope
[10-m in South Pole]
n_s=0.965±0.010
Atacama Cosmology Telescope
[6-m in Chile]
Amplitude of Waves [μK^2]
100
10
100
500
1000
2000
Multipole moment *l*

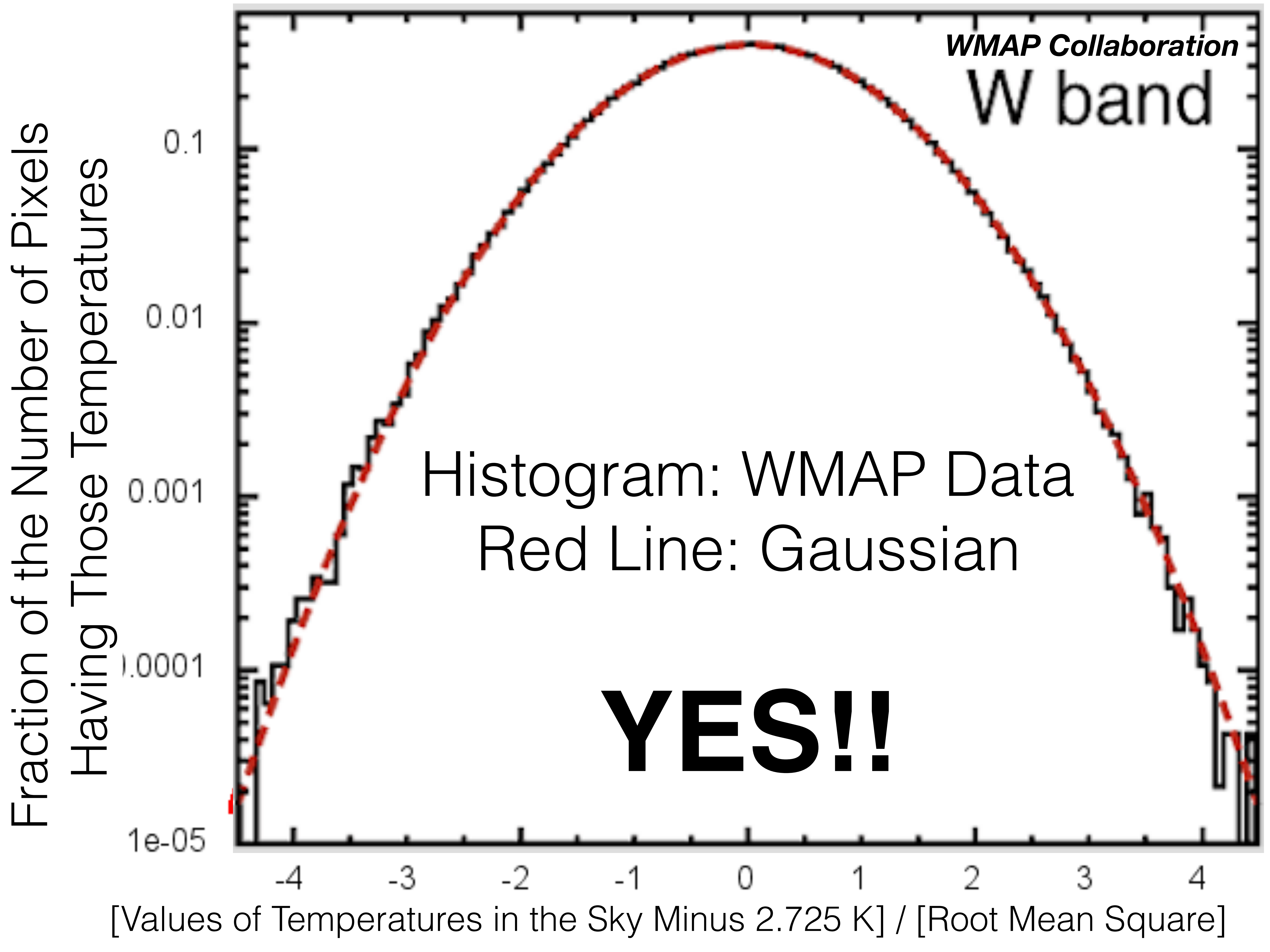


WMAP Collaboration
Angular scale
90°
2°
0.5°
0.2°
0.1°
2001–2010
South Pole Telescope
[10-m in South Pole]
n_s=0.961±0.008
~5σ discovery of n_s<1 from the
CMB data combined with the
distribution of galaxies
Atacama Cosmology Telescope
[6-m in Chile]
Amplitude of
100
10
100
500
1000
2000
Multipole moment l

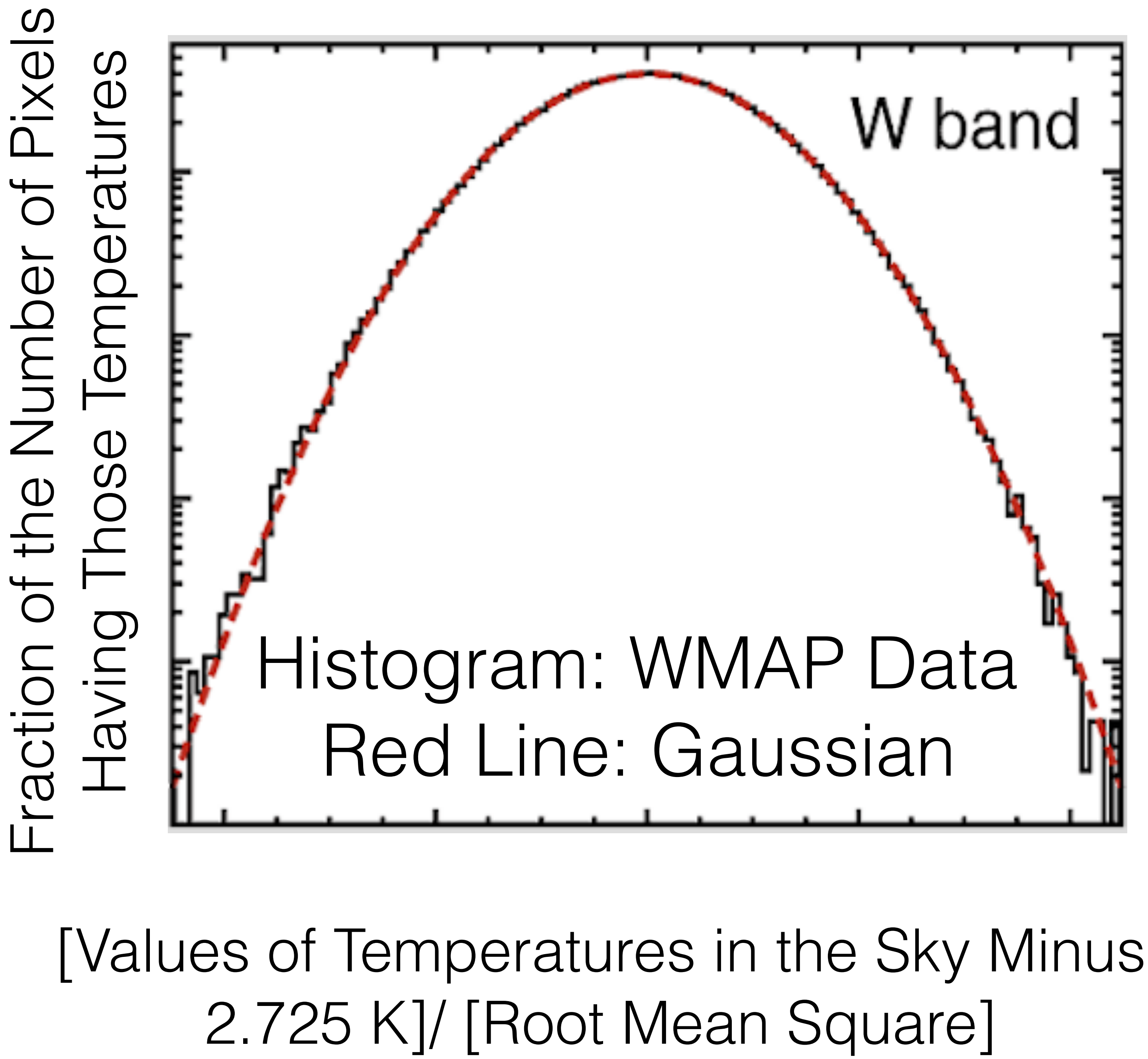


Planck 2013 Result!
n_s=0.960±0.007
First >5σ discovery of n_s<1
from the CMB data alone
[Planck+WMAP]
2009–2013
Amplitude of Waves [μK^2]
Residual
180 degrees/(angle in the sky)
6000
5000
4000
3000
2000
1000
0
80
40
0
−40
−80
500
1000
1500
2000
250


exp(-x**2/2)/sqrt(2*pi)
Quantum Fluctuations give a Gaussian distribution of temperatures.
Do we see this in the WMAP data?
Fraction of the Number of Pixels Having Those Temperatures
0.1
0.01
0.001
0.0001
1e-05
-4
-3
-2
-1
0
1
2
3
4
[Values of Temperatures in the Sky Minus 2.725 K] / [Root Mean Square]


WMAP Collaboration
W band
Fraction of the Number of Pixels
Having Those Temperatures
0.1
0.01
0.001
0.0001
1e-05
Histogram: WMAP Data
Red Line: Gaussian
YES!!
-4
-3
-2
-1
0
1
2
3
4
[Values of Temperatures in the Sky Minus 2.725 K] / [Root Mean Square]

Testing Gaussianity



- Since a Gauss distribution is symmetric, it must yield a vanishing **3-point function**

$$\langle \delta T^3 \rangle \equiv \int_{-\infty}^{\infty} d\delta T \; P(\delta T)\delta T^3$$

- More specifically, we measure this by averaging the product of temperatures at three different locations in the sky

$$\langle \delta T(\hat{n}_1)\delta T(\hat{n}_2)\delta T(\hat{n}_3)\rangle$$

Lack of non-Gaussianity

- The WMAP data show that the distribution of temperature fluctuations of CMB is very precisely Gaussian
  - with an upper bound on a deviation of **0.2%** (95%CL)

$$\zeta(\mathbf{x}) = \zeta_{\rm gaus}(\mathbf{x}) + \frac{3}{5} f_{\rm NL} \zeta_{\rm gaus}^2(\mathbf{x})$$ **with** $$f_{\rm NL} = 37 \pm 20 \ (68\% \ {\rm CL})$$

WMAP 9-year Result

- The Planck data improved the upper bound by an order of magnitude: deviation is <**0.03%** (95%CL)

$$f_{\rm NL} = 0.8 \pm 5.0 \ (68\% \ {\rm CL})$$

Planck 2015 Result

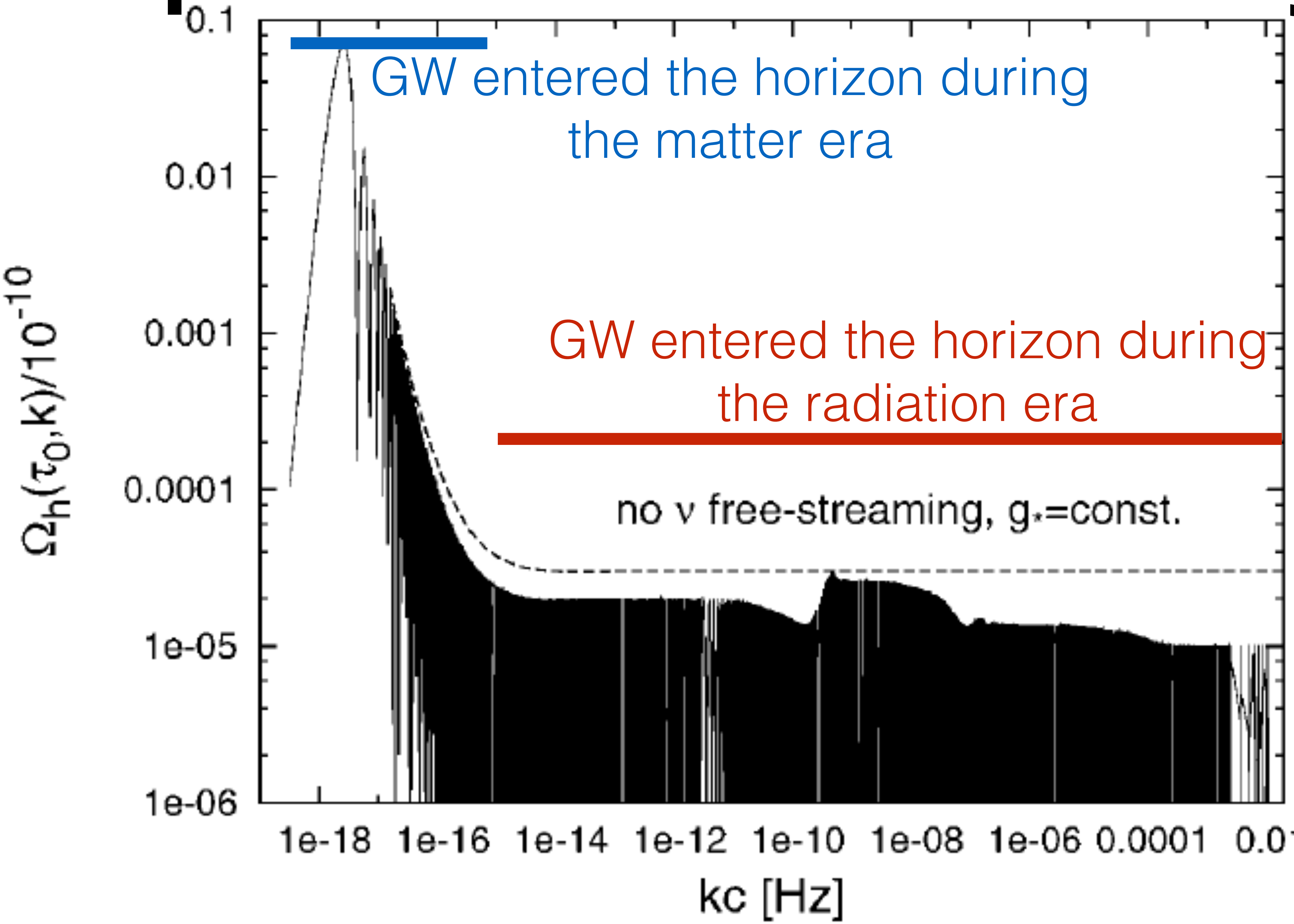
So, have we found inflation?

- Single-field slow-roll inflation looks remarkably good:
  - **Super-horizon fluctuation**
  - **Adiabaticity**
  - **Gaussianity**
  - **n_s<1**
- What more do we want? **Gravitational waves**. Why?
  - Because the “*extraordinary claim requires extraordinary evidence*”

Watanabe & EK (2006)

Theoretical energy density

Spectrum of GW today



Watanabe & EK (2006)

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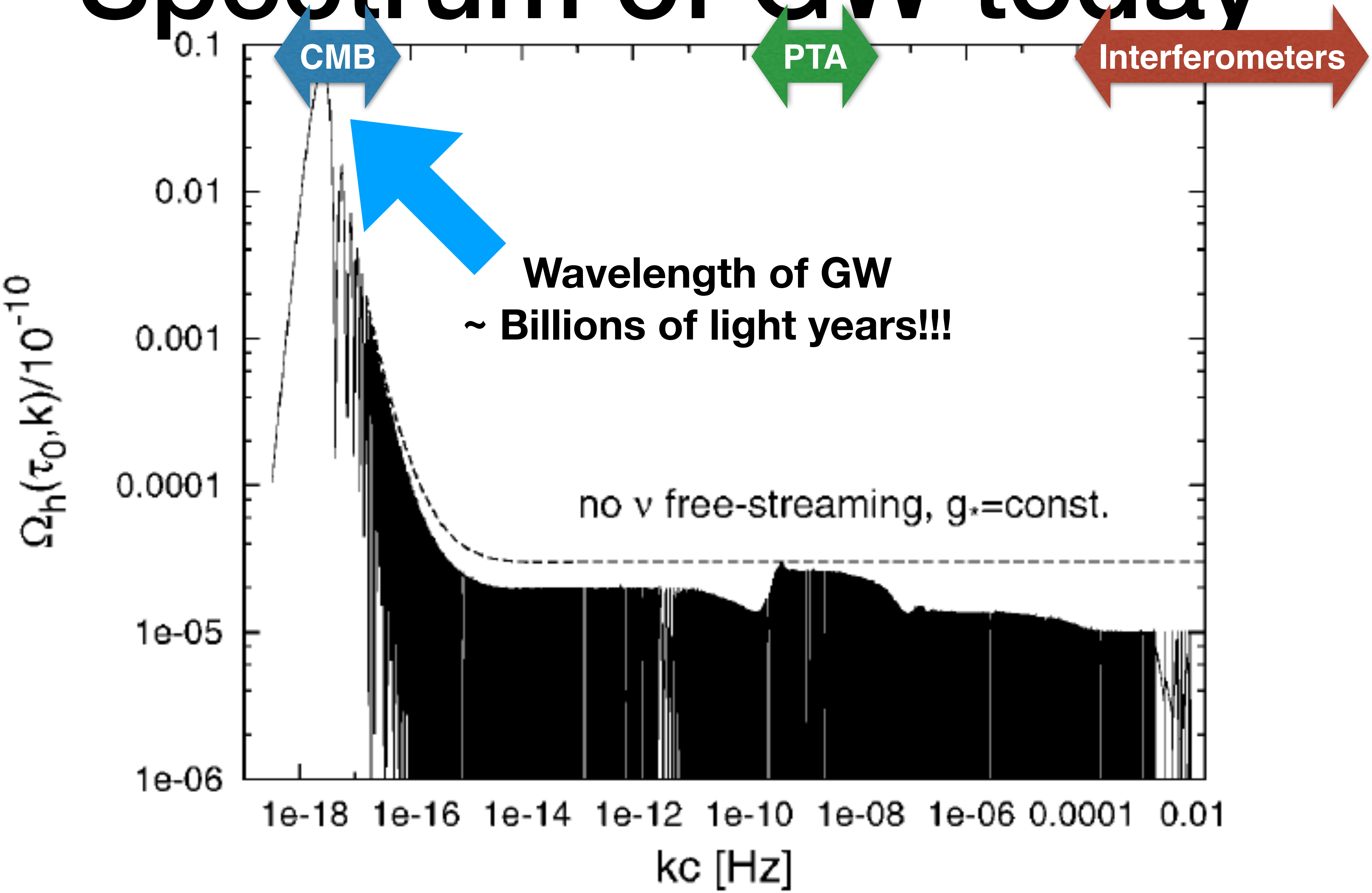


