

# Complex X-ray morphology of Abell 3128: A distant cluster behind a disturbed cluster

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**Abstract.** We present here the results of a detailed study of the X-ray properties of the cluster of galaxies Abell 3128 ( $z = 0.06$ ), based on the analysis of deep (100 ks) XMM-Newton data. The most obvious feature of the X-ray morphology of A3128 is the presence of two X-ray peaks separated by  $\sim 12'$ . By detecting the redshifted Fe K line, we find that the Northeast (NE) X-ray peak observed toward A3128 is a distant luminous cluster of galaxies at redshift  $z = 0.44$ . Our subsequent optical spectroscopic observation of a distant radio bright galaxy in the centre of the NE X-ray peak with the Magellan telescope also revealed a redshift of  $z = 0.44$ , confirming the association of the galaxy with the cluster seen in X-rays. We detect a gravitational arc around the galaxy. The properties of this galaxy indicate that it is the cD galaxy of the cluster in the background. The properties of the Southwest X-ray peak suggest that it is the core of a group merging with A3128 along our line of sight. Based on 2D maps of thermodynamic properties of the intra-cluster medium determined after subtracting a model for the background cluster, we conclude that an enhanced surface brightness region at a distance of  $\sim 2.8'$  from the centre of the galaxy distribution is the centre of the gravitational potential of the cluster A3128. The unrelaxed nature of A3128 can be attributed to its location in the high density environment of the Horologium-Reticulum supercluster.

**Key words.** galaxies: clusters: general – galaxies: clusters: individual: Abell 3128 – X-rays: galaxies: clusters

## 1. Introduction

The formation and growth of the largest structures in the Universe, clusters of galaxies and superclusters, is a still ongoing process. According to the standard cosmological scenario clusters of galaxies continuously form and grow through merging with groups and individual galaxies, which are falling in along filaments. In several cases clusters are found close to each other in the highest overdensity regions of the Universe: the superclusters. In the crowded environment of superclusters, we can often study processes related to the growth of structure, like mutual interaction between the member clusters and their accretion history.

The Horologium-Reticulum (H-R) supercluster ( $z = 0.06$ ), together with the better known Shapley Supercluster are the most massive structures in the local universe within a distance of 300 Mpc, with an estimated total mass of  $10^{17} M_{\odot}$ . Einasto et al. (2001) list 35 member clusters in H-R.

Abell 3128 is a rich, highly substructured cluster in the H-R supercluster. It has been well studied using redshift and radio

surveys, and by Chandra (20 ks; Rose et al. 2002). Based on  $N$ -body simulations, Caldwell & Rose (1997) proposed that in the past A3128 encountered a tidal interaction with the cluster A3125 (current separation  $1^{\circ}$ , corresponding to  $\sim 6$  Mpc). Rosat and Chandra observations of A3128 show a complex, disturbed X-ray morphology. Embedded in a  $\sim 20'$  diameter diffuse halo there are two cores in A3128, separated by  $12'$ . The Northeast (NE) and Southwest (SW) cores comprise  $\sim 85\%$  of the X-ray emission of A3128, the remainder being in a diffuse halo. While the more luminous SW core is centred on a compact group around a bright elliptical galaxy, the NE core does not coincide with any bright galaxy. The NE core is asymmetric and elongated, with the peak of the emission slightly displaced to the Northeast. With core radii of  $\sim 30$  kpc, both cores are unusually narrow as compared to other non-cooling core clusters with similar temperatures ( $\sim 3.5$  keV).

Rose et al. (2002) proposed a model to explain the double-peaked nature of the X-ray emission. In position-position and position-redshift diagrams they identified an infalling group and a candidate for a post-passage tidally distended group (filament), which they propose as candidates for producing the

complex structure seen in X-ray images of the intra-cluster medium (ICM). They explain the SW core as the still intact hot gas of an infalling group. The NE core represents according to Rose et al. (2002) the surviving ICM of a group of galaxies that has fallen supersonically (Mach number 6) into the cluster A3128 along the main filament connecting A3128 and A3125. The infalling group has, according to this scenario, passed through the core of A3128 and the galaxies of the group, moving ballistically through the cluster, are by now well ahead of the stripped gas. The high infall velocity is caused by the deep potential well of the H-R supercluster, which makes even an infall of a small group an energetic and interesting event in the life of a galaxy cluster. Rose et al. (2002) also report a possible detection of a radio arc, which lies slightly to the Northeast of the peak emission in the NE core, where the bow shock is expected if the gas is moving supersonically.

However, the Fe abundance of the NE core found by Rose et al. (2002) is low (0.13 solar) and barring conclusive redshift information its association with A3128 could not be confirmed. As an alternative scenario to their unified view of the merging events occurring in A3128/A3125, Rose et al. (2002) invoke the possibility that the NE emission peak is associated with a background cluster and the detected radio emission is associated with a radio bright central galaxy in a background cluster. In order to obtain a conclusive measure for placing of the NE component, a deeper observation for a redshift measurement using X-ray emission lines was needed.

Here we present the results of a deep 100 ks observation of Abell 3128 with the European Photon Imaging Cameras (EPIC) on XMM-Newton. The observation allows us to measure the redshift of the NE core and to unambiguously demonstrate that the dominant fraction of its emission originates in a background cluster. We also present here data obtained by the Magellan telescope, that allow us to confirm the association of the X-ray emission with a background cluster surrounding a radio bright cD galaxy. The large effective area of XMM-Newton combined with the deep exposure, provides us with sufficient statistics to map the 2D distribution of thermodynamic properties of the ICM and thus to study the merging history of the disturbed, complex cluster of galaxies A3128. The XMM-Newton Reflection Grating Spectrometer data have insufficient statistics to derive accurate spectral properties for the cluster.

Throughout the paper we use  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_M = 0.3$ ,  $\Omega_\Lambda = 0.7$ , which imply a linear scale of  $78 \text{ kpc arcmin}^{-1}$  at the redshift of A3128 ( $z = 0.06$ , Rose et al. 2002). Unless specified otherwise, all errors are at the 68% confidence level for one interesting parameter ( $\Delta\chi^2 = 1$ ). The elemental abundances are given with respect to the proto-solar values of Lodders (2003).

## 2. Observations and data analysis

### 2.1. XMM-Newton data

Abell 3128 was observed with XMM-Newton during two pointings on May 29–30 and between May 31 and June 1st 2006 (revolutions 1185 and 1186) for a total exposure time of  $\sim 100$  ks. The two pointings were centred on the NE (72 ks)

and on the SW core (32 ks), respectively. The EPIC/MOS detectors were operated in the full frame mode, while for the EPIC/pn detector the extended full frame window mode was employed. The observations were performed using the thin filter. The calibrated event files were obtained using the 7.0.0 version of the XMM-Newton Science Analysis System (SAS). For EPIC/MOS we keep only the single, double, triple, and quadruple pixel events ( $\text{PATTERN} \leq 12$ ), while for EPIC/pn, we make use of single and double events ( $\text{PATTERN} \leq 4$ ). The spectral redistribution and ancillary response files are created with the SAS tasks `rmfgen` and `arfgen` separately for each camera and spectral extraction region that we analyze.

Because of low number of counts in original bins, the extracted spectra are rebinned into bins with a minimum of 30 counts per bin. Our rebinned spectrum has bins of size at least  $1/4$  times the FWHM of the instrument. We fit the MOS1, MOS2, and pn spectra from both pointings simultaneously with the same model, with their relative normalizations left as free parameters. Our spectral analysis is restricted to the 0.4–7.0 keV band. In the spectral analysis we remove all bright point sources with a flux higher than  $4.8 \times 10^{-14} \text{ ergs s}^{-1} \text{ cm}^{-2}$ .

For the spectral analysis we use the SPEX package (Kaastra et al. 1996). We fix the Galactic absorption in our model to the value deduced from the H I data ( $N_{\text{H}} = 1.47 \times 10^{20} \text{ cm}^{-2}$ , Dickey & Lockman 1990). We fit the cluster spectra with a plasma model in collisionally ionized equilibrium (MEKAL). The free parameters in the MEKAL model are the normalisation, temperature, and the Fe abundance. Since for elements other than Fe we can not obtain accurate abundance values, we fix them to 0.4 solar in the model.

