WMAP 5-Year Results: Measurement of f_{NL}

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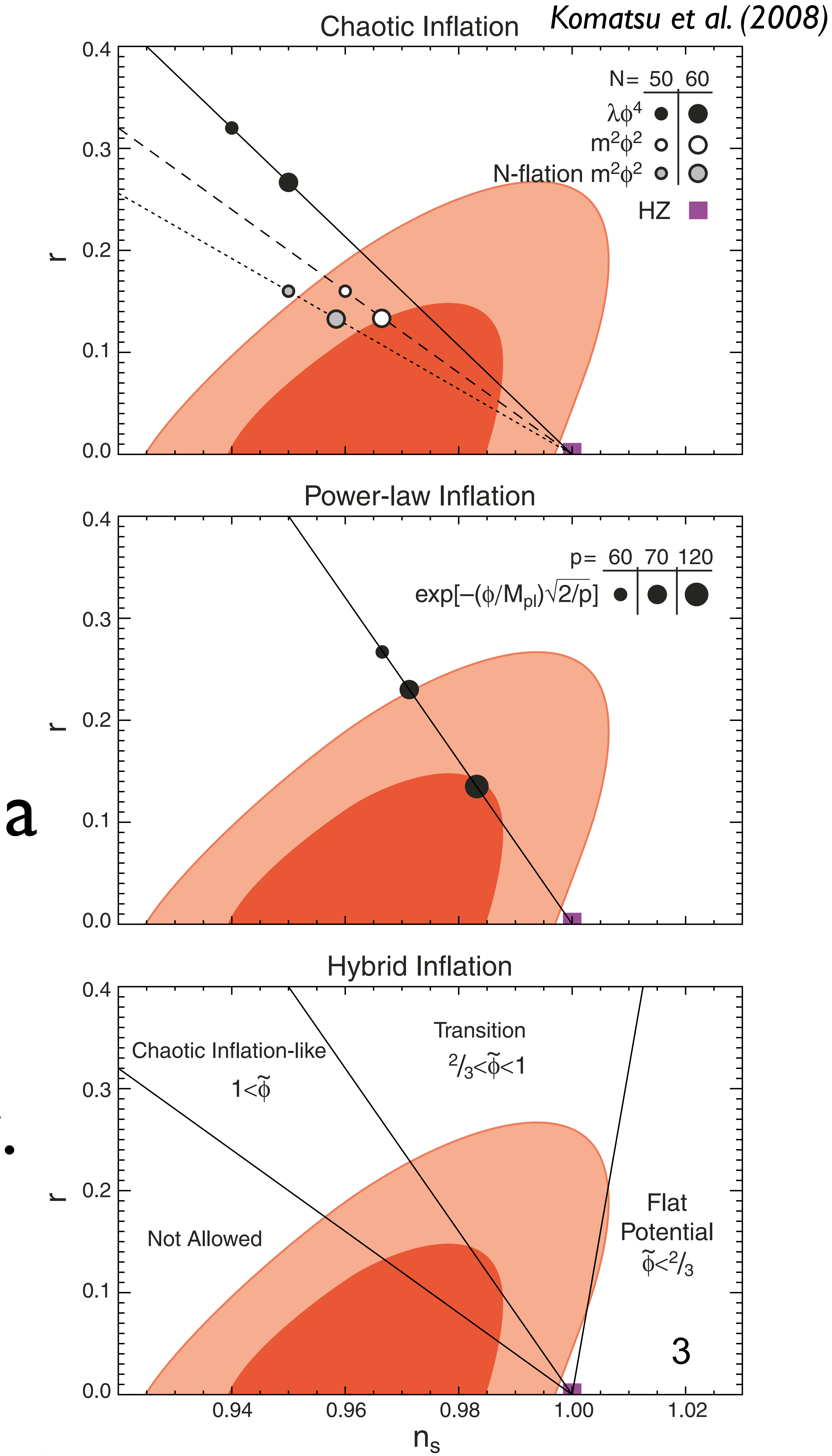
Non-Gaussianity From Inflation, Cambridge, September 8, 2008

Why is Non-Gaussianity Important?

- Because a detection of f_{NL} has a best chance of ruling out the **largest** class of early universe models.
- Namely, it will rule out inflation models based upon
  - a single scalar field with
  - the canonical kinetic term that
  - rolled down a smooth scalar potential slowly, and
  - was initially in the Banch-Davies vacuum.
- ***Detection of non-Gaussianity would be a major breakthrough in cosmology.***

We have *r* and *n_s*. Why Bother?

- While the current limit on the power-law index of the primordial power spectrum, **n_s**, and the amplitude of gravitational waves, **r**, have ruled out many inflation models already, many still survive (which is a good thing!)
- A convincing detection of f_{NL} would rule out most of them **regardless of n_s or r**.
- f_{NL} offers more ways to test various early universe models!

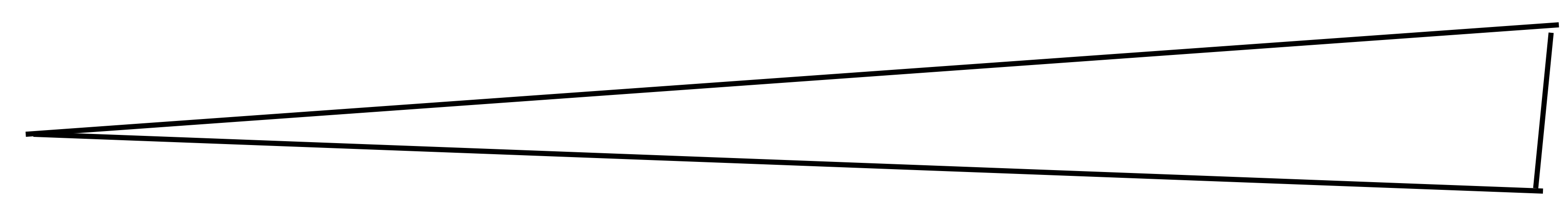
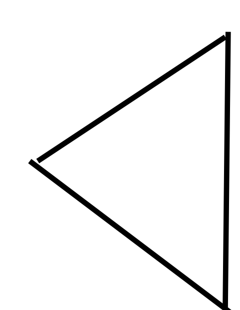


Why Bispectrum?

- **The bispectrum <u>vanishes</u>** for Gaussian fluctuations with random phases.
- Any non-zero detection of the bispectrum indicates the presence of (some kind of) non-Gaussianity.
- A sensitive tool for finding non-Gaussianity.

Two f_{NL}'s

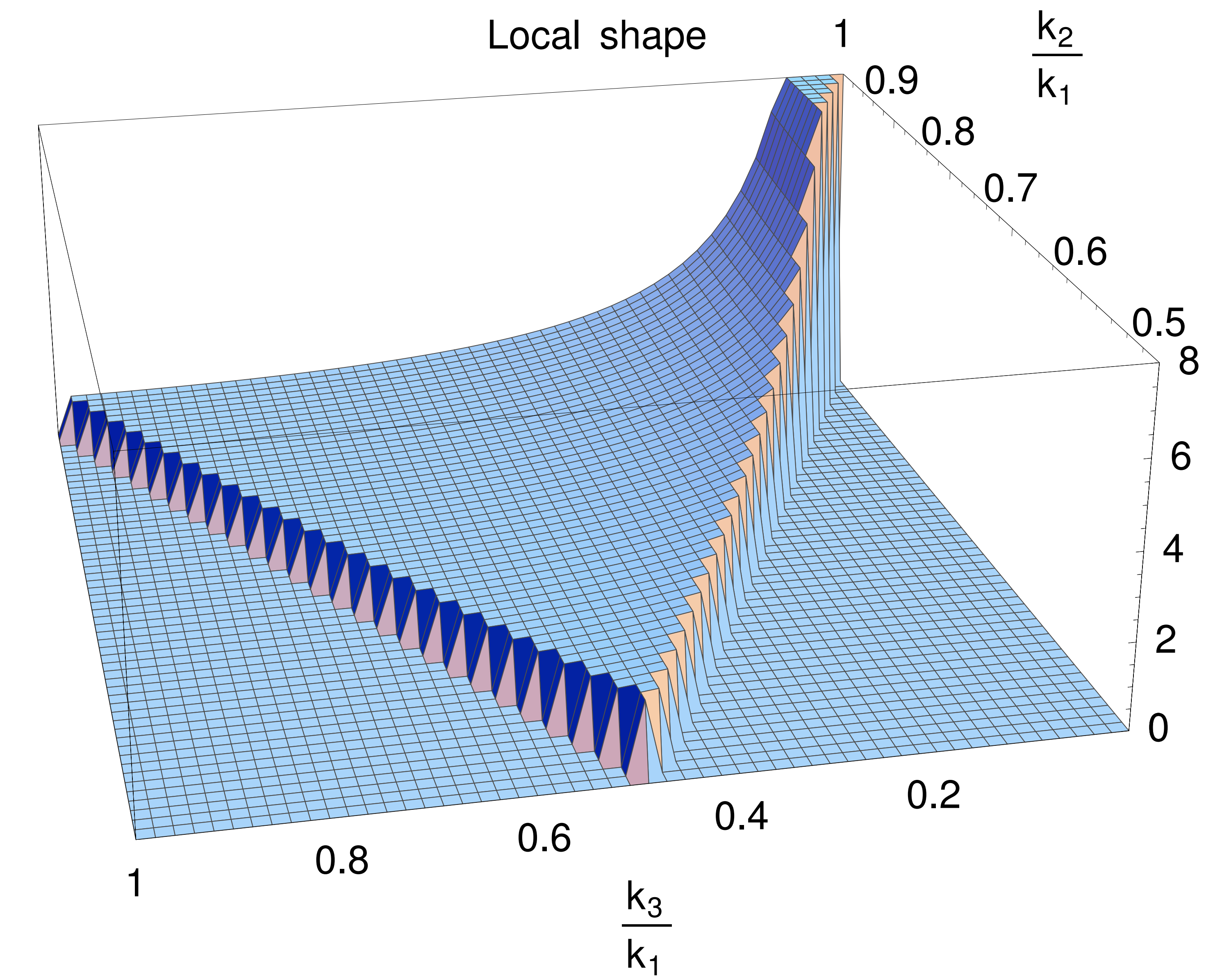
There are more than two; I will come back to that later.