Finding Signatures of Gravitational Waves in the CMB

- **Next frontier in the CMB research**

1. Find evidence for nearly scale-invariant gravitational waves

2. Once found, test Gaussianity to make sure (or not!) that the signal comes from the vacuum fluctuation in spacetime

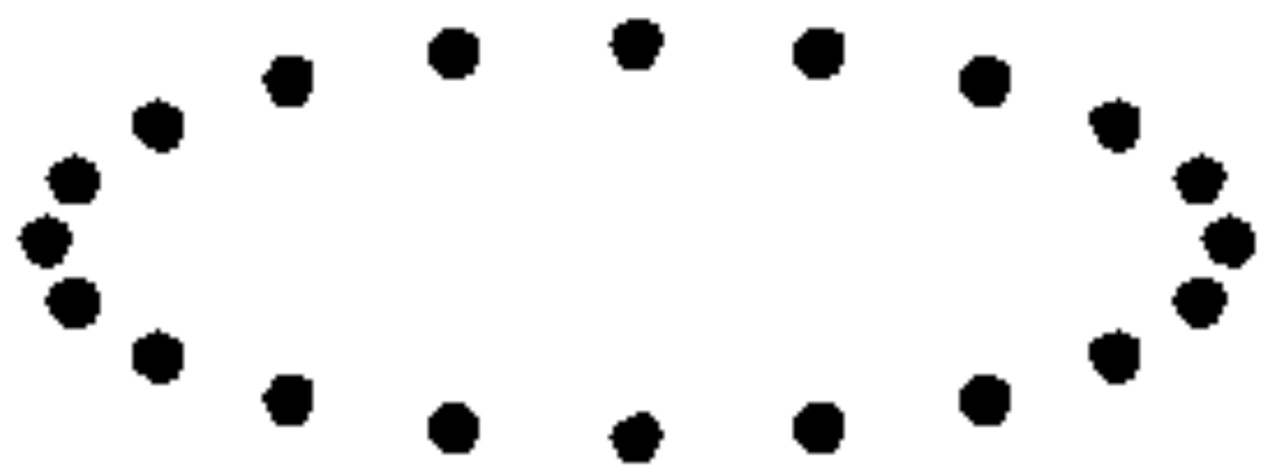
3. Constrain inflation models

Measuring GW

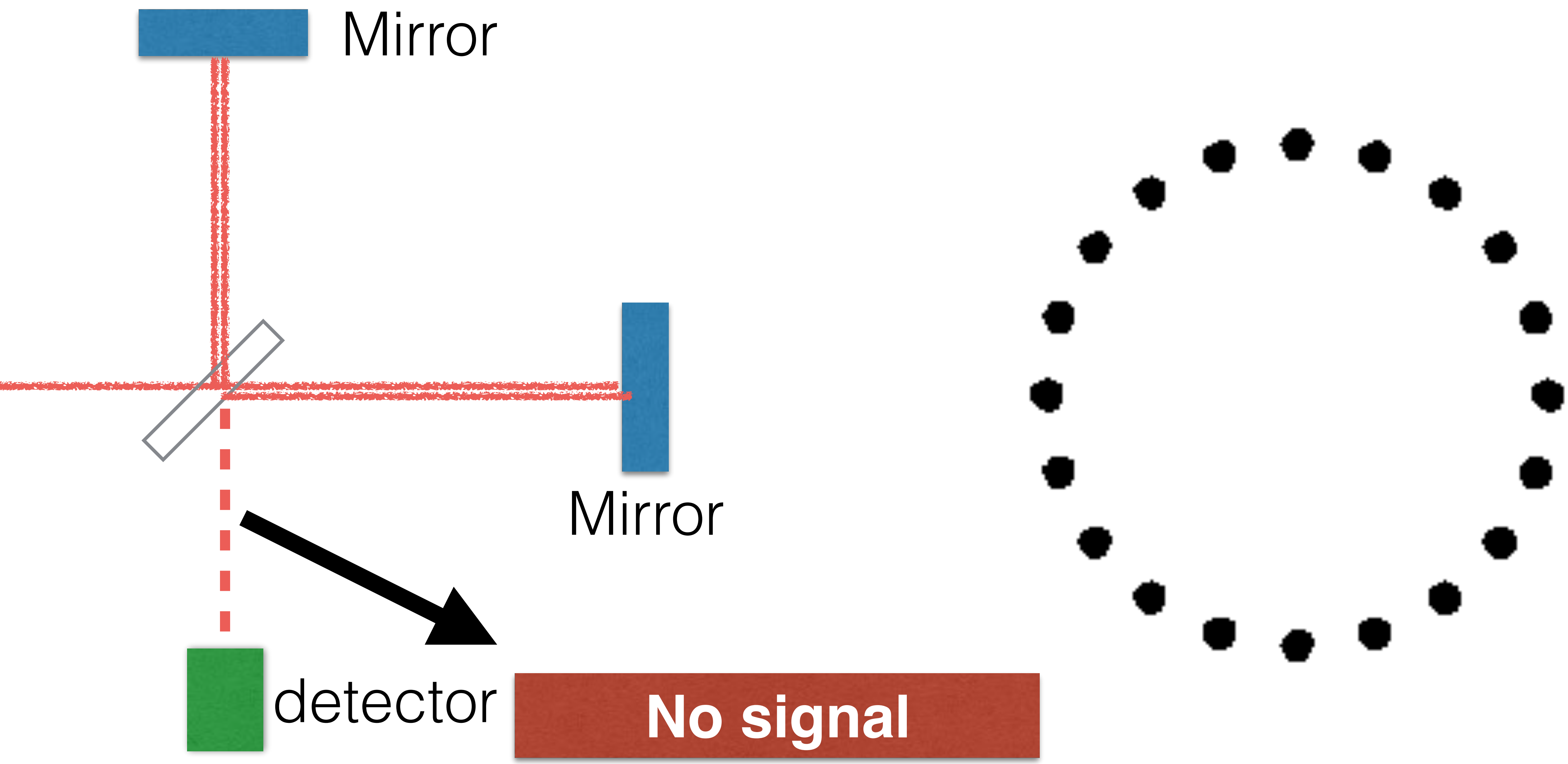
- GW changes distances between two points

$$d\ell^2 = d\mathbf{x}^2 = \sum_{ij} \delta_{ij} dx^i dx^j$$

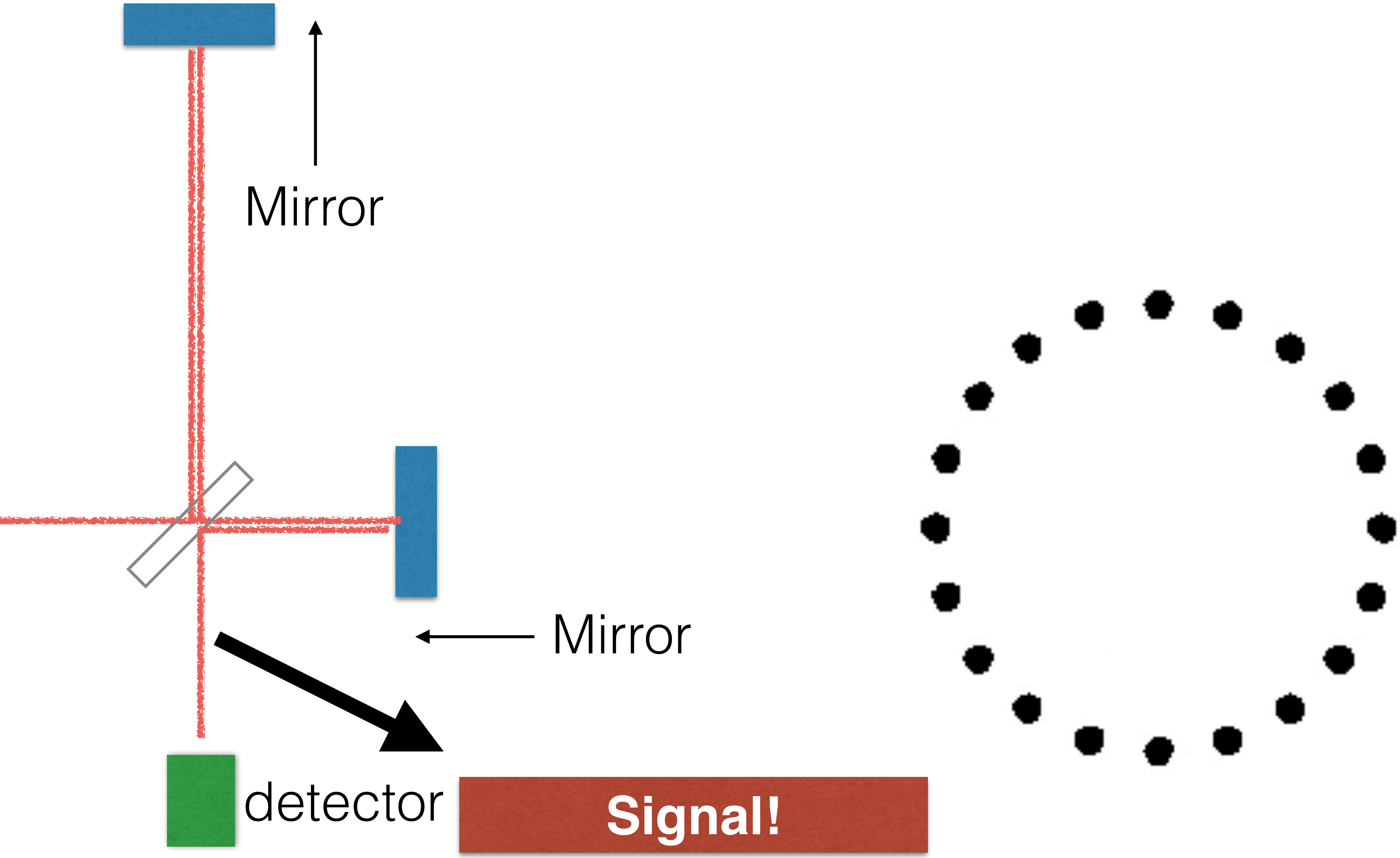
$$d\ell^2 = \sum_{ij} (\delta_{ij} + h_{ij}) dx^i dx^j$$



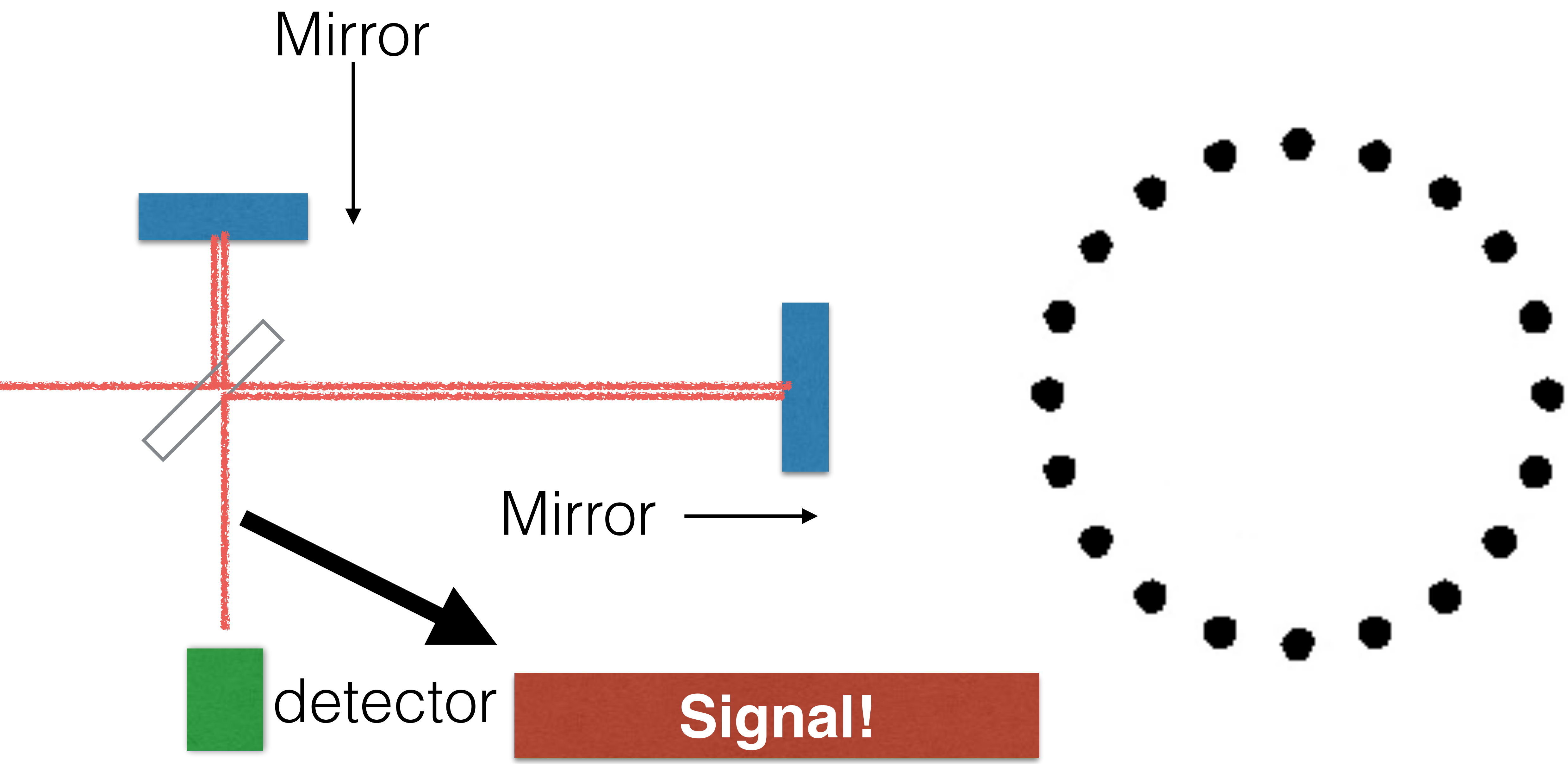
Laser Interferometer



Laser Interferometer



Laser Interferometer



LIGO detected GW from a binary blackholes, with the wavelength of thousands of kilometres

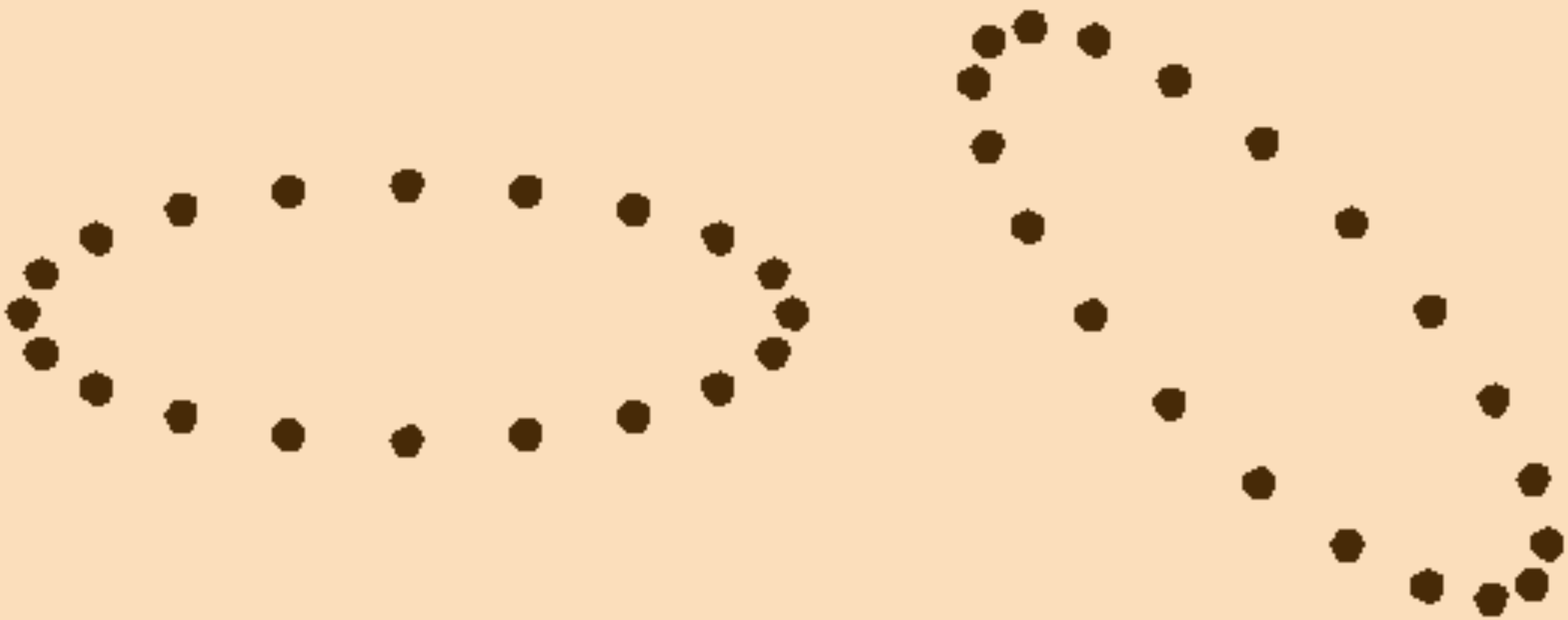
But, the primordial GW affecting the CMB has a wavelength of **billions of light-years**!! How do we find it?