### 2.2. Background modeling

#### 2.2.1. Particle and instrumental background

To clean the data from the soft proton induced events, we extract light curves for each detector separately in the 10–12 keV energy band where the cluster emission is negligible and the detected emission is dominated by the particle induced events. Since flares having a particularly soft spectrum may be missed when only the high energy part is studied, we also extract lightcurves in the 0.3–2.0 keV band. A visual inspection of the light curves reveals that the observation is not badly affected by the soft proton flares. After excluding the time periods when the count rate in the two considered energy ranges deviates from the mean by more than  $3\sigma$ , we are left for the first pointing with 71 ks and 67 ks, and for the second pointing with 31 ks and 25 ks for EPIC/MOS and EPIC/pn, respectively.

We subtract the EPIC instrumental background using closed-filter observations. For EPIC/MOS we use a closed filter observation (obs. ID 0150390101) with an exposure time of 200 ks. For EPIC/pn we use a closed filter observation (obs. ID 0106660401) with an exposure time of 120 ks. The instrumental background consists of fluorescence line emission, intrinsic instrumental noise, and particle induced noise caused by high-energy cosmic rays which are able to reach the detector even when the filter wheel is in closed position. The instrumental background varies from observation to observation. In order to

**Table 1.** The CXB components used in the fitting of the A3128 spectra. The power-law photon index is  $\Gamma = 1.41$ . The unabsorbed fluxes are determined in the 0.3–10.0 keV band.

Comp.	$kT$ (keV)	Flux ( $10^{-12}$ erg $\text{cm}^{-2}$ $\text{s}^{-1}$ $\text{deg}^{-2}$ )
LHB/SDC	0.08	$2.5 \pm 0.6$
HDC	$0.25 \pm 0.02$	$4.0 \pm 0.6$
EPL		$22.0 \pm 1.9$

scale the closed filter observations to the instrumental background during the source observation, we use the events registered outside the field of view (out-of-FOV) of the EPIC detectors, outside a radius of 15.4' from the FOV centre. Separately for each instrument, we scale the closed filter observation by the ratio of the 8–12 keV out-of-FOV count rate in our cluster observation and in the closed filter observation.

### 2.2.2. The cosmic X-ray background

We correct for the Cosmic X-ray Background (CXB) during spectral fitting. Kuntz & Snowden (2000) distinguish 4 different background/foreground components: the extragalactic power-law (EPL), the local hot bubble (LHB), the soft distant component (SDC), and the hard distant component (HDC). The EPL component is the integrated emission of faint discrete sources, mainly distant Active Galactic Nuclei (AGNs). The LHB is a local supernova remnant, in which our Solar System resides. It produces virtually unabsorbed emission at a temperature of  $\sim 10^6$  K. The soft and hard distant components originate at larger distances. They might be identified with the Galactic halo, Galactic corona or the Local group emission and are absorbed by almost the full Galactic column density. Using the spectral band above 0.4 keV we can not reliably distinguish the SDC emission from the LHB component. Therefore, at temperatures below 0.1 keV we only consider the contribution of one thermal component.

The large field of view, with regions where the contribution of cluster emission is small, enables us to determine the local properties of the background emission. We fit the spectra extracted from a region where the contribution of cluster emission is small and the emission is dominated by the X-ray background. We model the EPL emission with a power-law with a photon index of 1.41 (De Luca & Molendi 2004). We model the soft foreground emission (LHB/SDC and HDC) by 2 collisionally ionized plasma models (MEKAL). Since the fitted spectral band does not allow us to accurately constrain the LHB/SDC temperature, we fix its value to 0.08 keV (based on Kuntz & Snowden 2000). We leave the HDC temperature as a free parameter in the fit. To account for the remaining cluster emission in the extraction region, the contribution of which to the total flux is  $\sim 30\%$ , we use an additional thermal model. We find that both the temperature and the flux of the emission attributed to the HDC (0.25 keV and  $4.0 \times 10^{-12}$  erg  $\text{s}^{-1}$   $\text{cm}^{-2}$   $\text{deg}^{-2}$ , respectively) are higher than the values determined based on the results published by Kuntz & Snowden (2000) using the ROSAT All Sky Survey

data (0.127 keV and  $2.44 \times 10^{-12}$  erg  $\text{s}^{-1}$   $\text{cm}^{-2}$   $\text{deg}^{-2}$ , respectively). This difference might be due to the contribution of the intra-supercluster medium. The adopted values of the temperatures and of the 0.3–10.0 keV fluxes of the background components are given in Table 1. In our subsequent spectral fits we fix the parameters of the background components to these adopted values.

### 2.3. The X-ray images

The spatial resolution of XMM-Newton allows a robust image reconstruction down to 8'' scales. However, due to the large wings of the point spread function (PSF), there is a significant contamination from the emission on scales of 8'' or smaller to large spatial scales upto several arcminutes. For the purpose of cluster survey work (Finoguenov et al. in prep.), we have developed a procedure, which performs an image restoration using a symmetric model for the XMM PSF and the calibration of Ghizzardi (2001). The flux on small scales is estimated performing the scale-wise wavelet analysis, as described in Vikhlinin et al. (1998). Then the estimated flux is used to subtract the PSF model prediction on large scales and increase accordingly the flux on the small scales. The small scales in this procedure are a sum of 8'' and 16'' scales, which allows to avoid the variation in the PSF shape with off-axis angle, as described in Finoguenov et al. (2007, COSMOS special issue paper).

In Fig. 1 we display both the large scales of the emission starting at the 32'' scale together with smaller scales (8'' and 16'').

## 3. Results

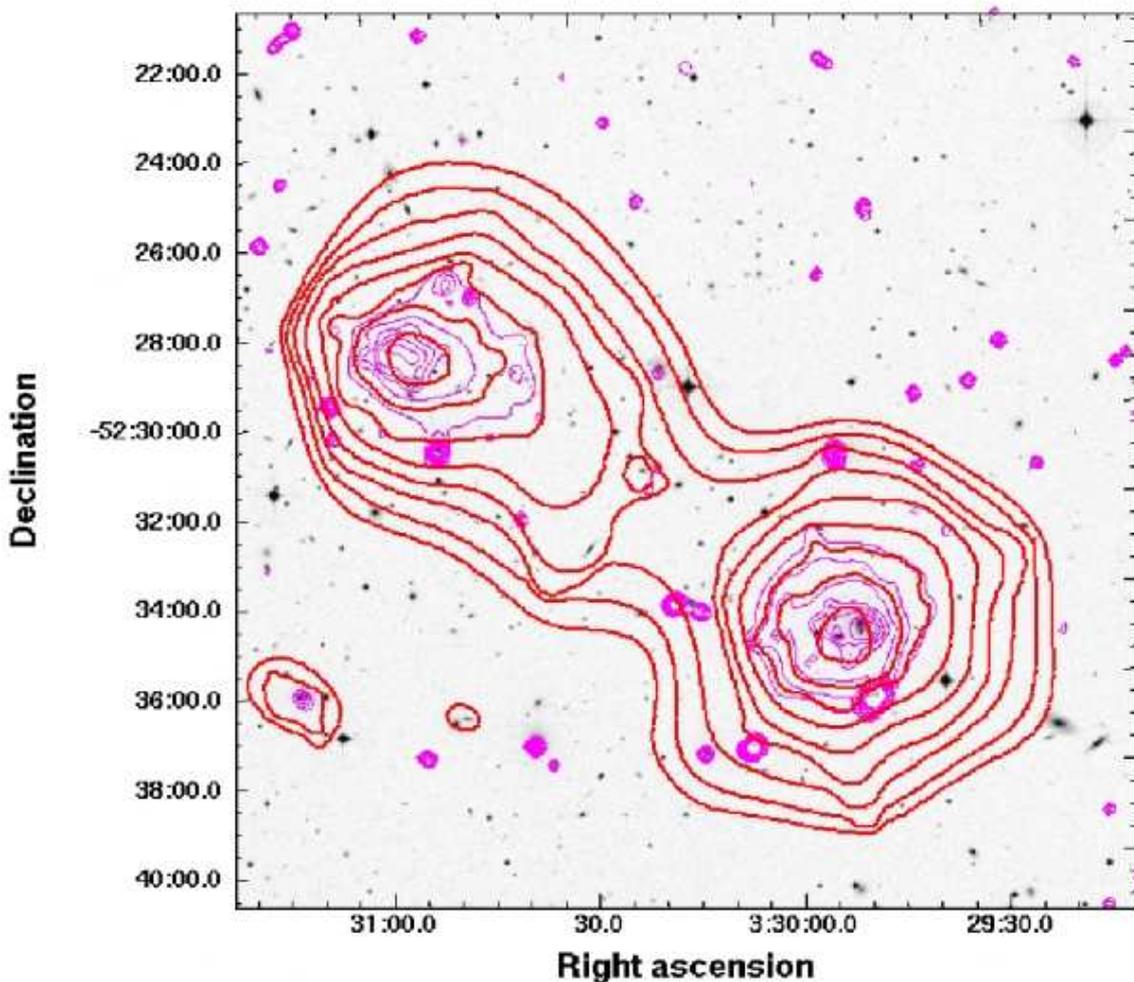
The most obvious feature of the X-ray morphology of A3128 is the presence of two X-ray peaks separated by  $\sim 12'$ . In Fig. 1 we show the X-ray contours superimposed on the Digitalized Sky Survey image of the cluster. The thick contours highlight the large scale (32'') features in X-ray morphology, the thin contours highlight the features on small spatial scales (8'' and 16'').