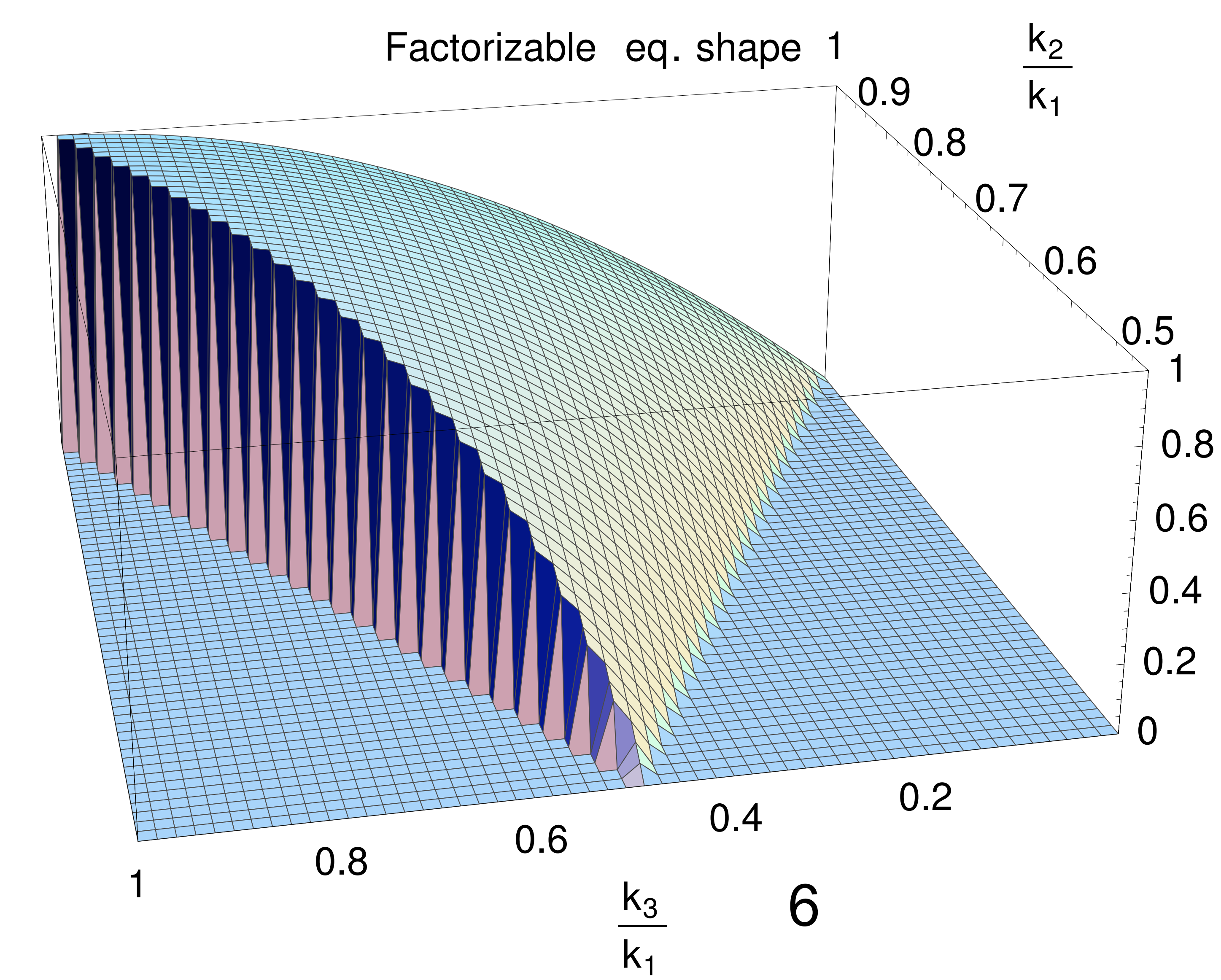
- Depending upon the shape of triangles, one can define various f_{NL}'s:
- "Local" form 
  - which generates non-Gaussianity locally in position space via $\Phi(x)=\Phi_{gaus}(x)+f_{NL}^{local}[\Phi_{gaus}(x)]^2$
- "Equilateral" form 
  - which generates non-Gaussianity locally in momentum space (e.g., k-inflation, DBI inflation)

Forms of $b(k_1,k_2,k_3)$

Earlier work on the local form:
Salopek&Bond (1990); Gangui et al. (1994);
Verde et al. (2000); Wang&Kamionkowski (2000)

- Local form (Komatsu & Spergel 2001)
  - $b^{local}(k_1,k_2,k_3) = 2[P(k_1)P(k_2)+cyc.]$
- Equilateral form (Babich, Creminelli & Zaldarriaga 2004)
  - $b^{equilateral}(k_1,k_2,k_3) = 6\{-[P(k_1)P(k_2)+cyc.] - 2[P(k_1)P(k_2)P(k_3)]^{2/3} + [P(k_1)^{1/3}P(k_2)^{2/3}P(k_3)+cyc.]\}$

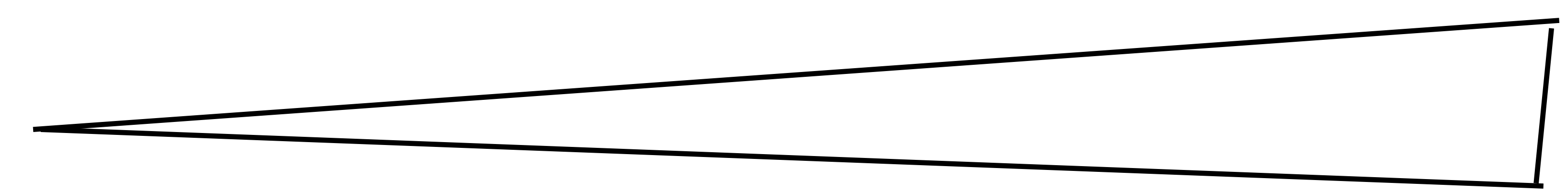




What if f_{NL} is detected?

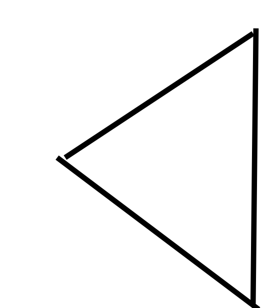
- A single field, canonical kinetic term, slow-roll, and/or Banch-Davies vacuum, must be modified.

Local
- Multi-field (curvaton);



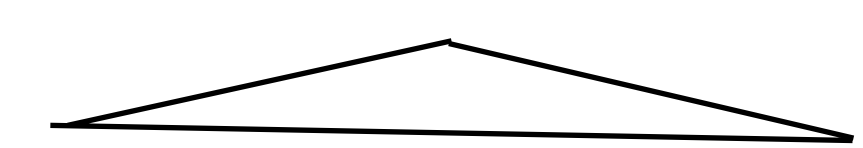
  Preheating (e.g., Chambers & Rajantie 2008)

Equil.
- Non-canonical kinetic term (k-inflation, DBI)



Bump +Osci.
- Temporary fast roll (features in potential; Ekpyrotic fast roll)

Folded/ Flat
- Departures from the Banch-Davies vacuum



- It will give us a lot of clues as to what the correct early universe models should look like.

Journal on f_{NL}

- Local
  - $-3500 < f_{NL}^{local} < 2000$ [COBE 4yr, l_{max}=20] Komatsu et al. (2002)
  - $-58 < f_{NL}^{local} < 134$ [WMAP 1yr, l_{max}=265] Komatsu et al. (2003)
  - $-54 < f_{NL}^{local} < 114$ [WMAP 3yr, l_{max}=350] Spergel et al. (2007)
  - **$-9 < f_{NL}^{local} < 111$ [WMAP 5yr, l_{max}=500]** Komatsu et al. (2008)
- Equilateral
  - $-366 < f_{NL}^{equil} < 238$ [WMAP 1yr, l_{max}=405] Creminelli et al. (2006)
  - $-256 < f_{NL}^{equil} < 332$ [WMAP 3yr, l_{max}=475] Creminelli et al. (2007)
  - **$-151 < f_{NL}^{equil} < 253$ [WMAP 5yr, l_{max}=700]** Komatsu et al. (2008)

Methodology

- A fast cubic statistics method developed over the years by: Komatsu, Spergel & Wandelt (2005); Creminelli et al. (2006); Yadav, Komatsu & Wandelt (2007)
  - Please read Appendix A of Komatsu et al., if you are interested in details.
  - Sub-optimal for f_{NL}^{local} in the noise dominated regime (l>500) if noise is inhomogeneous
    - Nearly optimal for $f_{NL}^{equilateral}$ and b_{src}
  - There is a room for improvement using the optimal C^{-1} weighting (Smith & Zaldarriaga 2006)

Data Combination

- We mainly use V band (61 GHz) and W band (94 GHz) data.
  - The results from Q band (41 GHz) are discrepant, probably due to a stronger foreground contamination
- These are *foreground-reduced maps*, delivered on the LAMBDA archive.
  - We also give the results from the raw maps.

Gold et al. (2008)

Mask

- We have upgraded the Galaxy masks for the 5-year analysis:
  - 1yr and 3yr release
    - "Kp0" mask for Gaussianity tests (76.5%)
    - "Kp2" mask for the C_l analysis (84.6%)
  - 5yr release
    - "KQ75" mask for Gaussianity tests (71.8%)
    - "KQ85" mask for the C_l analysis (81.7%)