Detecting GW by CMB

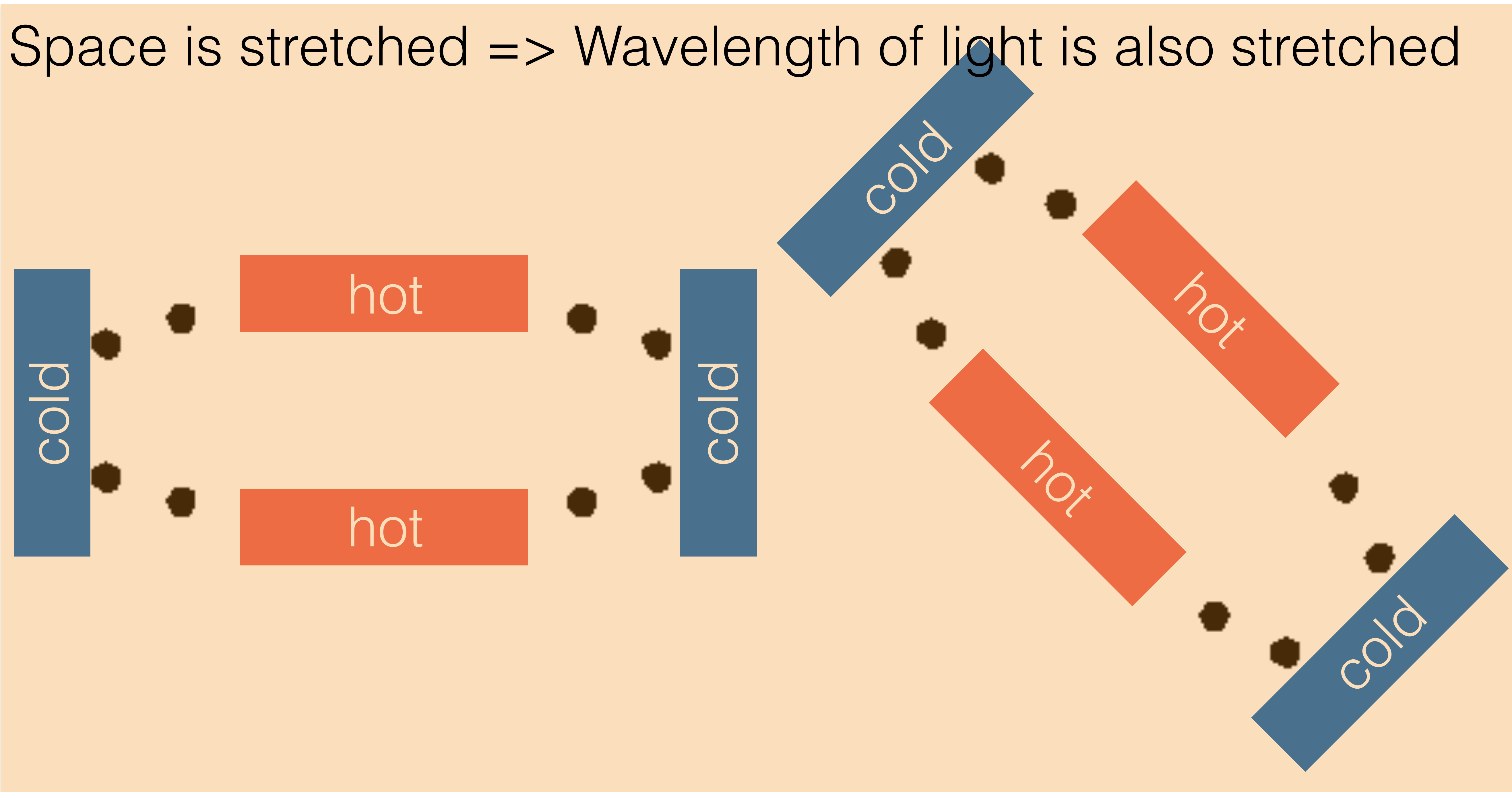
Isotropic electro-magnetic fields

Detecting GW by CMB

GW propagating in isotropic electro-magnetic fields



Detecting GW by CMB



Detecting GW by CMB
Polarisation

Space is stretched => Wavelength of light is also stretched

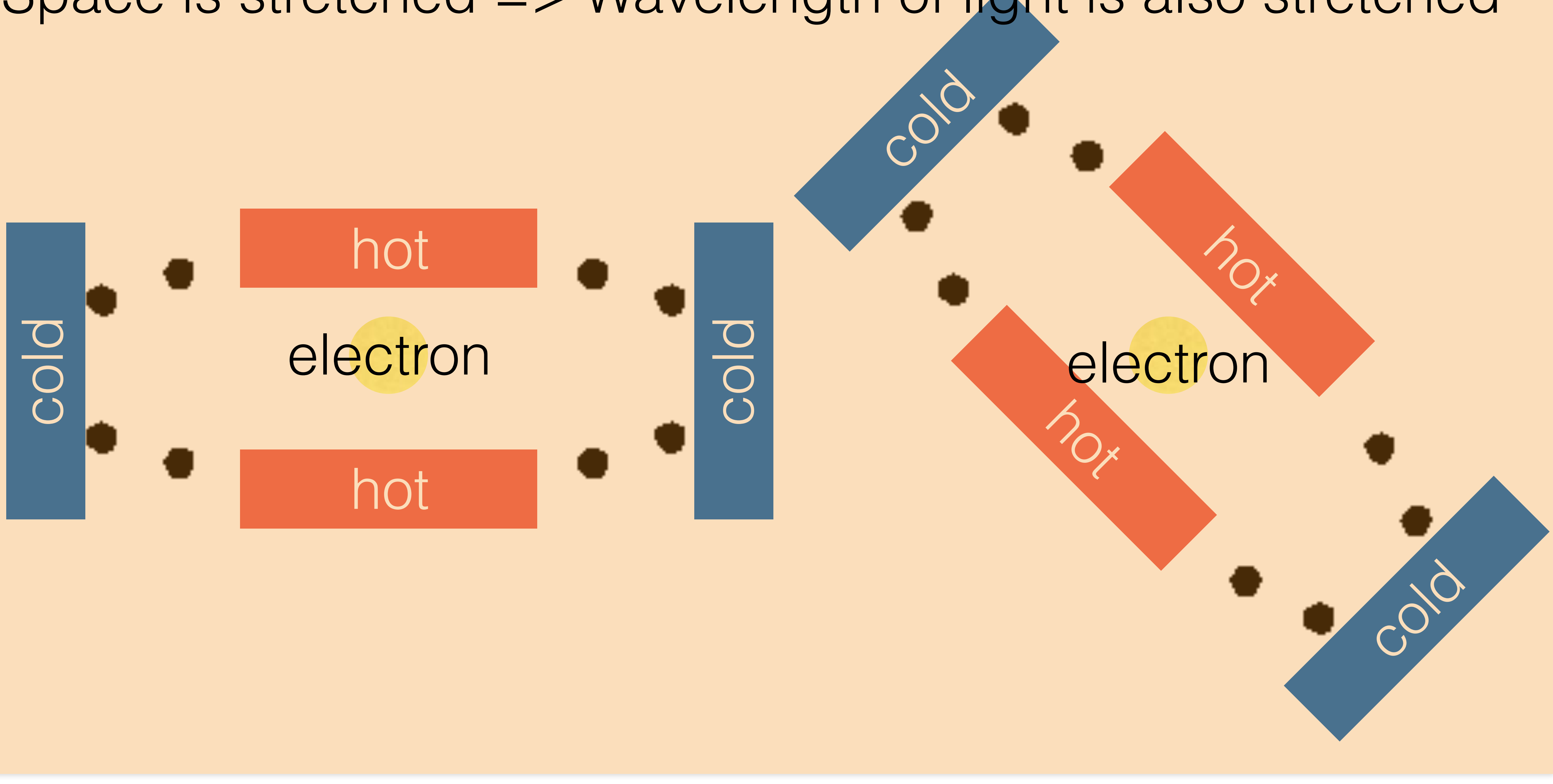
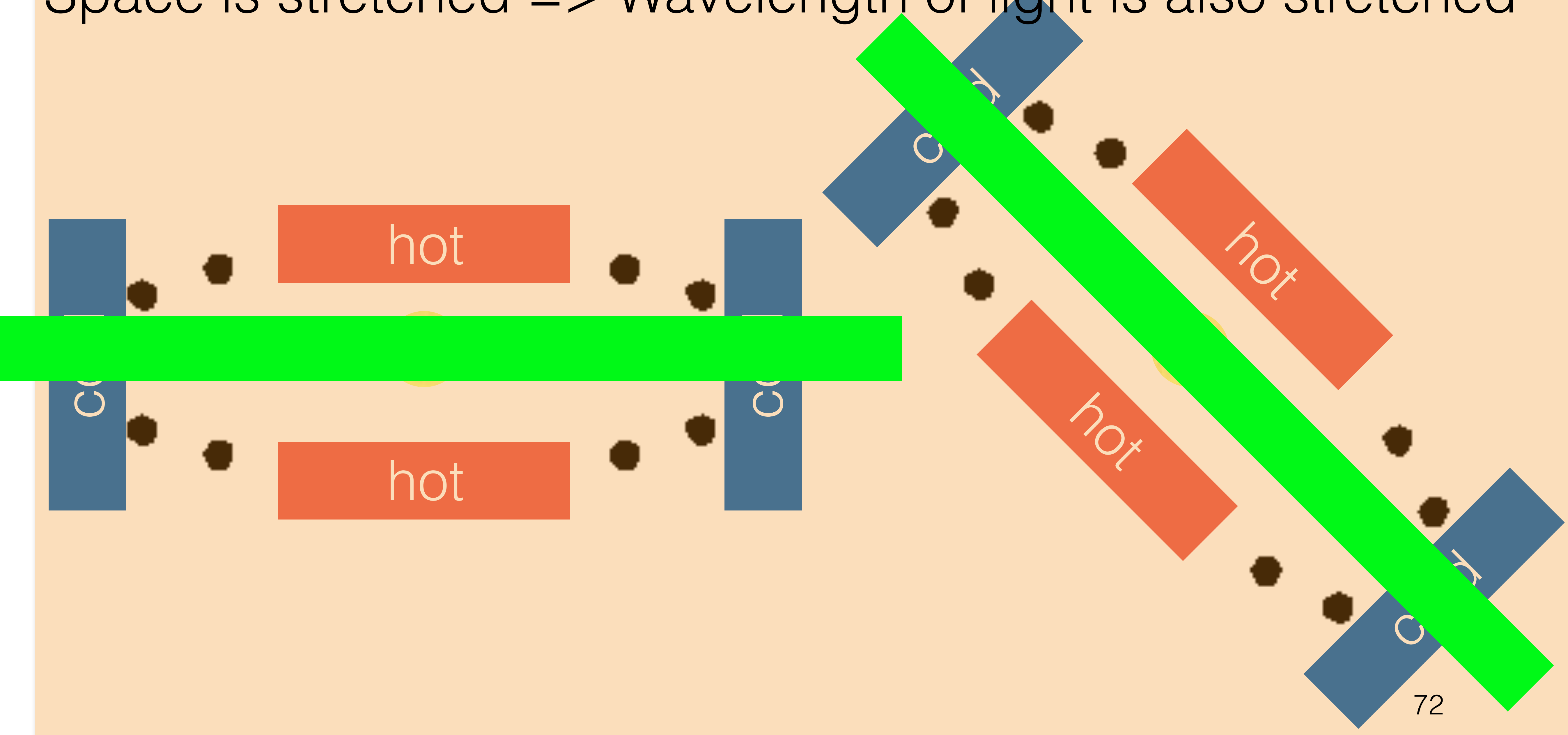



Photo Credit: TALEX
horizontally polarised

Photo Credit: TALEX

Detecting GW by CMB
Polarisation

Space is stretched => Wavelength of light is also stretched

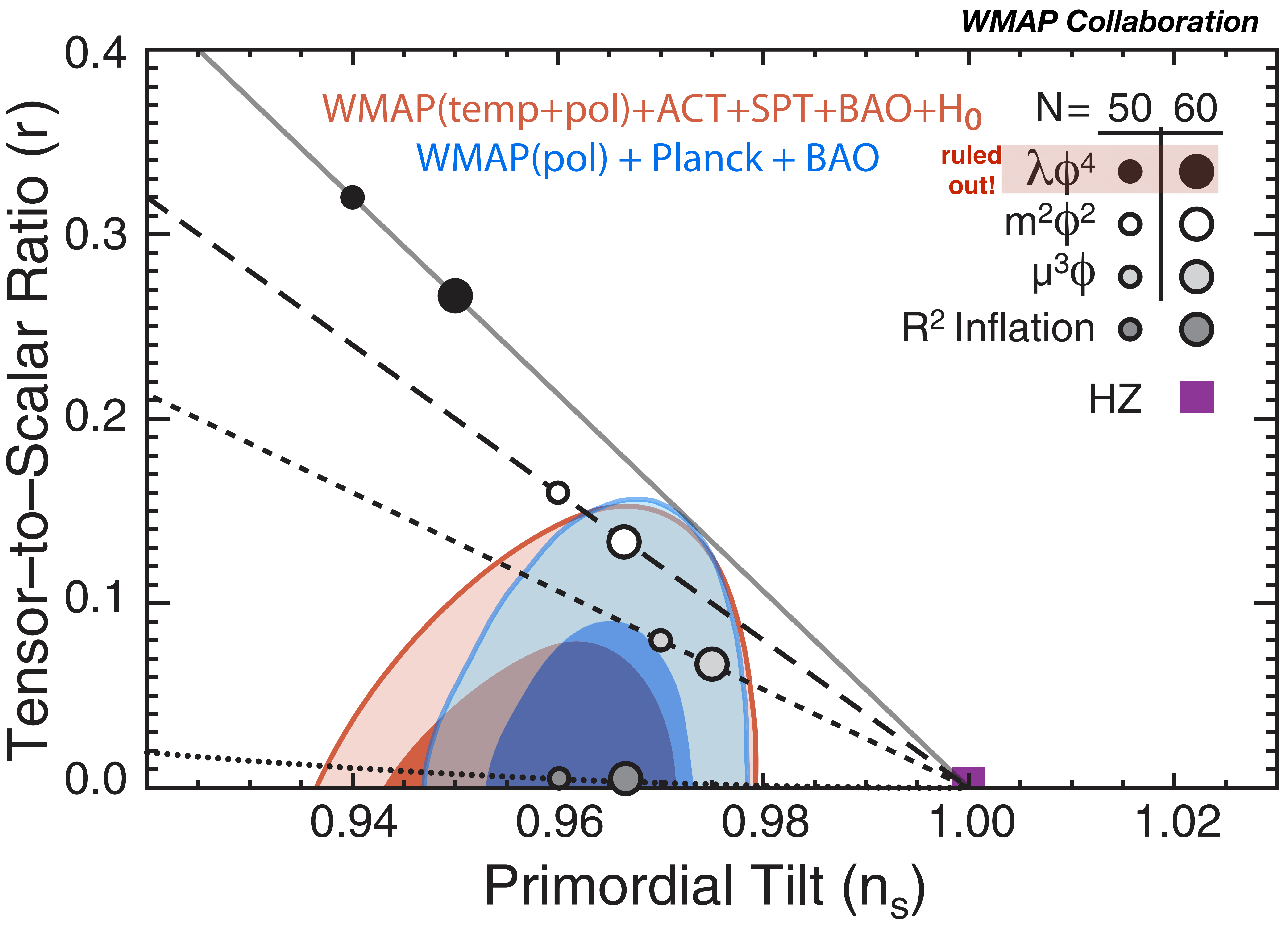


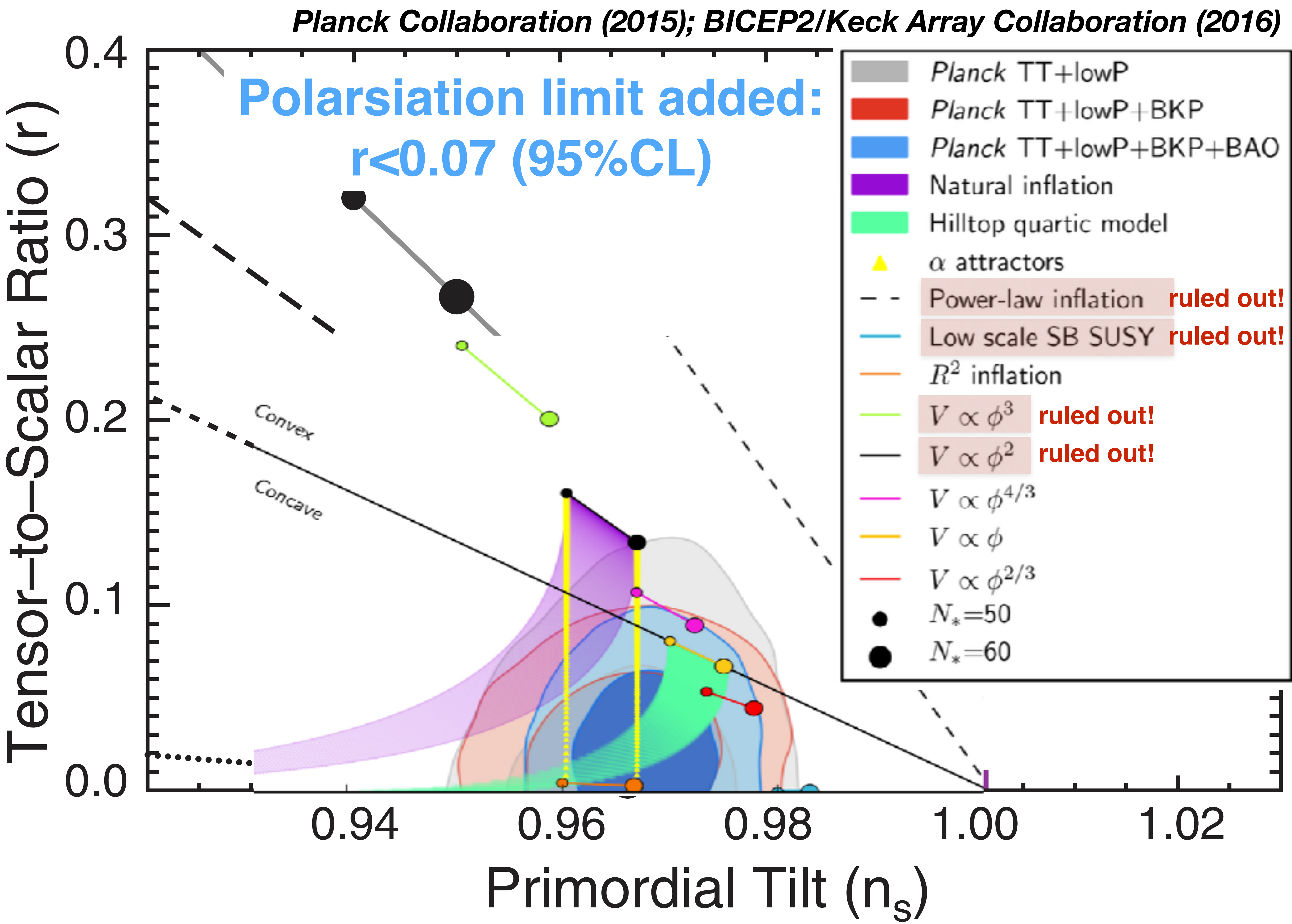
Tensor-to-scalar Ratio

$$r \equiv \frac{\langle h_{ij} h^{ij} \rangle}{\langle \zeta^2 \rangle}$$

- We really want to find this! The current upper bound is **$r<0.07$** (95%CL)

  BICEP2/Keck Array Collaboration (2016)


WMAP Collaboration
WMAP(temp+pol)+ACT+SPT+BAO+H0
WMAP(pol) + Planck + BAO
N= 50 60
ruled out!
λϕ4
m2ϕ2
μ3ϕ
R2 Inflation
HZ
Tensor–to–Scalar Ratio (r)
0.4
0.3
0.2
0.1
0.0
0.94
0.96
0.98
1.00
1.02
Primordial Tilt (ns)


Planck Collaboration (2015); BICEP2/Keck Array Collaboration (2016)
Polarsiation limit added:
r<0.07 (95%CL)
Tensor–to–Scalar Ratio (r)
Primordial Tilt (n_s)
0.0
0.1
0.2
0.3
0.4
0.94
0.96
0.98
1.00
1.02
Convex
Concave
Planck TT+lowP
Planck TT+lowP+BKP
Planck TT+lowP+BKP+BAO
Natural inflation
Hilltop quartic model
α attractors
Power-law inflation
ruled out!
Low scale SB SUSY
ruled out!
R² inflation
V ∝ φ³
ruled out!
V ∝ φ²
ruled out!
V ∝ φ^{4/3}
V ∝ φ
V ∝ φ^{2/3}
N*=50
N*=60

But, wait a minute…

Are GWs from vacuum fluctuation in spacetime, or from sources?

$$\underline{\Box h_{ij}} = -16\pi G \underline{\pi_{ij}}$$

- **Homogeneous solution**: “GWs from vacuum fluctuation”
- **Inhomogeneous solution**: “GWs from sources”
  - Scalar and vector fields cannot source tensor fluctuations at linear order (possible at non-linear level)
  - SU(2) gauge field can!