While the SW peak of the X-ray intensity coincides with a bright elliptical galaxy (ENACS 75) surrounded by a compact group of galaxies, the NE core does not coincide with any bright galaxy. On the small spatial scales the NE peak has a strongly elongated morphology and on larger scales it seems to have a tail toward the SW direction. The thick contours also reveal an association of the hot gas with two small groups of galaxies to the East of the main cluster contours.

### 3.1. Properties of the X-ray peaks

#### 3.1.1. The global properties

In order to study the global spectral properties of the two X-ray bright peaks, we extract their spectra from circular regions with a radius of 1.5'. In Fig. 2 we show the spectra with their best fit model. Previous Chandra data showed that while the SW X-ray core has a high Fe abundance, the Fe abundance of the NE



**Fig. 1.** X-ray contours superimposed on the Digitalized Sky Survey image of A3128. North is up, West is to the right. The thick contours highlight the large scale features in the X-ray morphology, the thin contours highlight the features at small spatial scales. The contour spacing is arbitrary.

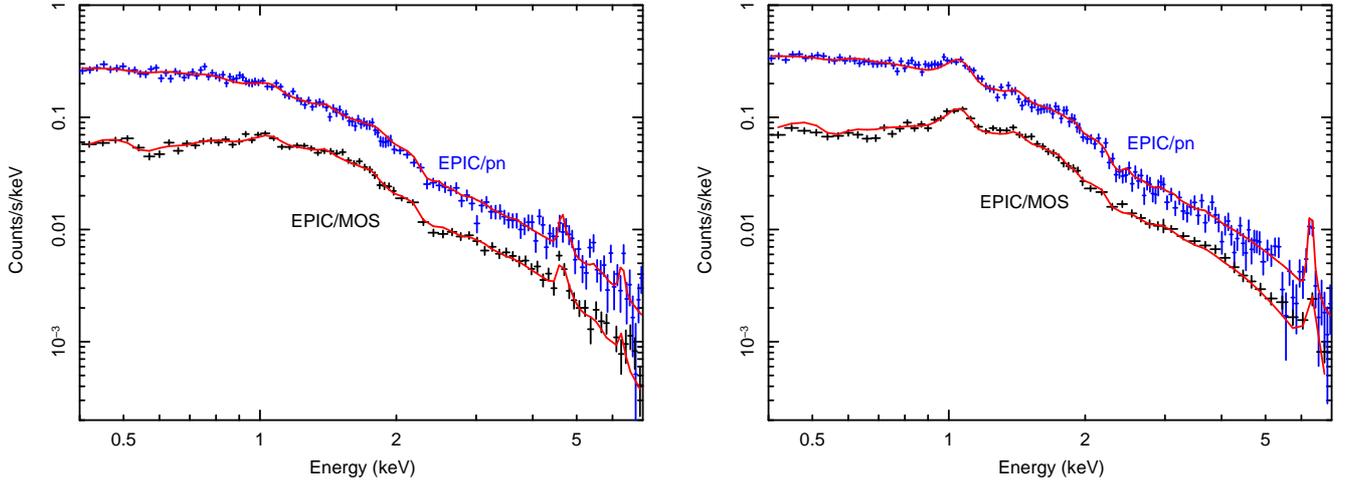
core is very low, and therefore it is also not possible to confirm its association with the cluster A3128 using the redshift of the Fe K line. Even using the deep XMM-Newton data we do not see a clear Fe K emission line at the expected energy for the redshift of A3128. However, we detect in all EPIC detectors for both XMM-Newton pointings a line at  $\sim 4.6$  keV. This corresponds to the Fe K line emission of a cluster at a redshift of  $\sim 0.45$ .

We also extract a spectrum from a circular region with a radius of  $2'$  centred on the apparent tail of the NE X-ray peak directed to the SW ( $\alpha = 3^{\text{h}}30^{\text{m}}30^{\text{s}}$ ,  $\delta = -52^{\circ}31'0''$ ). We immediately see an Fe K line at  $\sim 6.3$  keV and we see no obvious line feature around 4.6 keV. We fit the spectrum extracted from the “tail” region and from the SW X-ray peak with a thermal model. The free parameters of the fit are the redshift, temperature, Fe abundance, and emission measure. The best fit parameters are shown in Table 2.

We fit the spectrum of the NE peak with a combination of two thermal plasma models: one for the emission of A3128 and the other for the emission of the background cluster. We fix the parameters of the model of the foreground cluster to the

values determined for the tail region: redshift  $z = 0.06$ , metallicity  $Z = 0.49$ , and temperature  $kT = 3.5$  keV. The emission measure of the foreground cluster and the redshift, Fe abundance, temperature, and emission measure of the background cluster are free parameters in the spectral fit. The best fit redshift value for the background cluster is  $z = 0.444 \pm 0.003$ . The linear scale at this redshift for the  $\Lambda$ CDM cosmology is  $341 \text{ kpc arcmin}^{-1}$ . The best fit temperature, Fe abundance, and emission measure of the background cluster are shown in the column “NE” of Table 2. Our spectral analysis indicates that  $\sim 60\%$  of the X-ray emission in the extraction region is coming from the background cluster. The  $0.3\text{--}10.0$  keV luminosity of the background cluster within the radius of  $1.5'$  (512 kpc) is  $4.7 \times 10^{44} \text{ ergs s}^{-1}$ .

The best fit parameters of the background cluster depend on the assumptions about the emission of the foreground cluster. Therefore, the systematic uncertainties on the best fit parameters are larger than the quoted statistical uncertainties. For a lower assumed temperature for the foreground cluster, we fit a higher temperature for the background cluster. For example, if we assume that the temperature of the foreground cluster is



**Fig. 2.** The spectrum of the NE (left panel) and SW (right panel) X-ray peak extracted from circular region with a radius of  $1.5'$ . The continuous line represents the best fit model. The EPIC/pn and the coadded EPIC/MOS spectra are indicated. The position of the Fe K line in the spectrum from the NE peak shows that the emitting cluster is at a redshift of  $z = 0.44$ .

**Table 2.** Fit results for the spectra extracted from the SW X-ray peak, from the apparent “tail” of the NE peak toward SW, and from the NE X-ray peak. The spectra extracted from the SW peak and from the “tail” region were fitted with a thermal model with the  $N_{\text{H}}$  fixed to the Galactic value ( $N_{\text{H}} = 1.47 \times 10^{20} \text{ cm}^{-2}$ , Dickey & Lockman 1990). The NE X-ray peak was fitted with two thermal models: one for the emission from A3128 in the foreground and one for the background cluster. In this fit the redshift, temperature, and the Fe abundance of the foreground cluster were fixed to the best fit values from the “tail” region. The reported best fit parameter of the NE peak describe the properties of the background cluster. The temperature is given in keV, the emission measure  $Y$  is given in units of  $10^{65} \text{ cm}^{-3}$  and the Fe abundance is given with respect to the proto-solar values of Lodders (2003).