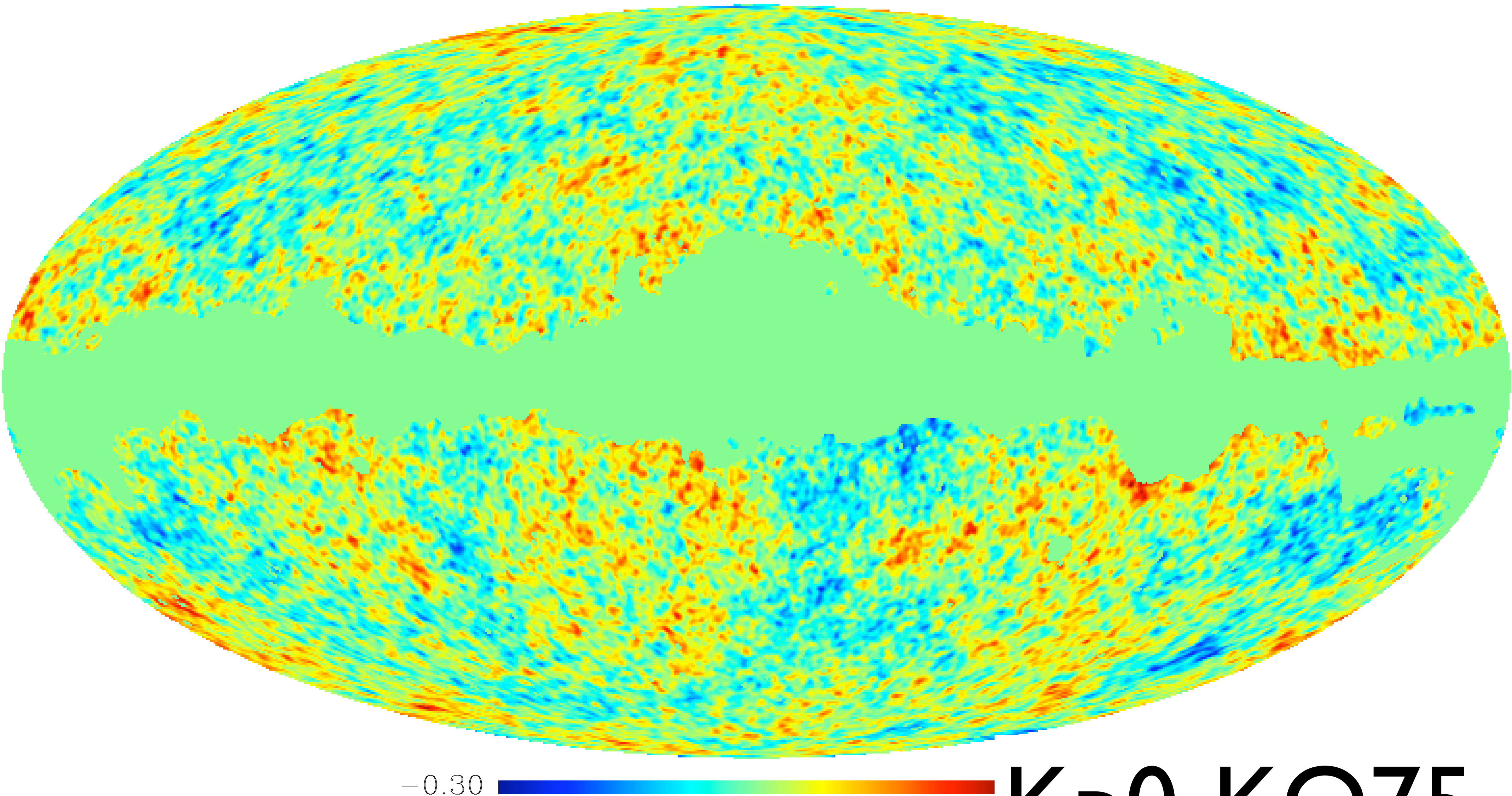
- What are the KQx masks? *Gold et al. (2008)*
  - The previous KpN masks identified the bright region in the K band data (22 GHz), which are contaminated mostly by the synchrotron emission, and masked them.
    - "p" stands for "plus," and N represents the brightness level above which the pixels are masked.
  - The new KQx masks identify the bright region in the K band minus the CMB map from Internal Linear Combination (the CMB picture that you always see), as well as the bright region in the Q band minus ILC.
  - Q band traces the free-free emission better than K.
  - x represents a fraction of the sky retained in K or Q.

Gold et al. (2008)

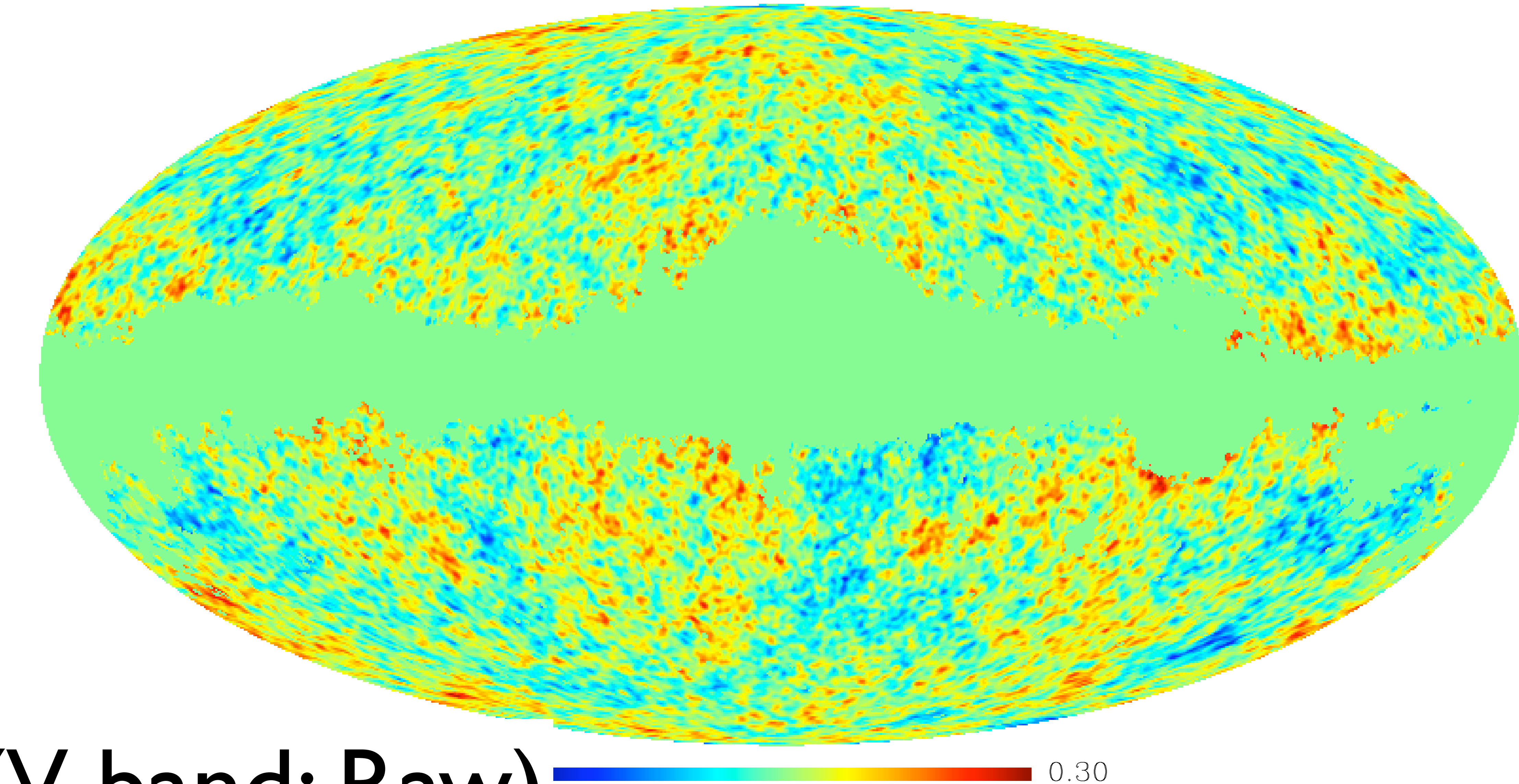
Why KQ75?

- The KQ75 mask removes the pixels that are contaminated by the free-free region better than the Kp0 mask.
- CMB was absent when the mask was defined, as the maske was defined by the K (or Q) band map minus the CMB map from ILC.
- The final mask is a combination of the K mask (which retains 75% of the sky) and the Q mask (which also retains 75%). Since K and Q masks do not always overlap, the final KQ75 mask retains less than 75% of the sky. (It retains 71.8% of the sky for cosmology.)

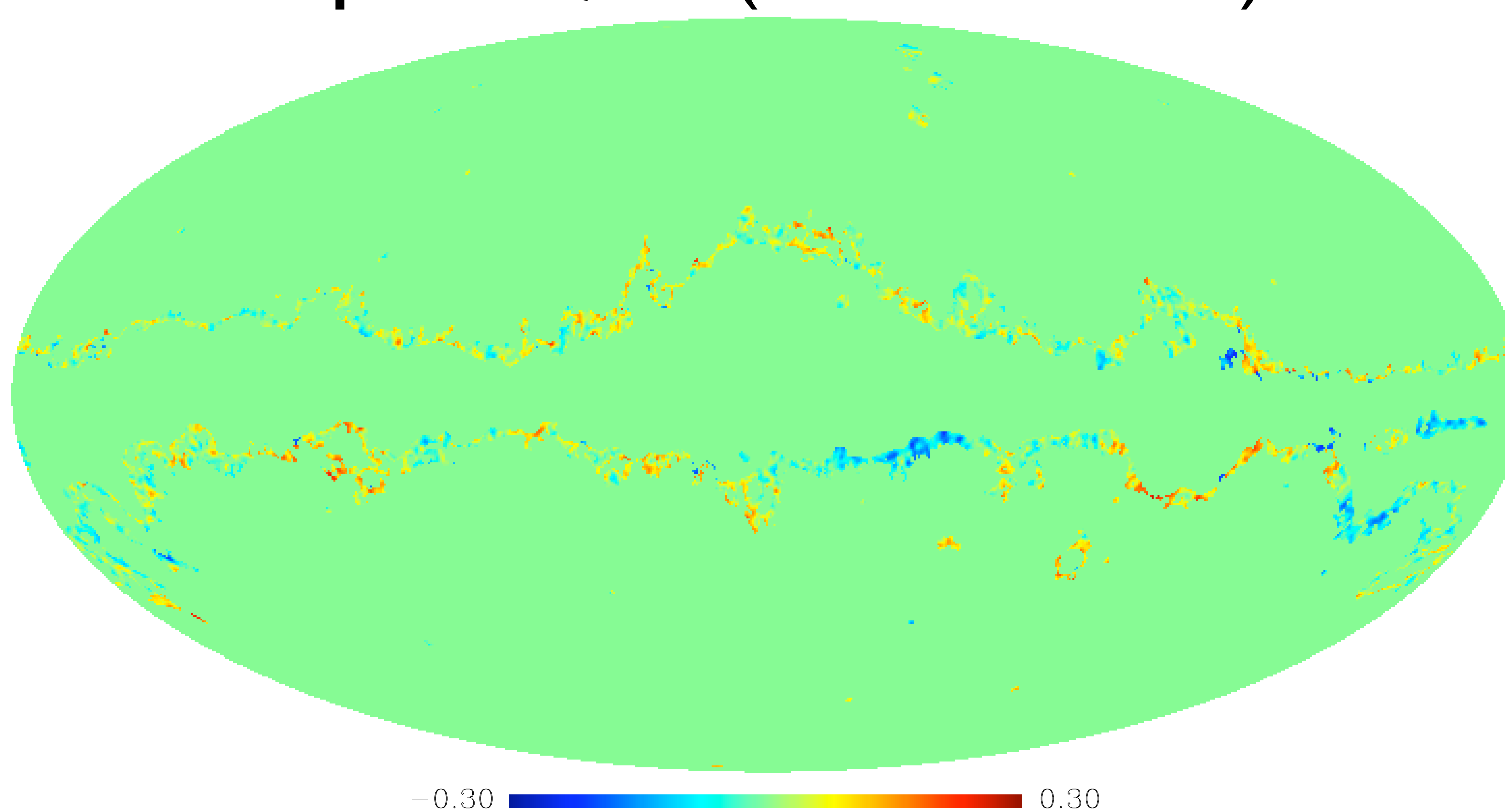
Kp0 (V band; Raw)