Maleknejad & Sheikh-Jabbari (2013); Dimastrogiovanni & Peloso (2013); Adshead, Martinec & Wyman (2013); Obata & Soda (2016); …

Important Message

$$\Box h_{ij} = -16\pi G \pi_{ij}$$

- **Do not take it for granted if someone told you that detection of the primordial gravitational waves would be a signature of "quantum gravity"!**
    - Only the homogeneous solution corresponds to the vacuum tensor metric perturbation. **There is no *a priori* reason to neglect an inhomogeneous solution!**
    - Contrary, we have several examples in which detectable B-modes are generated by **sources** [U(1) and SU(2)]

Experimental Strategy Commonly Assumed So Far

1. Detect CMB polarisation in multiple frequencies, to make sure that it is from the CMB (i.e., Planck spectrum)

2. Check for scale invariance: Consistent with a scale invariant spectrum?

   - Yes => Announce discovery of the vacuum fluctuation in spacetime

   - No => WTF?

New Experimental Strategy: New Standard!

1. Detect CMB polarisation in multiple frequencies, to make sure that it is from the CMB (i.e., Planck spectrum)

2. Consistent with a scale invariant spectrum?

3. Parity violating correlations consistent with zero?

4. Consistent with Gaussianity?

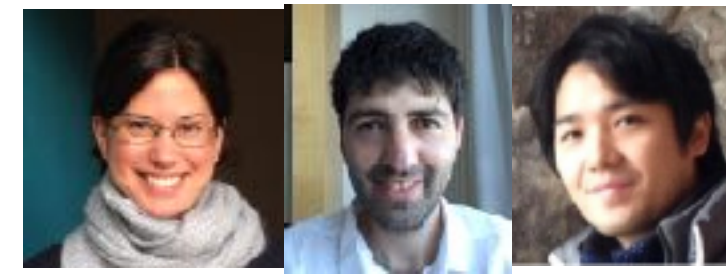
- If, and **ONLY IF** Yes to **all** => Announce discovery of the vacuum fluctuation in spacetime

If not, you may have just discovered new physics during inflation!

2. Consistent with a scale invariant spectrum?
3. Parity violating correlations consistent with zero?
4. Consistent with Gaussianity?

- If, and **ONLY IF** Yes to **all** => Announce discovery of the vacuum fluctuation in spacetime

Dimastrogiovanni, Fasielo & Fujita (2017)



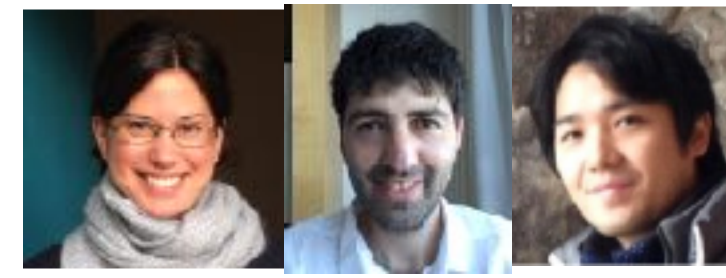
GW from Axion-SU(2) Dynamics

$$\mathcal{L} = \mathcal{L}_{GR} + \mathcal{L}_{\phi} + \mathcal{L}_{\chi} - \frac{1}{4} F^a_{\mu\nu} F^{a\mu\nu} + \frac{\lambda\,\chi}{4f} F^a_{\mu\nu} \tilde{F}^{a\mu\nu}$$

- ϕ: inflaton field => Just provides quasi-de Sitter background
- χ: pseudo-scalar "axion" field. Spectator field (i.e., negligible energy density compared to the inflaton)
- Field strength of an SU(2) field A^a_ν :

$$F^a_{\mu\nu} \equiv \partial_\mu A^a_\nu - \partial_\nu A^a_\mu - g\epsilon^{abc} A^b_\mu A^c_\nu$$

Dimastrogiovanni, Fasielo & Fujita (2017)



Background and Perturbation

- In an inflating background, the SU(2) field has a background solution:

$$A_i^a = [\text{scale factor}] \times Q \times \delta_i^a$$

$$Q \equiv (-f\partial_\chi U/3g\lambda H)^{1/3}$$

U: axion potential

- Perturbations contain a tensor mode (as well as S&V)

$$\delta A_i^a = t_{ai} + \cdots$$

$$t_{ii} = \partial_a t_{ai} = \partial_i t_{ai} = 0$$

Scenario

- The SU(2) field contains tensor, vector, and scalar components

- The tensor components are amplified strongly by a coupling to the axion field

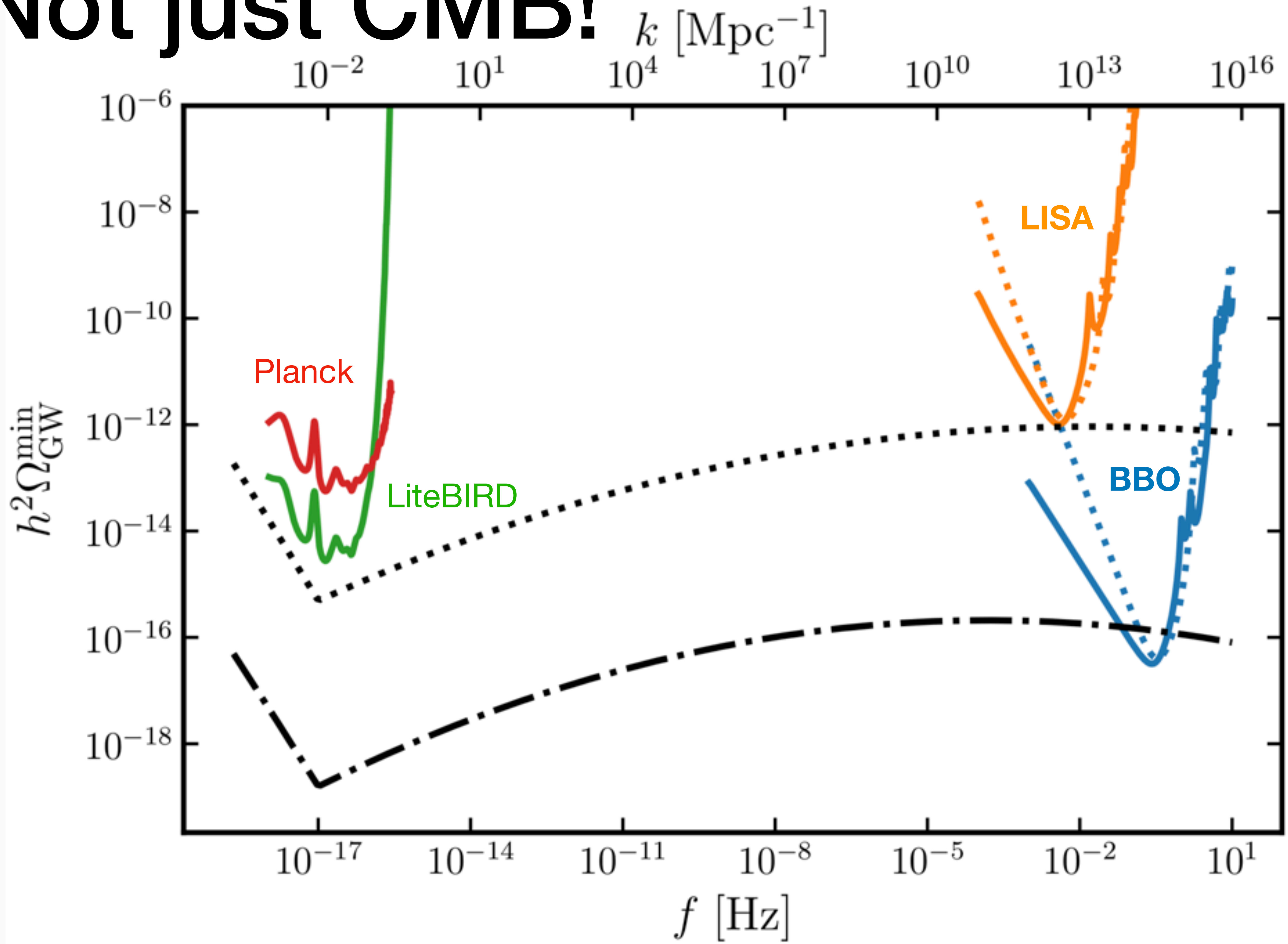
  - But, **only one helicity is amplified** => GW is **chiral** (well-known result)

- Brand-new result: **GWs sourced by this mechanism are strongly non-Gaussian!**

  Agrawal, Fujita & EK (2017)

Thorne, Fujita, Hazumi, Katayama, EK & Shiraishi, arXiv:1707.03240

Not just CMB!



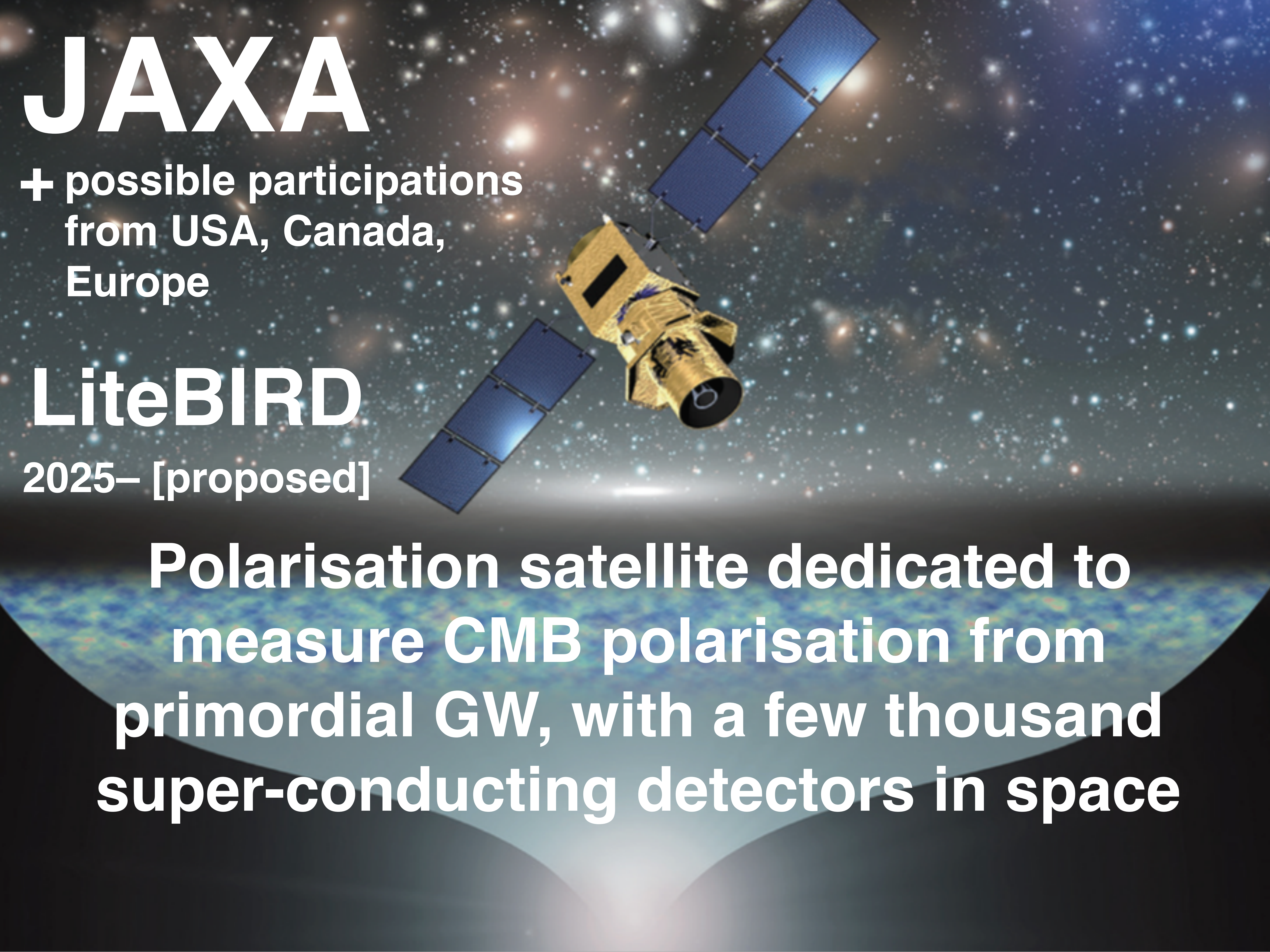
JAXA

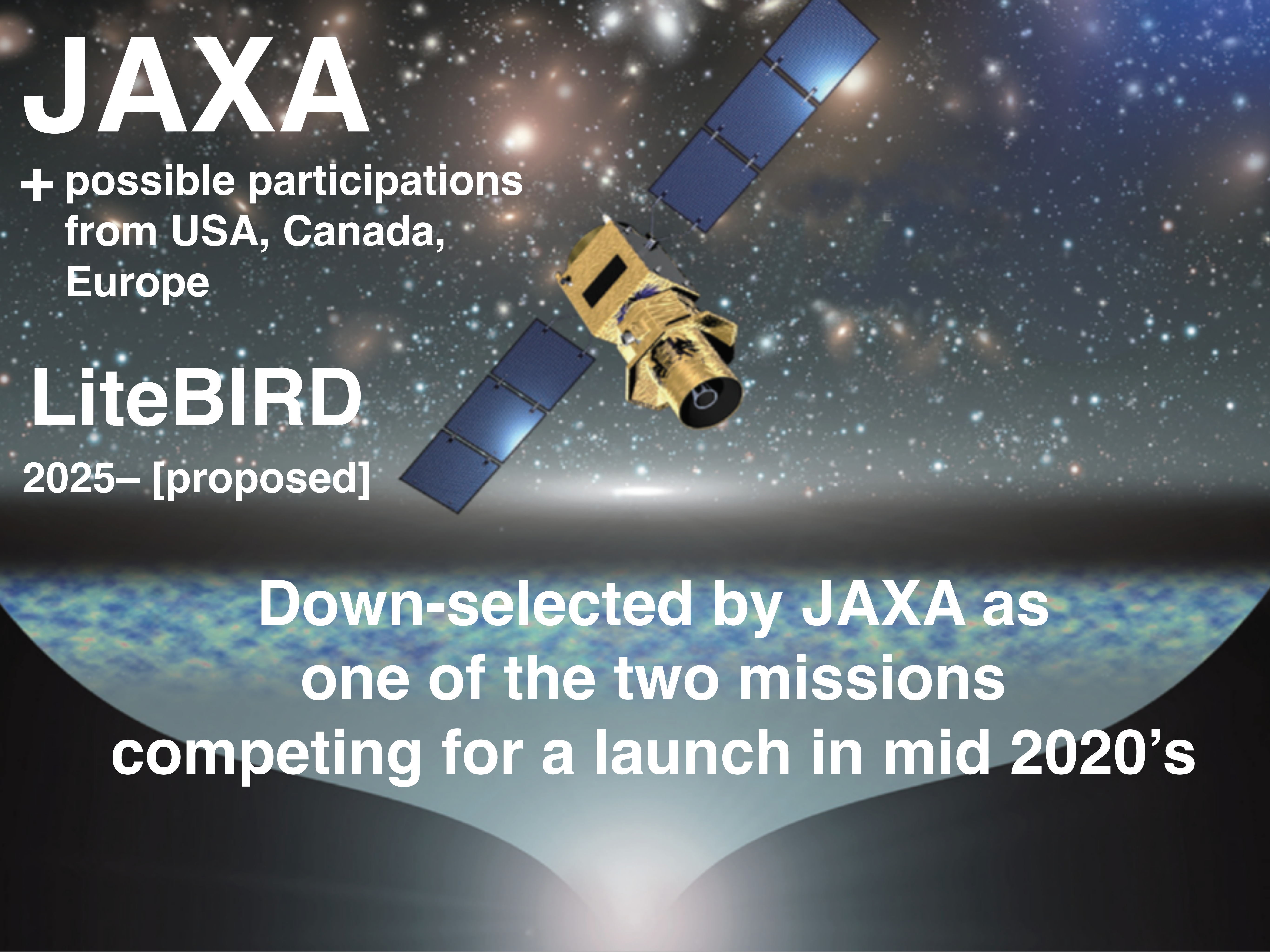
+ possible participations from USA, Canada, Europe