Par.	SW	“tail”	NE
$Y_{z=0.06}$	$9.72 \pm 0.10$	$6.80 \pm 0.07$	$3.4 \pm 0.4$
$Y_{z=0.44}$	–	–	$303 \pm 34^{\dagger}$
$kT$	$3.36 \pm 0.05$	$3.50 \pm 0.08$	$5.14 \pm 0.15^{\dagger}$
Fe	$0.69 \pm 0.04$	$0.49 \pm 0.04$	$0.47 \pm 0.08^{\dagger}$
$z$	$0.060 \pm 0.001$	$0.058 \pm 0.003$	$0.444 \pm 0.003^{\dagger}$
$\chi^2 / \nu$	652/501	435/398	943/800

$^{\dagger}$  values determined for the background cluster.

3.0 keV instead of 3.5 keV, the best fit temperature of the background cluster will be 5.8 keV instead of 5.1 keV.

We note, that there is no evidence of emission from the background cluster in the “tail” region. From the spectral analysis we find that its contribution in this region is less than 1%.

### 3.1.2. Radial profiles

We determine the emission measure, temperature, and Fe abundance profiles for both X-ray peaks. We extract the spectra from circular annuli with outer radii of  $0.5'$ ,  $1.5'$ ,  $2.5'$ , and  $3.5'$ .

**Table 3.** Radial profiles determined for the SW X-ray peak. The temperature is given in keV, the emission measure  $Y$  is given in units of  $10^{65} \text{ cm}^{-3}$ , and the Fe abundance is given with respect to the proto-solar values of Lodders (2003).

Par.	$0.0'-0.5'$	$0.5'-1.5'$	$1.5'-2.5'$	$2.5'-3.5'$
$Y$	$1.99 \pm 0.05$	$7.91 \pm 0.09$	$8.67 \pm 0.10$	$9.57 \pm 0.12$
$kT$	$3.36 \pm 0.11$	$3.36 \pm 0.06$	$3.17 \pm 0.06$	$3.38 \pm 0.07$
Fe	$0.87 \pm 0.10$	$0.67 \pm 0.04$	$0.48 \pm 0.06$	$0.46 \pm 0.04$
$\chi^2 / \nu$	269/229	1125/959	1183/1072	1273/1137

**Table 4.** Radial profiles determined for the NE X-ray peak. The metallicity and the temperature of the foreground cluster were fixed to 0.49 solar and 3.5 keV, respectively. The emission measure  $Y$  is given in units of  $10^{65} \text{ cm}^{-3}$  and the temperature is given in keV. The Fe abundance is given with respect to the proto-solar values of Lodders (2003). The  $f_{z=0.44}/f_{z=0.06}$  indicates the ratio of the fluxes from the background cluster and A3128 in the given extraction region. The superscript “ $f$ ” indicates that the parameter value was fixed during the fitting.

Par.	$0.0'-0.5'$	$0.5'-1.5'$	$1.5'-2.5'$	$2.5'-3.5'$
$Y_{z=0.06}$	$0.56 \pm 0.20$	$2.7 \pm 0.4$	$4.7 \pm 0.4$	$5.6 \pm 0.3$
$Y_{z=0.44}$	$72 \pm 15$	$178 \pm 27$	$154 \pm 34$	$< 140$
$kT$	$5.0 \pm 0.3$	$5.7 \pm 0.3$	$3.5 \pm 0.3$	$4^f$
Fe	$0.56 \pm 0.18$	$0.49 \pm 0.09$	$0.21 \pm 0.11$	$0.2^f$
$\frac{f_{z=0.44}}{f_{z=0.06}}$	2.1	1.2	0.55	0.09
$\chi^2 / \nu$	260/266	827/749	986/841	1109/870

In Table 3 we show the best fit parameters determined for the SW X-ray peak. The temperature profile of this peak is flat and there is no evidence for a cool core. However, the Fe abundance has a strong peak in the centre of the X-ray emission, which coincides with an elliptical galaxy surrounded by a group of galaxies. The abundance peaks within the central  $1.5'$ , which corresponds to  $\sim 120$  kpc.

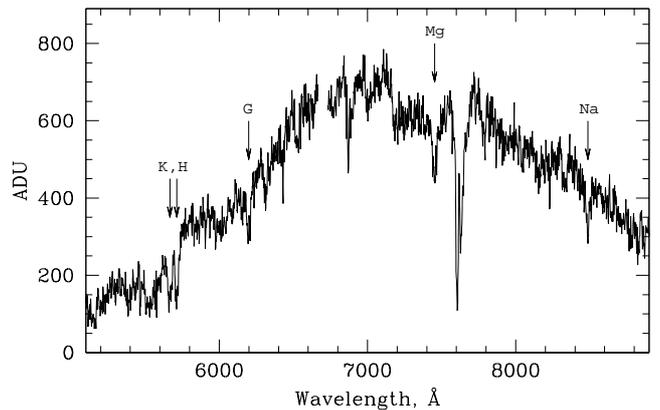
Fitting a beta model (Cavaliere & Fusco-Femiano 1978) to the radial surface brightness distribution of the SW X-ray peak we find  $\beta = 0.3$  and a core radius of  $r_c = 30$  kpc. As already noted by Rose et al. (2002) this core radius is about an order of magnitude lower than the more typical values of  $\sim 250$  kpc that are observed for non-cooling core clusters of galaxies. The  $\beta$  value is also low compared to the typical value of  $\sim 0.7$ . The small value of  $\beta$  shows that the surface brightness distribution of the core is broader than that usually observed for clusters, which indicates that we might see the emission of a group superimposed on the diffuse cluster emission.

Using the parameters of the  $\beta$  model fit, and a global temperature value of  $kT = 3.4$  keV, we find that the mass enclosed within the radius of 120 kpc is  $1.2 \times 10^{13} M_\odot$ . The total gas mass within the same volume, determined using a central electron density of  $1.2 \times 10^{-2} \text{ cm}^{-3}$ , is  $\sim 1.3 \times 10^{12} M_\odot$ . Using the same parameters, the estimated total mass within a radius of 1 Mpc is  $\approx 1.1 \times 10^{14} M_\odot$ . A radius of 1 Mpc is approximately the virial radius of a 3.4 keV cluster (e.g. Finoguenov et al. 2001).

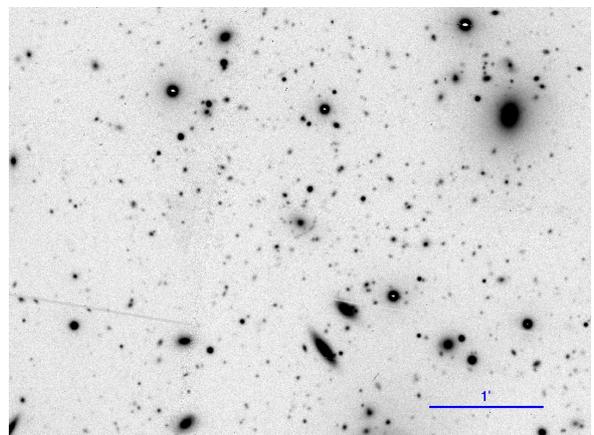
The total Fe mass enclosed within the radius of 120 kpc, the region within which the Fe abundance peaks, is  $\sim 1.5 \times 10^9 M_\odot$ . Assuming a flat Fe abundance distribution in A3128 of 0.45 solar, the excess Fe mass in the SW X-ray peak is  $\approx 5 \times 10^8 M_\odot$ .

The radial profiles determined for the background cluster seen as the NE X-ray surface brightness peak are shown in Table 4. For the temperature and Fe abundance of the foreground cluster we assume 3.5 keV and 0.49 solar, respectively. The free parameters in the fit are the temperature and the Fe abundance of the background cluster, and the emission measures of both clusters. We see an indication of a temperature drop in the inner 170 kpc, then the temperature peaks, and outside 500 kpc it drops again. The spectra indicate that the Fe abundance peaks in the cluster core. In Table 4 we also show for each annulus the ratio of the flux of the background cluster with respect to the flux of the foreground cluster. In the innermost extraction region the emission of the background cluster clearly dominates, between  $0.5' - 1.5'$  the contributions of both clusters are similar. Between  $1.5' - 2.5'$  the emission of the foreground cluster dominates, and between  $2.5' - 3.5'$  we can only determine an upper limit for the emission of the background cluster. The total 0.3–10.0 keV luminosity of the  $z = 0.44$  cluster within the radius of  $3.5'$  is  $6.9 \times 10^{44} \text{ ergs s}^{-1}$ .