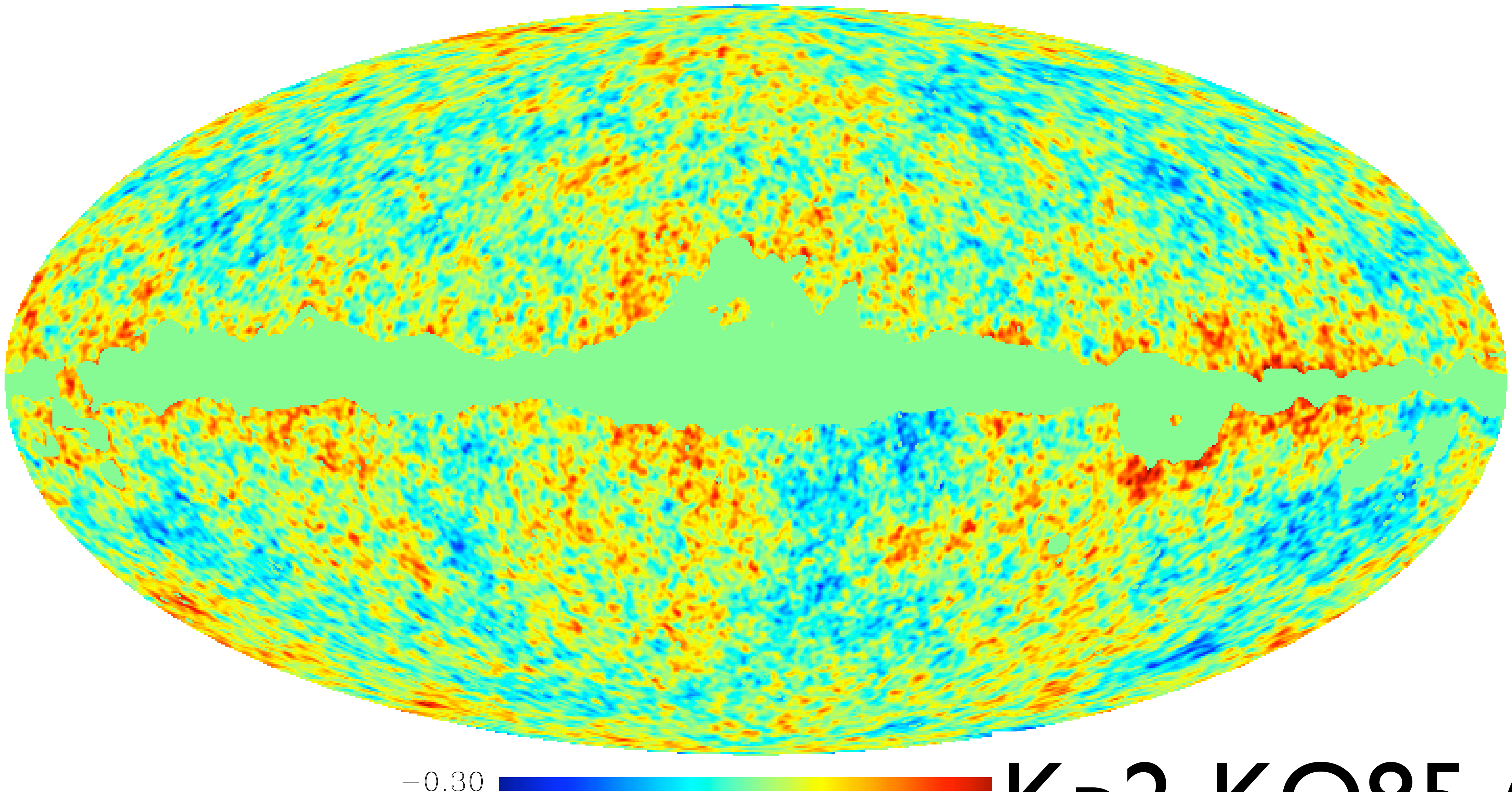
KQ75 (V band; Raw)



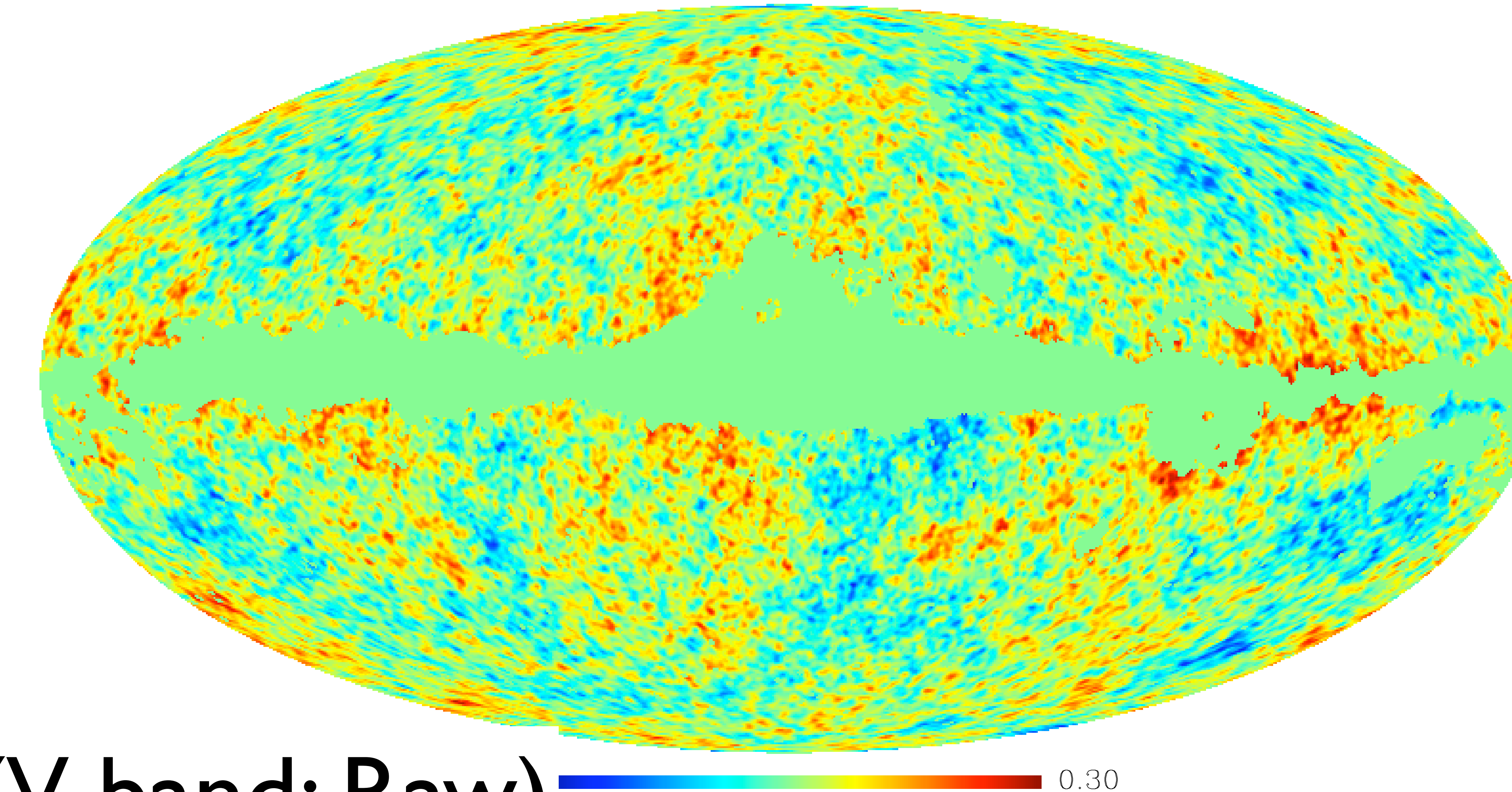
Kp0-KQ75 (V band; Raw)



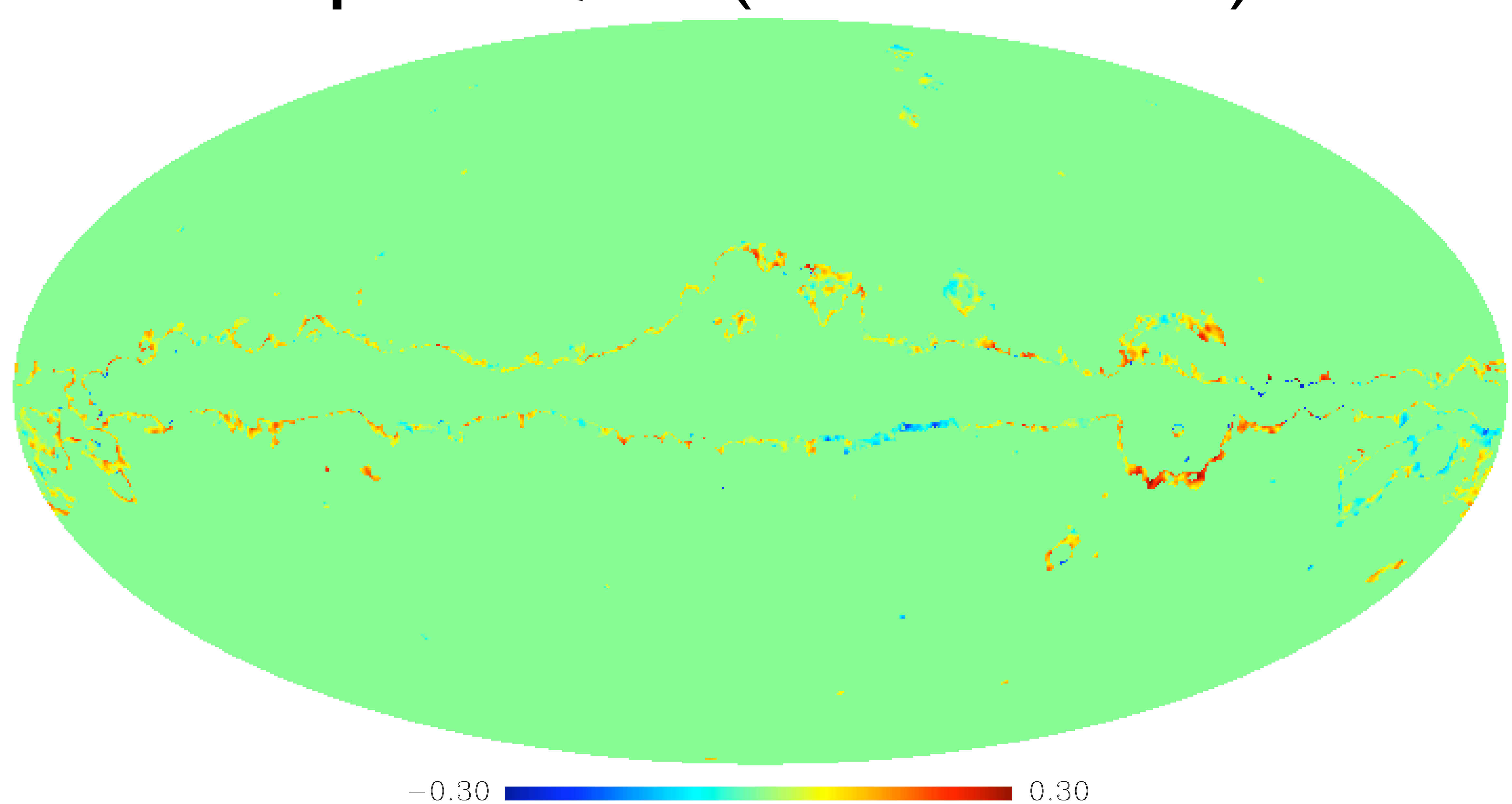
Kp2 (V band; Raw)