LiteBIRD

2025– [proposed]

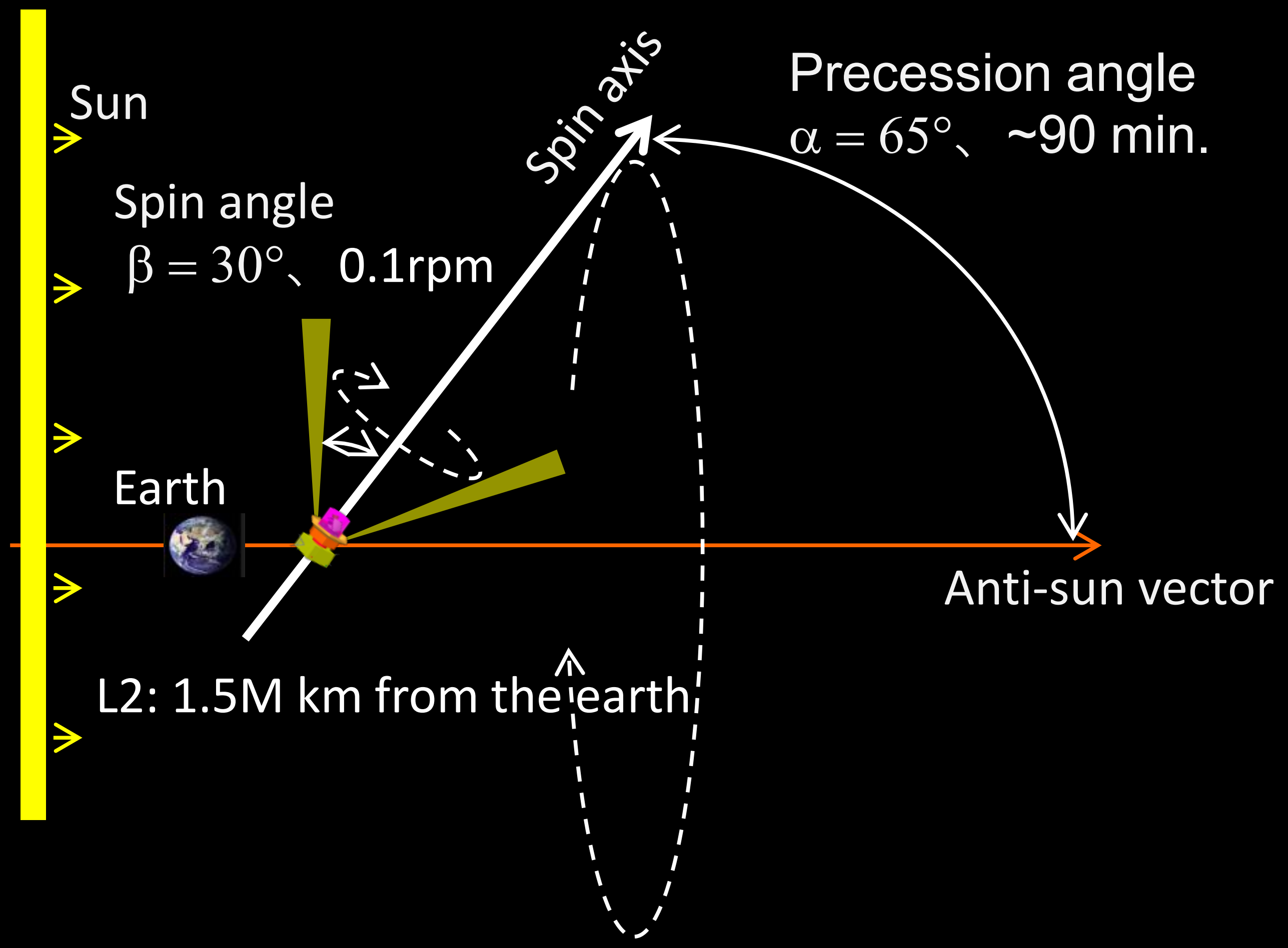
Target: δr<0.001


JAXA
+ possible participations from USA, Canada, Europe
LiteBIRD
2025– [proposed]
Polarisation satellite dedicated to measure CMB polarisation from primordial GW, with a few thousand super-conducting detectors in space


JAXA
+ possible participations from USA, Canada, Europe
LiteBIRD
2025– [proposed]
Down-selected by JAXA as one of the two missions competing for a launch in mid 2020’s

Observation Strategy

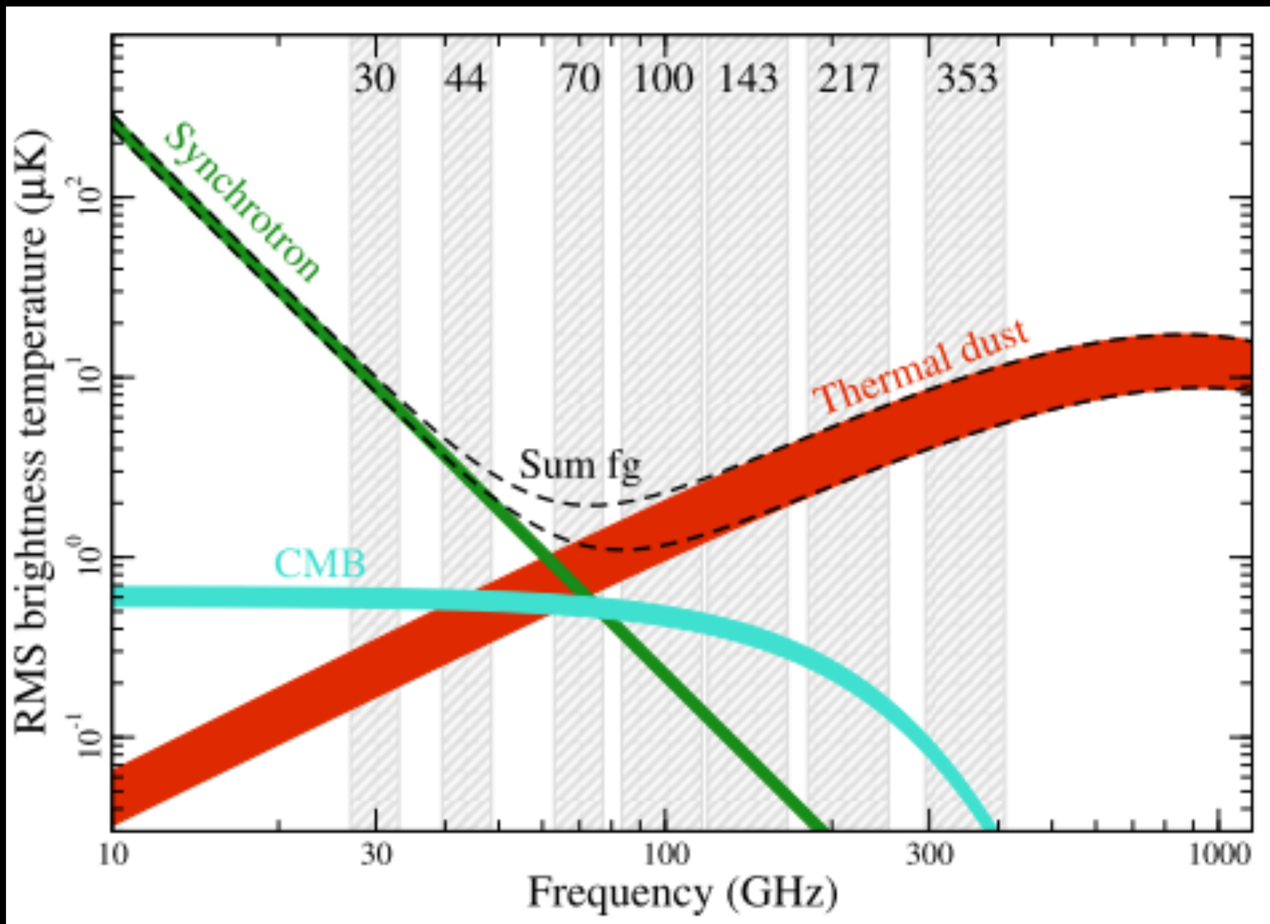
JAXA H3 Launch Vehicle (JAXA)



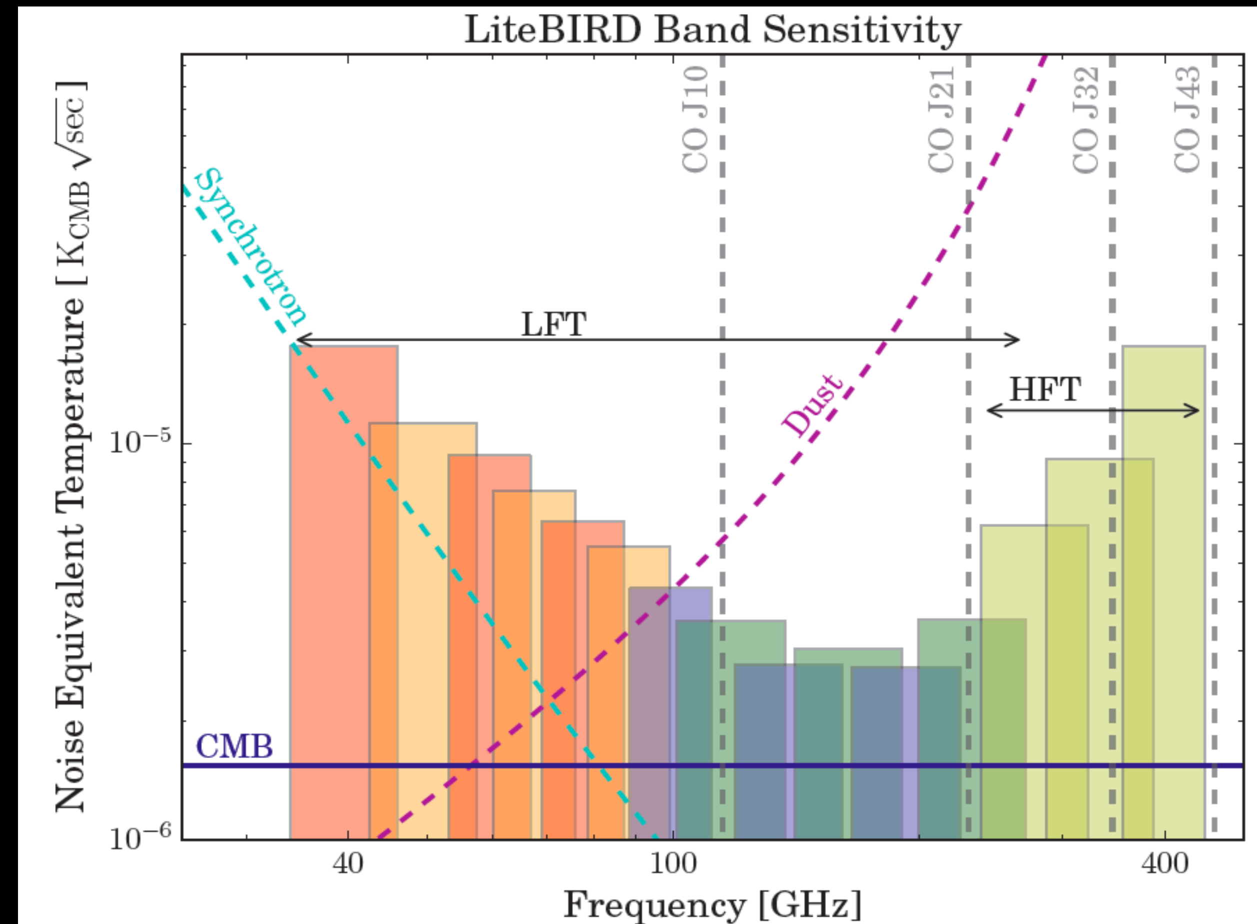
- Launch vehicle: JAXA H3
- Observation location: Second Lagrangian point (L2)
- Scan strategy: Spin and precession, full sky
- Observation duration: 3-years
- Proposed launch date: Mid 2020's

Slide courtesy Toki Suzuki (Berkeley)

Foreground Removal



Polarized galactic emission (Planck X)



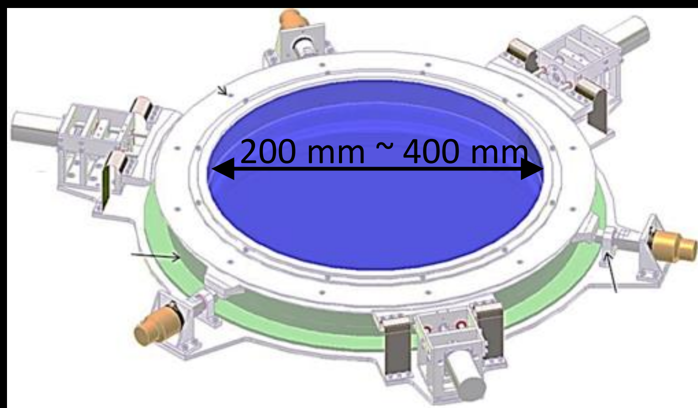
LiteBIRD: 15 frequency bands

- Polarized foregrounds
  - Synchrotron radiation and thermal emission from inter-galactic dust
  - Characterize and remove foregrounds
- 15 frequency bands between 40 GHz - 400 GHz
  - Split between Low Frequency Telescope (LFT) and High Frequency Telescope (HFT)
  - LFT: 40 GHz – 235 GHz
  - HFT: 280 GHz – 400 GHz

Slide courtesy Toki Suzuki (Berkeley)

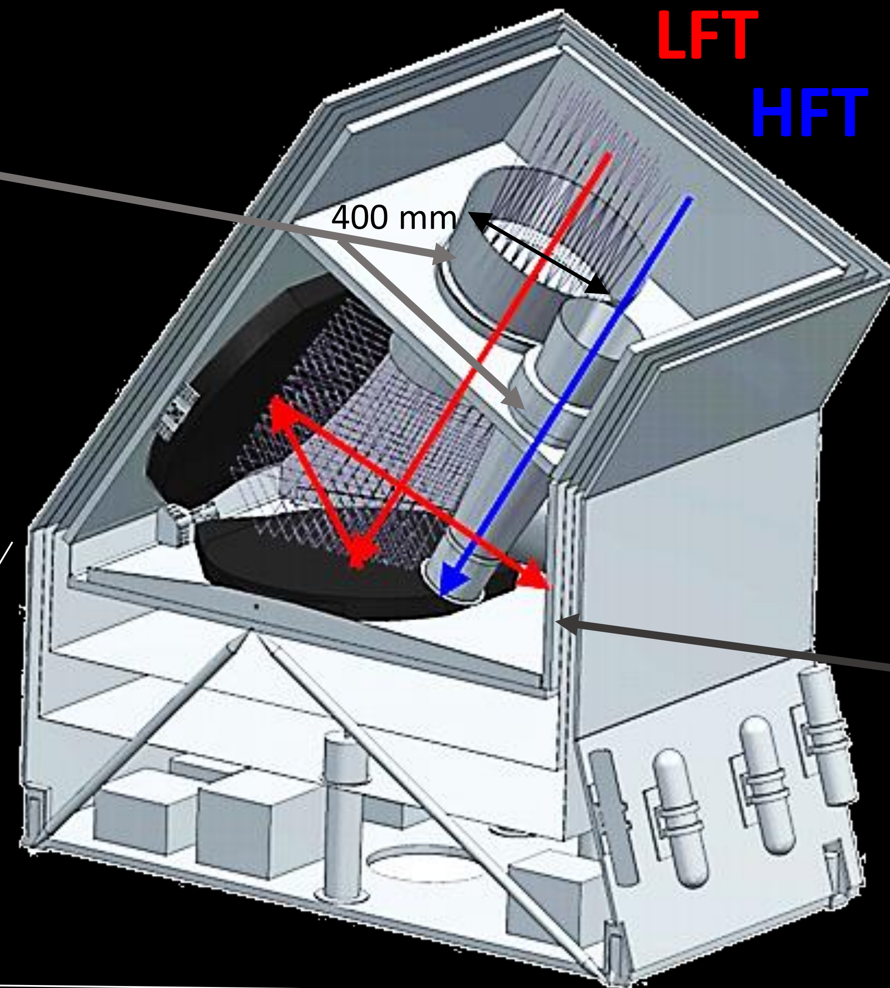
Slide courtesy Toki Suzuki (Berkeley)