To obtain a rough estimate of the total mass of the background cluster, we fit a beta model to the radial profile of its best fit emission measure per arcminute. The best-fit parameter values are  $r_c = 157$  kpc and  $\beta = 0.41$ . Assuming a global temperature of 5.14 keV (the best fit global temperature from Sect. 3.1.1) we obtain a total mass within a radius of 1.5 Mpc (estimated  $r_{500}$  of a 5.1 keV cluster, e.g. see Finoguenov et al. 2001) of  $3.4 \times 10^{14} M_\odot$ . We note that due to the uncertainties associated with the subtraction of the foreground cluster, and because the core of the background cluster is far from being relaxed and it is not single-temperature, this derived mass is only a very rough estimate.



**Fig. 3.** Spectrum of the galaxy associated with the radio source SUMSS J033057-522811 in the centre of the NE X-ray peak. On the y-axis we plot arbitrary count units (ADU). Lines, used for redshift determination, are shown with arrows.



**Fig. 4.** Combined R-band image of the galaxy associated with the radio source SUMSS J033057-522811 in the centre of the NE X-ray peak. The total exposure time is 360 s. Note the arc to the Southwest of the galaxy. The scale of  $1'$  is indicated on the image.

### 3.2. Optical observation of the background cluster

Optical spectroscopic observations of the galaxy associated with the radio source SUMSS J033057-522811 (Mauch et al. 2003) were done with the 6.5 m Magellan I Baade telescope, using the Inamori Magellan Areal Camera and Spectrograph (IMACS). We used the short f/2 camera with a  $27' \times 27'$  FOV and  $0.2''$  pixel scale. In spectroscopic mode this setup with 300 l/mm grism and  $1.2''$  slit gives a spectral resolution of  $\approx 6 \text{ \AA}$ .

The spectrum obtained is shown in Fig. 3. It is typical for an early type galaxy and does not show detectable emission lines. Using the standard set of absorption lines, marked in Fig. 3 with arrows, the redshift  $z = 0.43961 \pm 0.00014$  was derived. This value is in a very good agreement with the X-ray determined redshift thus providing immediate confirmation that the observed X-ray emission is due to the emission of a distant cluster.

**Table 5.** The best fit values determined for the X-ray bright low entropy region and for the region at the centre of the galaxy distribution. The emission measure  $Y_S$  per square arcminute is given in units of  $10^{64} \text{ cm}^{-3} \text{ arcmin}^{-2}$  and the temperatures are in keV. The Fe abundance is given with respect to the proto-solar values of Lodders (2003). The entropy  $s$  is given in  $\text{keV cm}^2$  and the pressure  $P$  is given in  $10^{-12} \text{ dyne cm}^{-2}$ .

Par.	X-ray bright reg.	centre of gal. dist.
$Y_S$	$7.54 \pm 0.08$	$4.03 \pm 0.06$
$kT$	$3.30 \pm 0.08$	$3.79 \pm 0.13$
Fe	$0.43 \pm 0.04$	$0.53 \pm 0.07$
$s$	223	335
$P$	9.5	7.3
$\chi^2 / \nu$	563/569	448/460

During the slit alignment procedure we also obtained a few direct R-band images of this field. The combined image is shown in Fig. 4. The field was photometrically calibrated using observations of Landolt standard stars (Landolt 1992). We measured the magnitude of the galaxy to be  $m_R = 18.78$  (in  $4''$  aperture), which is approximately what is expected for a brightest cluster cD galaxy at this redshift (e.g. Vikhlinin et al. 1998). The absolute magnitude of the galaxy in R was estimated to be  $-23.35$  (adopting K and evolution corrections from Poggianti 1997). For comparison the absolute magnitudes of M 87 and Cygnus A are approximately  $-23.2$  and  $-23.3$  respectively - close to the absolute magnitude of SUMSS J033057-522811.

The gravitational lensing arc is clearly seen around this galaxy (Fig. 4) at a radial distance of  $\sim 6.2''$ . The enclosed mass within this radius can be estimated (e.g. Narayan & Bartelmann 1996) as

$$M = 1.1 \times 10^{14} M_\odot \left( \frac{\theta}{30''} \right)^2 \left( \frac{D}{1 \text{ Gpc}} \right) \quad (1)$$

where  $\theta = 6.17''$  is the arc radius and  $D = D_d D_{ds} / D_s$  is the combination of (angular diameter) distances from the observer to the lens  $D_d$  and to the source  $D_s$ , and from the lens to the source  $D_{ds}$ . In our case only  $D_d = 1.17 \text{ Gpc}$  is known. For  $D = 1 \text{ Gpc}$  we obtain an upper limit to the total mass of  $M < 5 \times 10^{12} M_\odot$  within  $\sim 35 \text{ kpc}$  of SUMSS J033057-522811. This value is about a factor of 2 higher than the total mass (within a similar distance) of M 87 (Matsushita et al. 2002), derived using X-ray data.

One can therefore conclude that optical data strongly suggest that SUMSS J033057-522811 is a massive elliptical galaxy with the parameters characteristic for most bright cD galaxies in the local Universe.

In the Magellan image we also see an excess of the surface density of faint galaxies near this cD galaxy, most of which are probably members of the distant cluster.

### 3.3. 2D maps of thermodynamic properties

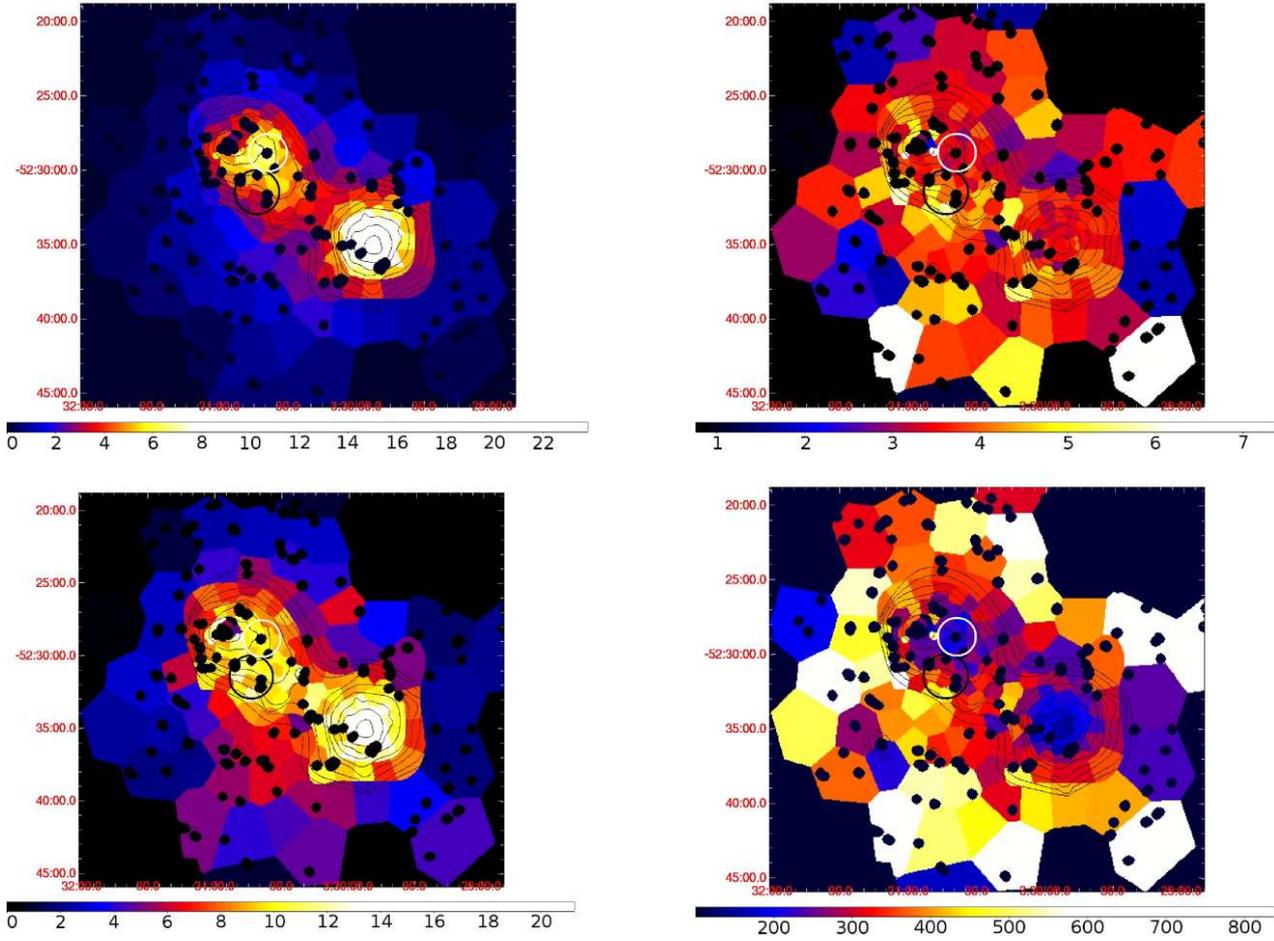
For making 2D maps of spectral parameters, we select regions according to the cluster surface brightness in the 3.5–7.5 keV