KQ85 (V band; Raw)



Kp2-KQ85 (V band; Raw)



Why Use KQ75?

- Because WE KNOW that Kp0 leaves some free-free emission unmasked.
- KQ75 is completely free from any potential contamination of CMB.
- Note that the mask was defined before Gaussianity tests.
- Drawback: KQ75 cuts more sky than Kp0.
  - Kp0 retains 76.5% of the sky for cosmological analysis, whereas KQ75 retains 71.8%.
  - 3% increase in the uncertainty of f_{NL} expected

Komatsu et al. (2008)

Main Result (Local)

Band	Mask	$l_{\max}$	$f_{NL}^{\rm local}$	$\Delta f_{NL}^{\rm local}$	b_{src}
V+W	*KQ85*	400	50 ± 29	1 ± 2	0.26 ± 1.5
V+W	*KQ85*	500	61 ± 26	2.5 ± 1.5	0.05 ± 0.50
V+W	*KQ85*	600	68 ± 31	3 ± 2	0.53 ± 0.28
V+W	*KQ85*	700	67 ± 31	3.5 ± 2	0.34 ± 0.20
V+W	*Kp0*	500	61 ± 26	2.5 ± 1.5	
V+W	*KQ75p1*[a]	500	53 ± 28	4 ± 2	
V+W	*KQ75*	400	47 ± 32	3 ± 2	-0.50 ± 1.7
V+W	*KQ75*	500	55 ± 30	4 ± 2	0.15 ± 0.51
V+W	*KQ75*	600	61 ± 36	4 ± 2	0.53 ± 0.30
V+W	*KQ75*	700	58 ± 36	5 ± 2	0.38 ± 0.21

- ~ 2 sigma "hint": **f_{NL}^{local} ~ 60 +/- 30 (68% CL)**
- 1.8 sigma for KQ75; 2.3 sigma for KQ85 & Kp0

Komatsu et al. (2008)

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- The results are not sensitive to the maximum multipoles used in the analysis, $l_{\rm max}$.

Komatsu et al. (2008)

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Band	Mask	$l_{\max}$	f_{NL}^{local}	$\Delta f_{NL}^{\text{local}}$	b_{src}
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- The estimated contamination from the point sources is small, if any. (Likely overestimated by a factor of ~2.)

Komatsu et al. (2008)

Null Tests

Band	Foreground	Mask	f_{NL}^{local}
Q−W	Raw	$KQ75$	-0.53 ± 0.22
V−W	Raw	$KQ75$	-0.31 ± 0.23
Q−W	Clean	$KQ75$	0.10 ± 0.22
V−W	Clean	$KQ75$	0.06 ± 0.23

- **No signal in the difference of cleaned maps.**

Komatsu *et al.* (2008)

Frequency Dependence

Band	Foreground	Mask	$f_{NL}^{\rm local}$
Q	Raw	$KQ75$	-42 ± 48
V	Raw	$KQ75$	41 ± 35
W	Raw	$KQ75$	46 ± 35
Q	Clean	$KQ75$	10 ± 48
V	Clean	$KQ75$	50 ± 35
W	Clean	$KQ75$	62 ± 35

- **Q is very sensitive to the foreground cleaning.**

Komatsu et al. (2008)

V+W: Raw vs Clean (l_{max}=500)

Band	Foreground	Mask	f_{NL}^{local}
V+W	Raw	*KQ85*	9 ± 26
V+W	Raw	*Kp0*	48 ± 26
V+W	Raw	*KQ75p1*	41 ± 28
V+W	Raw	*KQ75*	43 ± 30

- Clean-map results:
  - KQ85; 61 +/- 26
  - Kp0; 61 +/- 26
  - KQ75p1; 53 +/- 28
  - KQ75; 55 +/- 30

Foreground contamination is not too severe.

The Kp0 and KQ85 results may be as clean as the KQ75 results.

Komatsu et al. (2008)

Our Best Estimate

- Why not using Kp0 or KQ85 results, which have a higher statistical significance?
- Given the profound (i.e., game-chaning) implications and impact of non-zero f_{NL}^{local}, we have chosen a conservative limit from the KQ75 with the point source correction ($\Delta f_{NL}^{local}=4$, which is also conservative) as our best estimate.
  - The 68% limit: $f_{NL}^{local} = 51 +/- 30$ [1.7 sigma]
  - **The 95% limit: $-9 < f_{NL}^{local} < 111$**

Effect of Mask?

- The best-fitting value of f_{NL} shifted from 61 to 55 (for l_{max}=500) by changing KQ85 (81.7% retained) to KQ75 (71.8% retained). Is this shift expected?
- Monte Carlo simulations show that the r.m.s. difference in f_{NL} between these masks is Δf_{NL}=12; thus, the observed change is consistent with a statistical fluctuation.
- The change for Kp0->KQ75 (f_{NL}=61 -> 55) is also consistent: Δf_{NL}=9.7.

Yadav & Wandelt (2008)

Comparison with Y&W

- Yadav and Wandelt used the raw V+W map from the 3-year data.
  - 3yr: f_{NL}^{local} = 68 +/- 30 for l_{max}=450 & Kp0 mask
  - 3yr: f_{NL}^{local} = 80 +/- 30 for l_{max}=550 & Kp0 mask
- Our corresponding 5-year raw map estimate is
  - 5yr: f_{NL}^{local} = 48 +/- 26 for l_{max}=500 & Kp0 mask
  - C.f. clean-map estimate: f_{NL}^{local} = 61 +/- 26
- With more years of observations, the values have come down to a lower significance.

Komatsu et al. (2008)

Main Result (Equilateral)

Band	Mask	l_{max}	f_{NL}^{equil}	$\Delta f_{NL}^{\mathrm{equil}}$
V+W	*KQ75*	400	77 ± 146	9 ± 7
V+W	*KQ75*	500	78 ± 125	14 ± 6
V+W	*KQ75*	600	71 ± 108	27 ± 5
V+W	*KQ75*	700	73 ± 101	22 ± 4

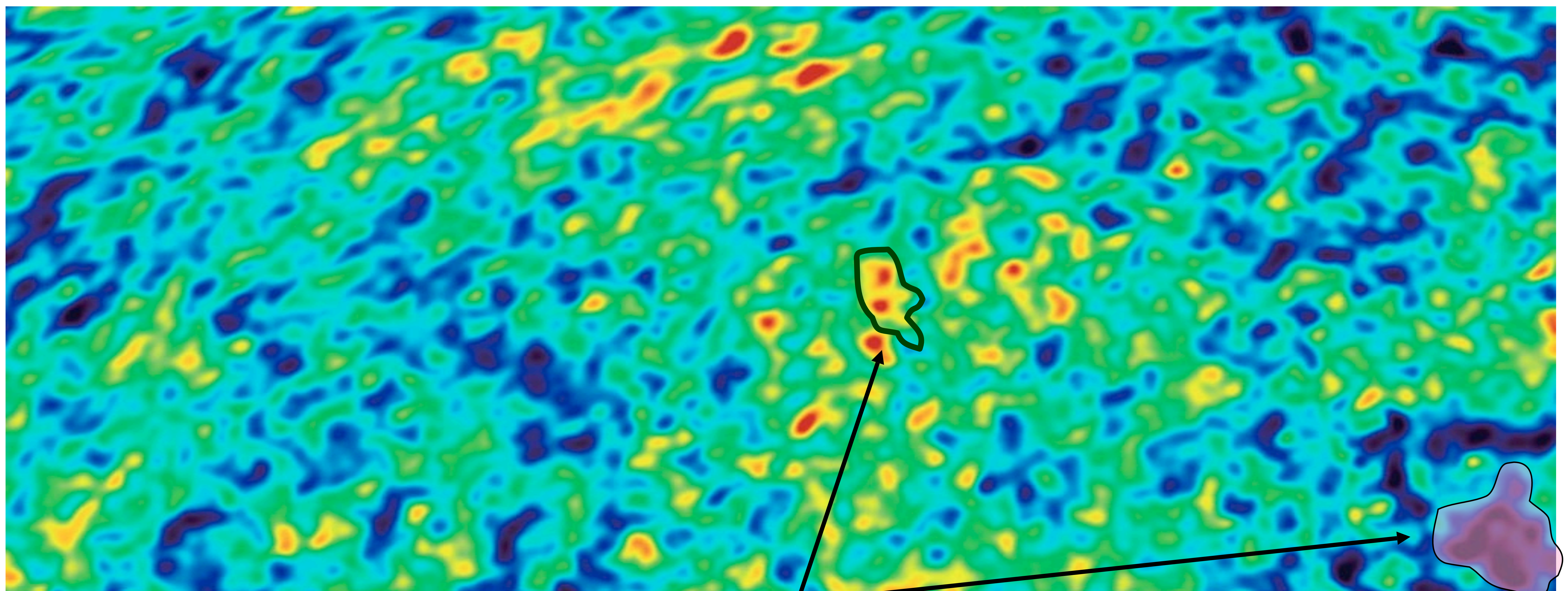
- The point-source correction is much larger for the equilateral configurations.
- Our best estimate from l_{max}=700:
  - The 68% limit: f_{NL}^{equil} = 51 +/- 101
  - **The 95% limit: $-151 < f_{NL}^{equil} < 253$**