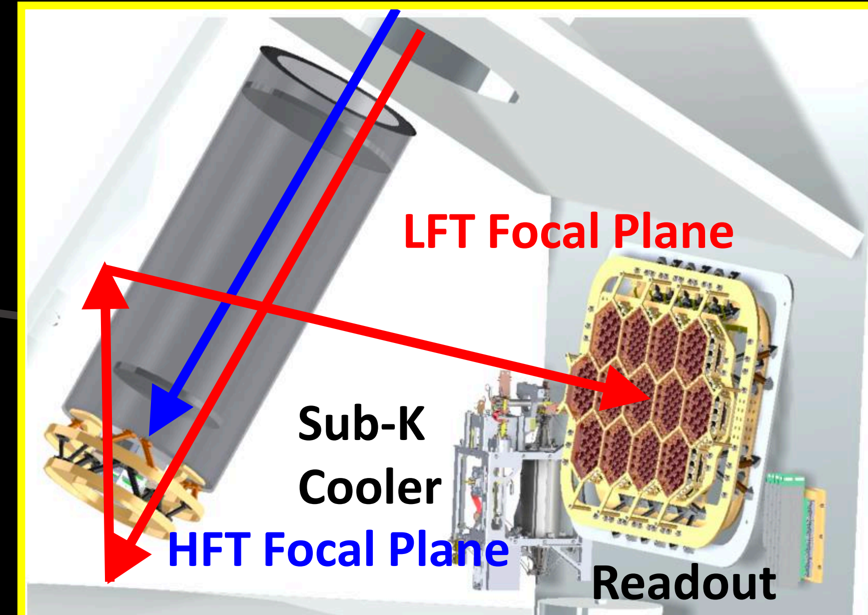
Instrument Overview



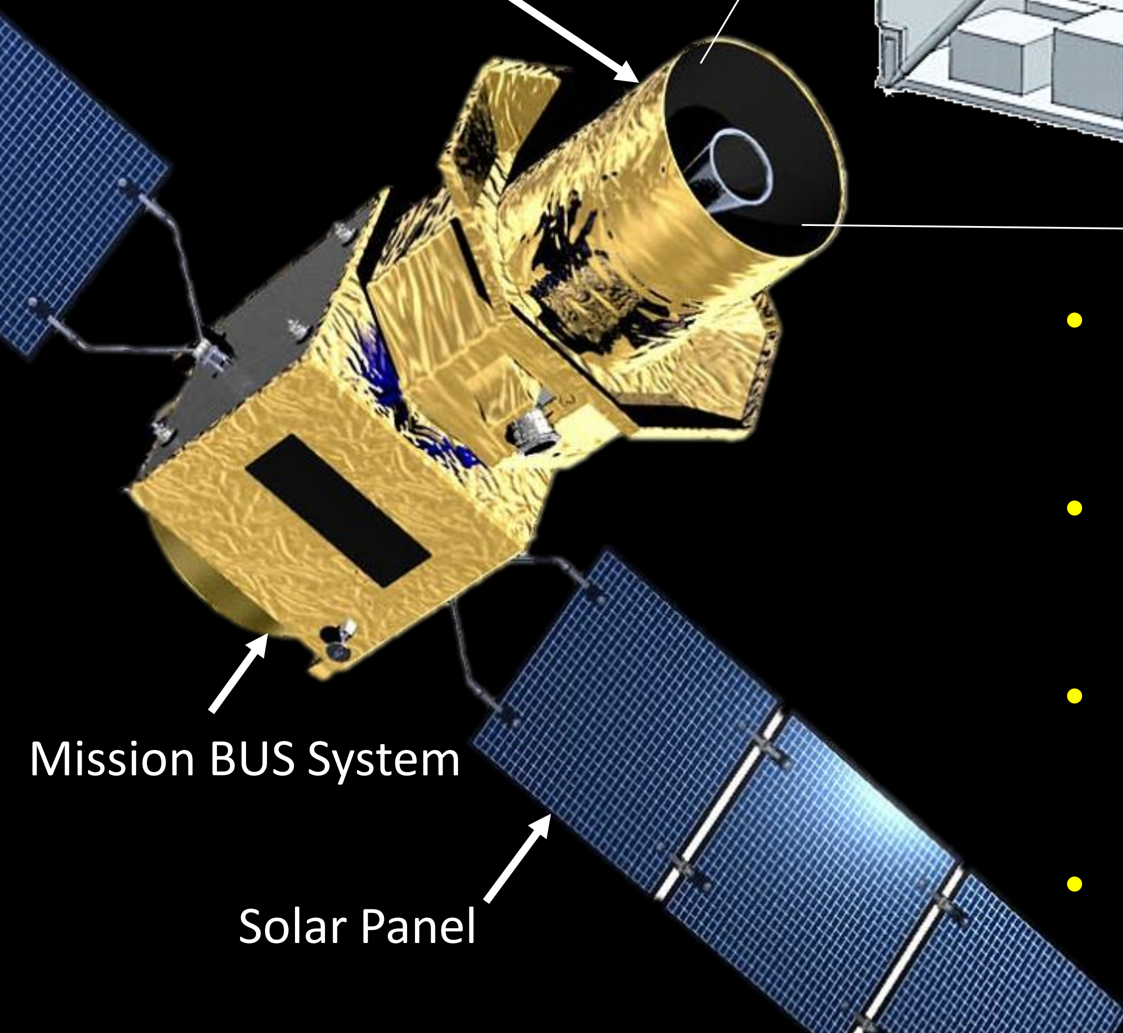
Half-wave plate

Stirling & Joule Thomson Coolers





Sub-Kelvin Instrument



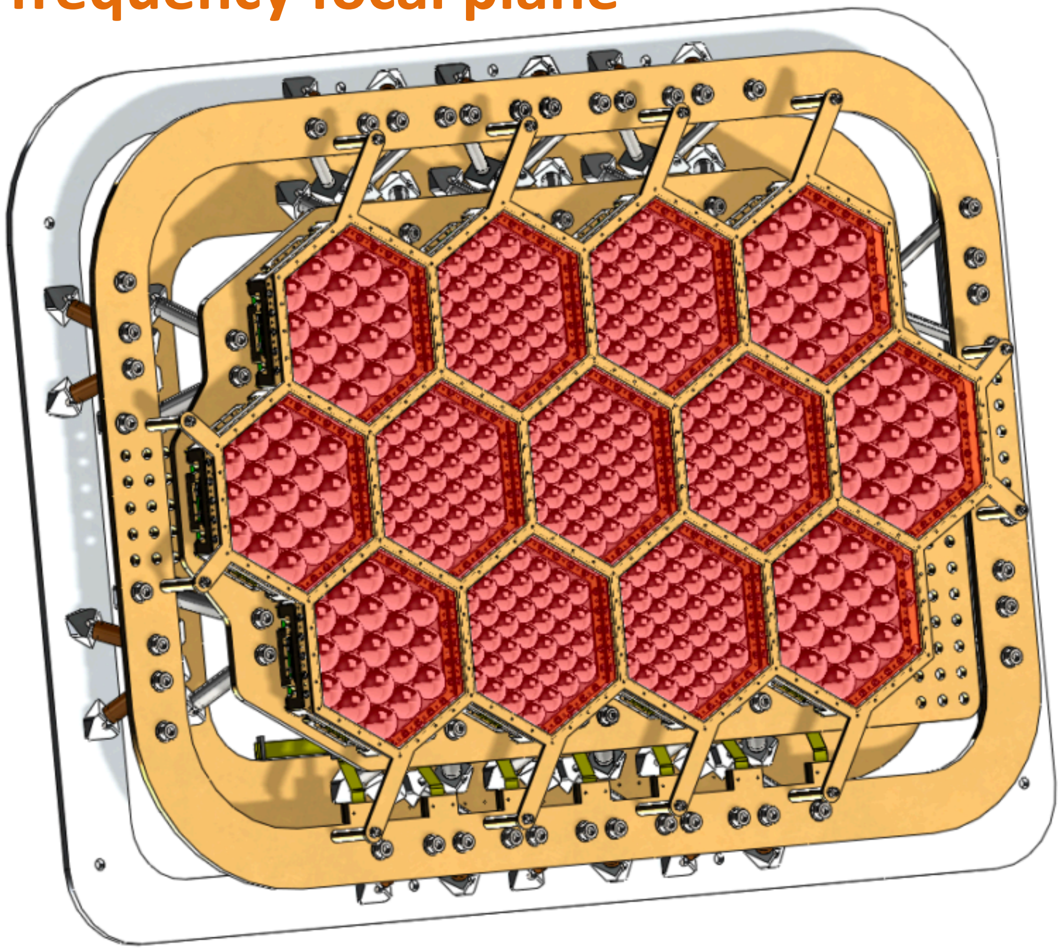
- Two telescopes
  - Crossed-Dragone (LFT) & on-axis refractor (HFT)
- Cryogenic rotating achromatic half-wave plate
  - Modulates polarization signal
- Stirling & Joule Thomson coolers
  - Provide cooling power above 2 Kelvin
- Sub-Kelvin Instrument
  - Detectors, readout electronics, and a sub-kelvin cooler

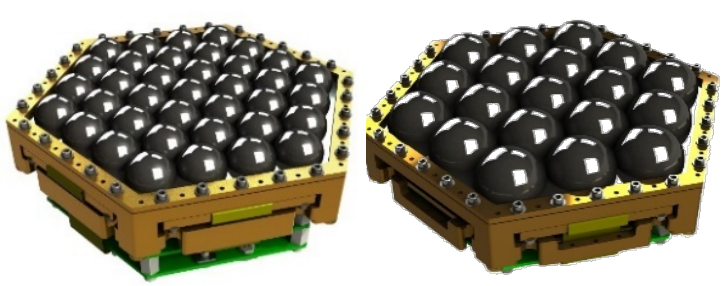
Summary

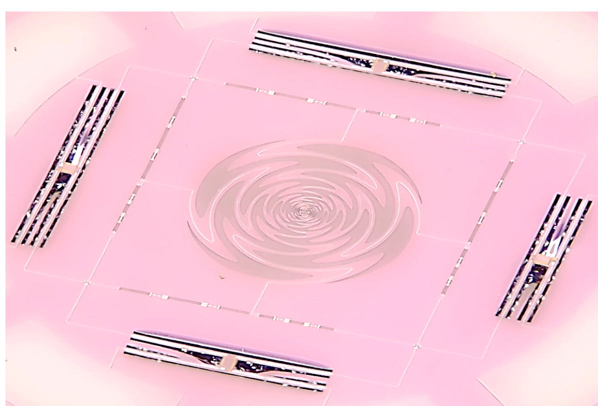
- Inflation looks good: all the CMB data support it
- **Next frontier**: Using CMB polarisation to find GWs from inflation. **Definitive evidence for inflation!**
  - With LiteBIRD we plan to reach $r \sim 10^{-3}$, i.e., 100 times better than the current bound
  - GW from vacuum or sources? An exciting window to new physics

LFT and HFT focal plane units using TES

Low frequency focal plane

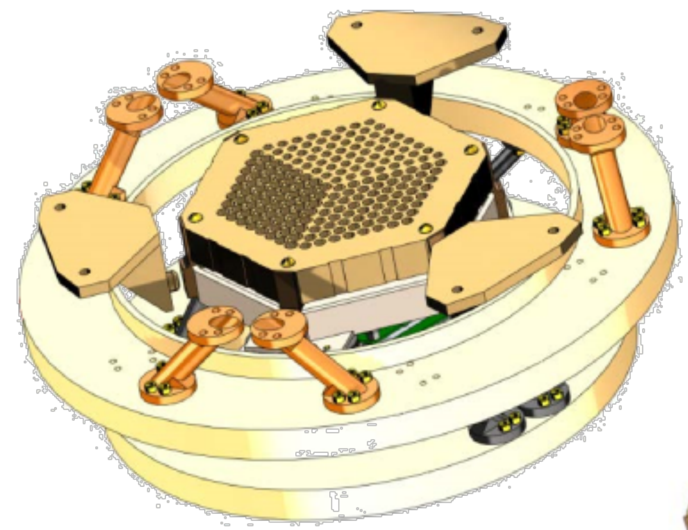




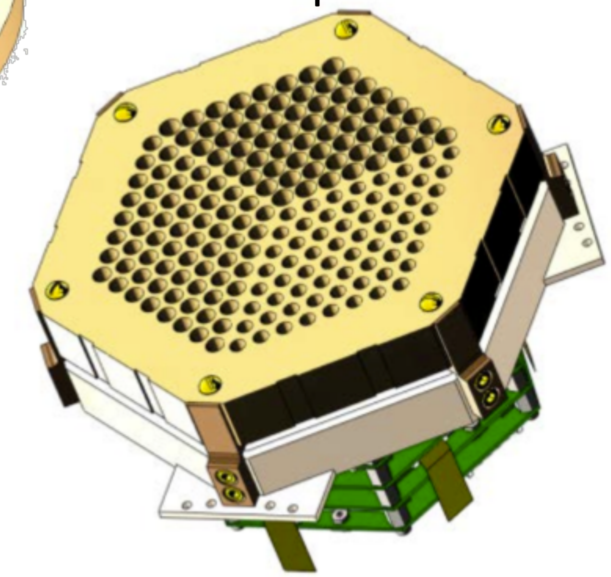


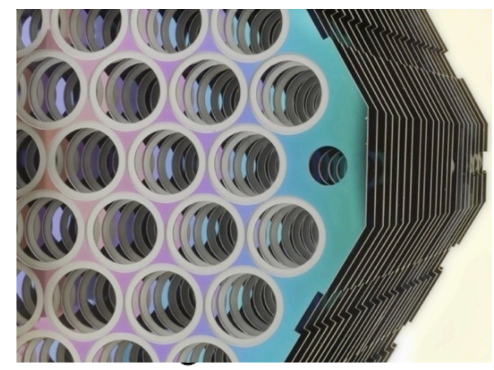
Three colors per pixel with a lenslet coupling.

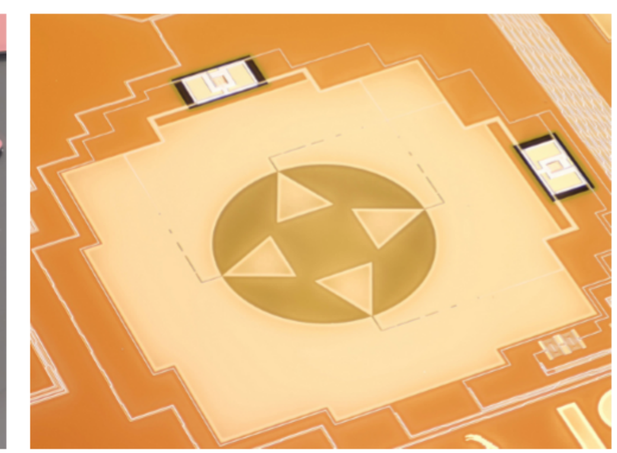
High frequency focal plane



Each color per feed, and three colors within one focal plane.



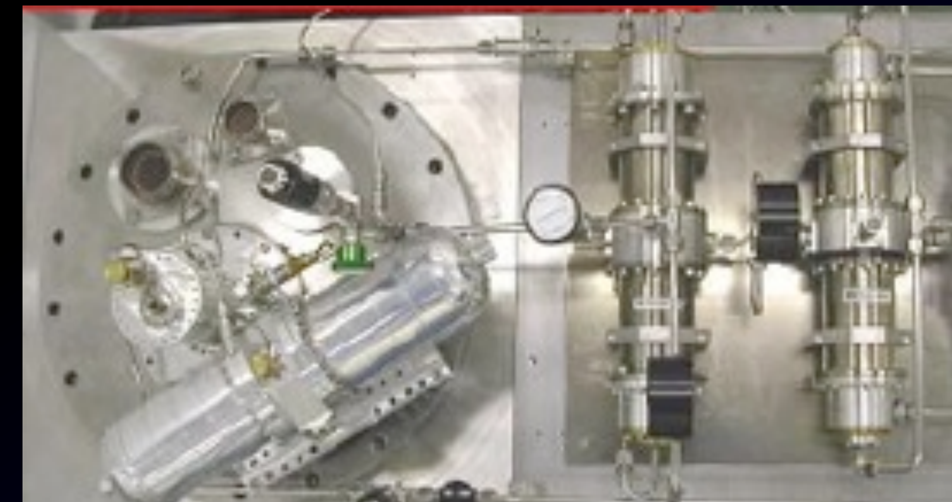




- The current baseline design uses a single ADR to cool the both focal planes.
- The LF focal plane has ** TESs and the HF focal plane has ** TESs.
- The TES is read by SQUID together with the readout electronics is based on the digital frequency multiplexing system.
- The effect of the cosmic ray is evaluated by building a model. The irradiation test is in plan.

Slide courtesy Tomo Matsumura (Kavli IPMU)

Cooling system



SHI/JAXA

Cryogenics

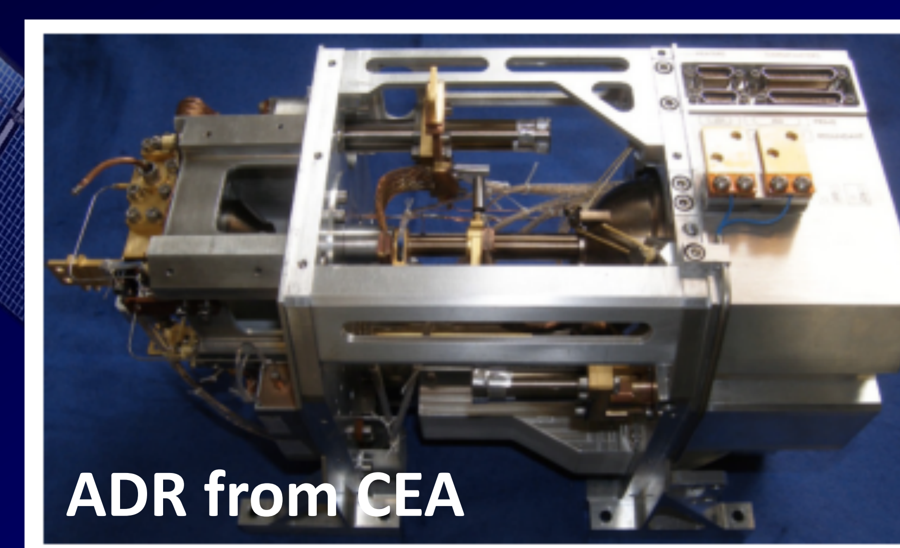
- Warm launch
- 3 years of observations
- 4 K for the mission instruments (optical system)
- 100 mK for the focal plane

Mechanical cooler

- The 2-stage Stirling cooler and 4K-JT cooler from the heritage of the JAXA satellites, Akari (Astro-F), JEM-SMILES and Astro-H.
- The 1K-JT provides the 1.7 K interface to the sub-Kelvin stage.

Sub-Kelvin cooler

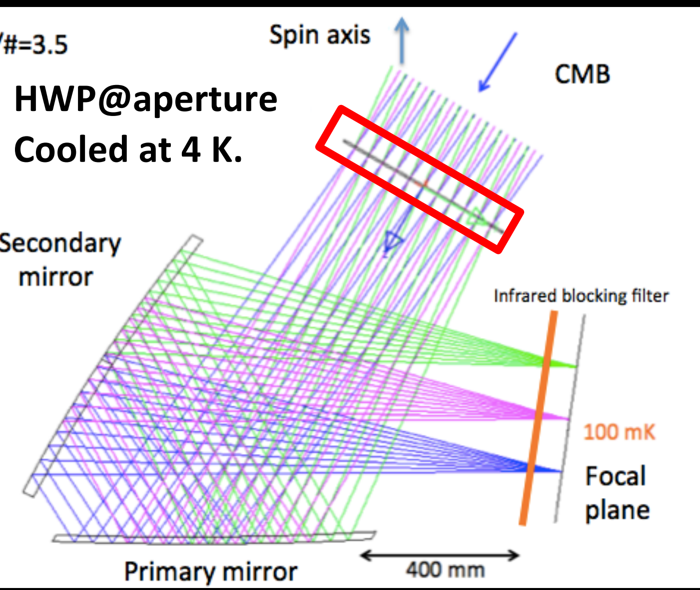
- ADR has a high-TRL and extensive development toward Astro-H, SPICA, and Athena.
- Closed dilution with the Planck heritage is also under development.