band. Considering an optically thin plasma model Forman et al. (2007) showed that for Chandra ACIS-I spectral response the flux  $F$  in the 3.5–7.5 keV band coming from a unit volume with a given pressure depends only weakly on the gas temperature (over the 1–3 keV temperature range and for metallicity of  $\sim 0.7$  solar). For XMM-Newton EPIC and for the range of temperatures characteristic for A3128 (3–5 keV) this is also approximately correct: a factor  $\eta$  relating the pressure and the square root of the flux  $\eta = (\epsilon(T)/T^2)^{1/2}$  decreases monotonically as a function of temperature by about 20% (for metallicity of 0.45 solar). Here  $\epsilon(T)$  is the gas emissivity in the 3.5–7.5 keV band (with account for the spectral response of the instrument). Therefore, areas with similar surface brightness in the 3.5–7.5 keV band, provide us with contiguous regions without strong pressure discontinuities. We note here that our final maps of spectral parameters depend very weakly on the procedure of region selection.

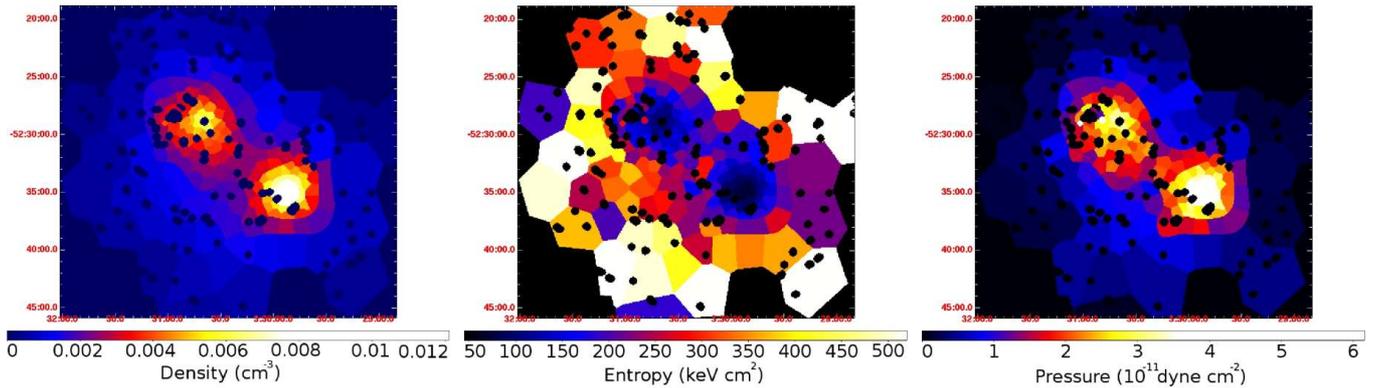
We then use the Voronoi tessellation method (Cappellari & Copin 2003; Diehl & Statler 2006) to further bin the selected areas with a stable minimum signal-to-noise ratio of  $S/N = 33$ , which is needed for relatively accurate temperature determination. We identify 213 independent regions from which we extract the spectra. For each region we compute a spectral redistribution file and an ancillary response file. We fit the spectrum of each bin individually with a single temperature thermal plasma model. The abundances of all elements except Fe in our model are fixed to 0.4 times the solar value. The emission measure, temperature, and Fe abundance are free parameters in the fit. From the best fit emission measure and temperature, we calculate for each extraction region the density  $n$ , entropy  $s \equiv kT/n^{2/3}$ , and the pressure  $P \equiv nkT$ , assuming a length scale along our line of sight of 1 Mpc. As a first step, we use a constant length scale in order to calculate the projected values of the thermodynamic properties without making any assumptions about the cluster centre and gas distribution.

In order to subtract the emission of the background cluster from the 2D maps, we include in the fitted model for each bin with a mean distance smaller than  $3.5'$  from the NE X-ray surface brightness peak, an additional thermal component. We fix its temperature, metallicity, and redshift to the mean values determined for the background cluster from the global fit: 5.14 keV and 0.47 solar, respectively (see Sect. 3.1.1). Using the best fit beta model to the background cluster (see Sect. 3.1.2) we calculate the normalization of the emission of the background cluster for each bin. By including this component in the fitted model, we effectively subtract its emission and determine the best fit parameters for the foreground cluster. However, since the core of the background cluster is not radially symmetric and it has a strongly elongated morphology, at radii smaller than  $0.8'$  from the NE peak its emission cannot be properly subtracted. Therefore, we exclude the bins with a distance smaller than  $0.8'$  from the NE peak from the analysis.

At the top left panel of Fig. 5 we show the map of the projected emission measure per square arcminute integrated along the line of sight, after the emission of the background cluster was subtracted. We see, that the surface brightness of the X-ray emission of A3128 has a peak at about  $1.5'$  West-Southwest



**Fig. 5.** *Top left:* The projected emission measure per arcminute, in units of  $10^{64} \text{ cm}^{-3} \text{ arcmin}^{-2}$ , integrated along the line of sight, after the emission of the background cluster was subtracted. *Top right:* Temperature map derived from spectral fitting in units of keV. *Bottom left:* Map of the pressure derived from the best fit temperature and emission measure, assuming a length scale along our line of sight of 1 Mpc. The units are  $10^{-12} \text{ dyne cm}^{-2}$ . *Bottom right:* Map of the entropy derived from the best fit temperature and emission measure, assuming a length scale along our line of sight of 1 Mpc. The units are  $\text{keV cm}^2$ . The white and black circles indicate the high surface brightness low entropy region and the region at the centre of the galaxy distribution, respectively. These we use for subsequent spectral analysis. The X-ray isophotes from the wavelet decomposed image are overlotted.



**Fig. 6.** *Left panel:* Density map of the ICM. *Central panel:* Map of the entropy of the ICM. *Right panel:* The pressure map of the ICM. The thermodynamic properties are calculated using volume estimates for the spectral extraction regions described in Henry et al. (2004) and Mahdavi et al. (2005). The volumes are calculated assuming two centres, one at the SW X-ray peak, and one at the X-ray bright low entropy region at the possible centre of the gravitational potential of A3128.

from the original NE peak. This region, as we show in the top right panel of Fig. 5 has a temperature between 3.0–3.5 keV. While the pressure is relatively high in this region (see the lower left panel of Fig. 5), the entropy is low (lower right panel of Fig. 5). This indicates that this region may be the centre of the gravitational potential of A3128. The centre of the galaxy distribution of A3128 is about 2.8′ South of this region (Dalton et al. 1997; Rose et al. 2002).

In order to better compare the spectral properties of these two regions: the X-ray bright low entropy region and the region at the centre of the galaxy distribution, we extract a spectrum from a circular area with a radius of 1.26′ centred on the new X-ray brightness peak ( $\alpha = 3^{\text{h}}30^{\text{m}}40^{\text{s}}$ ,  $\delta = -52^{\circ}28'50''$ ) and a spectrum from a circular area with a radius of 1.5′ centred on the optical centre ( $\alpha = 3^{\text{h}}30^{\text{m}}43^{\text{s}}$ ,  $\delta = -52^{\circ}31'30''$ ) reported by Rose et al. (2002). The spectral extraction regions are indicated by circles on the maps in Fig. 5. The best-fit values of the single temperature thermal fits to the spectra are shown in Table 5. We also show in the table the entropy and pressure values calculated for the volume of a sphere with a radius corresponding to the radius of the extraction region. The results confirm that the high surface brightness region has a factor of  $\sim 1.5$  lower entropy than the area at the optical centre of the galaxy distribution. The pressure at the high surface brightness region is a factor of  $\sim 1.3$  higher than the pressure at the centre of the galaxy distribution. The fit results indicate that the Fe abundance of the X-ray bright region is lower than the Fe abundance at the centre of the galaxy distribution. Under the assumption of hydrostatic equilibrium, the pressure peak combined with low entropy are unambiguous signs of the centre of the gravitational potential.