Forecasting 9-year Data

- The WMAP 5-year data do not show any evidence for the presence of f_{NL}^{equil}, but *do* show a (~2-sigma) hint for f_{NL}^{local}.
- Our best estimate is probably on the conservative side, but our analysis clearly indicates that more data are required to claim a firm evidence for $f_{NL}^{local}>0$.
- The 9-year error on f_{NL}^{local} should reach $\Delta f_{NL}^{local}=17$
  - If $f_{NL}^{local}\sim 50$, **we would see it at 3 sigma by 2011.**

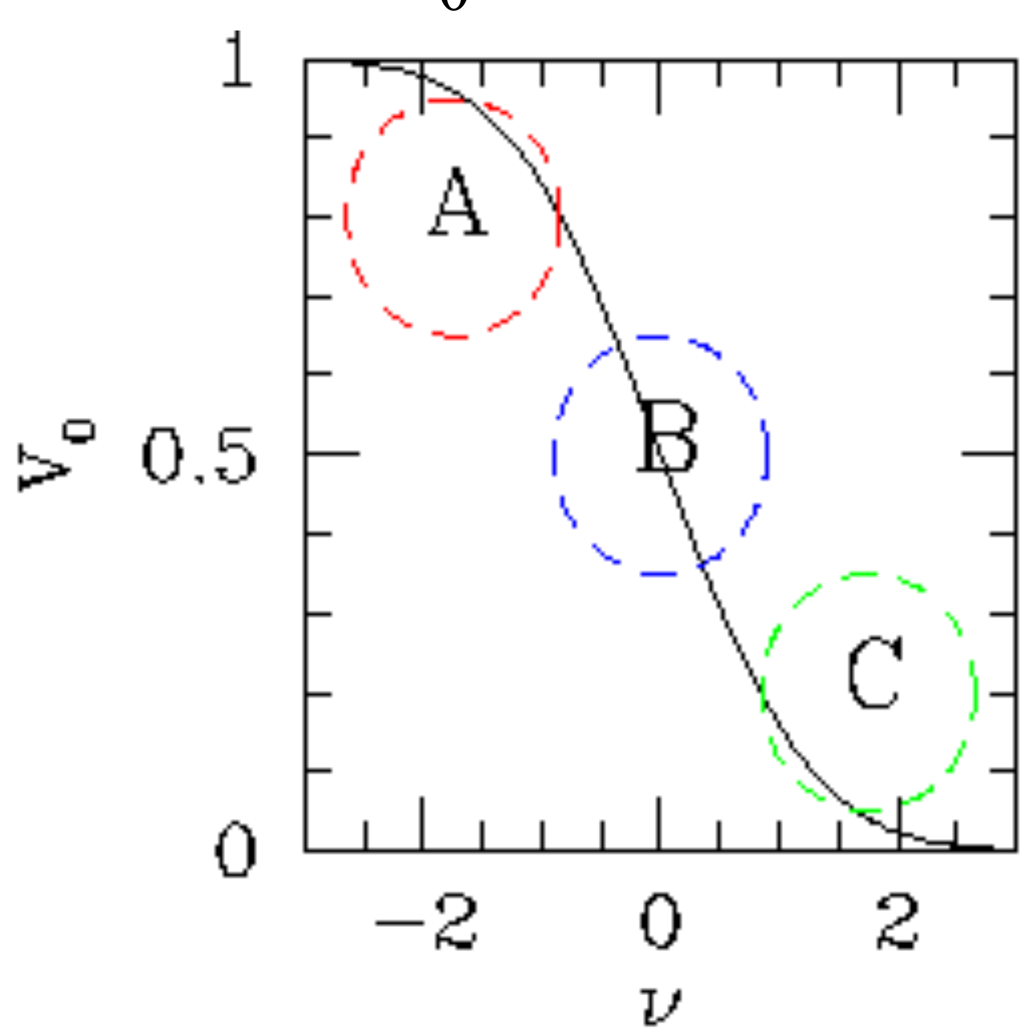
(The WMAP 9-year survey, **recently funded**, will be complete in August 2010.)

Minkowski Functionals (MFs)

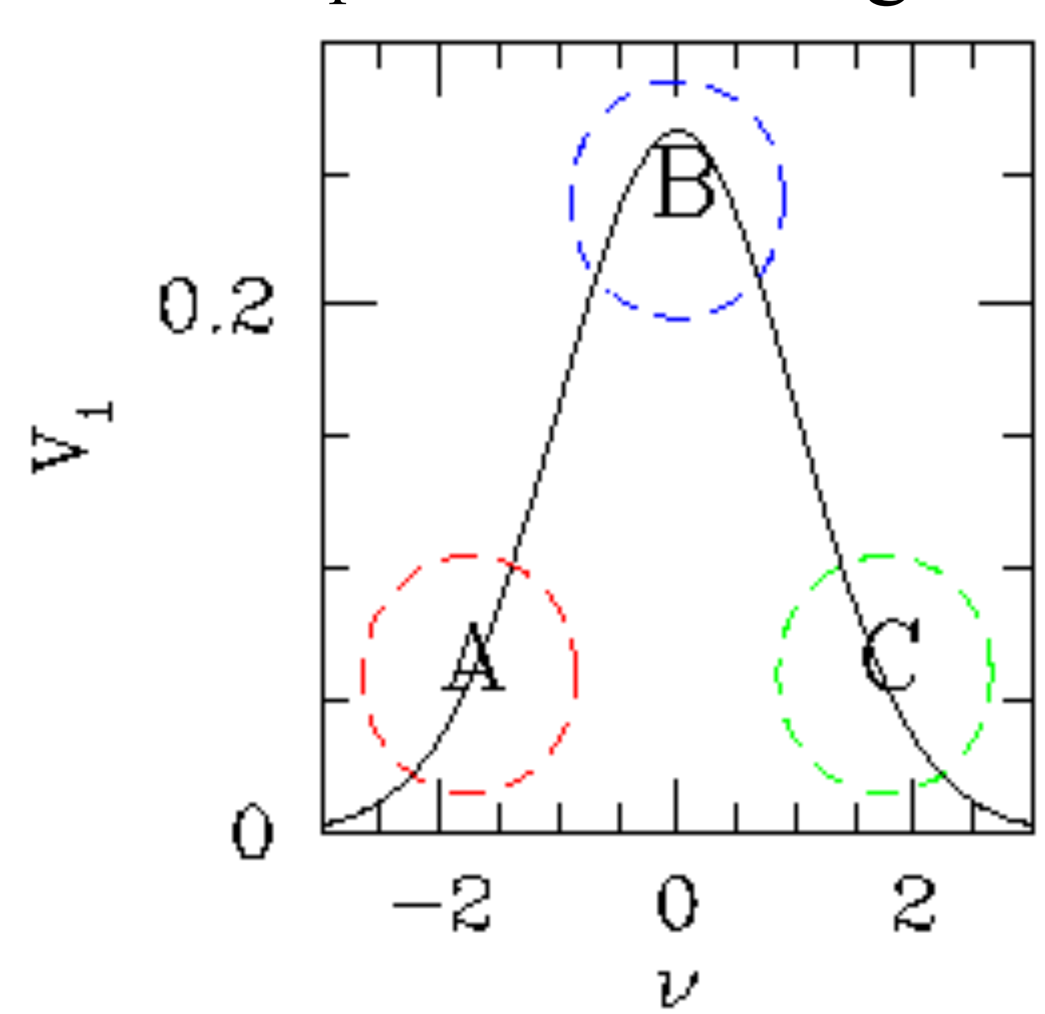


The number of hot spots minus cold spots.