ADR from CEA

Slide courtesy Tomo Matsumura (Kavli IPMU)

Polarization modulator

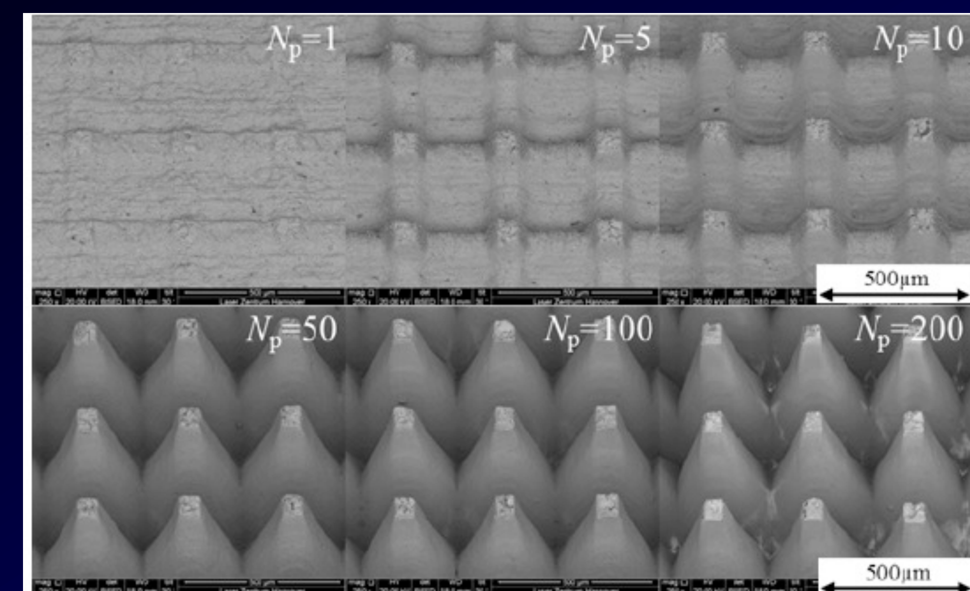


- Due to our focus on the primordial signal at low ell, we employ the continuously rotating achromatic half-wave plate (HWP).
- The HWP modulator suffices mitigating the 1/f noise and the differential systematics.

ote: we also employ the
olarization modulator for HFT.

Broadband coverage

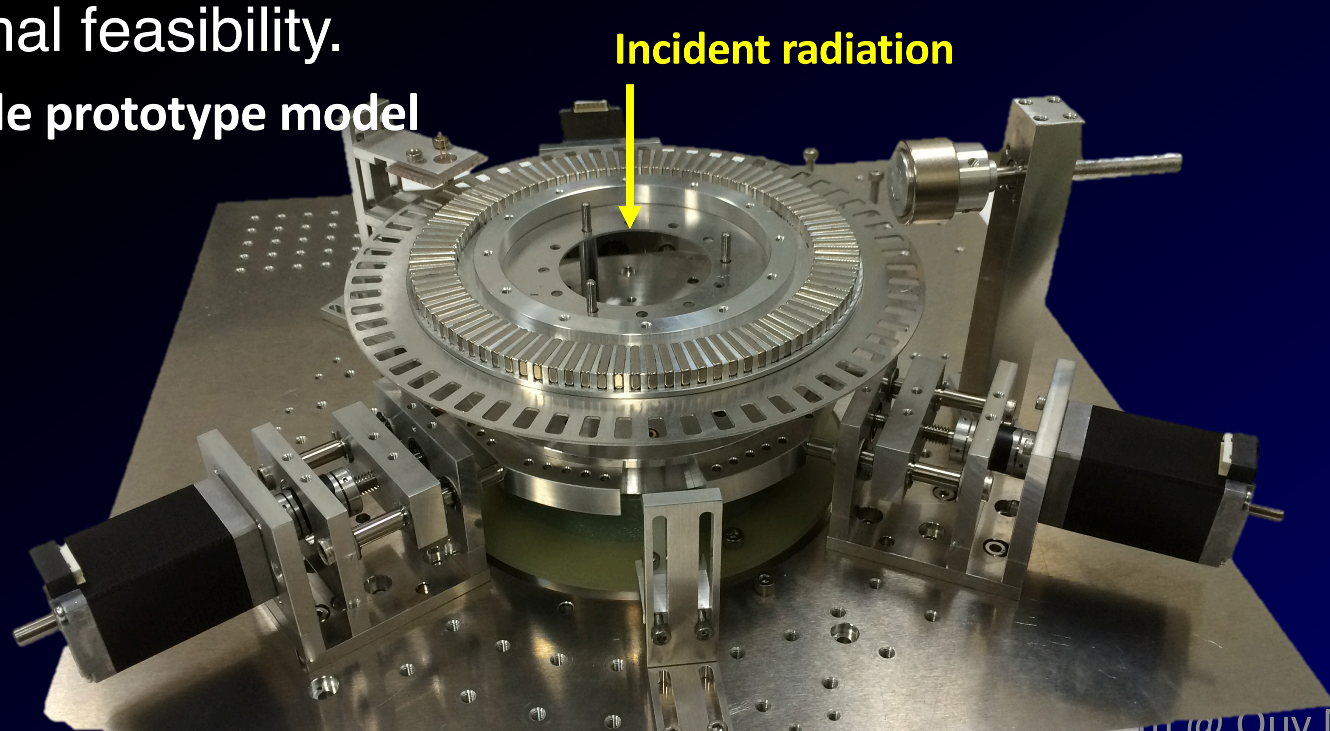
- The broadband coverage is done by the sub-wavelength anti-reflection structure.
- The broadband modulation efficiency is achieved by using 9-layer achromatic HWP.



otational mechanism

he continuous rotation is achieved by employing the uperconducting magnetic bearing. This system has a heritage om EBEX. The prototype system has built and test the kinetic nd thermal feasibility.

he 1/9 scale prototype model



The proton irradiation test is conducted to key components, including sapphire, YBCO, and magnets. We have not found the no-go results. And the further test is in progress.

Slide courtesy Tomo Matsumura (Kavli IPMU)

Agrawal, Fujita & EK, arXiv:1707.03023

Large bispectrum in GW from SU(2) fields

Aniket Agrawal (MPA)

$$\frac{B_h^{RRR}(k,k,k)}{P_h^2(k)} \approx \frac{25}{\Omega_A}$$

Tomo Fujita (Kyoto)

$$\langle \hat{h}_R(\boldsymbol{k}_1)\hat{h}_R(\boldsymbol{k}_2)\hat{h}_R(\boldsymbol{k}_3)\rangle = (2\pi)^3\delta\left(\sum_{i=1}^{3}\boldsymbol{k}_i\right) B_h^{RRR}(k_1,k_2,k_3)$$

- $\Omega_A << 1$ is the energy density fraction of the gauge field
- **B_h/P_h^2 is of order unity for the vacuum contribution**
  [Maldacena (2003); Maldacena & Pimentel (2011)]
- ***Gaussianity offers a powerful test of whether the detected GW comes from the vacuum or sources***

Agrawal, Fujita & EK, arXiv:1707.03023

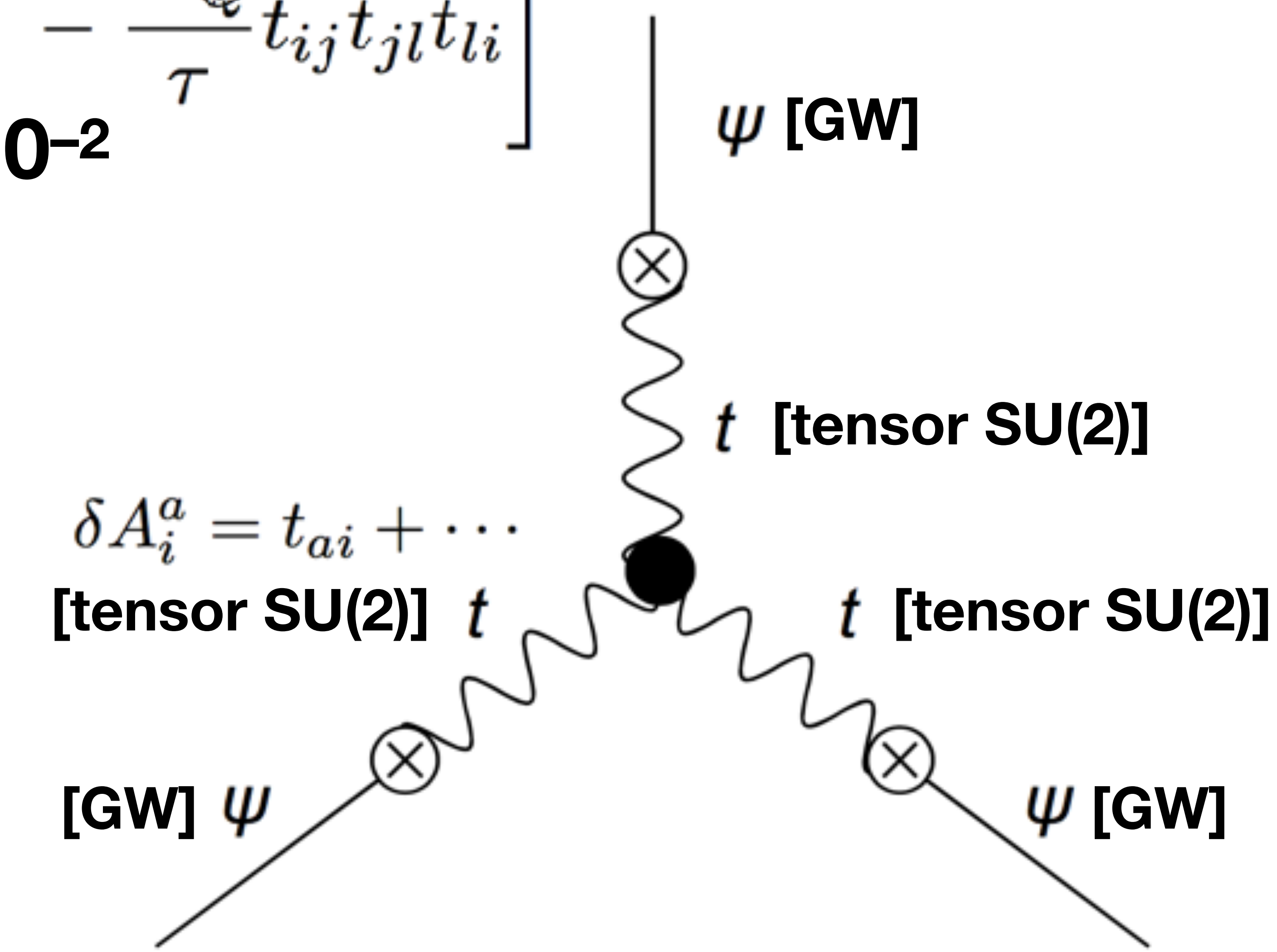
NG generated at the tree level

$$L_3^{(i)} = c^{(i)} \left[\epsilon^{abc} t_{ai} t_{bj} \left(\partial_i t_{cj} - \frac{m_Q^2 + 1}{3 m_Q \tau} \epsilon^{ijk} t_{ck} \right) - \frac{m_Q}{\tau} t_{ij} t_{jl} t_{li} \right]$$

$$c^{(i)} = g = m_Q^2 H / \sqrt{\epsilon_B} M_{\rm Pl}$$ **~10^{-2}**

$$\epsilon_B \equiv \frac{g^2 Q^4}{H^2 M_{\rm Pl}^2} \simeq \frac{2\Omega_A}{1 + m_Q^{-2}} \ll 1$$

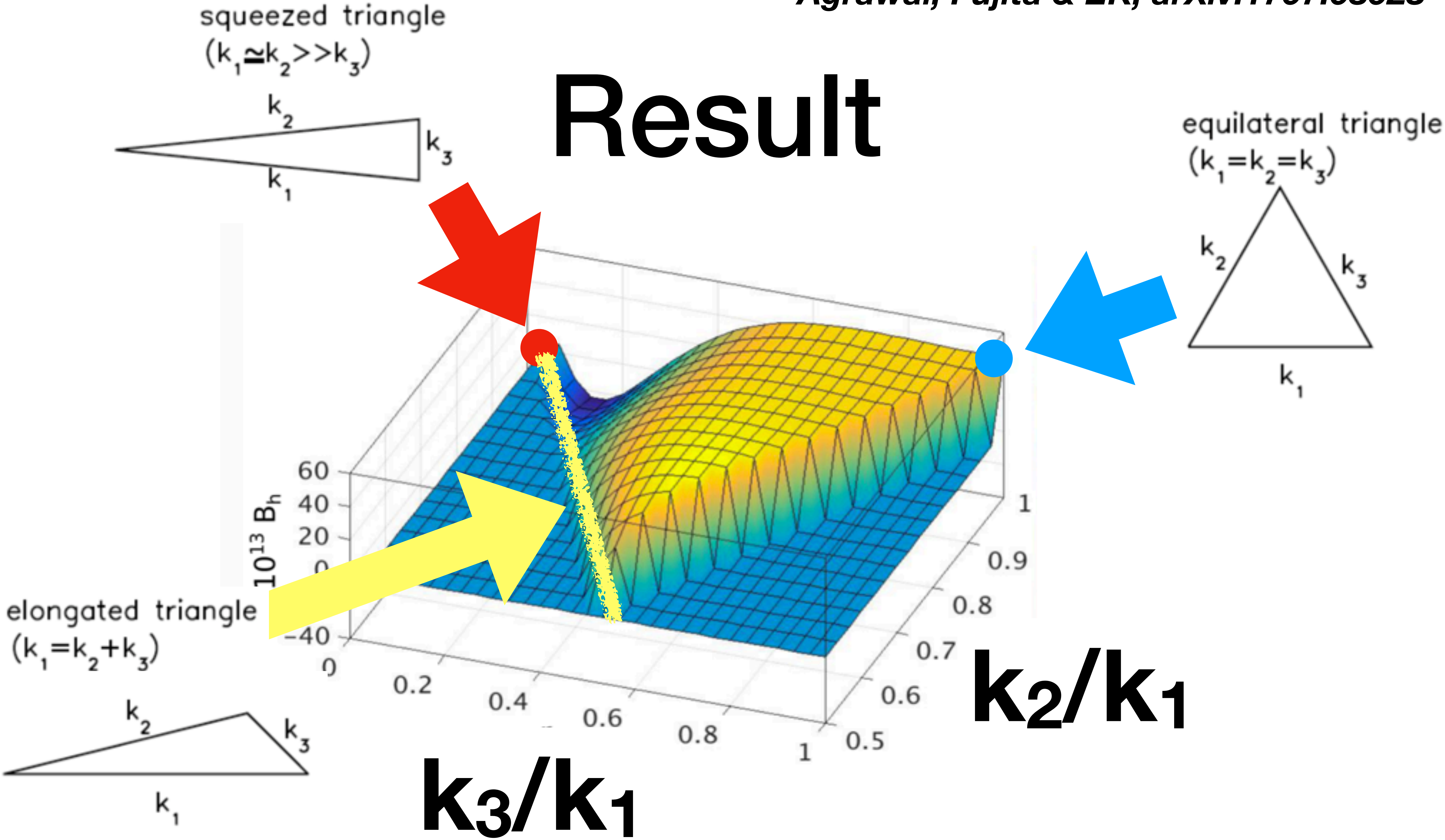
$m_Q \equiv gQ/H$ **[m_Q ~ a few]**



- This diagram generates second-order equation of motion for GW

Agrawal, Fujita & EK, arXiv:1707.03023

Result



- This shape is similar to, but not exactly the same as, what was used by the Planck team to look for tensor bispectrum

Planck Collaboration (2015)

Current Limit on Tensor NG

- The Planck team reported a limit on the tensor bispectrum in the following form:

$$f_{\rm NL}^{\rm tens} \equiv \frac{B_h^{+++}(k,k,k)}{F_{\rm scalar}^{\rm equil.}(k,k,k)}$$

- The denominator is the **scalar** equilateral bispectrum template, giving $F_{\rm scalar}^{\rm equil.}(k,k,k) = (18/5)P_{\rm scalar}^2(k)$

- The current 68%CL constraint is $f_{\rm NL}^{\rm tens} = 400 \pm 1500$

Agrawal, Fujita & EK, arXiv:1707.03023

SU(2), confronted

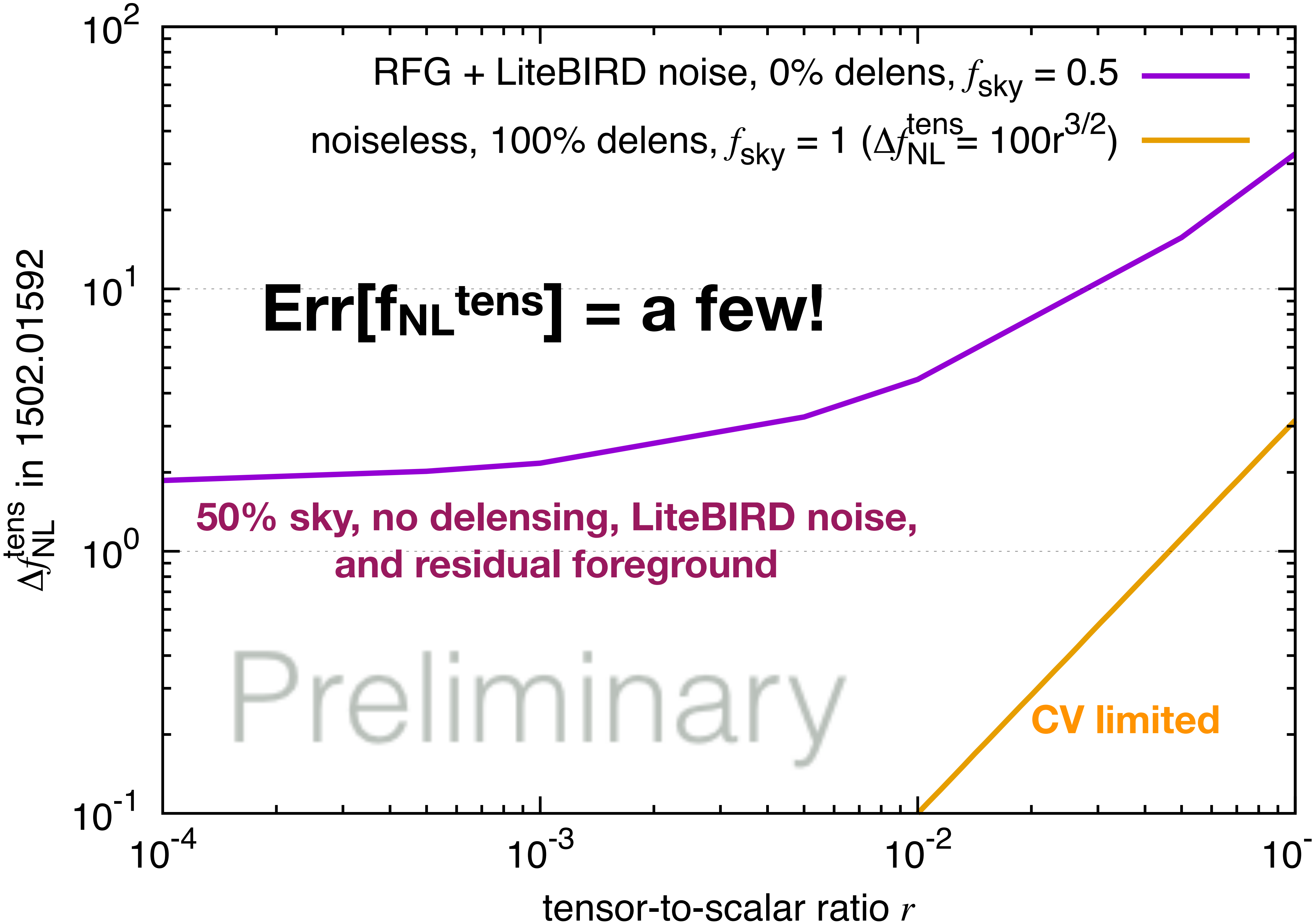
- The SU(2) model of Dimastrogiovanni et al. predicts:

$$f_{\rm NL}^{\rm tens} \approx \frac{125}{18\sqrt{2}}\frac{r^2}{\epsilon_B} \approx 2.5\frac{r^2}{\Omega_A}$$

- The current 68%CL constraint is $f_{\rm NL}^{\rm tens} = 400 \pm 1500$
    - This is already constraining!