The X-ray bright region at the possible centre of the gravitational potential of A3128 in the NE is separated from the bright SW core by a surface brightness depression. The pressure map indicates that the dark matter potential well associated with the SW core is connected with the dark matter potential in the NE. This is possibly the region where the dark matter potentials of the SW core and of A3128 overlap. We note that the “bridge” between the SW and NE in the pressure map coincides with an apparent chain of galaxies seen in the Digitalised Sky Survey image. Unfortunately, non of these galaxies has a known radial velocity.

In Fig. 6, we show density, entropy, and pressure maps determined assuming the X-ray emitting gas is associated with two gravitational potentials. The thermodynamic properties are calculated using volume estimates for the spectral extraction regions described in Henry et al. (2004) and Mahdavi et al. (2005). The volumes are calculated assuming two centres, one at the SW X-ray peak, and one at the X-ray bright low entropy region at the possible centre of the gravitational potential of A3128. For each polygon the closest of the two centres is assumed to be the centre in the volume calculation.

## 4. Discussion

### 4.1. The background cluster associated with the NE X-ray peak

While at the expected energy we do not detect the strong Fe K line in the spectrum extracted from the NE X-ray peak, we detect line emission at the energy corresponding to the line energy of Fe K emission redshifted by  $z = 0.44$ . Therefore we conclude that the NE X-ray peak observed toward A3128 is not associated with the surviving ICM of a group falling in supersonically to the cluster as previously thought (Rose et al. 2002), but it is a distant luminous cluster of galaxies at redshift  $z = 0.44$ . Subsequent optical spectroscopic observation of the distant radio bright galaxy in the centre of the NE X-ray peak (radio source SUMSS J033057-522811) with the Magellan telescope revealed a redshift of  $z = 0.44$  confirming its association with the cluster seen in X-rays.

The red magnitude of the galaxy  $m_{\text{R}} = 18.5$  is typical for cD galaxies at this redshift (e.g. Vikhlinin et al. 1998). The Magellan observation also confirmed the gravitational arc around the galaxy, the presence of which was previously suggested by the images obtained with the 0.9 m CTIO telescope (Rose et al. 2002). We estimate the enclosed total mass within the radius of the lens to be  $< 5 \times 10^{12} M_{\odot}$ , which is within a factor of 2 consistent with the total mass of M 87 at a similar radius. The optical data thus strongly suggest that SUMSS J033057-522811 with properties characteristic for cD galaxies is the dominant galaxy of the cluster in the background.

The X-ray morphology of the background cluster is strongly elongated showing that its core is not relaxed and that it is possibly undergoing a merger. However, its observed X-ray luminosity of  $L = 6.9 \times 10^{44} \text{ erg s}^{-1}$ , temperature  $kT = 5.14 \text{ keV}$ , and derived mass  $M = 3.4 \times 10^{14} M_{\odot}$  agree well with the generally observed  $L_{\text{X}}-T$  and  $M-T$  scaling relations (e.g. Wu et al. 1999; Pratt 2006). Further deep optical observations are needed to characterise the properties of the background cluster.

### 4.2. The nature of the SW surface brightness peak

The SW surface brightness peak is centred on an apparent compact group of galaxies. Compared to other clusters of galaxies with similar temperatures, its core radius is small. Such small core radii are observed only in clusters with cooling cores, or in clusters undergoing merging events where the cluster profile is truncated. While the radial Fe abundance distribution has a strong peak at the centre of the group, the temperature distribution is flat, with no indication for a cool core. The cooling time in the centre of the SW core, assuming the gas cools isobarically, is  $4.5 \times 10^9 \text{ yr}$ . Its estimated total mass of  $M_{500} = 1.1 \times 10^{14}$  is a factor of 2 lower than the expected mass of a 3.4 keV cluster (e.g. Pratt 2006). But since the value of  $\beta = 0.3$  was determined by fitting the surface brightness profile at radii much smaller than  $r_{500}$ , the estimated  $M_{500}$  is highly uncertain. The radial velocity of the dominant galaxy of this apparent group (ENACS 75) is  $V = 19252 \text{ km s}^{-2}$  (Katgert et al.

1998), which is higher by  $\approx 1500 \text{ km s}^{-1}$  than the mean radial velocity of galaxies in A3128. Unfortunately, we know the radial velocity for only one more galaxy of the apparent group (ENACS 78), which is lower than that of the dominant galaxy:  $V = 18380 \text{ km s}^{-1}$  (Katgert et al. 1998), but still higher by  $\approx 500 \text{ km s}^{-1}$  than the mean radial velocity of the cluster. The sound speed corresponding to the mean cluster temperature of  $\sim 3.5 \text{ keV}$  is  $960 \text{ km s}^{-1}$ , which means that if the SW peak is the remaining core of a group or a cluster merging with A3128 at a relative velocity of  $\approx 1500 \text{ km s}^{-1}$ , then the merger is supersonic.

On the smaller spatial scales, the X-ray emission of the SW peak appears to follow the galaxy mass distribution as it is slightly displaced toward the West and clearly peaks on the galaxy ENACS 75. However, the association of the galaxy ENACS 75 with the SW X-ray peak in radial velocity is difficult. The redshift of the SW peak determined from the energy centroid of the Fe K line corresponds to a radial velocity of  $V = 18000 \pm 300 \text{ km s}^{-1}$ , and the redshift of the “tail” region, which we might consider as the diffuse emission of A3128 corresponds to a radial velocity of  $V = 17400 \pm 900 \text{ km s}^{-1}$ . However, the systematic uncertainties on these values are larger than the quoted statistical errors. The EPIC detectors are known to have gain problems which are both time and position dependent. The absolute error in the redshift determination can be as large as  $1500 \text{ km s}^{-1}$  (Simionescu et al. in prep.). Therefore, the radial velocity of the hot gas is within the systematic uncertainties consistent with the radial velocity of the galaxies.

If the SW peak is a core of a group or of a cluster that moves through the ICM of A3128, then its gas is being stripped by ram pressure. The gas of the infalling cluster is stripped at radii where the thermal pressure of the hot gas in the infalling group is too small to balance the sum of the thermal and ram pressure of the cluster ICM. This happens at radii larger than the radius where the thermal pressure of the infalling cluster is equal to the thermal pressure at the stagnation point (e.g. Markevitch & Vikhlinin 2007). The ratio of thermal pressures at the stagnation point  $p_0$  and in the free stream,  $p_1$ , for Mach numbers relative to the sound speed in the free stream region  $M > 1$  is (Landau & Lifshitz 1959):

$$\frac{p_0}{p_1} = \left( \frac{\gamma + 1}{2} \right)^{\frac{\gamma+1}{\gamma-1}} M^2 \left( \gamma - \frac{\gamma-1}{2M^2} \right)^{-\frac{1}{\gamma-1}}, \quad (2)$$

where  $\gamma = 5/3$  is the adiabatic index of the gas. For the pressure in the free stream region,  $p_1$ , we assume the value determined assuming a density of  $n = 1 \times 10^{-3} \text{ cm}^{-3}$  and a temperature of  $3.5 \text{ keV}$ . For the pressure of the SW core at the given radius, we assume a density profile of  $n(r) = n_0(1 + (r/r_c)^2)^{-\frac{3}{2}\beta}$ , where  $n_0 = 1.2 \times 10^{-2} \text{ cm}^{-3}$ ,  $r_c = 30 \text{ kpc}$ , and  $\beta = 0.3$ . We find that at a Mach number of  $M = 1.5$  (inferred from the radial velocity difference of  $\sim 1500 \text{ km s}^{-1}$  between ENACS 75 and A3128, and from the sound speed  $v_s = 960 \text{ km s}^{-1}$  in the ICM), the gas outside the radius of  $r = 90 \text{ kpc}$  should be stripped, and only the ICM of the merging core inside of this radius should be surviving.