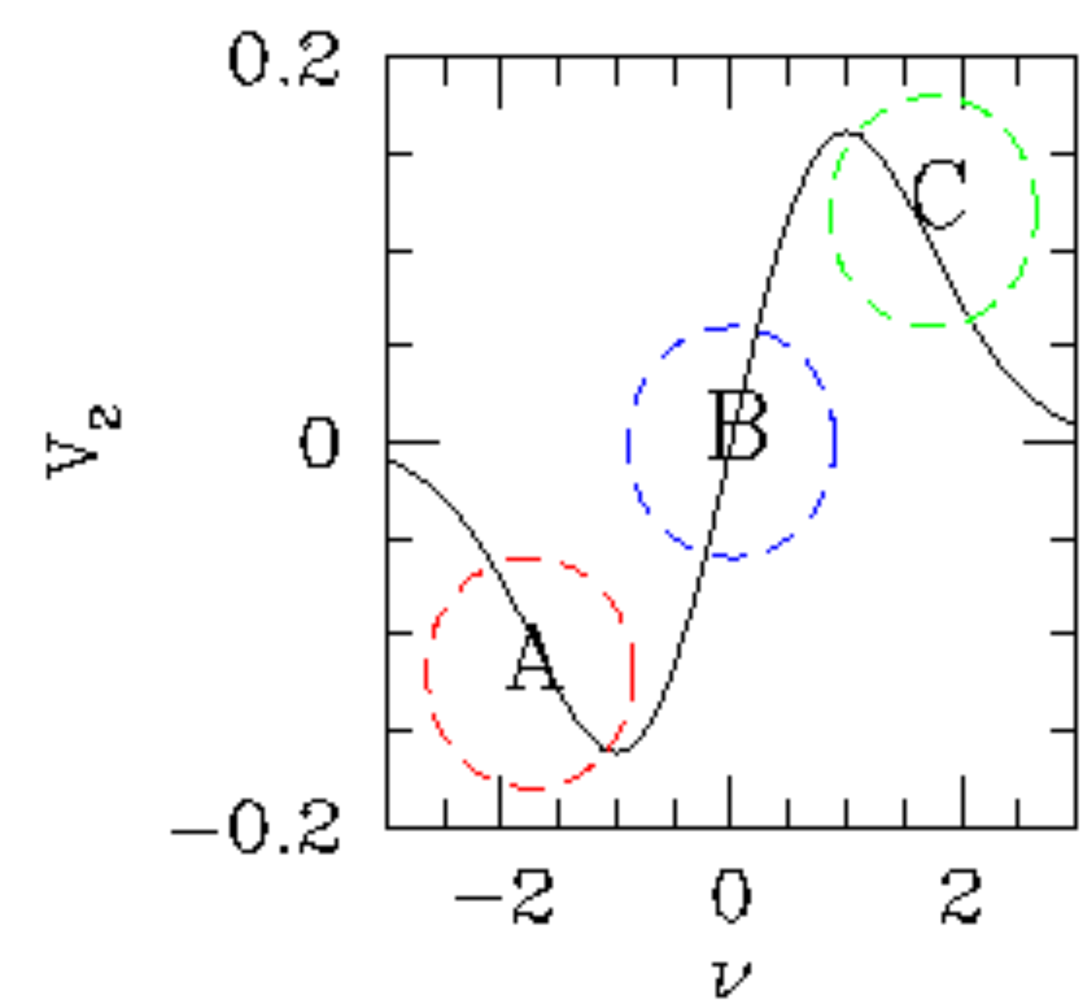
V_0:surface area



V_1: Contour Length

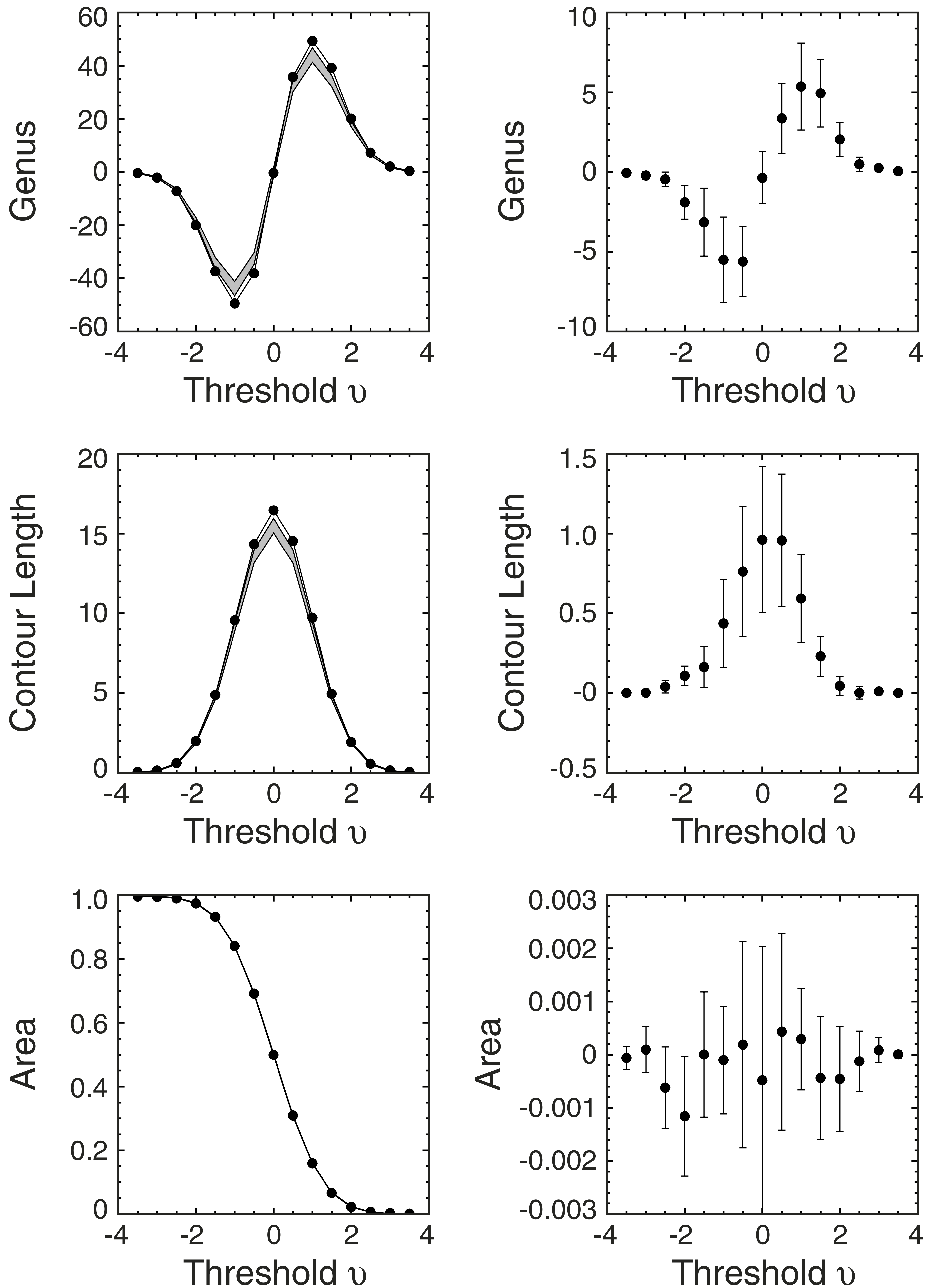


V_2: Euler Characteristic



Komatsu et al. (2008)

MFs from *WMAP* *5-Year Data (V+W)*



Result from a single resolution
(N_{side}=128; 28 arcmin pixel)
[*analysis done by Al Kogut*]

f_{NL}^{local} = -57 +/- 60 (68% CL)

-178 < f_{NL}^{local} < 64 (95% CL)

See Chiaki Hikage's Talk for an extended analysis of MFs from the 5-year data.

Summary

- The best estimates of primordial non-Gaussian parameters from the bispectrum analysis of the WMAP 5-year data are
  - $-9 < f_{NL}^{local} < 111$ (95% CL)
  - $-151 < f_{NL}^{equil} < 253$ (95% CL)
- **9-year data are required to test $f_{NL}^{local} \sim 50$!**

Future Prospects

- Future is always bright, right?

Gaussianity vs Flatness: Future

- **Flatness will never beat Gaussianity.**
  - In 5-10 years, we will know **flatness** to **0.1%** level.
  - In 5-10 years, we will know **Gaussianity** to **0.01%** level (f_{NL}~10), or even to **0.005%** level (f_{NL}~5), at 95% CL.

- However, a real potential of Gaussianity test is that **we might detect something at this level** (multi-field, curvaton, DBI, ghost cond., new ekpyrotic…)
  - Or, we might detect curvature first?
  - Is 0.1% curvature interesting/motivated?

Beyond Bispectrum: Trispectrum of Primordial Perturbations

- Trispectrum is the Fourier transform of four-point correlation function.
- Trispectrum(k_1,k_2,k_3,k_4)

  $=\langle\Phi(k_1)\Phi(k_2)\Phi(k_3)\Phi(k_4)\rangle$

  which can be sensitive to the higher-order terms:

$$\Phi(\boldsymbol{x}) = \Phi_{\rm L}(\boldsymbol{x}) + f_{\rm NL}\left[\Phi_{\rm L}^2(\boldsymbol{x}) - \langle\Phi_{\rm L}^2(\boldsymbol{x})\rangle\right] + f_2\Phi_{\rm L}^3(\boldsymbol{x})$$

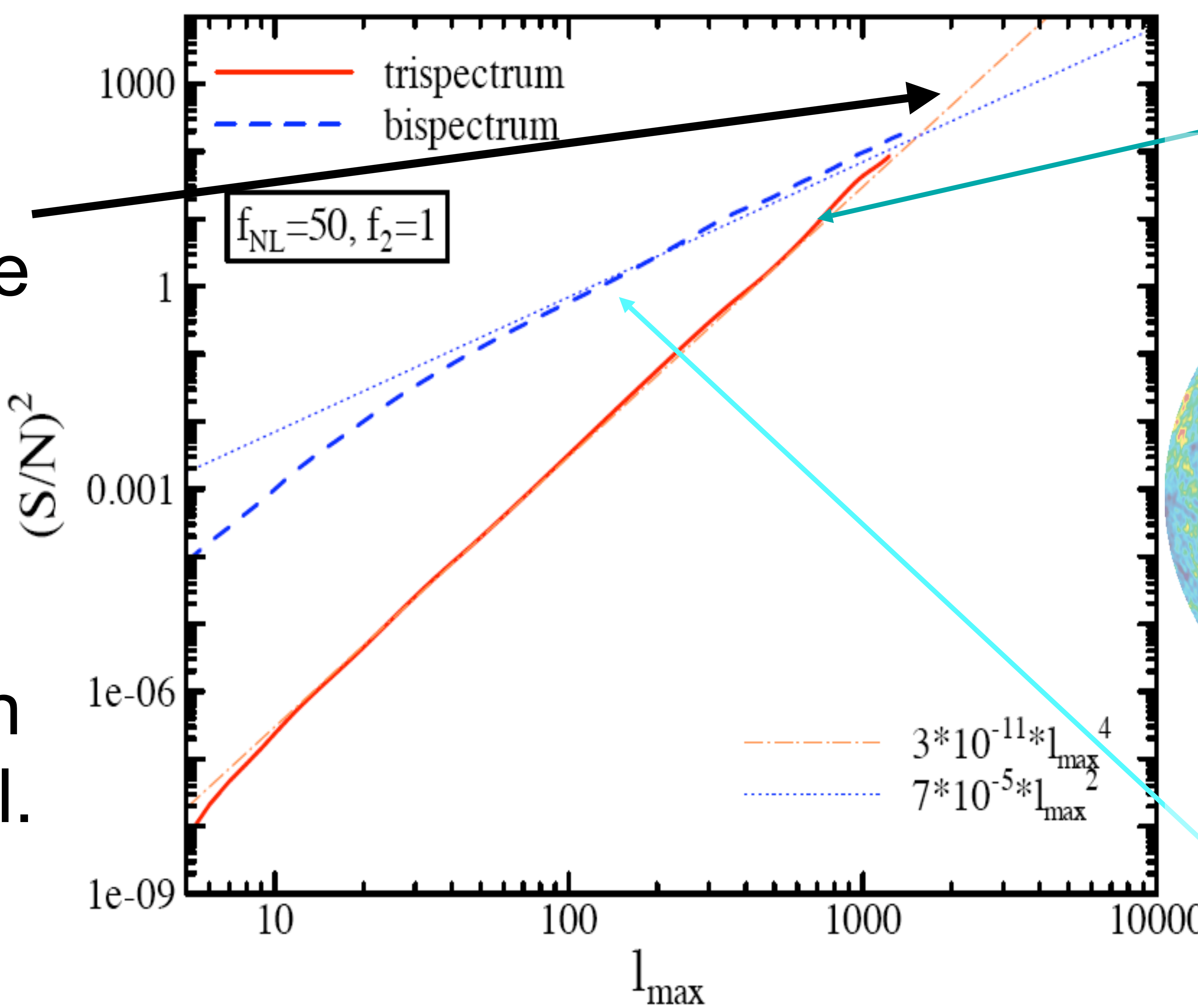
Measuring Trispectrum

- It's pretty painful to measure all the quadrilateral configurations.
  - Measurements from the COBE 4-year data (Komatsu 2001; Kunz et al. 2001)
- Only limited configurations measured from the WMAP 3-year data
  - Spergel et al. (2007)
- No evidence for non-Gaussianity, but f_{NL} has not been constrained by the trispectrum yet. (Work in progress: Dore, Smith & EK)