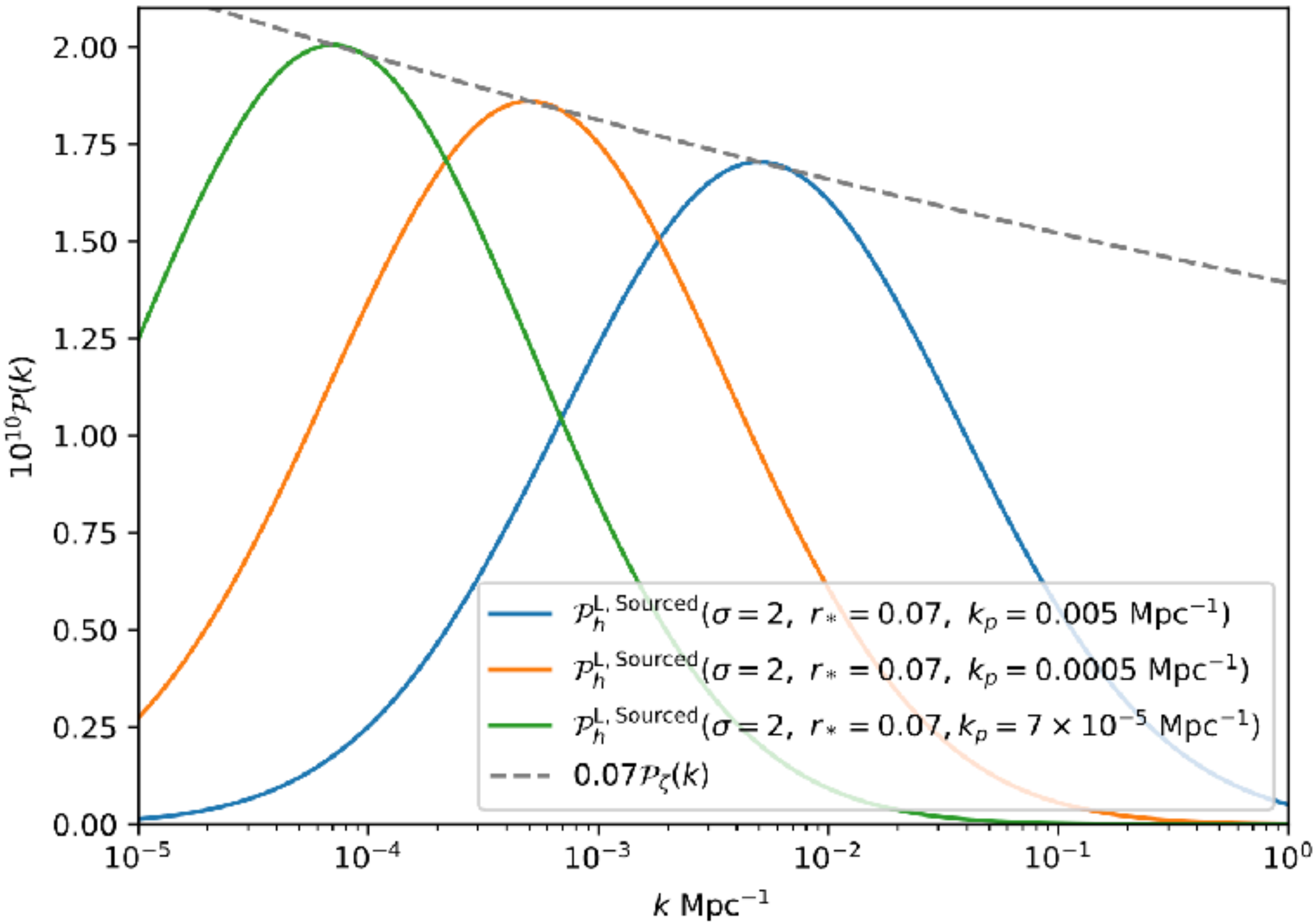
Courtesy of Maresuke Shiraishi

LiteBIRD would nail it!



Dimastrogiovanni, Fasiello & Fujita (2017)
Thorne, Fujita, Hazumi, Katayama, EK & Shiraishi, arXiv:1707.03240

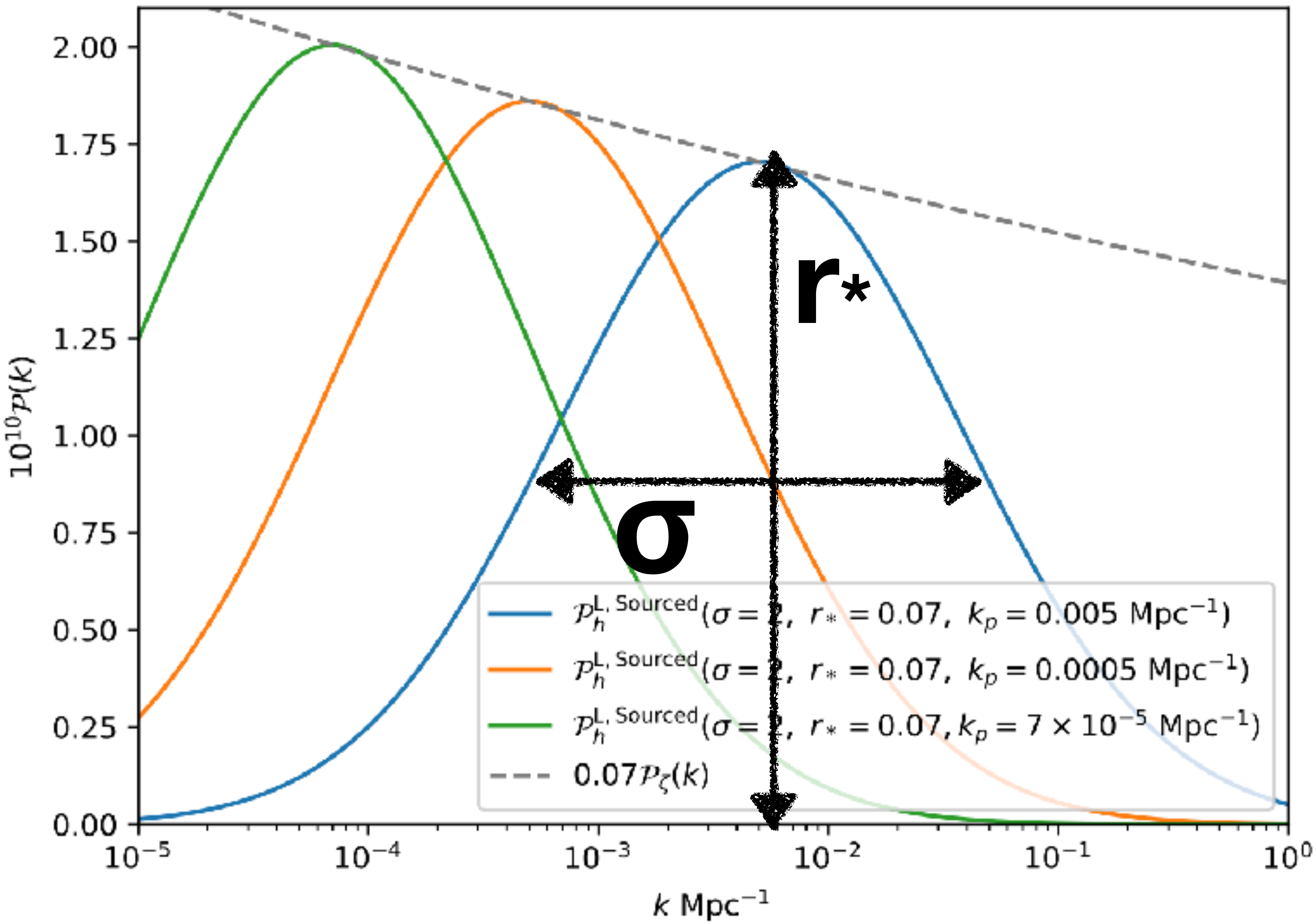
Example Tensor Spectra



- Sourced tensor spectrum can be close to scale invariant, but can also be bumpy

Dimastrogiovanni, Fasiello & Fujita (2017)
Thorne, Fujita, Hazumi, Katayama, EK & Shiraishi, arXiv:1707.03240

Example Tensor Spectra



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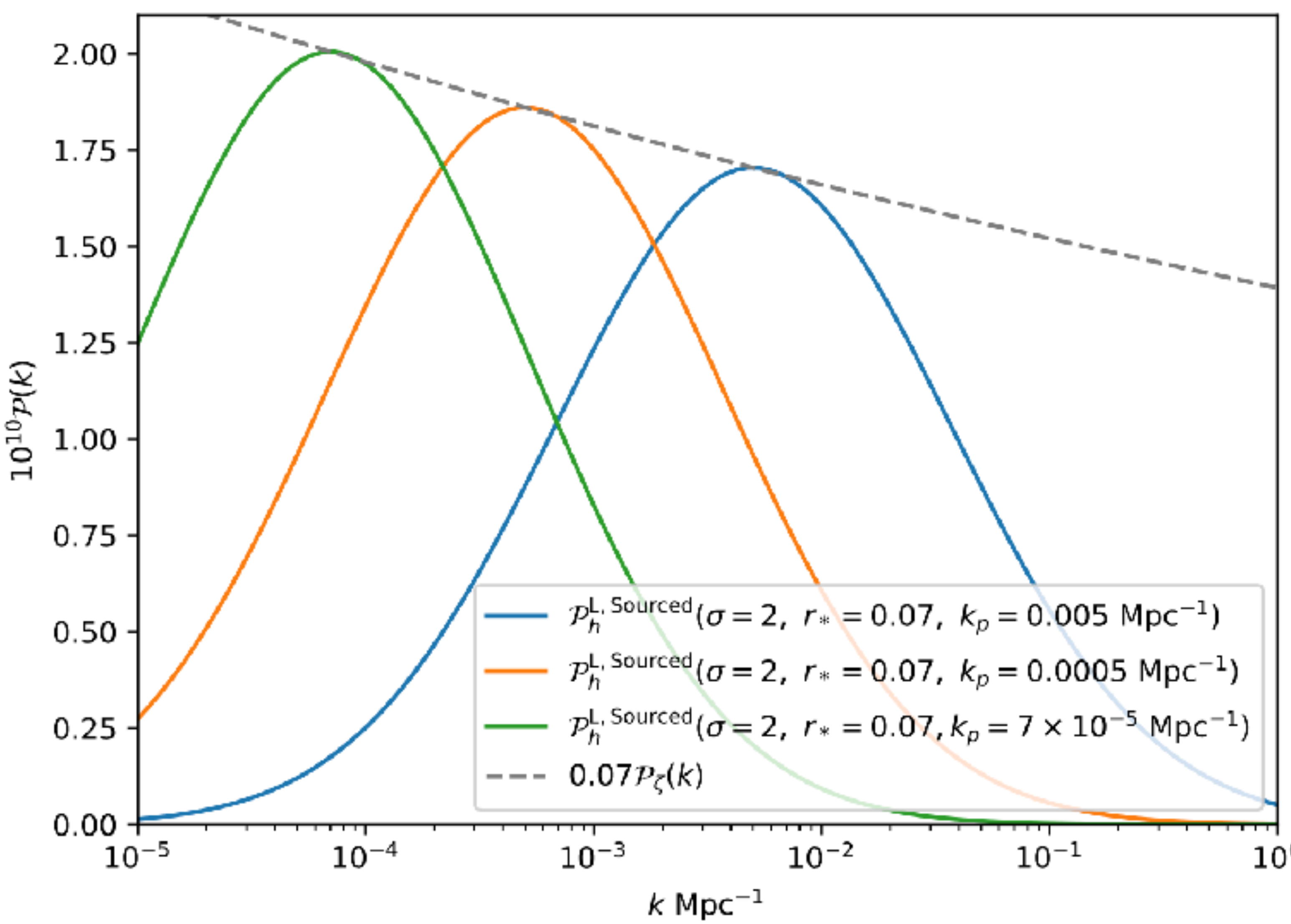
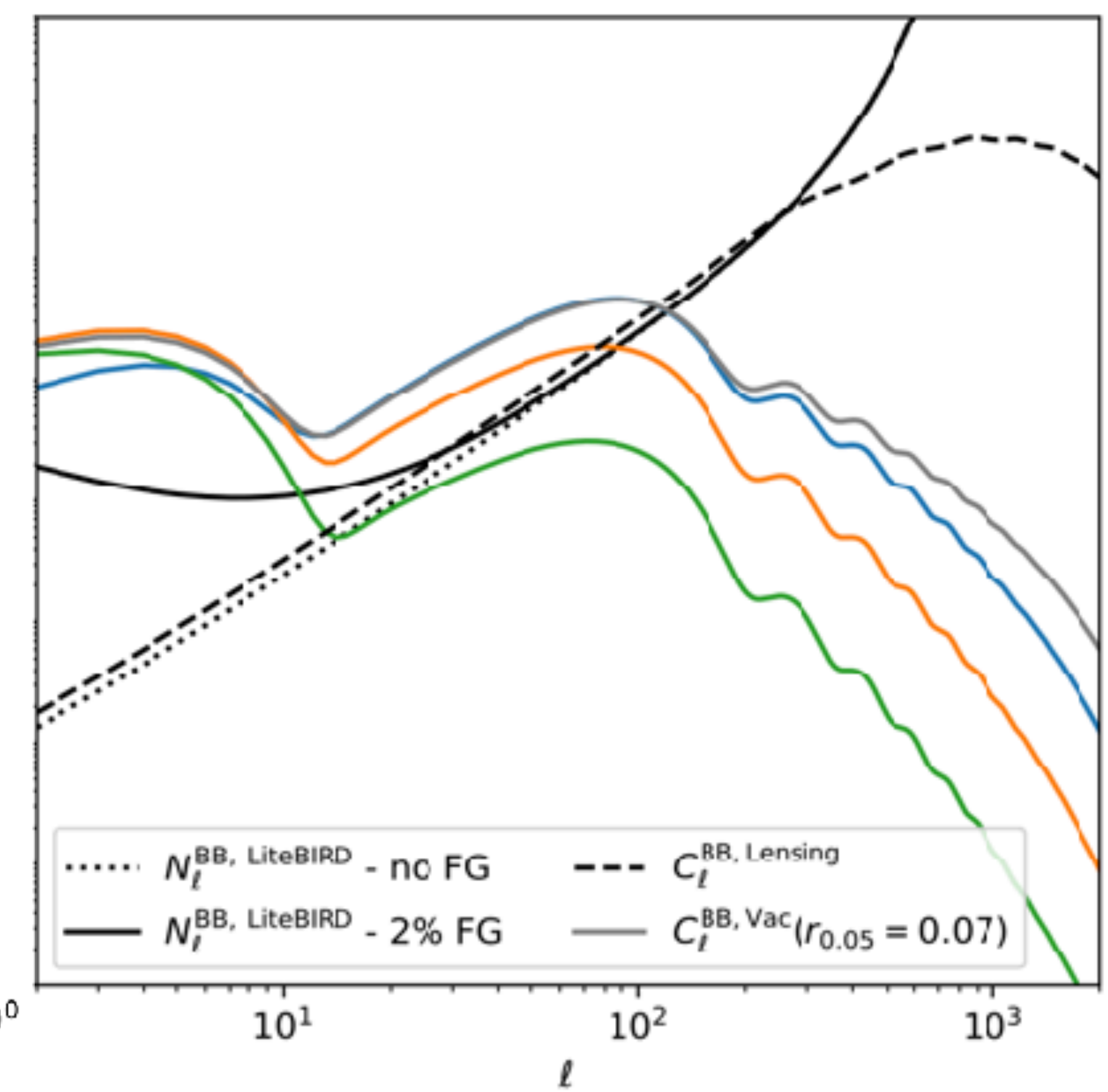
Dimastrogiovanni, Fasiello & Fujita (2017)
Thorne, Fujita, Hazumi, Katayama, EK & Shiraishi, arXiv:1707.03240

Example Tensor Spectra

Tensor Power Spectrum, P(k)



B-mode CMB spectrum, C_l^{BB}



- Sourced tensor spectrum can be close to scale invariant, but can also be bumpy