Such a merger would produce a tail of low entropy and high Fe abundance. However, in the images and in the maps

of the thermodynamic properties we do not observe any obvious signs of a merger for the SW peak. Moreover, we verified that the Fe abundance distribution is consistent with being symmetric and is not enhanced in any direction. These observed properties might indicate that the merger is occurring just along our line of sight. The merger is compressing and heating the ICM, which might explain the observed lack of a cool core. A shock with a Mach number  $M = 1.5$  heats the gas by a factor of 1.5, which means that for the cluster temperature of  $3.5 \text{ keV}$  the gas temperature at the stagnation point should be  $5.25 \text{ keV}$ . If we fit the spectrum extracted from the circular region with a radius of  $1.5'$  centred on the SW core with two thermal models, with the temperature of one thermal component fixed to  $kT = 5.25 \text{ keV}$ , the fit improves compared to the single temperature model (see Sect. 3.1.1). The reduced  $\chi^2$  improves from 1.30 to 1.16. We obtain a best fit temperature of  $kT = 2.23 \pm 0.17 \text{ keV}$  for the cooler component, and best emission measures of  $Y = (5.1 \pm 0.5) \times 10^{65} \text{ cm}^{-3}$  and  $Y = (4.7 \pm 0.5) \times 10^{65} \text{ cm}^{-3}$  for the hotter and cooler gas, respectively. This indicates that much of the emission of the SW core might be due to the shocked gas from a merger happening along our line of sight. As the enriched stripped gas is trailing behind the core we are seeing it in projection which explains the enhanced Fe abundance within the projected distance of  $\sim 120 \text{ kpc}$ . The velocity difference between the two galaxies of the compact group (ENACS 75, ENACS 78) may also be explained by tidal stretching of the group along our line of sight during the passage of the group through the cluster.

Alternatively, the SW core might be a group still falling toward A3128, which we observe projected on the ICM of the cluster. In this case the ICM of the group still did not start to interact with the ICM of the cluster, and ENACS 78 and 75 are only projected close to each other by chance. However, this scenario does not explain the lack of the temperature gradient and of a cool core in the group.

### 4.3. The diffuse ICM

After subtracting the emission of the  $z = 0.44$  cluster, we identified a new region with an enhanced surface brightness to the West-Southwest of the centre of the background cluster. This region has also an enhanced pressure and low entropy. It is at a distance of  $\sim 2.8'$  from the centre of the galaxy distribution given by Rose et al. (2002). However, the entropy of the ICM at the centre of the galaxy distribution is a factor of 1.4 higher and the pressure is lower than that of the region with the enhanced X-ray surface brightness. Because of the great dynamical complexity of the system, with infalling groups and filaments identified in position-position and position-redshift diagrams, the determination of the cluster centre from the galaxy distribution is highly uncertain. Based on the thermodynamic properties of the ICM, we conclude that the enhanced surface brightness region, centred at  $\alpha = 3^{\text{h}}30^{\text{m}}40^{\text{s}}$ ,  $\delta = -52^{\circ}28'50''$ , is the centre of the gravitational potential of the cluster A3128. The position of this region is also more consistent with the centre of the extended low surface brightness X-ray emission.

The images and the maps of thermodynamic properties do not reveal any obvious shocks. We only see two candidates, one to the Northeast of the background cluster and one to the Southwest of SW core. Both shock candidates exhibit temperature drops by  $\sim 1$  keV associated with a drop in entropy, and a factor of 2 drop in pressure. However, the surface brightness in these two areas is low and we cannot confirm the presence of surface brightness discontinuities in the regions where the maps indicate the presence of the shocks.

As proposed by Caldwell & Rose (1997), A3128 may have had encountered a merger with A3125 in the past, and as discussed by Rose et al. (2002) the optical redshift data reveal a number of groups, some of them tidally distended into filaments after a close passage through A3128. The X-ray data combined with the optical redshifts suggest an ongoing merger with a group. Moreover, the inconsistency between the cluster centre determined based on the distribution of galaxies and based on the thermodynamic properties of the ICM, and the lack of bright galaxies at the newly identified centre of gravitational potential, further highlight the view that A3128 is a dynamically young, unrelaxed system. The unrelaxed nature of A3128 can be attributed to its location in the high density environment of the Horologium-Reticulum supercluster.

## 5. Conclusions

We have analyzed new deep XMM-Newton EPIC data of the cluster of galaxies Abell 3128 located in the Horologium-Reticulum supercluster. We found that:

- The Northeast X-ray peak observed toward A3128 is a distant luminous cluster of galaxies at redshift  $z = 0.44$ .
- The properties of the distant radio bright galaxy in the centre of the NE X-ray peak indicate that it is the cD galaxy of the cluster in the background. We detect a gravitational arc around the galaxy.
- The properties of the Southwest X-ray peak suggest that it is the core of a group merging with A3128 along our line of sight.
- Based on 2D maps of thermodynamic properties of the ICM determined after subtracting a model for the background cluster, we conclude that the enhanced surface brightness region at a distance of  $\sim 2.8'$  from the centre of the galaxy distribution is the centre of the gravitational potential of the cluster A3128. The inconsistency between the cluster centres determined based on the distribution of galaxies and based on the thermodynamic properties of the ICM further highlights the view that A3128 is a dynamically young, unrelaxed system.

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## References

- Caldwell, N. & Rose, J. A. 1997, *AJ*, 113, 492  
 Cappellari, M. & Copin, Y. 2003, *MNRAS*, 342, 345  
 Cavaliere, A. & Fusco-Femiano, R. 1978, *A&A*, 70, 677  
 Dalton, G. B., Maddox, S. J., Sutherland, W. J., & Efstathiou, G. 1997, *MNRAS*, 289, 263  
 De Luca, A. & Molendi, S. 2004, *A&A*, 419, 837  
 Dickey, J. M. & Lockman, F. J. 1990, *ARA&A*, 28, 215  
 Diehl, S. & Statler, T. S. 2006, *MNRAS*, 368, 497  
 Einasto, M., Einasto, J., Tago, E., Müller, V., & Andernach, H. 2001, *AJ*, 122, 2222  
 Finoguenov, A., Arnaud, M., & David, L. P. 2001, *ApJ*, 555, 191  
 Finoguenov, A., Guzzo, L., Hasinger, G., et al. 2007, *ApJS*, in press, astro-ph/0612360  
 Forman, W., Churazov, E., Jones, C., et al. 2007, *ApJ*, in press, astro-ph/0604583  
 Ghizzardi, S. 2001, <http://xmm.vilspa.esa.es/docs/documents/CAL-TN-0022-1-0.ps.gz>  
 Henry, J. P., Finoguenov, A., & Briel, U. G. 2004, *ApJ*, 615, 181  
 Kaastra, J. S., Mewe, R., & Nieuwenhuijzen, H. 1996, in *UV and X-ray Spectroscopy of Astrophysical and Laboratory Plasmas* p.411, K. Yamashita and T. Watanabe. Tokyo : Universal Academy Press  
 Katgert, P., Mazure, A., den Hartog, R., et al. 1998, *A&AS*, 129, 399  
 Kuntz, K. D. & Snowden, S. L. 2000, *ApJ*, 543, 195  
 Landau, L. D. & Lifshitz, E. M. 1959, *Fluid mechanics* (Course of theoretical physics, Oxford: Pergamon Press, 1959)  
 Landolt, A. U. 1992, *AJ*, 104, 340  
 Lodders, K. 2003, *ApJ*, 591, 1220  
 Mahdavi, A., Finoguenov, A., Böhringer, H., Geller, M. J., & Henry, J. P. 2005, *ApJ*, 622, 187  
 Markevitch, M. & Vikhlinin, A. 2007, *Phys. Rep.*, 443, 1  
 Matsushita, K., Belsole, E., Finoguenov, A., & Böhringer, H. 2002, *A&A*, 386, 77  
 Mauch, T., Murphy, T., Buttery, H. J., et al. 2003, *MNRAS*, 342, 1117  
 Narayan, R. & Bartelmann, M. 1996, astro-ph/9606001  
 Poggianti, B. M. 1997, *A&AS*, 122, 399  
 Pratt, G. W. 2006, astro-ph/0607199  
 Rose, J. A., Gaba, A. E., Christiansen, W. A., et al. 2002, *AJ*, 123, 1216  
 Vikhlinin, A., McNamara, B. R., Forman, W., et al. 1998, *ApJ*, 502, 558  
 Wu, X.-P., Xue, Y.-J., & Fang, L.-Z. 1999, *ApJ*, 524, 22