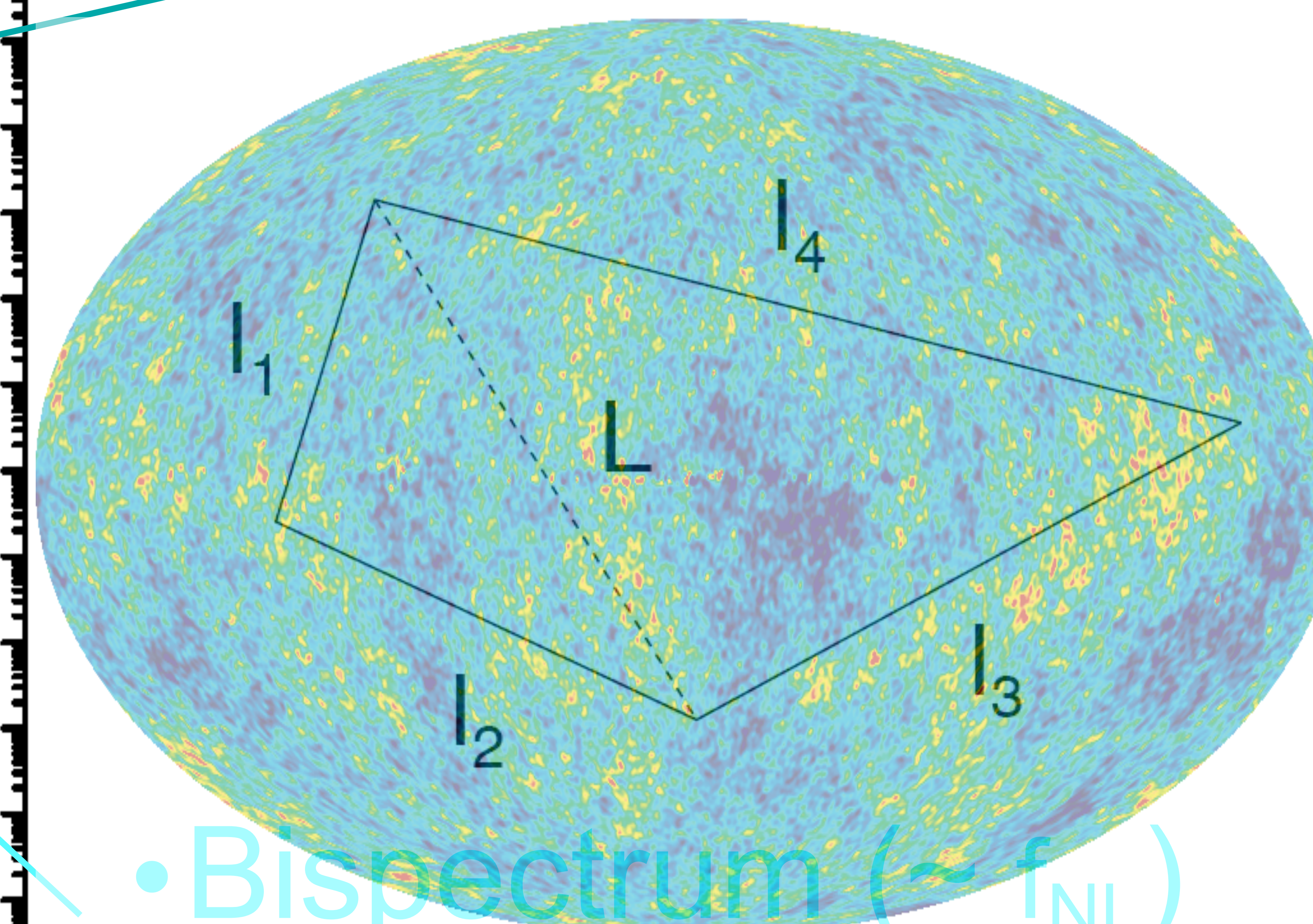
Kogo & Komatsu (2006)

Trispectrum: Not useful for WMAP, but maybe useful for Planck, if f_{NL} is greater than ~50: Excellent Cross-check!

These thin dotted lines are wrong
-- Thanks to Paolo Creminelli for point this out in Creminelli et al.



- Trispectrum (~ f_{NL}^2)



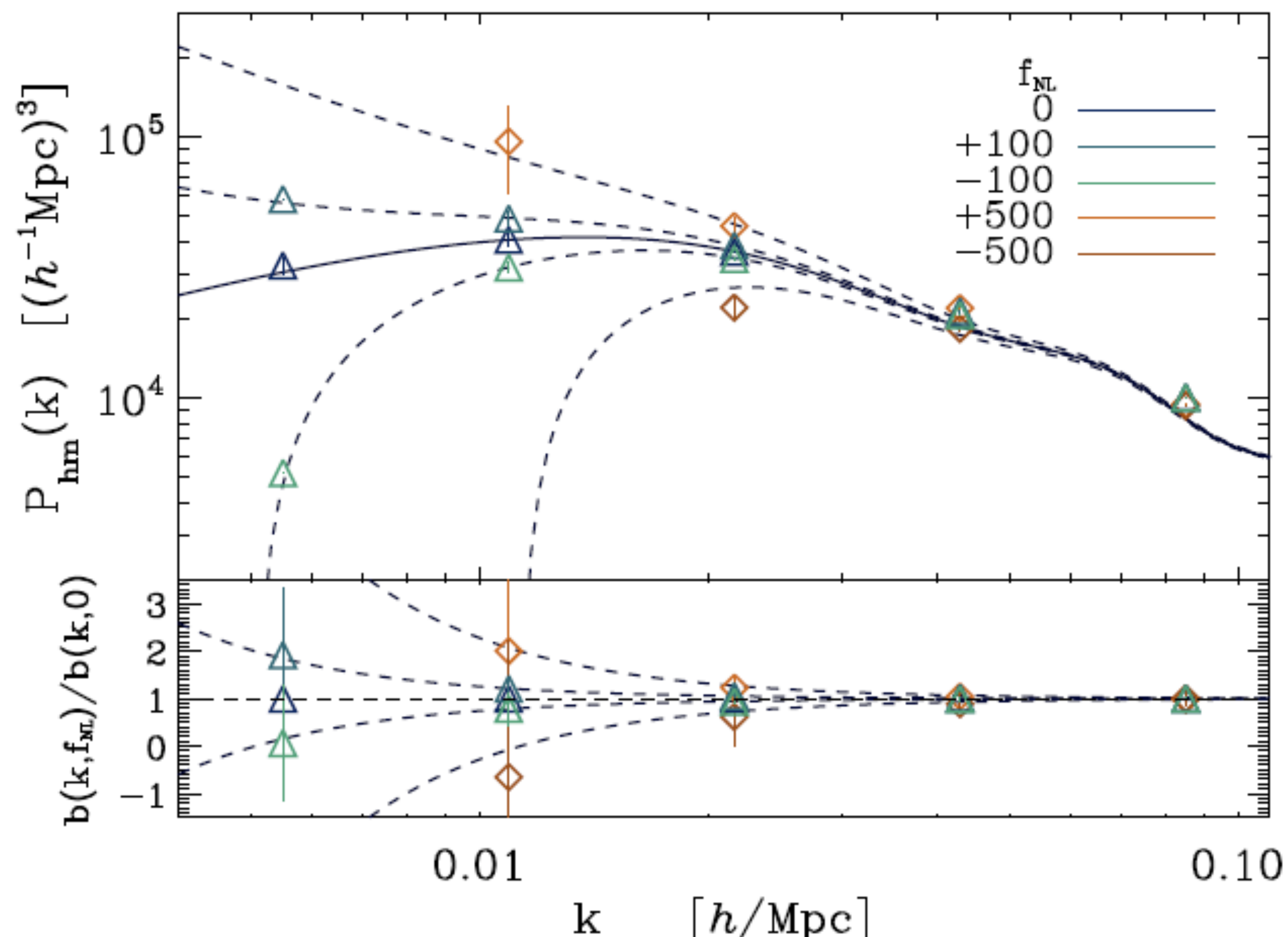
- Bispectrum (~ f_{NL})

More On Future Prospects

- CMB: Planck (temperature + polarization): **Δf_{NL}(local)=6 (95%)**
  - –Yadav, Komatsu & Wandelt (2007)
- Large-scale Structure: e.g., ADEPT, CIP: **Δf_{NL}(local)=7 (95%); Δf_{NL}(equilateral)=90 (95%)**
  - –Sefusatti & Komatsu (2007)
  - –This estimate is based upon the assumption of "local galaxy bias," which needs to be modified for f_{NL}(local) according to the recent findings (Licia Verde's Talk)
- CMB and LSS are independent. By combining these two constraints, we get **Δf_{NL}(local)=4.5**.

New, Powerful Probe of f_{NL}

- f_{NL} modifies the galaxy bias with a unique scale dependence
  - –Dalal et al.; Matarrese & Verde
  - –McDonald; Afshordi & Tolley
- The statistical power of this method is promising:
  - –SDSS: $-29 < f_{NL} < 70$ (95%CL); Slosar et al.
  - –Comparable to the WMAP limit already ($-9 < f_{NL} < 111$)
  - –Combined limit (SDSS+WMAP):
    - **$-1 < f_{NL} < 70$** (95%CL)



Where Should We Be Going?

- Explore different statistics (both CMB and LSS)
  - Minkowski functionals, trispectrum, wavelets and others
  - Purpose: Checking for systematic errors

- Go for the large-scale structure
  - The large-scale structure of the Universe at high redshifts offers a <u>*definitive*</u> cross-check for the presence of primordial non-Gaussianity.
  - If CMB sees primoridial non-Gaussianity, the **same** non-Gaussianity must also be seen by the large-scale structure!